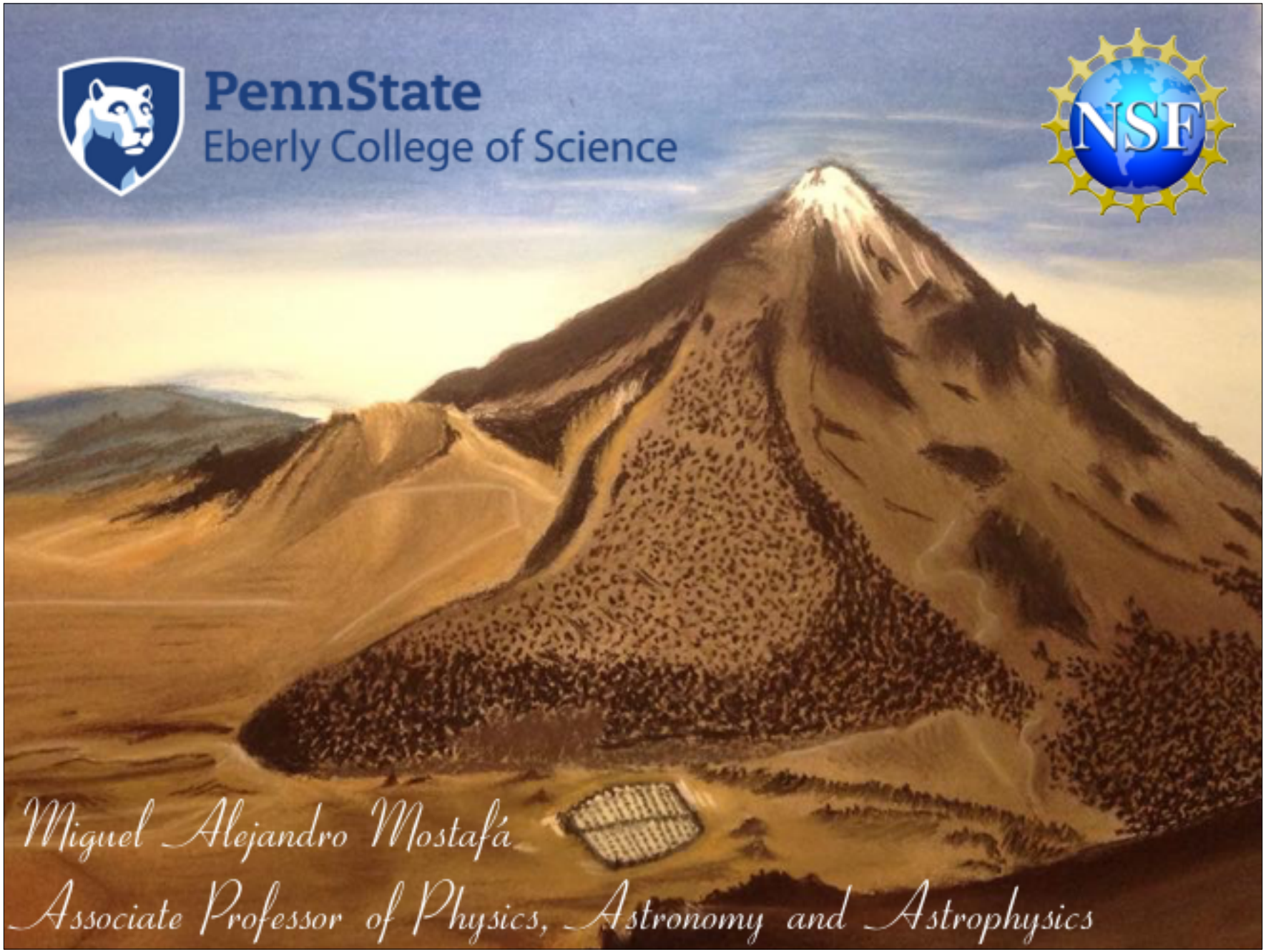

PennState
Eberly College of Science
NSF
Miguel Alejandro Mostafá
Associate Professor of Physics, Astronomy and Astrophysics

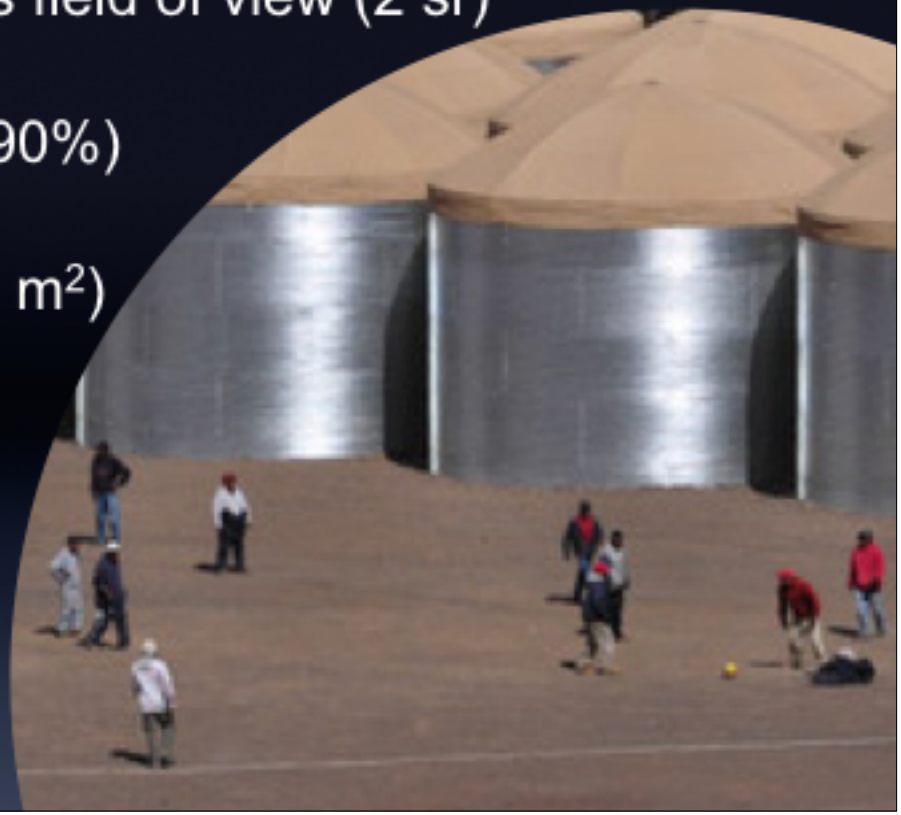
Introduction

Outline

1. Motivation for VHE γ rays
2. The HAWC Observatory
3. Preliminary first results
4. Outlook

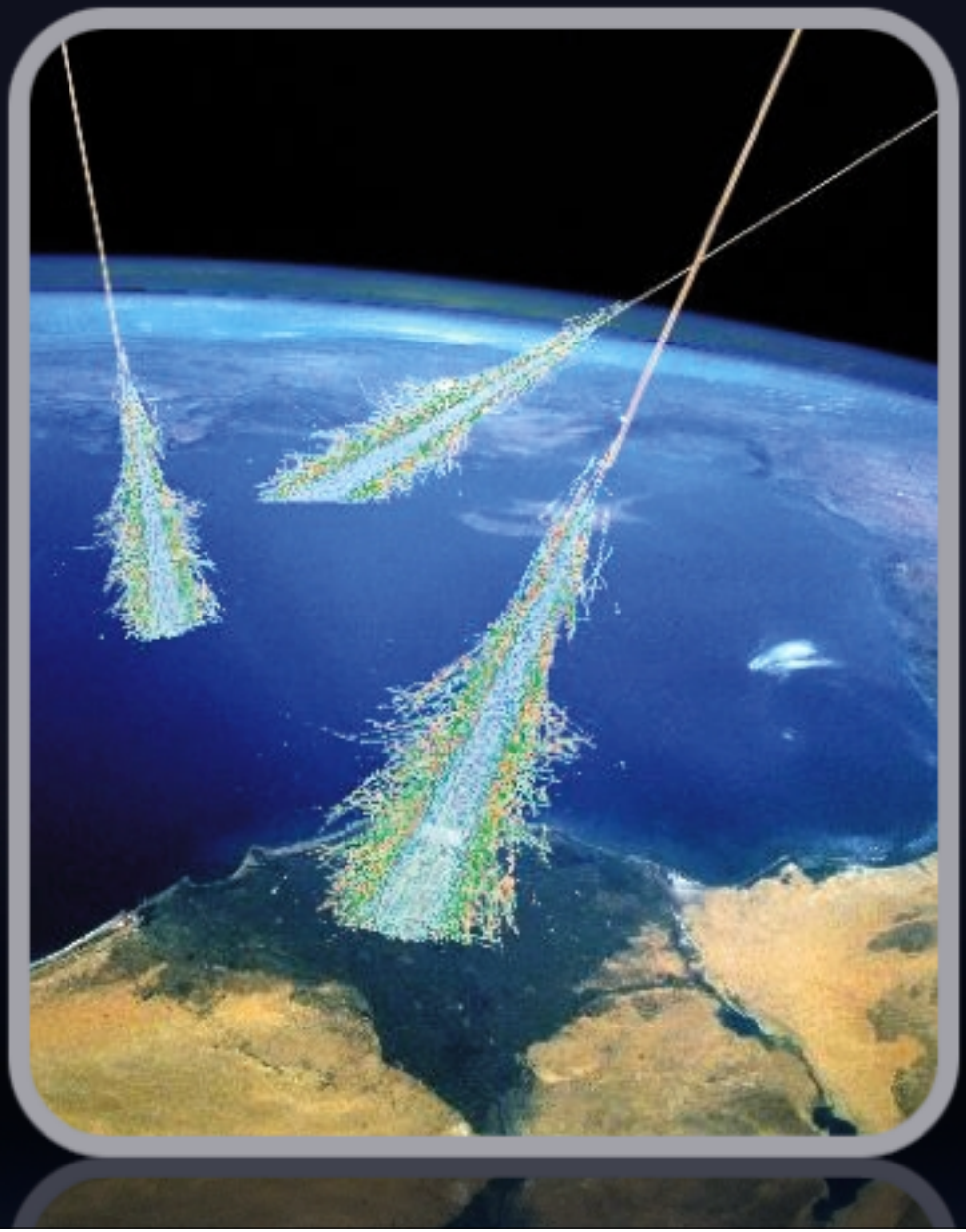
Introduction

- 2nd generation water Cherenkov
- Wide instantaneous field of view (2 sr)
- High duty cycle (> 90%)
- Large area (22,000 m^2)



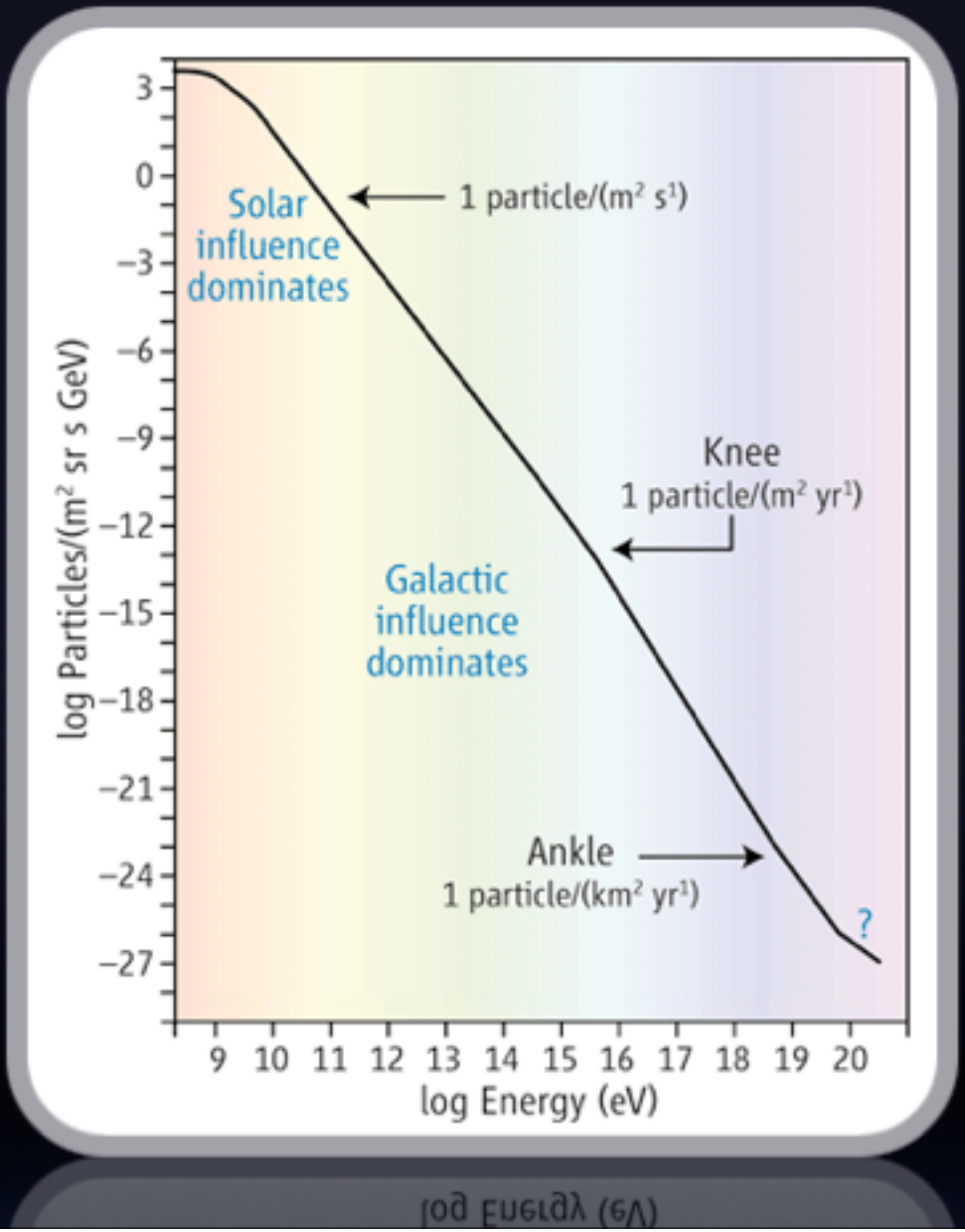
Cosmic Rays

- *All the time* you are being hit by a **flux of high energy particles** from above!



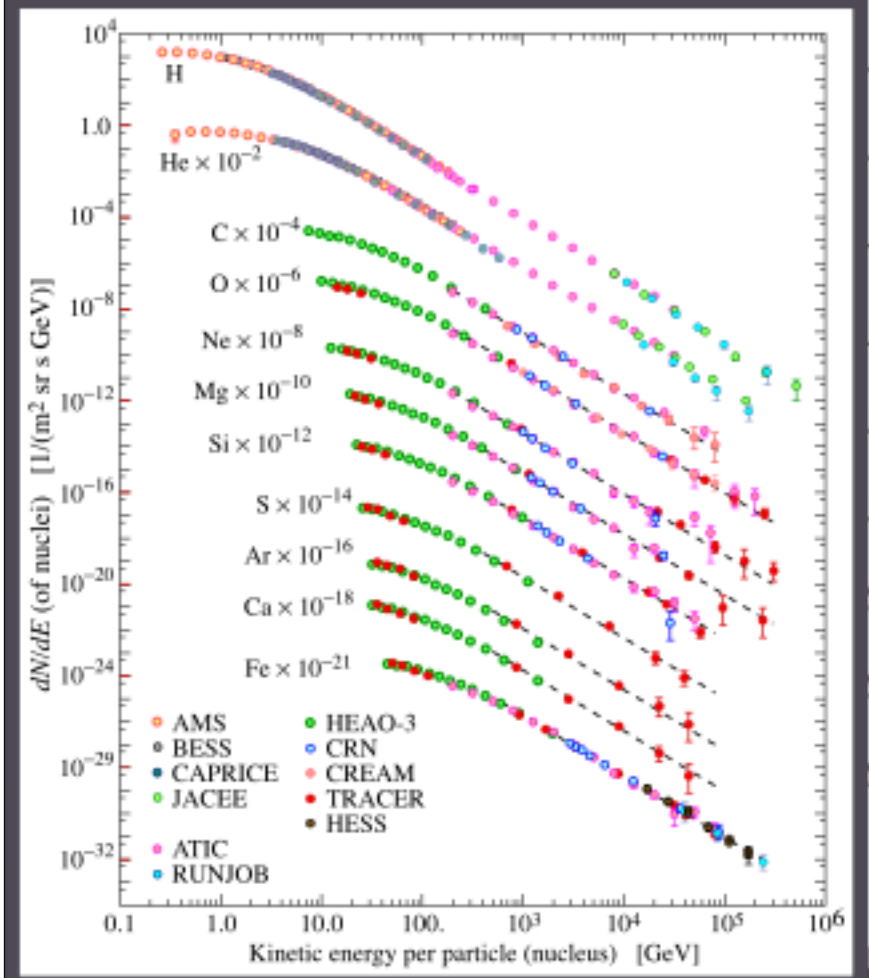
Cosmic Ray Flux

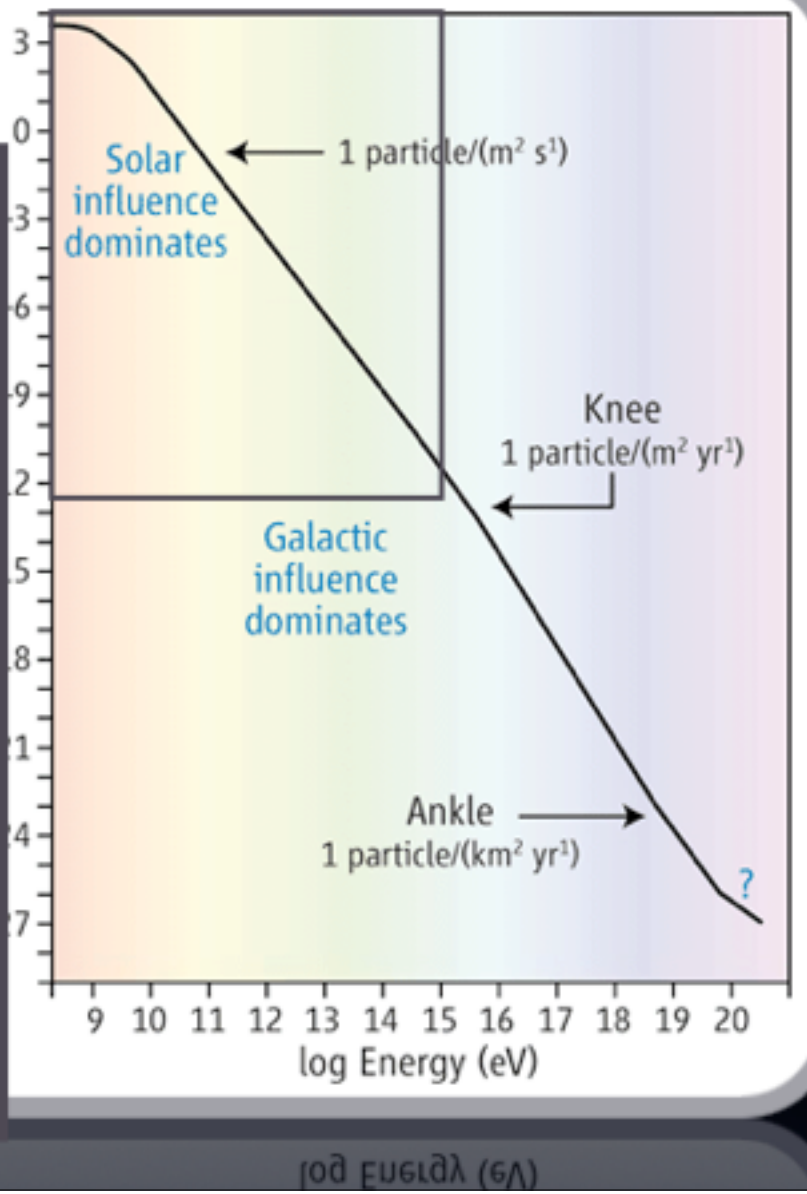
- At low E, solar magnetic fields strongly influence CR propagation
- At high E, the minimally deflected CRs are far less intense and thus much harder to detect.

