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Turbulence and structures in dispersive MHD.

T. Passot

Observatoire de la Cote d'Azur
Nice
France

Strada Costiera 11, 34014 Trieste, Italy - Tel. +39 040 2240 111; Fax +39 040 224 163 - sci_info@ictp.it, www.ictp.it

Turbulence and structures in dispersive MHD

T. Passot, D. Borgogno, D. Laveder, P.L. Sulem,
L. Marradi
Observatoire de la Côte d'Azur, Nice, France

OUTLINE

- Introduction : evidence of dispersive Alfvén waves in the solar wind and the terrestrial magnetosheath and of the signature of turbulence and coherent structures

- Break of the spectrum at the ion gyroscale: tentative explanations

- The simple but nevertheless complex case of the forced 1D DNLS equation

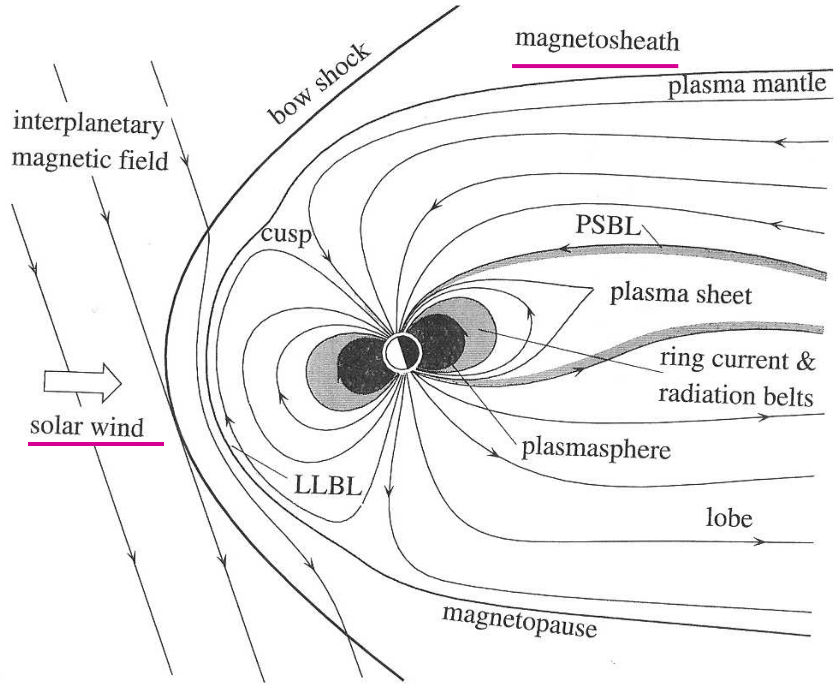
- The Landau fluid model as a tool for investigating dispersive turbulence

- Preliminary 1D Landau fluid simulations