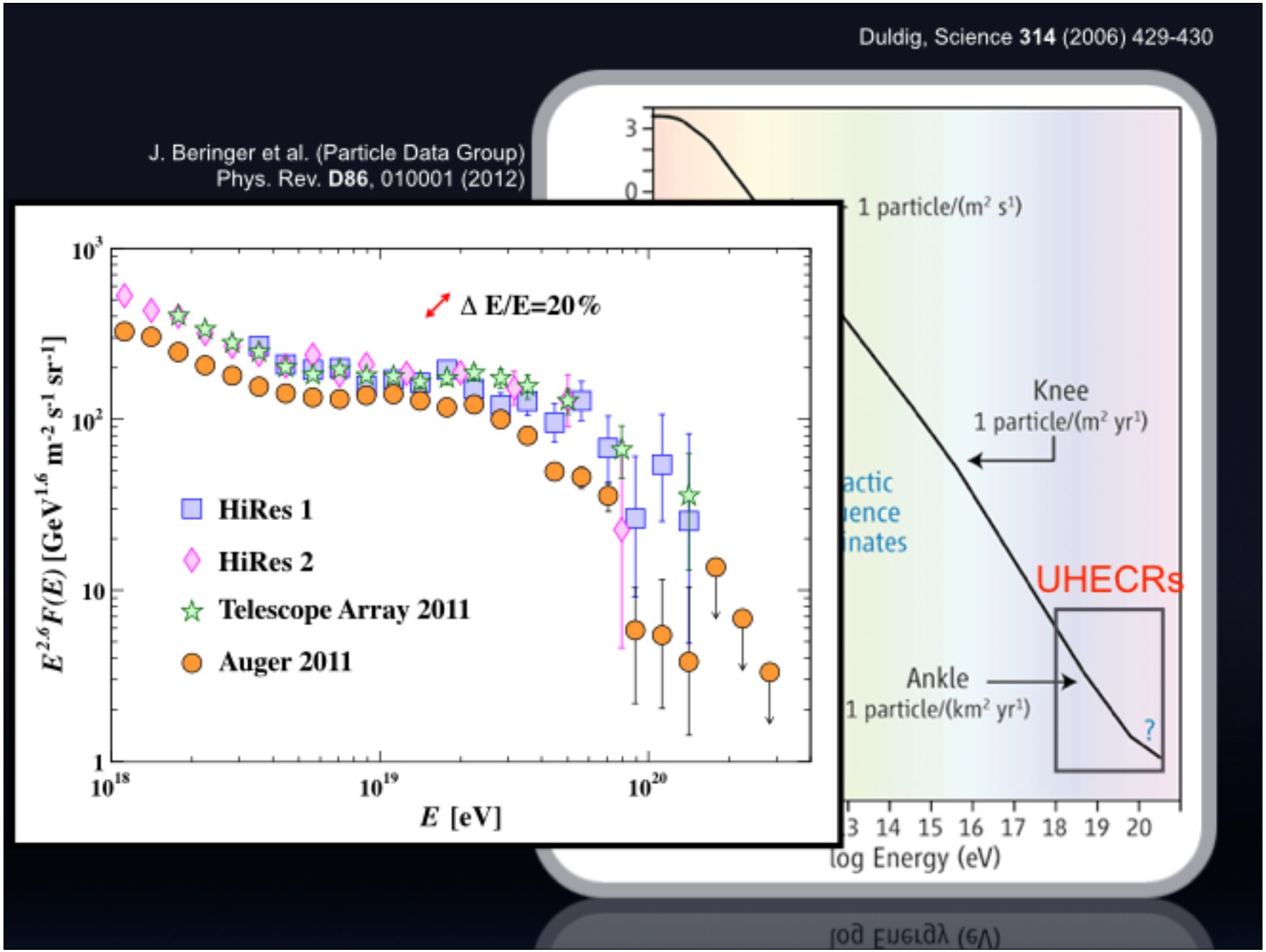


Duldig, Science **314** (2006) 429-430

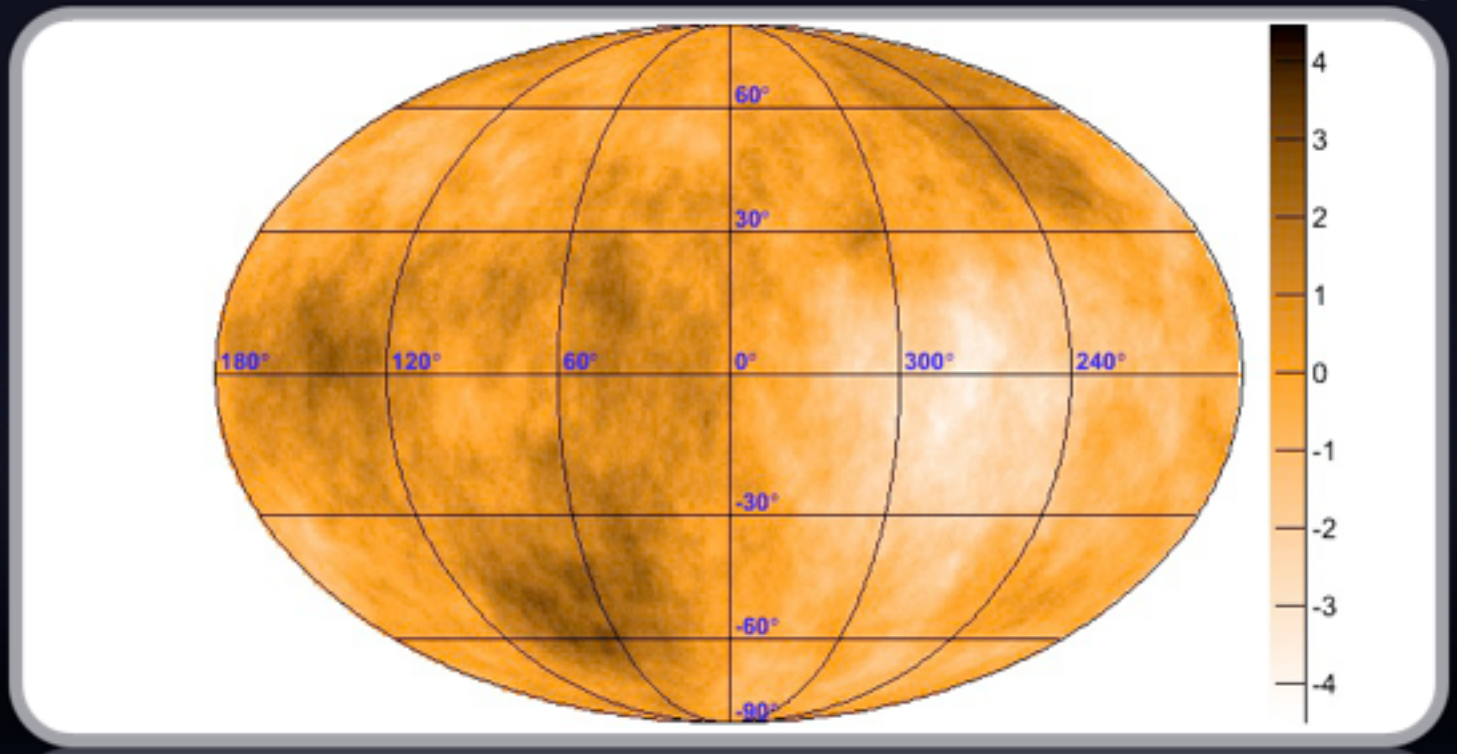
J. Beringer et al. (Particle Data Group)
Phys. Rev. **D86**, 010001 (2012)



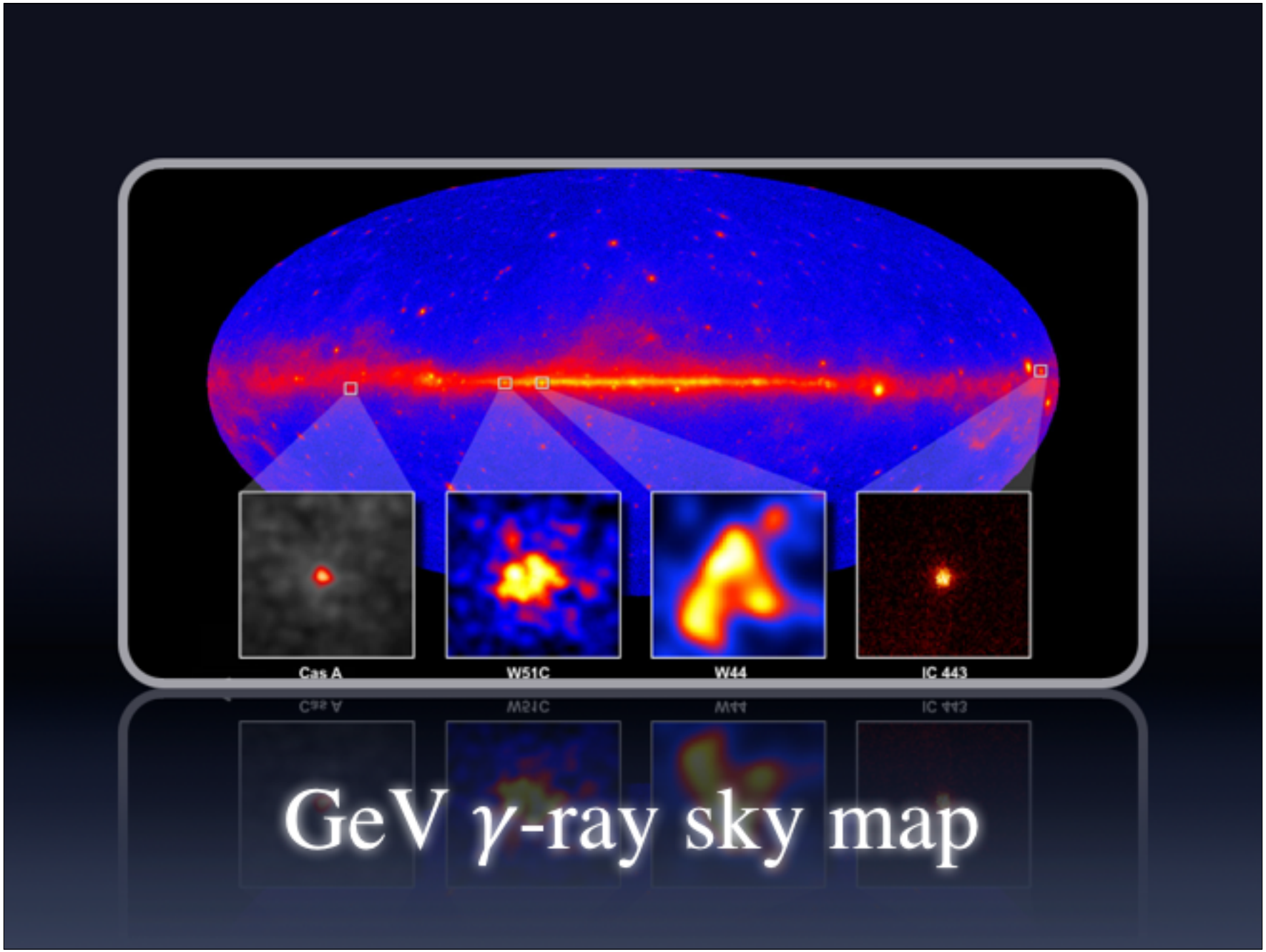



Duldig, Science 314 (2006) 429-430
J. Beringer et al. (Particle Data Group)
Phys. Rev. D86, 010001 (2012)
$E^{2.6}F(E)$ [GeV$^{1.6}$ m^{-2} s^{-1} sr^{-1}]
10^3
10^2
10
1
10^{18}
10^{19}
10^{20}
E [eV]
Δ E/E=20%
HiRes 1
HiRes 2
Telescope Array 2011
Auger 2011
1 particle/(m² s¹)
Knee
1 particle/(m² yr¹)
UHECRs
Ankle
1 particle/(km² yr¹)
?
3 14 15 16 17 18 19 20
log Energy (eV)

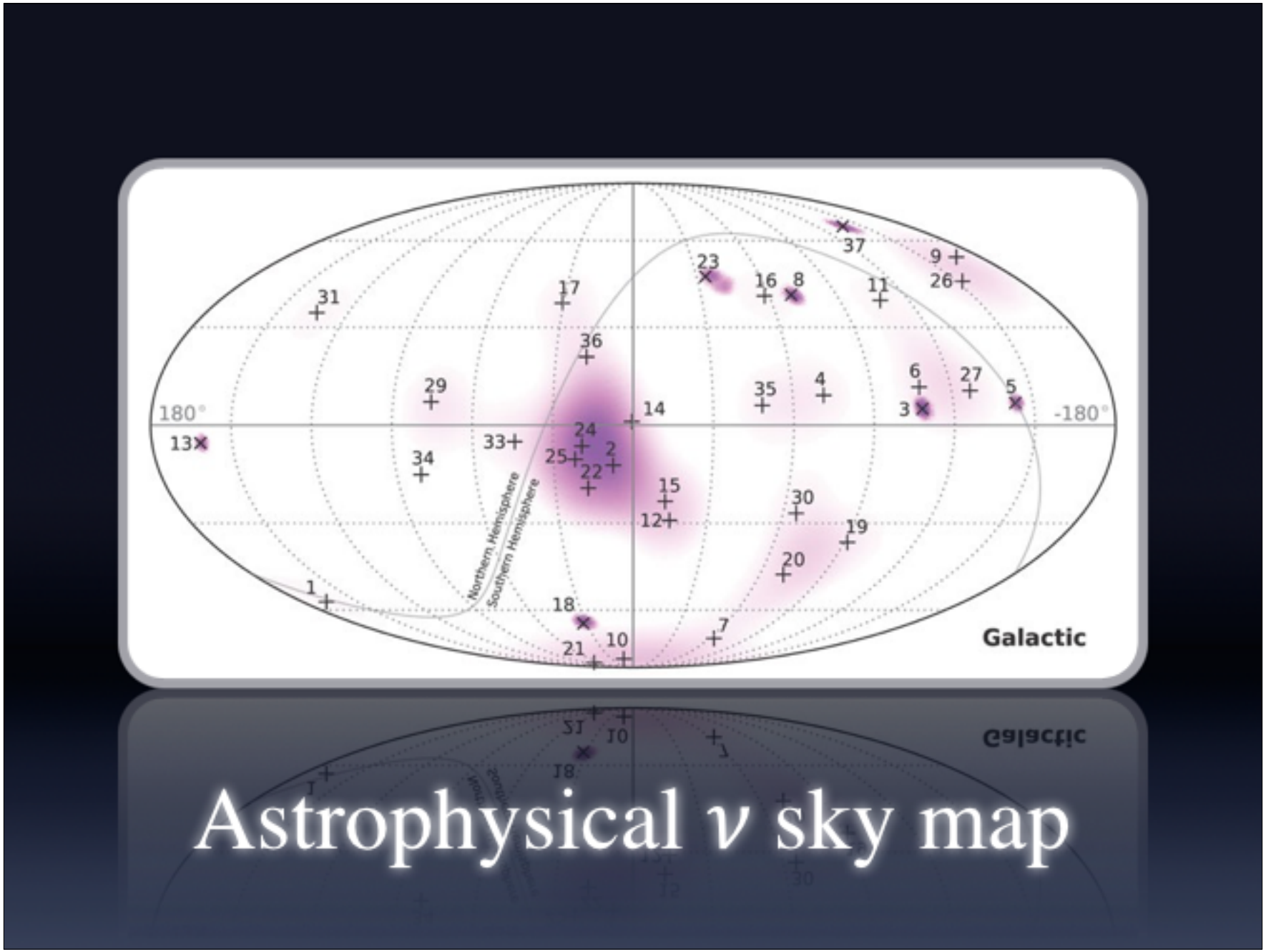
E > 10 EeV; 30° smoothing



UHECR sky map

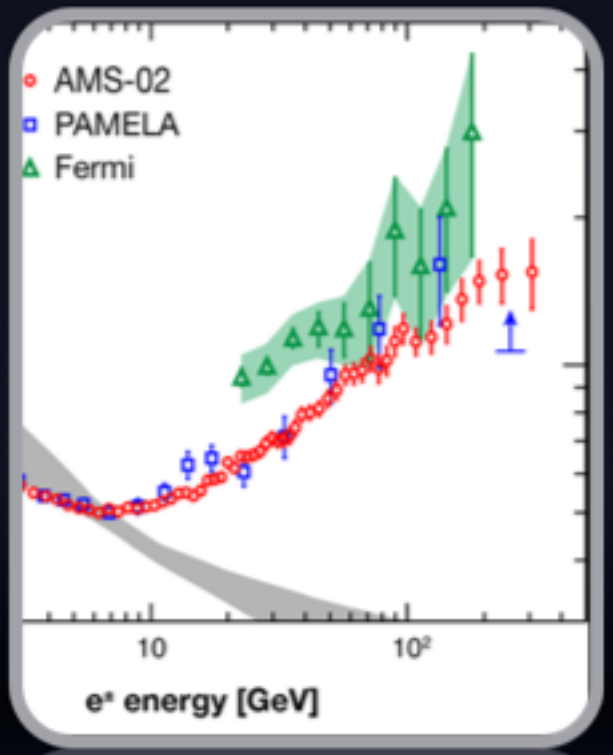


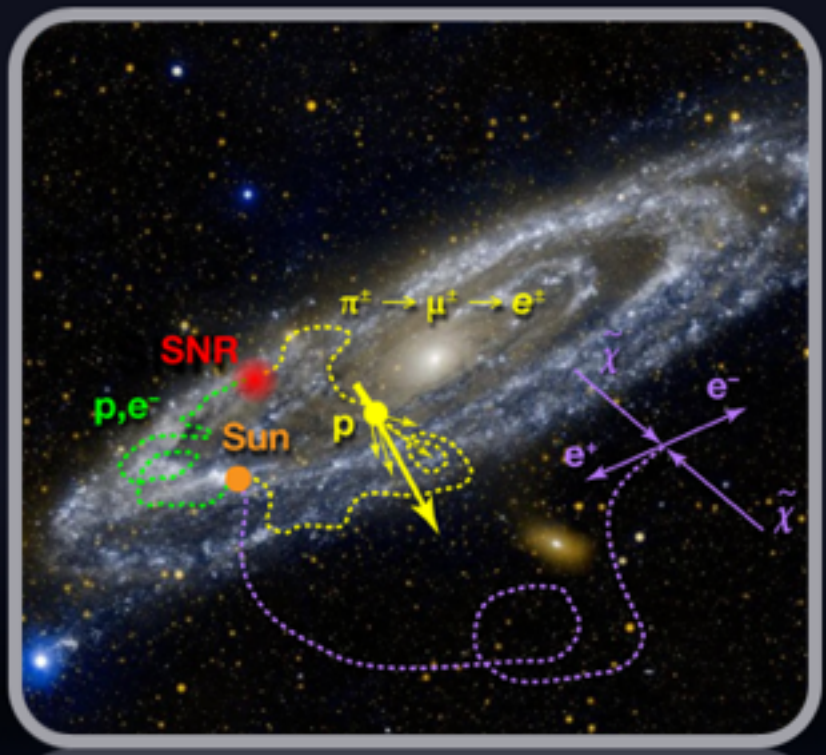
GeV γ-ray sky map



Astrophysical ν sky map

Positron Excess





Credit: Image: GALEX, JPL-Caltech, NASA;
Drawing: APS/**Alan Stonebraker**

Scientific Motivation

- Constrain the **origin of cosmic rays** by measuring gamma-ray **spectra to 100 TeV**.
- Probe **particle acceleration** in astrophysical jets with **wide field of view, high duty factor** observations.
- Explore **new physics** with an **unbiased survey** of the **TeV sky**.

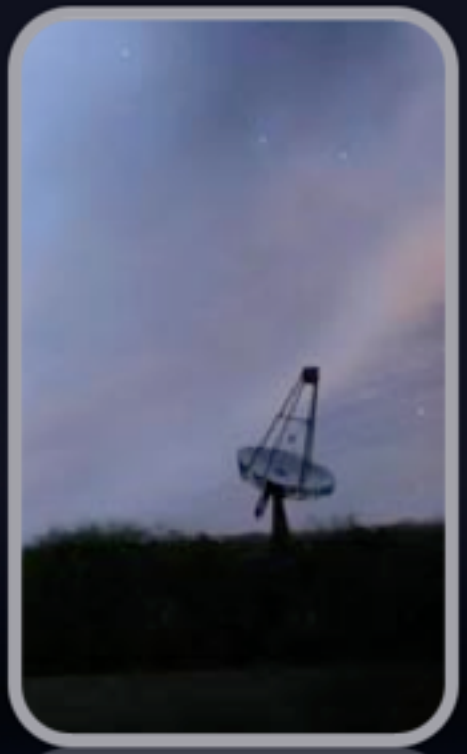
Experimental Techniques

- ✓ Background free
- ✓ Large duty cycle
- ✓ Large aperture

- \- Small area

- **Space-based detectors**
  - Low energy threshold
  - EGRET, Fermi-LAT

Experimental Techniques



- ✓ Large effective area
- ✓ Excellent background rejection

- - Small aperture
- - Low duty cycle

- **Imaging Atmospheric Cherenkov Telescopes**
  - High sensitivity
  - HESS, MAGIC, VERITAS

Experimental Techniques

- ✓ Large aperture
- ✓ Excellent background rejection
- ✓ Large duty cycle

- \- Moderate area

- **Ground array of air-shower particle detectors**
  - Large aperture + High duty cycle
  - Milagro, Tibet, ARGO, HAWC

The HAWC Collaboration

The HAWC Observatory

300 - 7 m x 5 m steel Water Cherenkov Detectors
(a.k.a. tanks) with 4 PMTs at 4,100 m a.s.l. in Mexico

Water Cherenkov Detectors

300 - 7 m x 5 m steel Water Cherenkov Detectors
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Water Cherenkov Detectors

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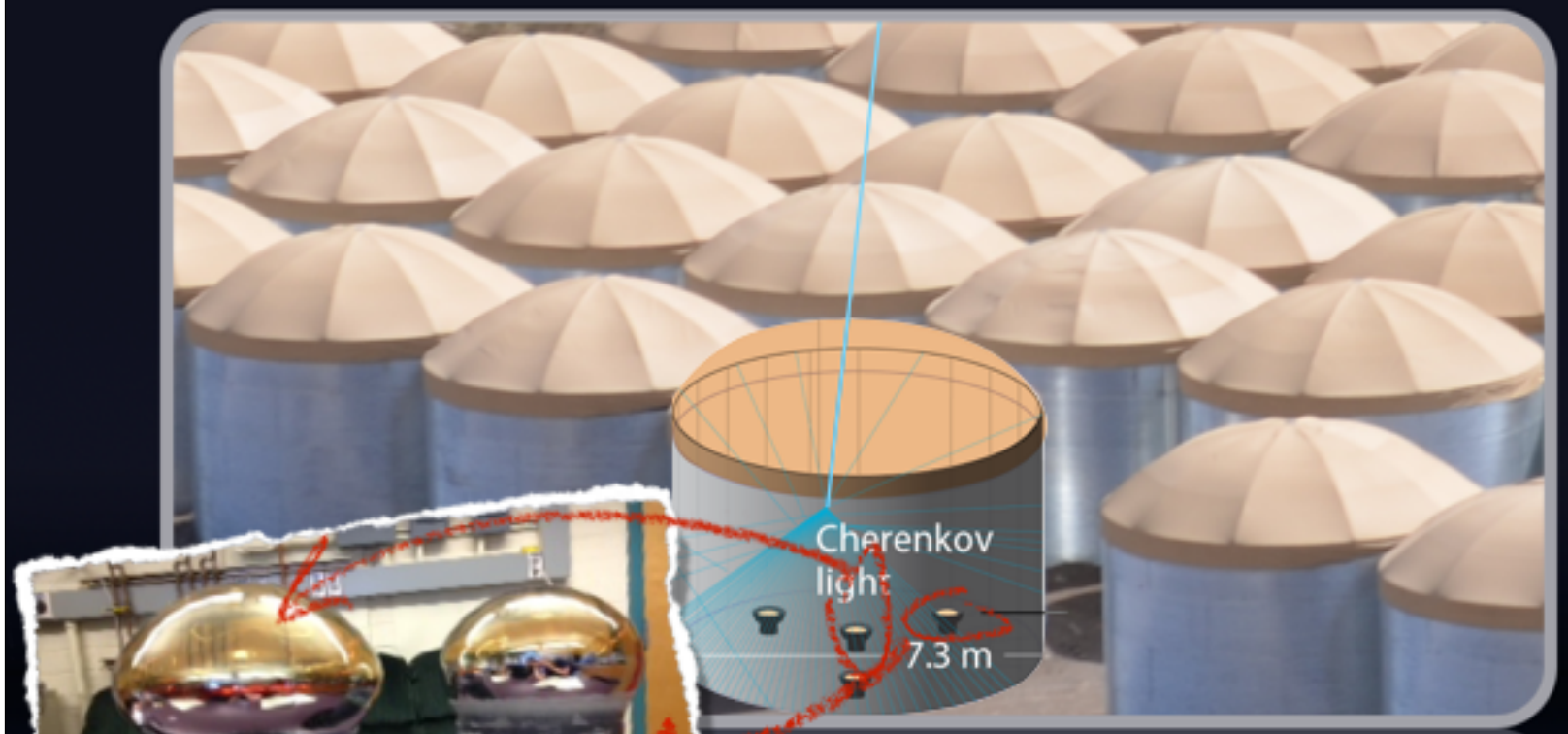
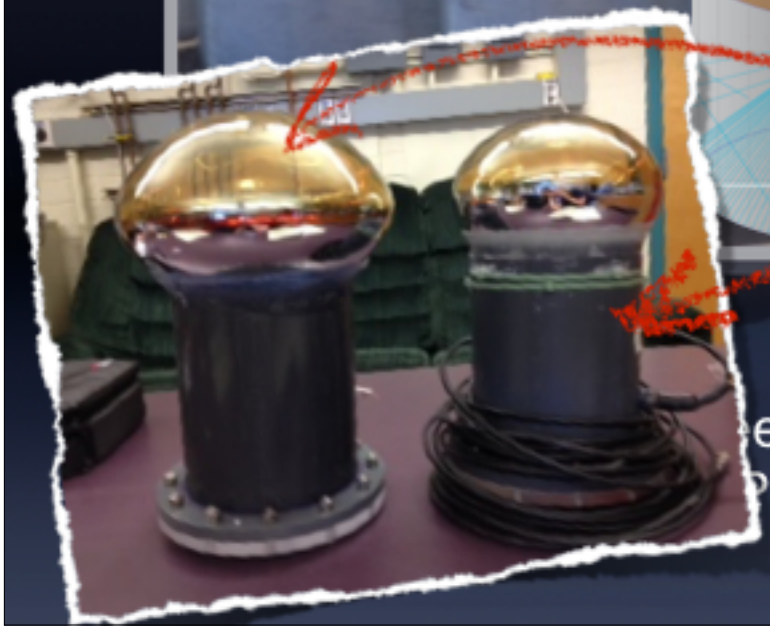
Water Cherenkov Detectors



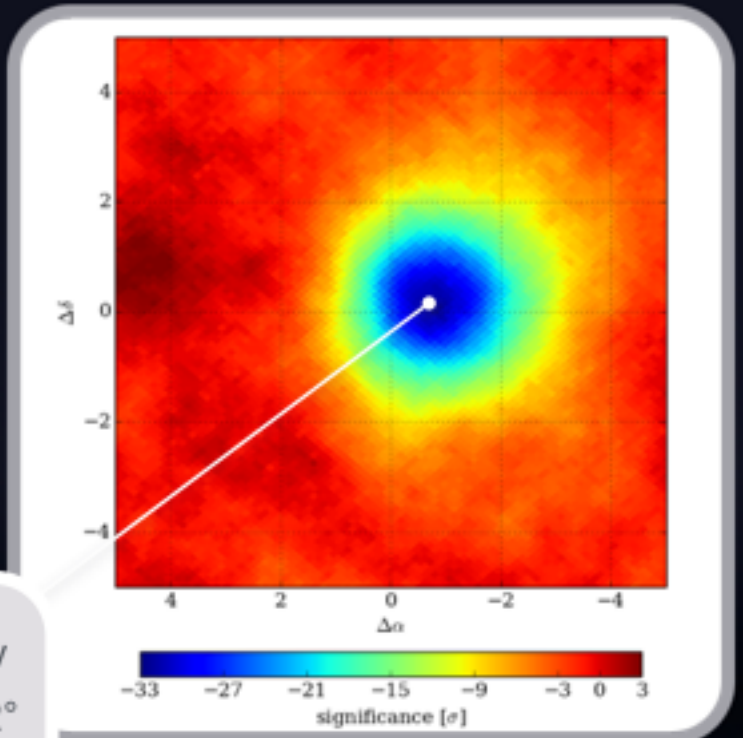


el Water Cherenkov Detectors
MTs at 4,100 m a.s.l. in Mexico

Water Cherenkov Detectors

Water Cherenkov Detectors

Effect of the laser calibration on the observation of the shadow of the Moon

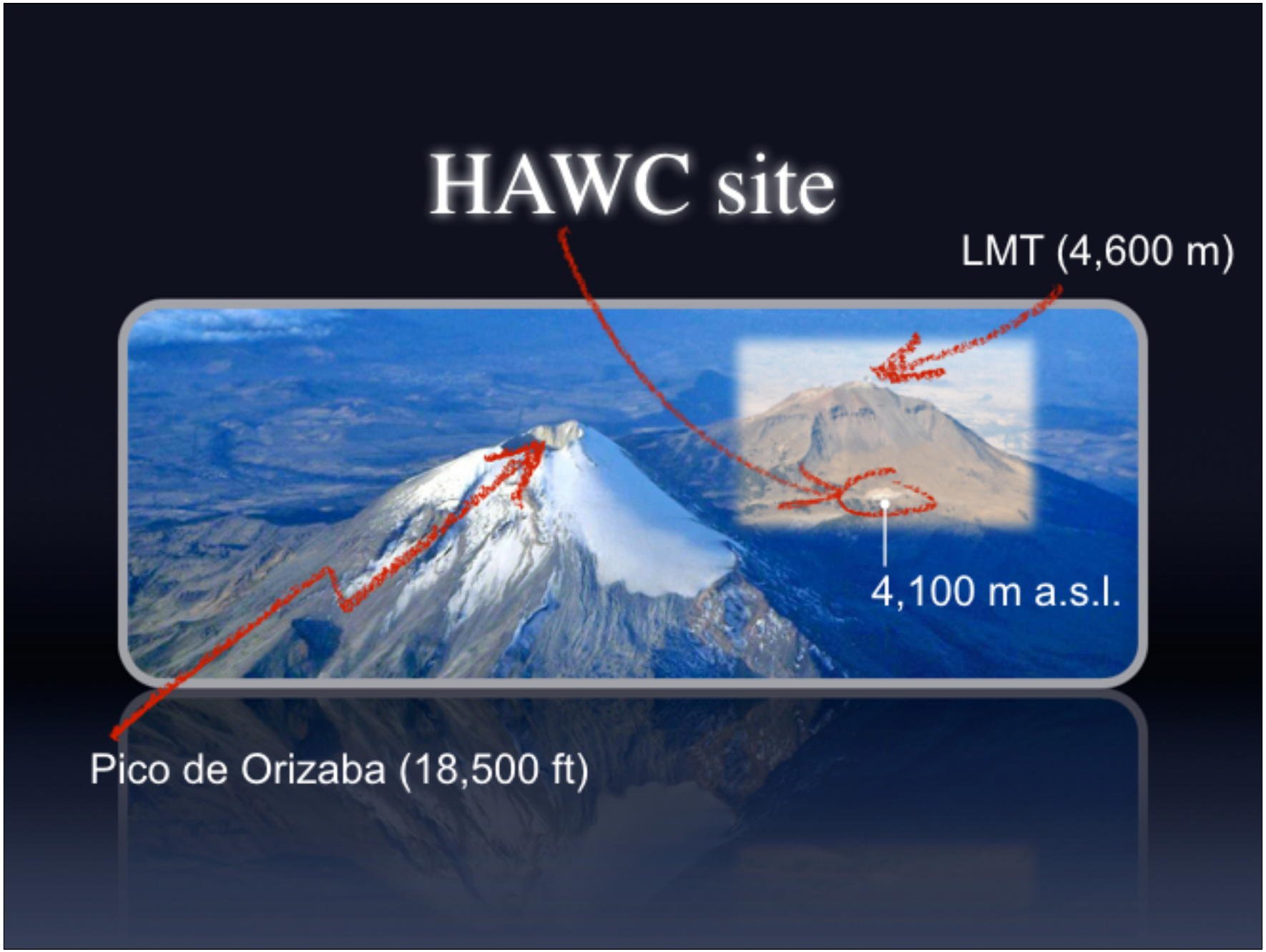


- deflection matches 2 TeV median energy
- angular resolution < shadow width of 1.2°
- position verifies pointing

Dedicated laser calibration system

HAWC site


Jalapa
Veracruz
Veracruz
Tlaxcala
HAWC Site
Puebla
Cuernavaca


HAWC site
LMT (4,600 m)
4,100 m a.s.l.
Pico de Orizaba (18,500 ft)

Design improvements

Go higher (altitude a.s.l.)

Milagro

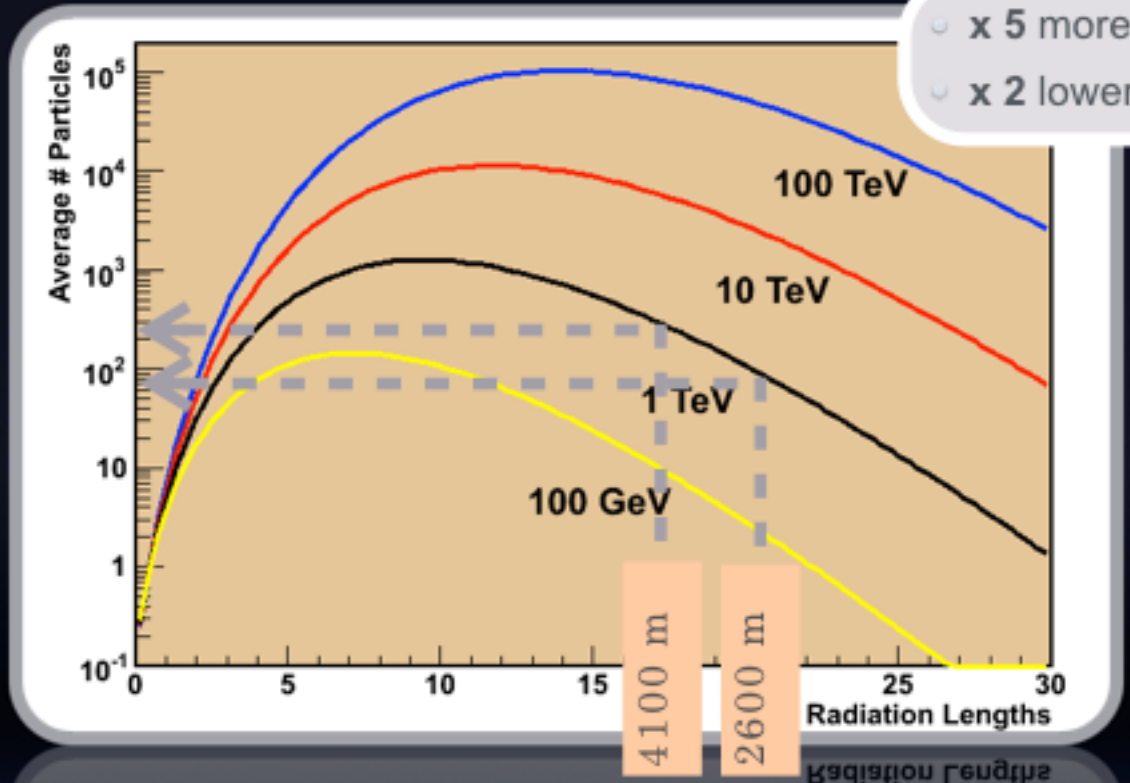
2,650 m a.s.l.

HAWC

4,100 m a.s.l.

Design improvements

- **x 5** more particles
- **x 2** lower energy threshold



Milagro

2,650 m a.s.l.

HAWC

4,100 m a.s.l.

Design improvements

Optical isolation

Milagro

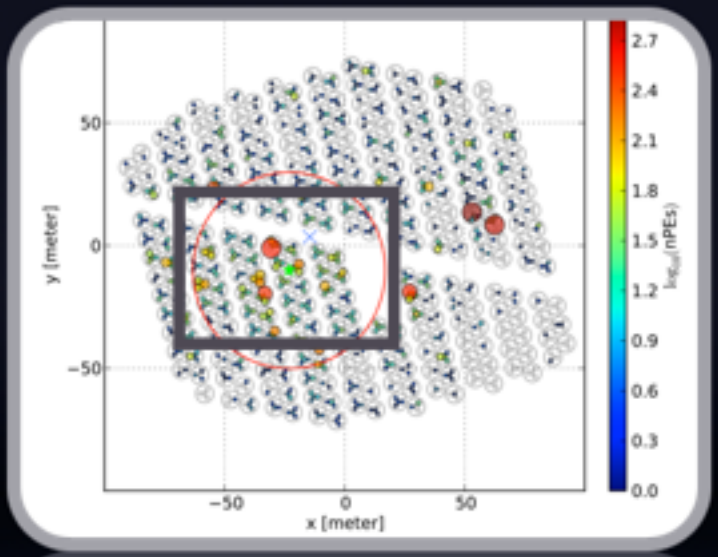
large pond

HAWC

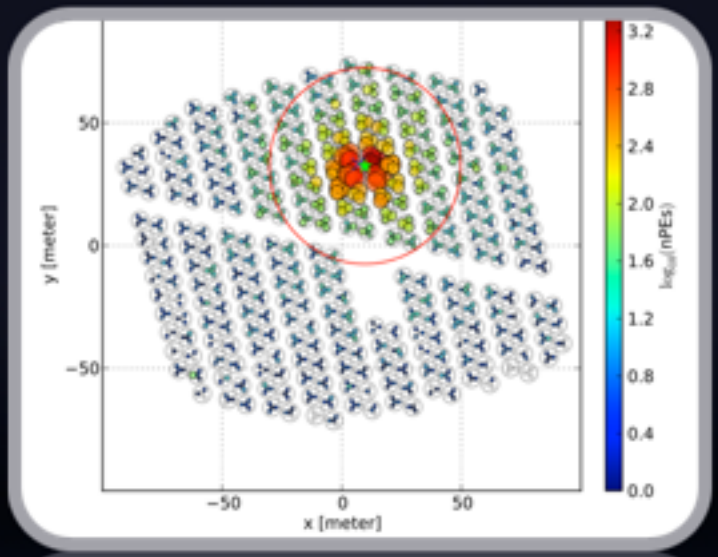
individual tanks

Design improvements

proton-shower; 24 TeV; 43°



gamma-shower; 20 TeV; 21°



Milagro

large pond

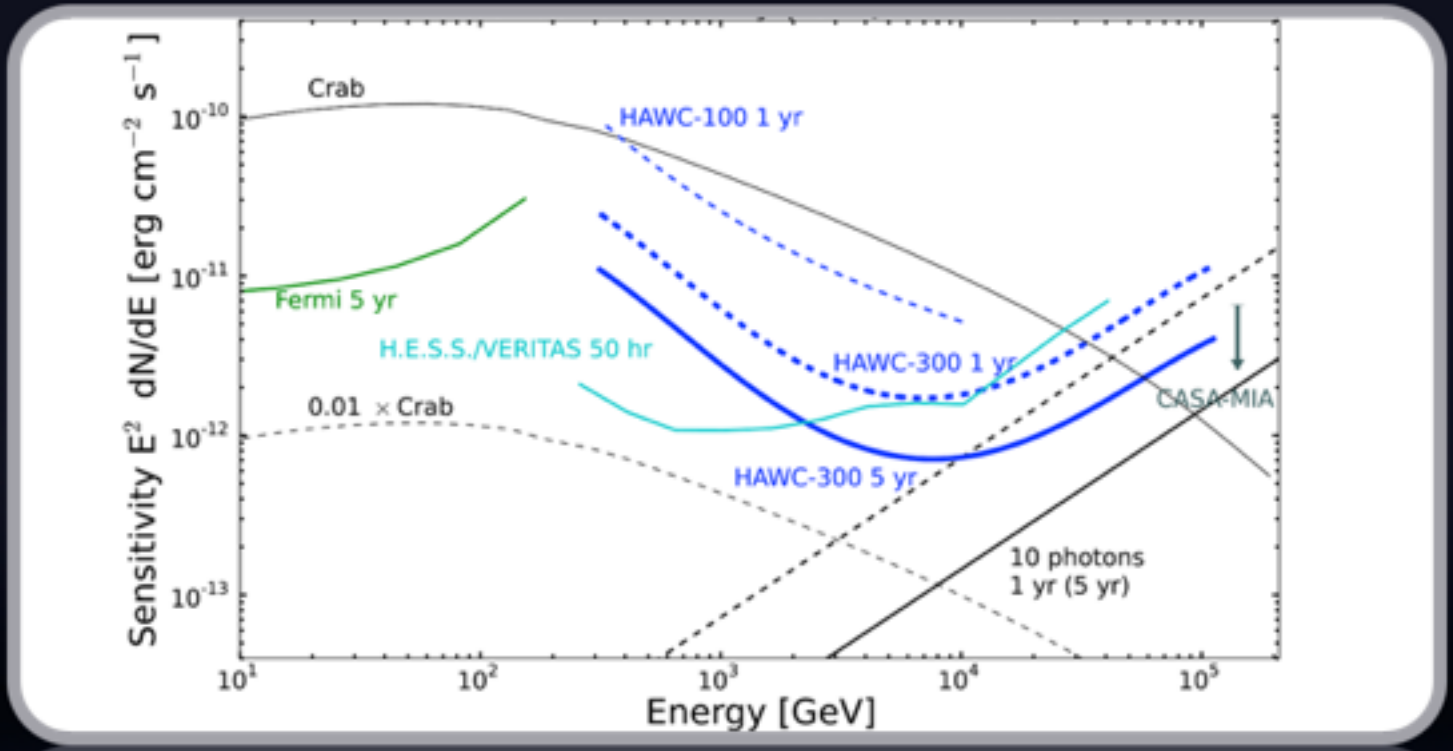
HAWC

individual tanks

Design improvements

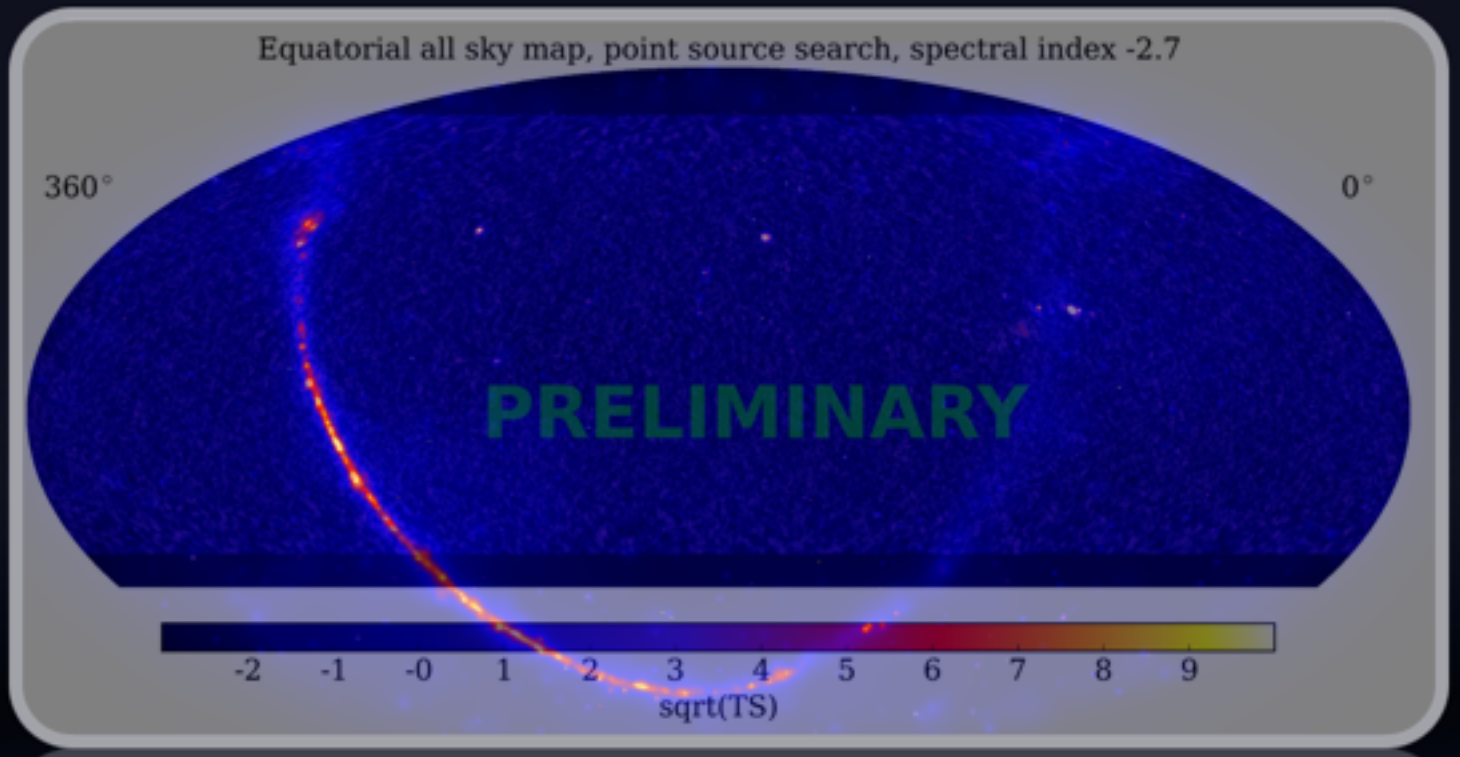
Milagro	HAWC
4,000 m^2	**22,000 m^2**

Design improvements



x15 more sensitive than Milagro

Design improvements



Fermi-LAT sky smoothed map
E > 50 GeV (Pass 8 - 6 years of data)
(courtesy of M. Ajello)

Preliminary HAWC smoothed map
E > 300 GeV (~1 year of data)
Full array

Deployment status

February 2011

Deployment status

February 2012 — VAMOS

Deployment status

September 2012 — HAWC-30

Deployment status

May 2013 — HAWC-111

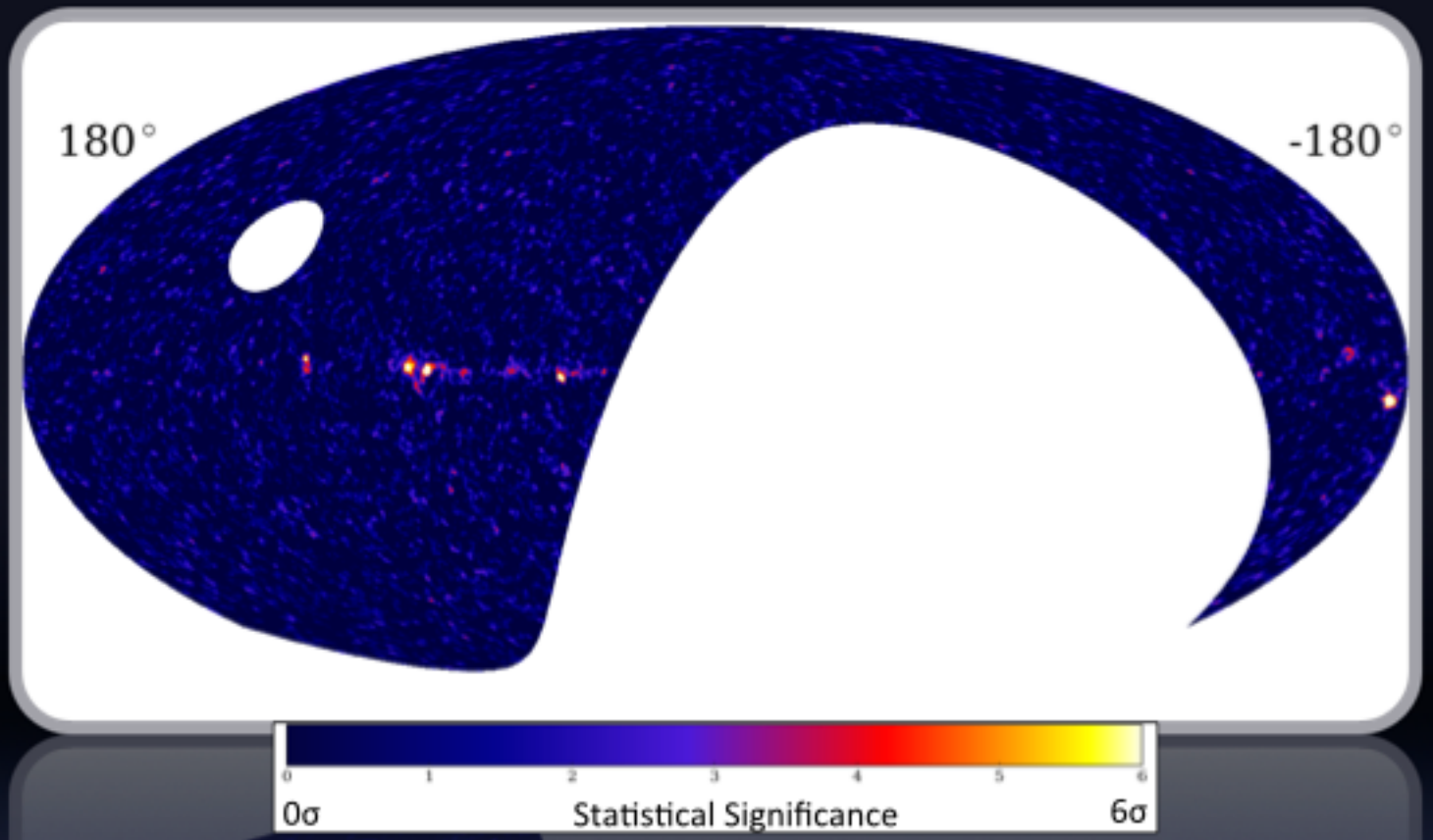
Deployment status

May 2014 — HAWC-250

Deployment status

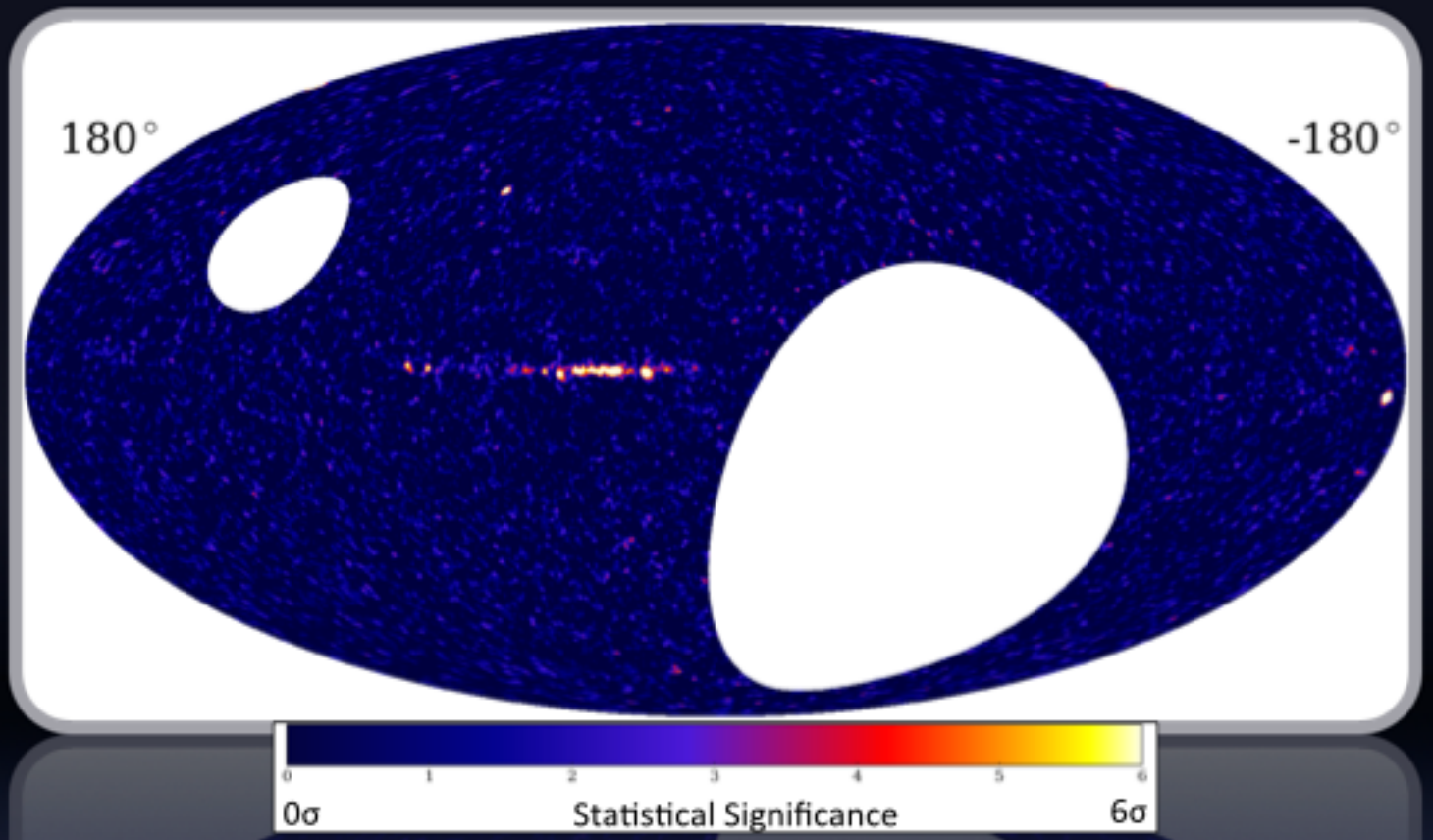
December 2014 — HAWC

Milagro — 8-year TeV sky survey (17σ Crab)



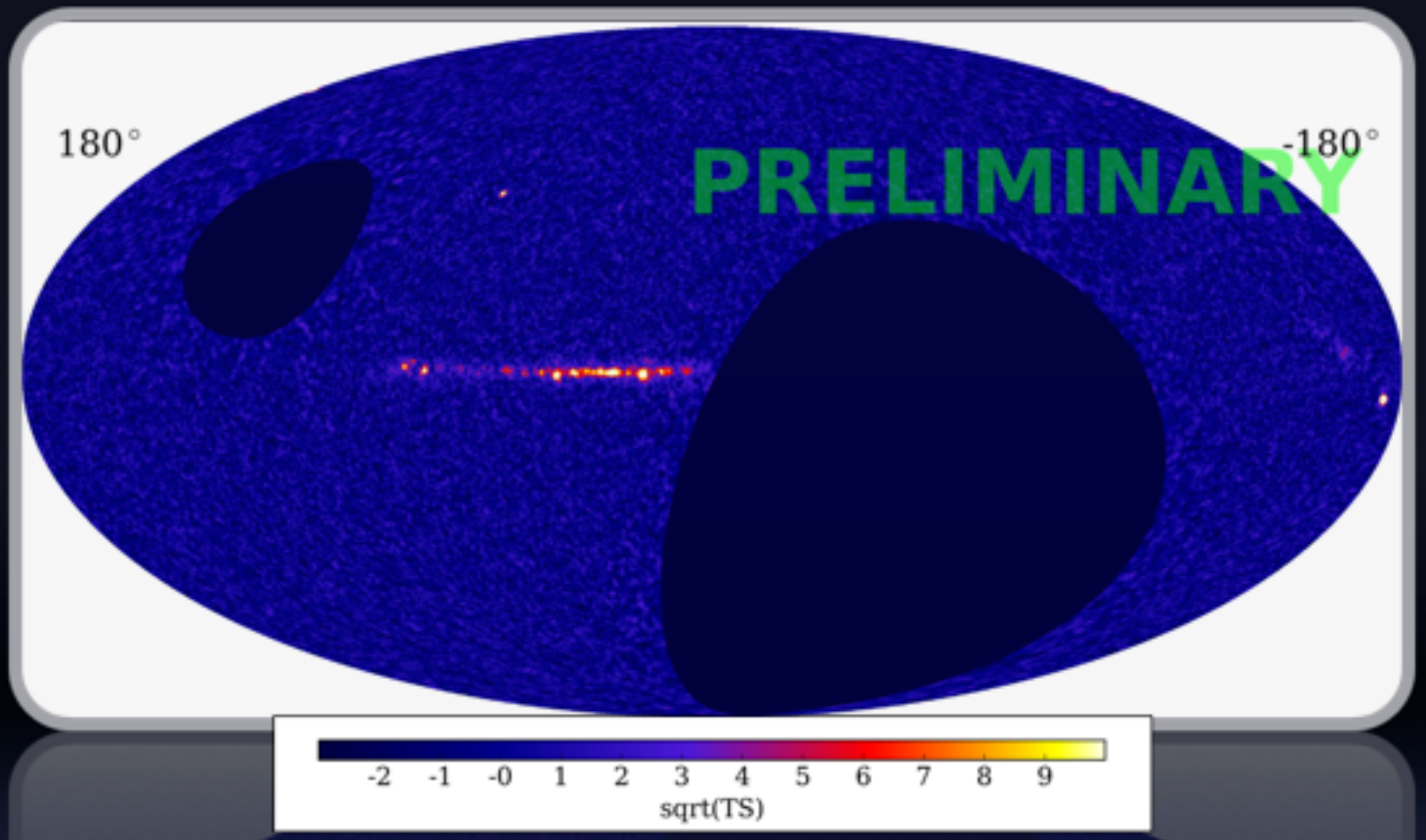
Multi-TeV sky

HAWC-250 — 150-day TeV sky survey (38σ Crab)



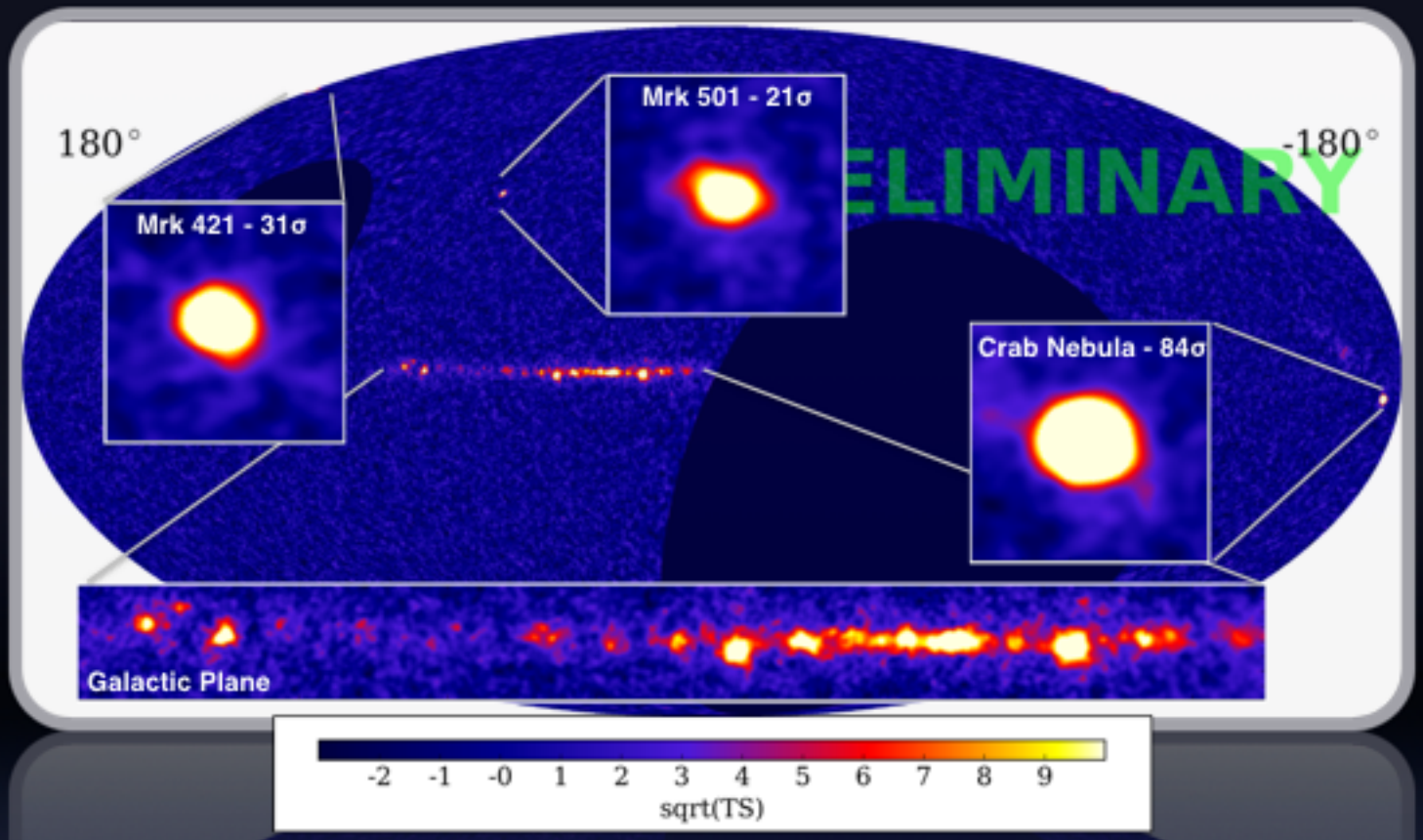
Multi-TeV sky

HAWC — 340-day TeV sky survey (>80σ Crab)



Multi-TeV sky

HAWC — 340-day TeV sky survey (>80σ Crab)

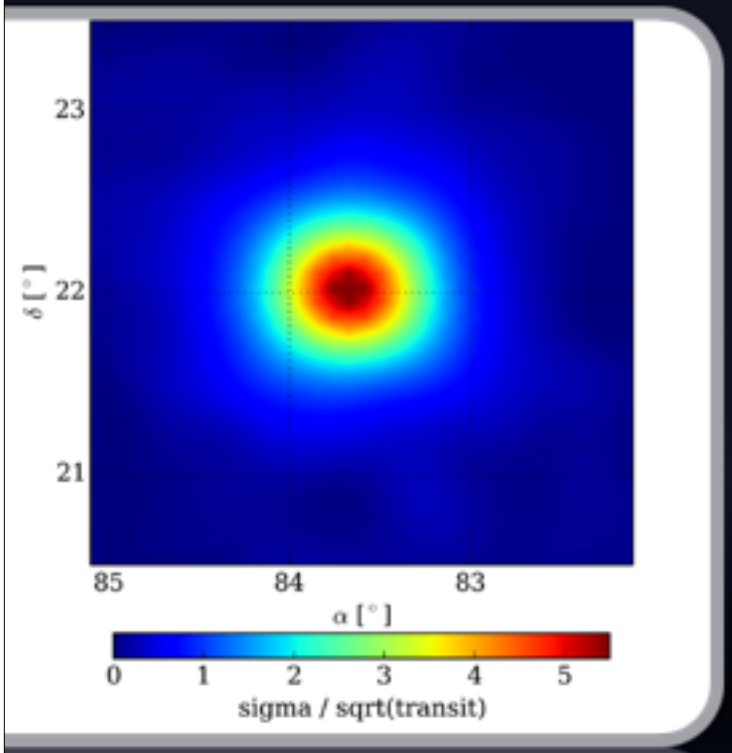


PRELIMINARY

Multi-TeV sky

reconstruction improvements

Crab Nebula

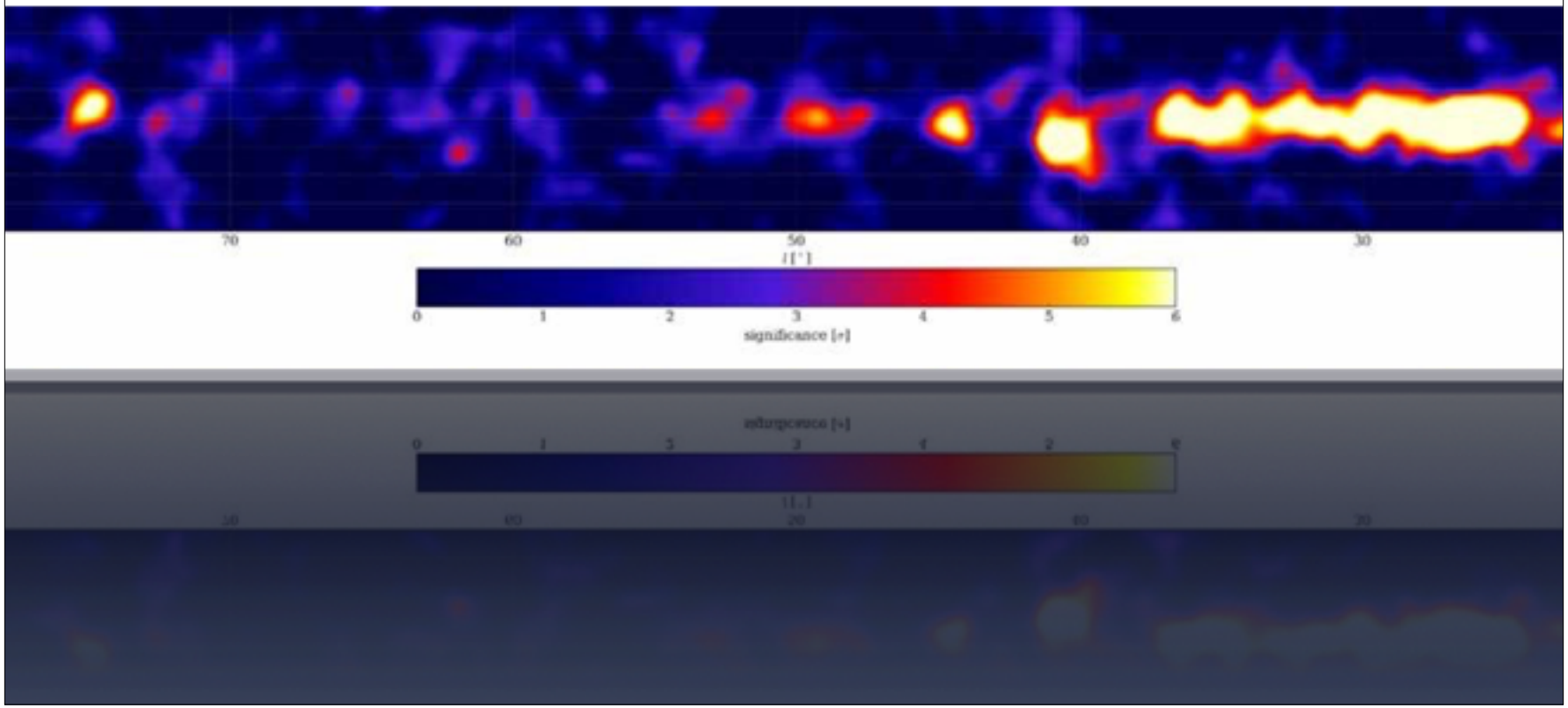


- A factor of 2 over previous “pass”
- New shower front curvature

reconstruction improvements

Preliminary results…

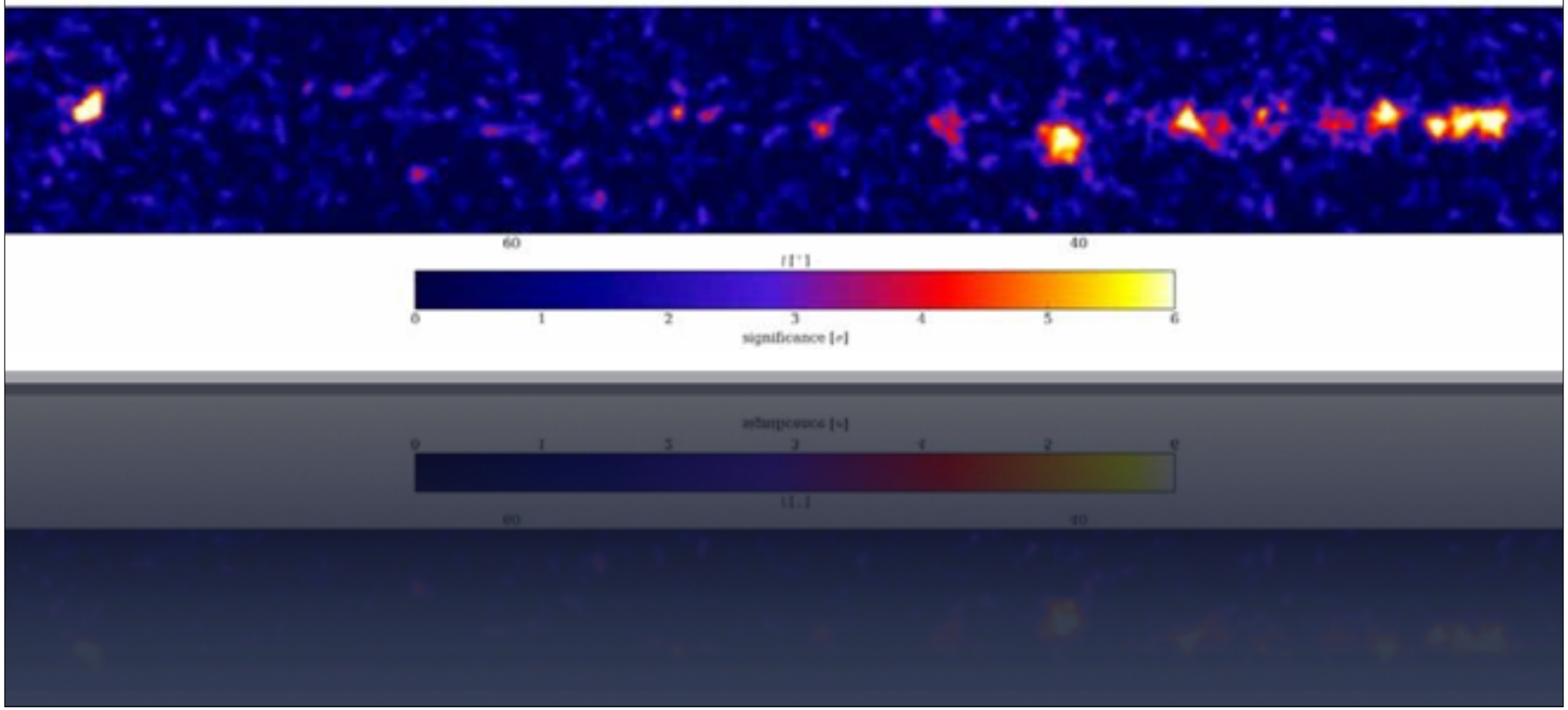
Galactic Plane



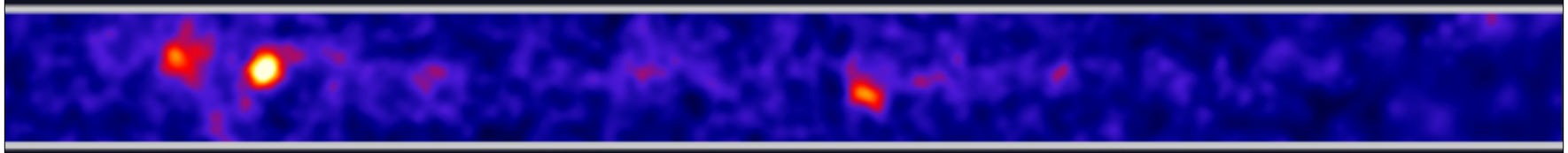
reconstruction improvements

Preliminary results…

Galactic Plane

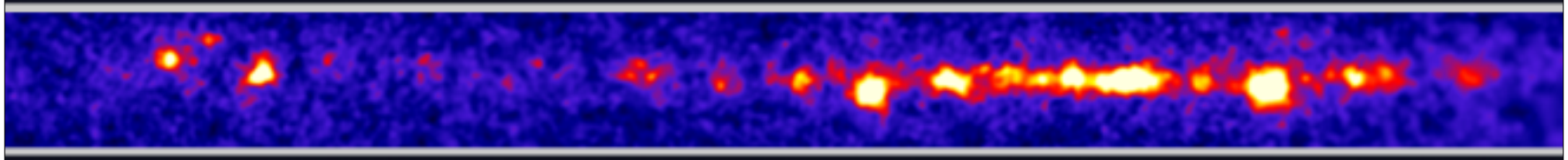


from Milagro to HAWC



- Inner Galaxy: Milagro (8 years) vs. 1st year of HAWC

from Milagro to HAWC



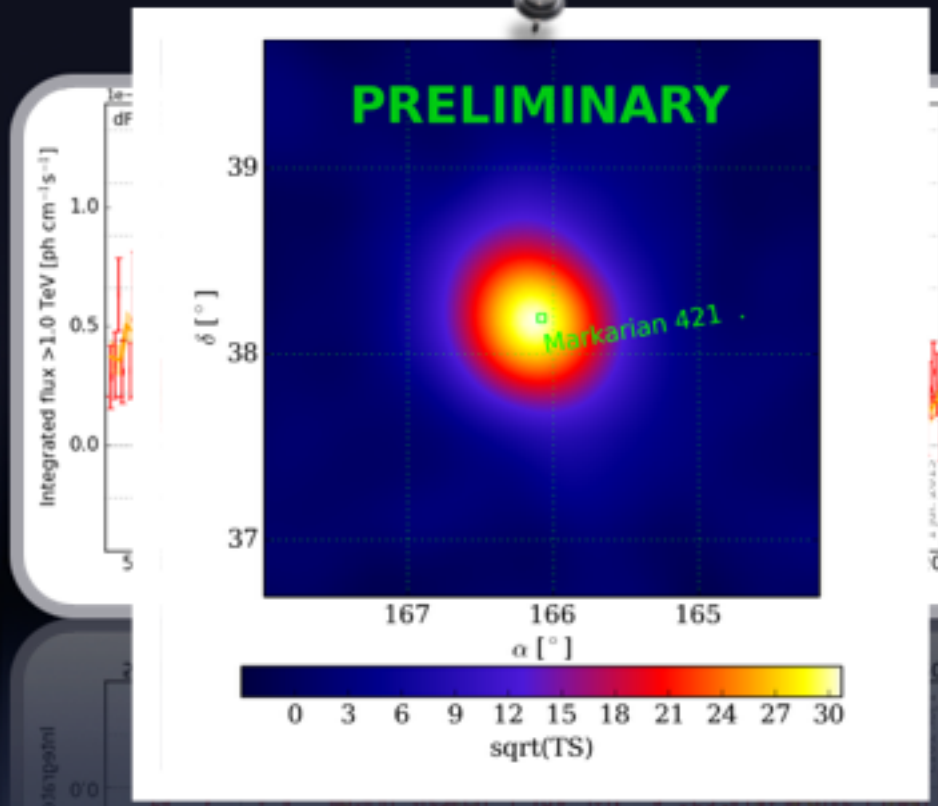
- Inner Galaxy: Milagro (8 years) vs. 1st year of HAWC

transient searches

Crab

- Pass 4 data from 26 Nov 2014 to 9 Dec 2015
- >80σ in 315 transits
- lightcurve binned in sidereal day
- consistent with constant flux

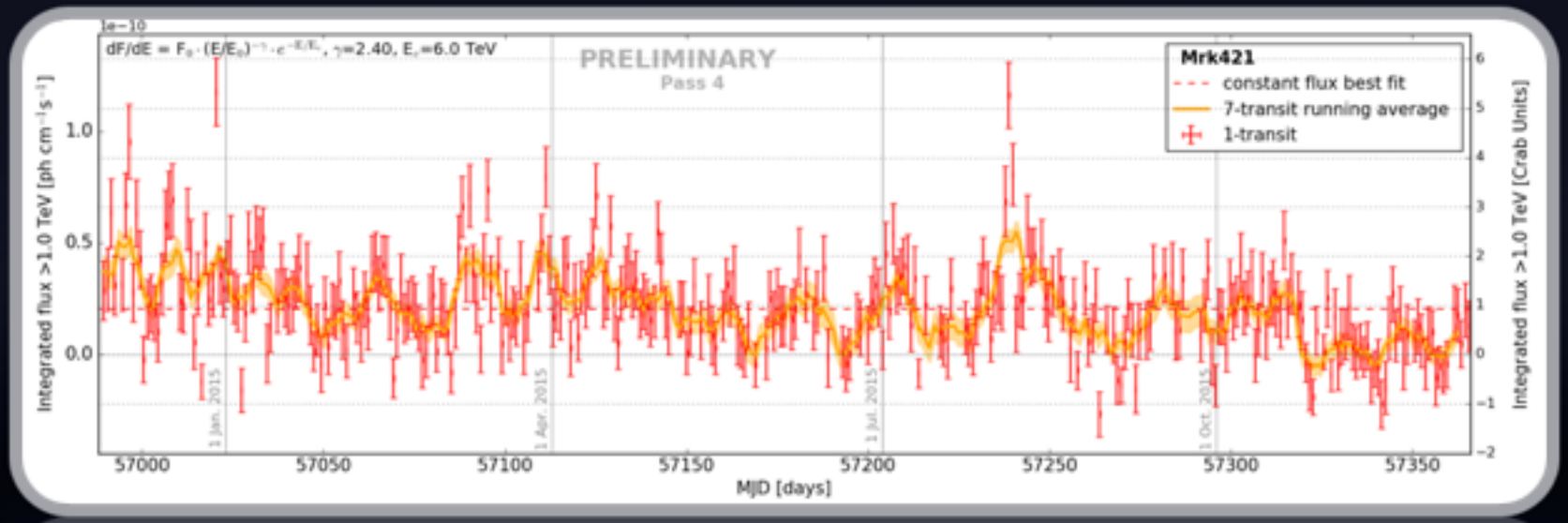
AGN flares with HAWC



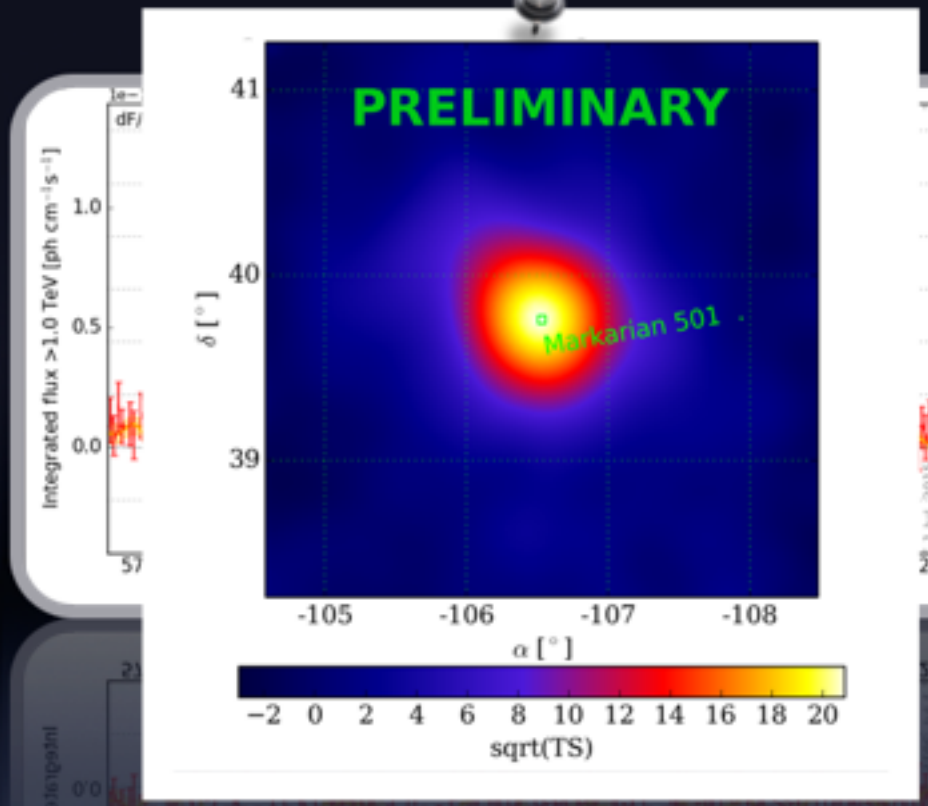
Nov 26, 2014 – Dec 9, 2015

AGN flares with HAWC

Nov 26, 2014 – Dec 9, 2015



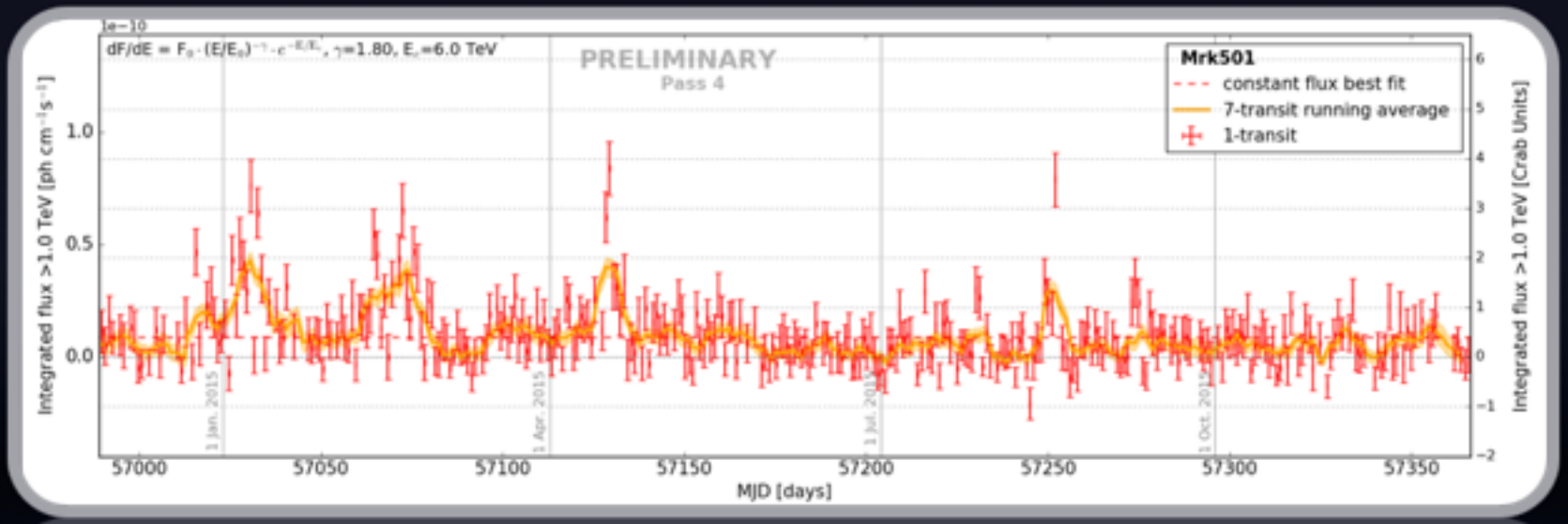
AGN flares with HAWC



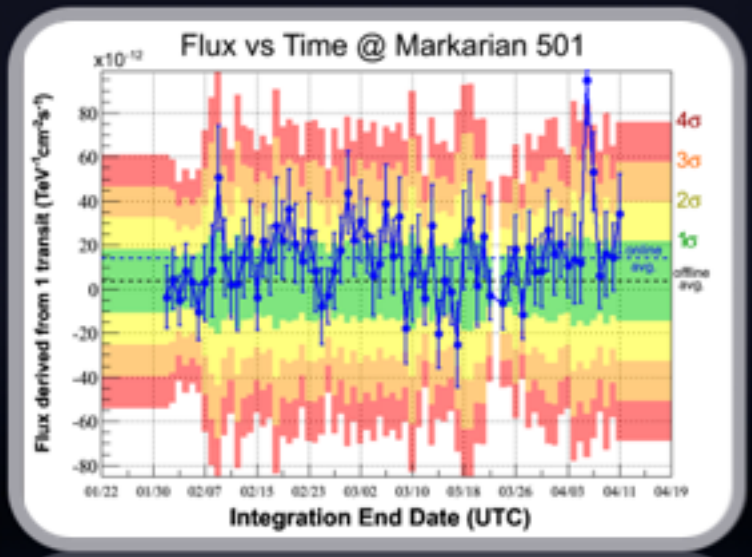
Nov 26, 2014 – Dec 9, 2015

AGN flares with HAWC

Nov 26, 2014 – Dec 9, 2015



AGN flares with HAWC



HAWC detection of increased TeV flux state for Markarian 501

ATel #8922; *Andrés Sandoval (IF-UNAM), Robert Lauer (UNM), Joshua Wood (UMD) on behalf of the HAWC collaboration*
on 7 Apr 2016; 23:38 UT
Credential Certification: C. Michelle Hui (c.m.hui@nasa.gov)

Subjects: Gamma Ray, TeV, VHE, Request for Observations, AGN, Blazar

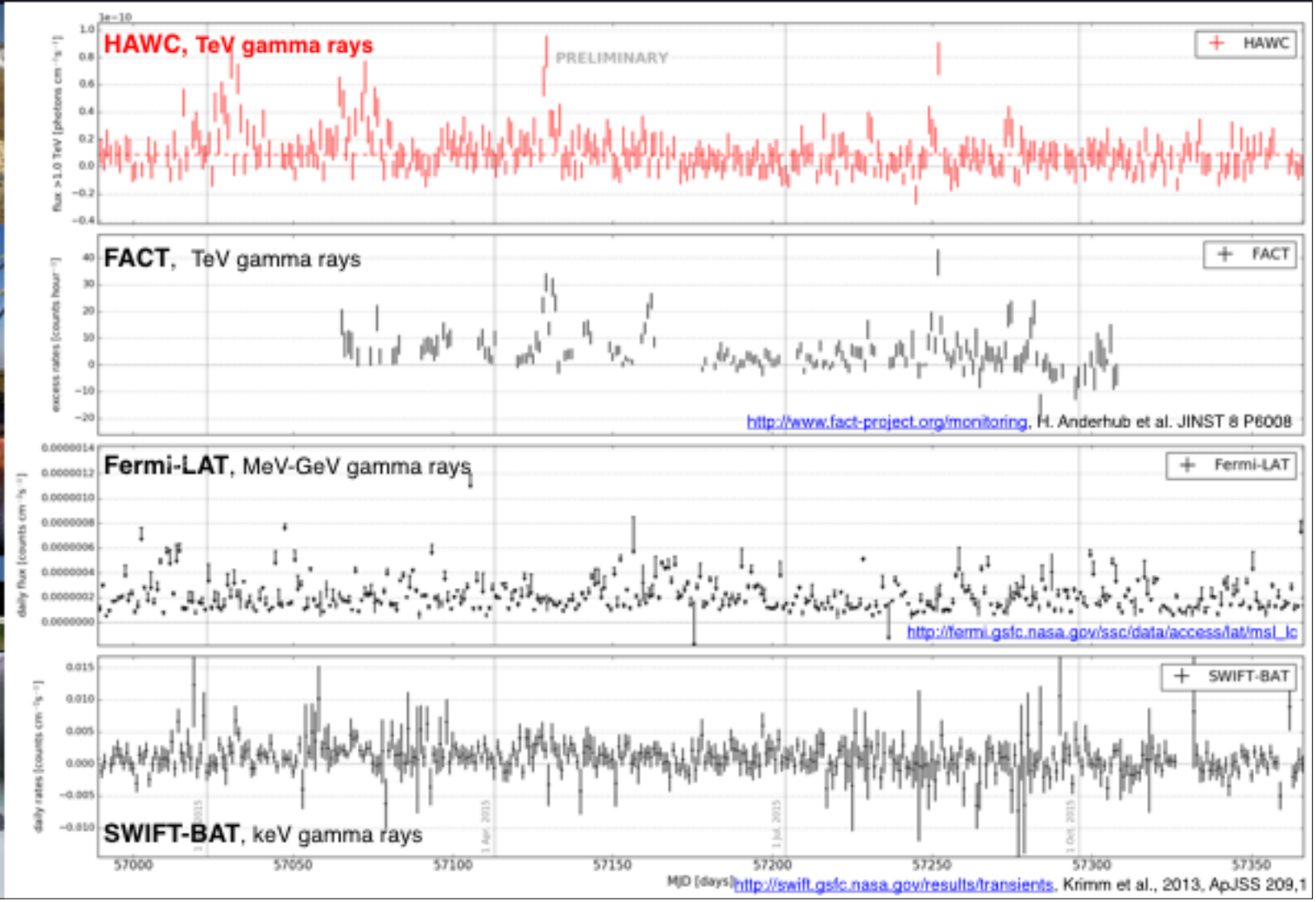
Tweet Recommend 15

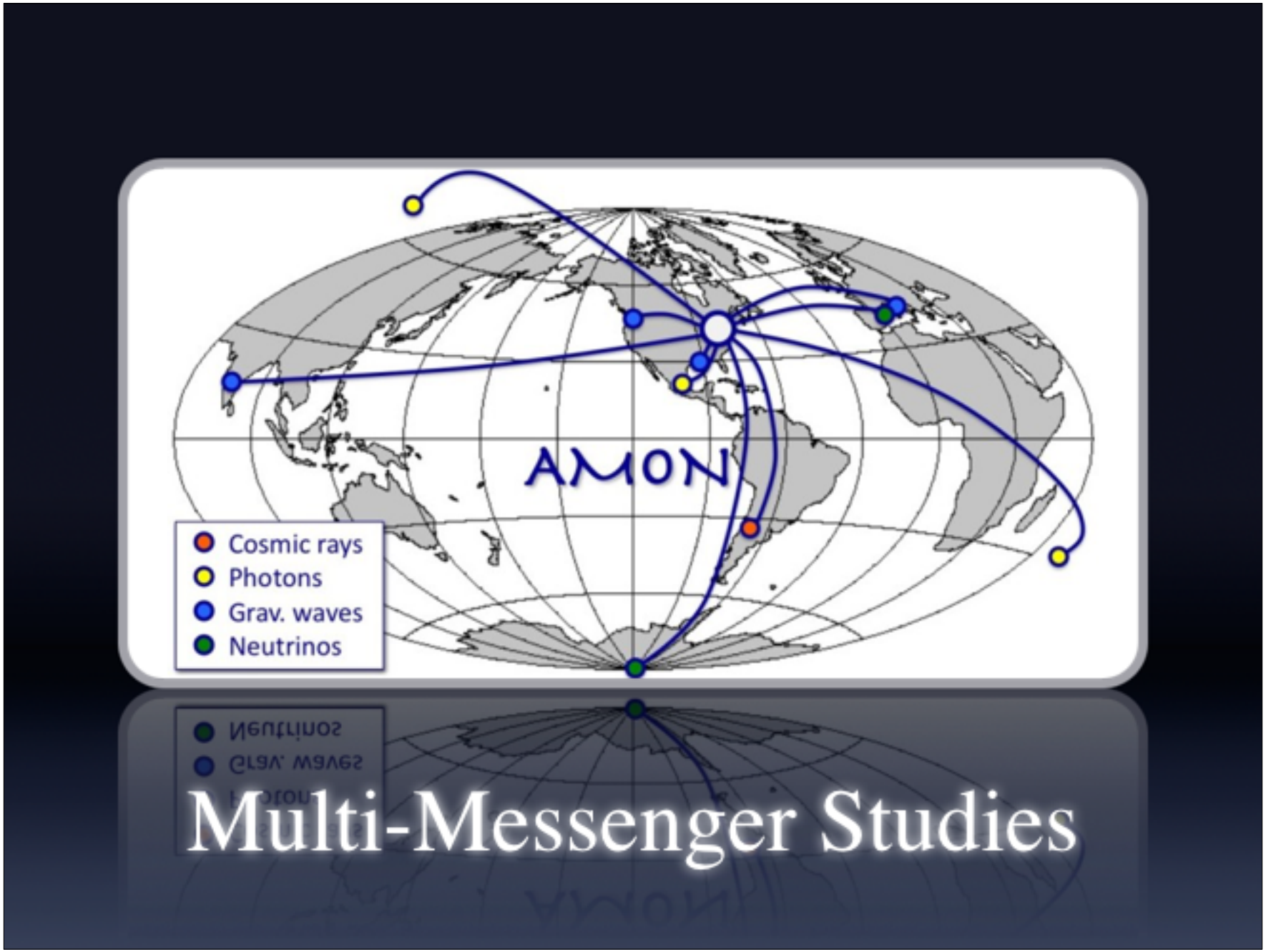
The HAWC Observatory measured an increased gamma-ray flux from the direction of the BL Lac Markarian 501 (z=0.033) at the level of (4.88 +/- 1.05) x 10e-11 photons cm^-2 s^-1 above 1 TeV when averaged during the 6 hour transit over HAWC on April 6, 2016 (MJD 57484.31 - 57484.56) which is 2.2 times the average Crab flux observed by HAWC. For the following transit on April 7, 2016 (MJD 57485.30 - 57485.55), a decreased but still above-average flux of (2.78 +/- 0.09) x 10e-11 photons cm^-2 s^-1 was observed, 1.3 times the Crab flux seen by HAWC. The flux on April 6 lies 4 sigma above the average flux of 0.89 x 10e-11 photons cm^-2 s^-1 that was measured for this source by HAWC during the previous year. The flux level on April 7 is 2 sigma above this average and seems to indicate a declining but on-going high flux state. All flux values are obtained from a maximum likelihood fit under the assumption of a fixed spectral shape with power law index of 1.8 and exponential cut-off at 6 TeV. These spectral parameters are the best fit results for HAWC data from Markarian 501 collected between November 2014 and December 2015. HAWC is a TeV gamma ray water Cherenkov array located in the state of Puebla, Mexico that monitors 2/3 of the sky every day with an instantaneous field of view of ~2 sr. The HAWC contact people for this analysis are Robert Lauer (University of New Mexico, rjlauer@unm.edu) and Michelle Hui (Marshall Space Flight Center, c.m.hui@nasa.gov).

First TeV flare alert from HAWC!

multi-wavelength studies

Daily Monitoring of Markarian 501

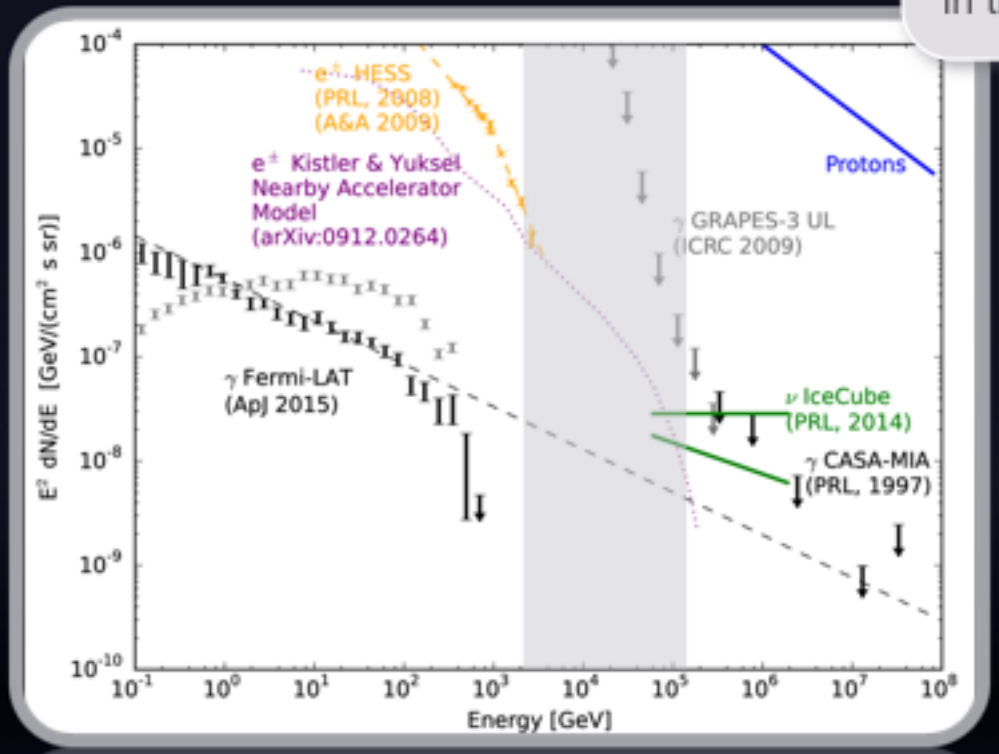




Multi-Messenger Studies

Isotropic Diffuse Emission

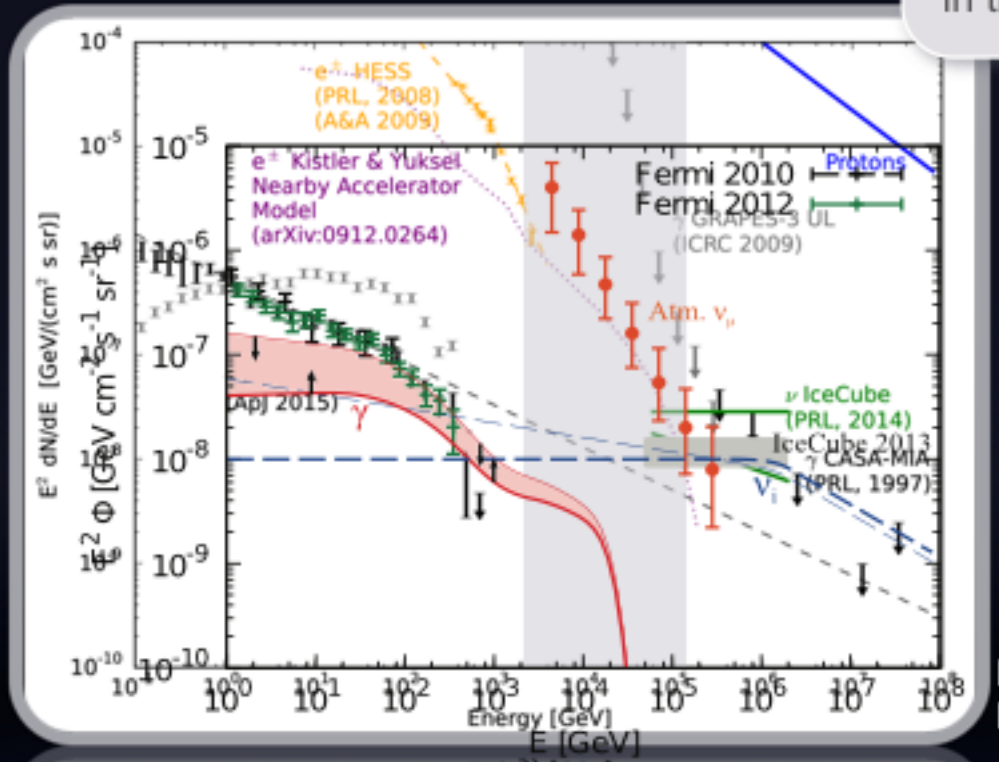
Neither measurements nor limits in the **10-100 TeV** range



connecting HAWC with IC

Isotropic Diffuse Emission

Neither measurements nor limits in the **10-100 TeV** range

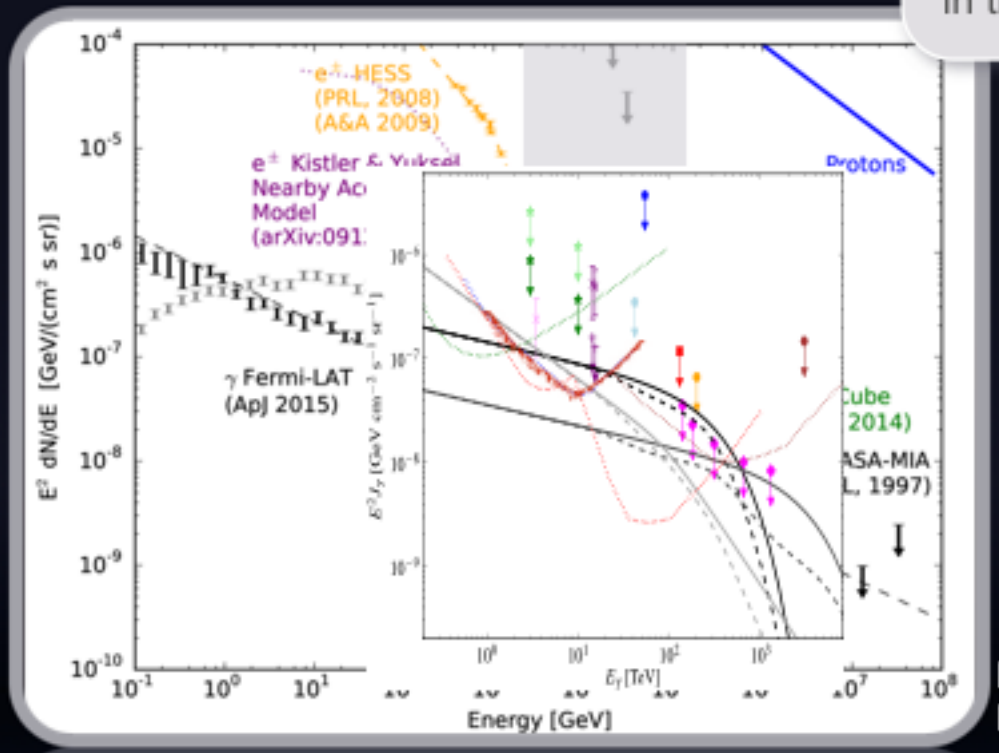


Kohta Murase, M. Ahlers, B.C. Lacki, Phys. Rev. D 88 (2013) 121301

connecting HAWC with IC

Isotropic Diffuse Emission

Neither measurements nor limits in the **10-100 TeV** range

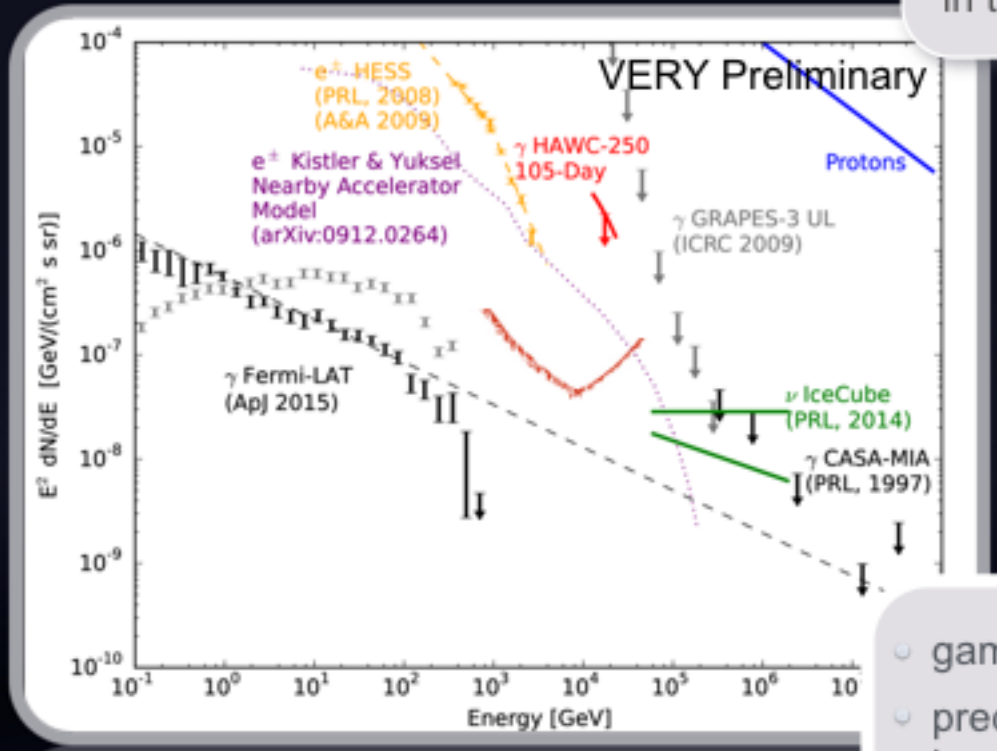


M. Ahlers, Kohta Murase,
Phys. Rev. D 90 (2014) 023010

connecting HAWC with IC

Isotropic Diffuse Emission

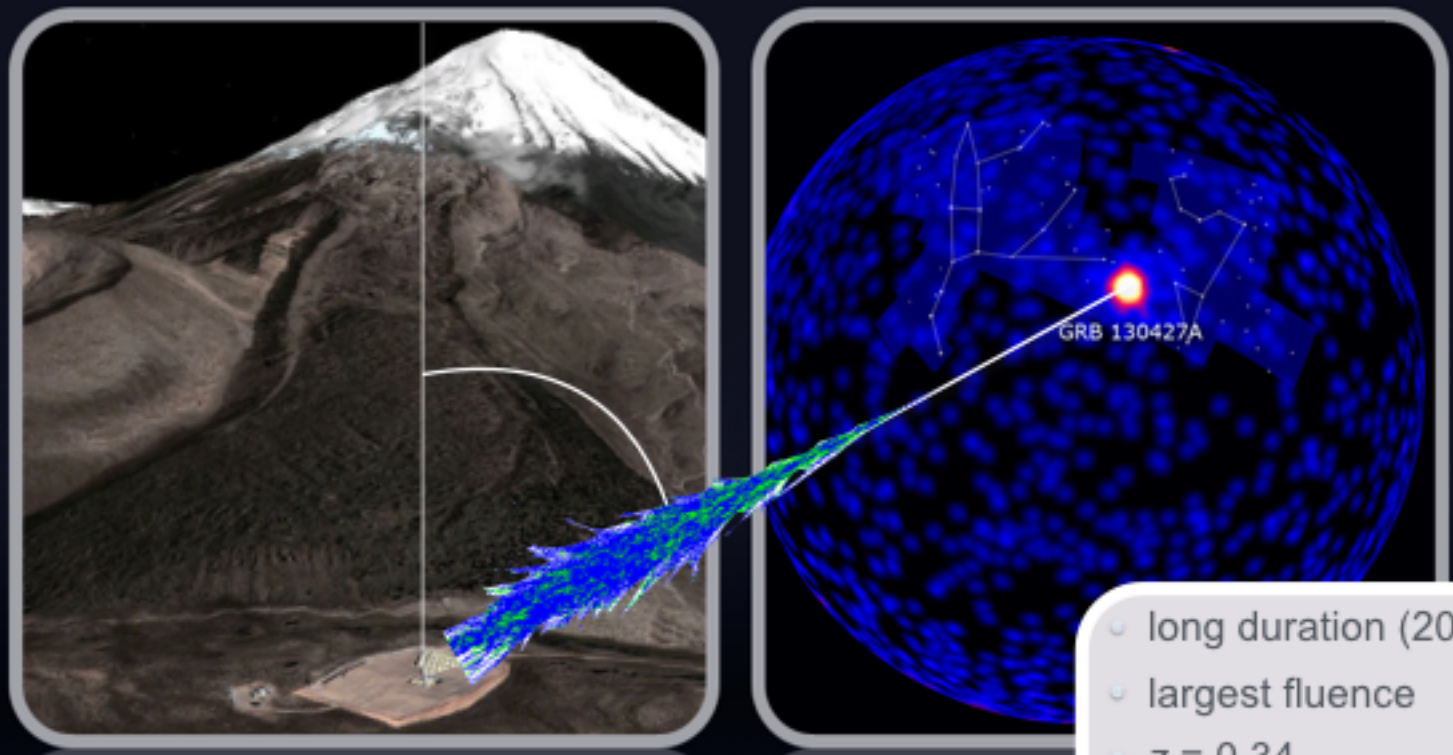
Neither measurements nor limits in the **10-100 TeV** range



- gamma/hadron separation
- precise understanding of background efficiency

connecting HAWC with IC

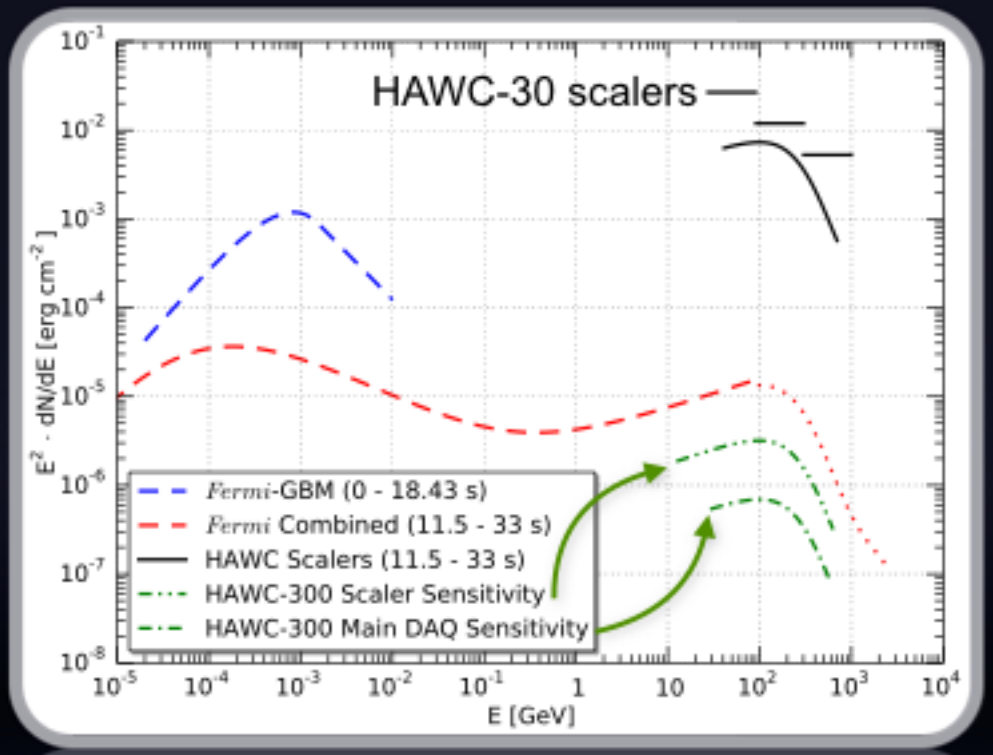
GRB 130427A

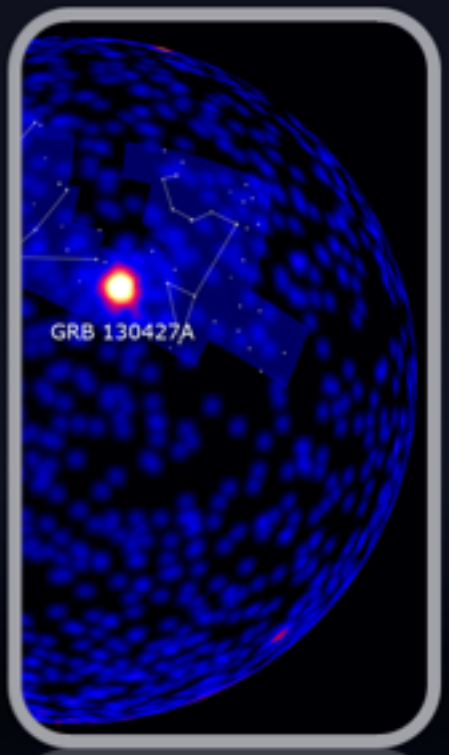


- HAWC-30
- θ = 57°
- only scalers DAQ up

- long duration (20 h)
- largest fluence
- z = 0.34
- highest energy photon (95 GeV)

GRB 130427A



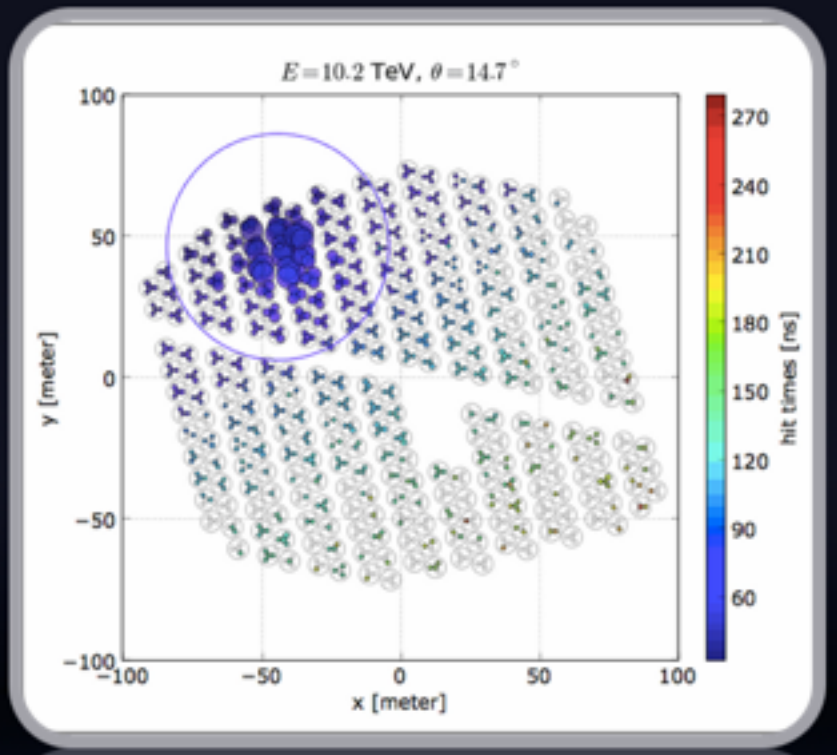


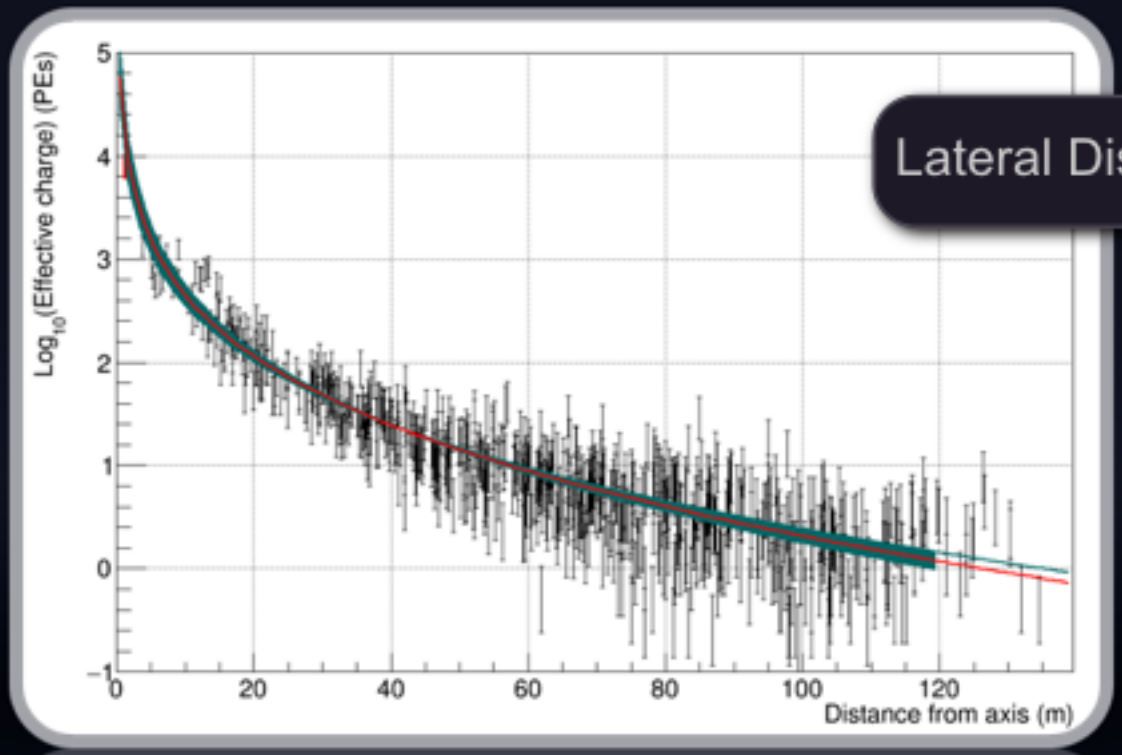
- HAWC-30
- $\theta = 57°$
- only scalers DAQ up

"Search for gamma-rays from the unusually bright GRB 130427A with the HAWC Gamma-ray Observatory," *The Astrophysical Journal*, Volume 800 (2015) Number 2, p78

Energy measurement

Number of triggered PMTs is the energy proxy



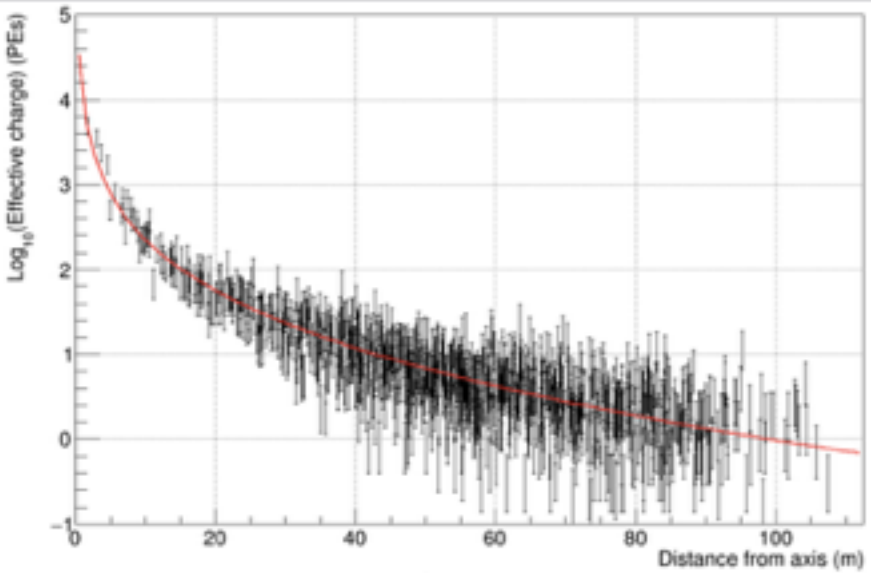


Lateral Distribution Function

Energy measurement

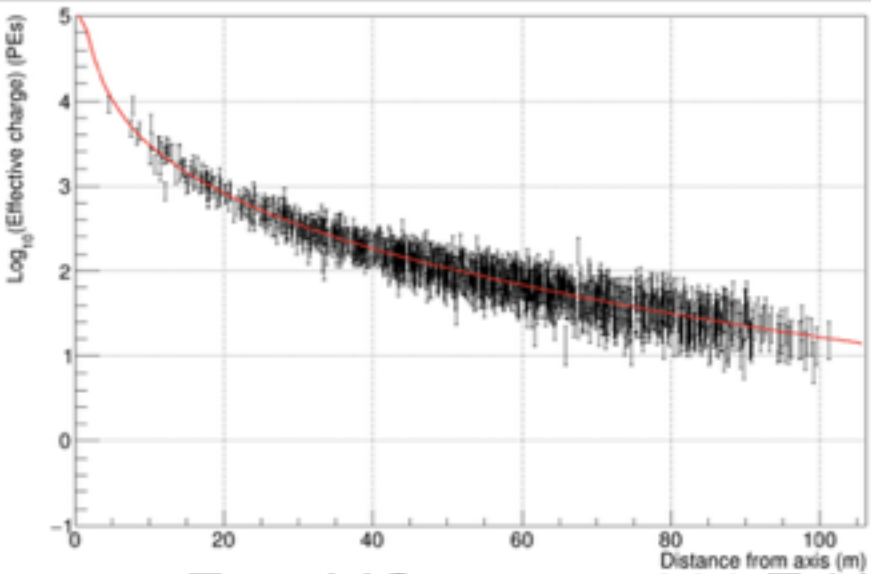
Energy measurement

Preliminary results!



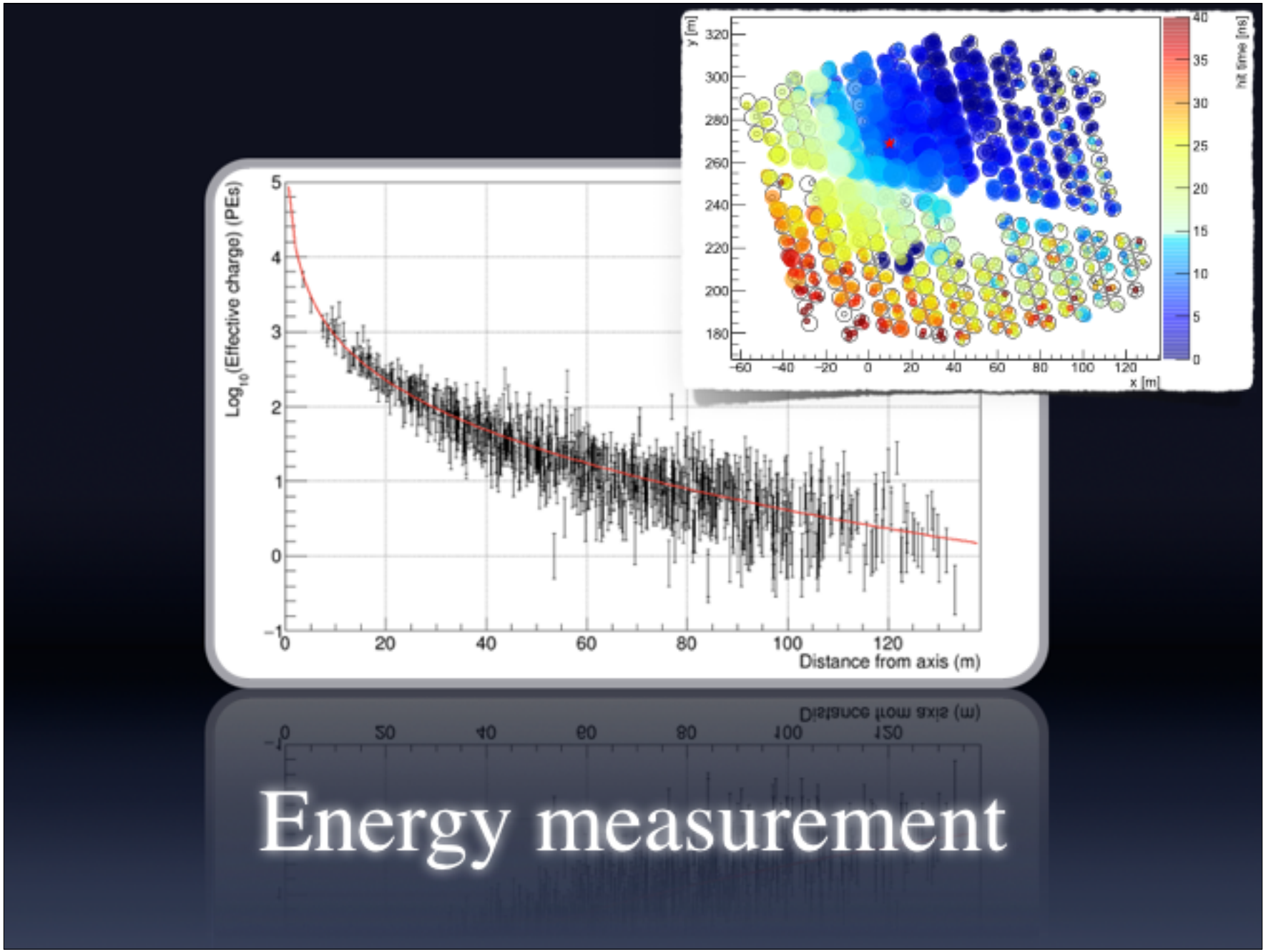
True MC energy 18 TeV

Reconstructed energy 17 TeV

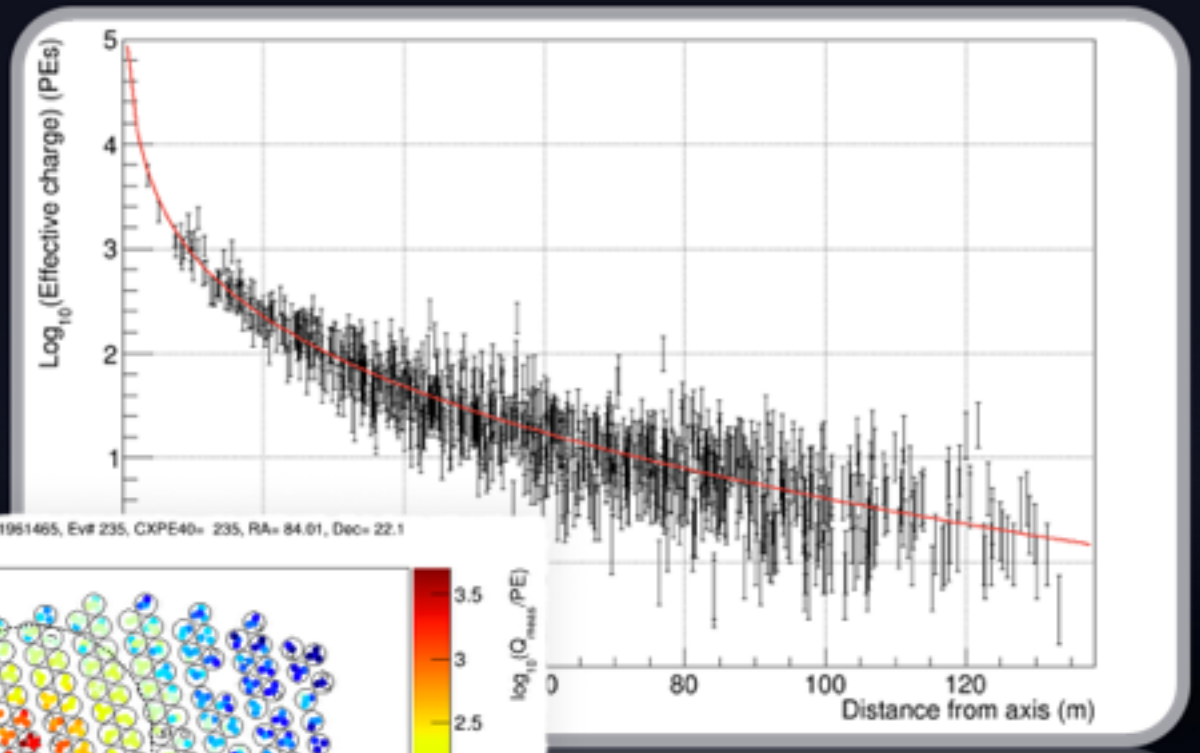


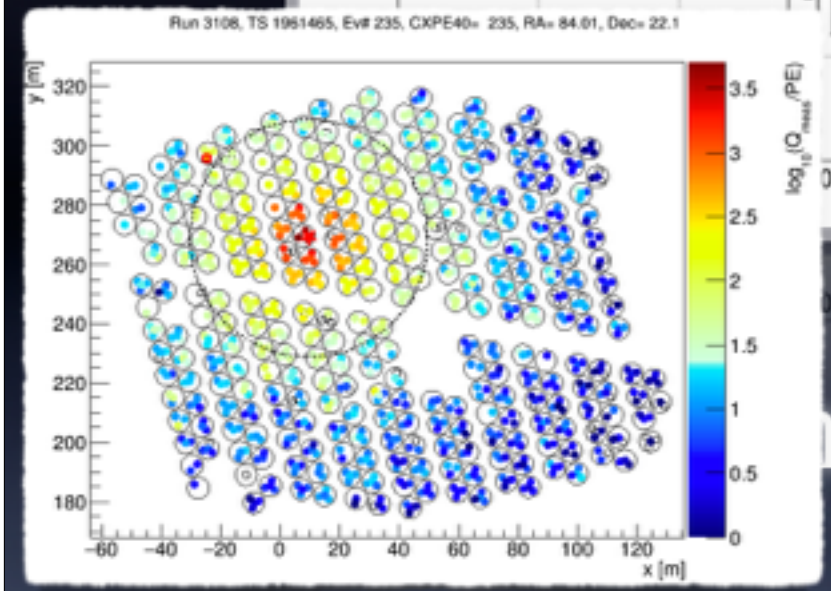
True MC energy 380 TeV

Reconstructed energy 350 TeV


Log10(Effective charge) (PEs)
Distance from axis (m)
y [m]
x [m]
hit time [ns]
Energy measurement

Preliminary results!





easurement

Outlo

Other results

- Dark matter, extended regions, Cosmic rays, …
- EBL, solar physics, …

Multi-wavelength physics

- MoUs with IceCube, IACTs, etc
- AMON
- HAWC alerts

Enhancements

- Array of Outriggers

Outlook

Other results

- Dark matter, extended regions, Cosmic rays, …
- EBL, solar physics, …

Multi-wavelength physics

- MoUs with IceCube, IACTs, etc
- AMON
- HAWC alerts

Enhancements

- Array of Outriggers

Future Experiment

- Southern Observatory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H A W C
Thank you very much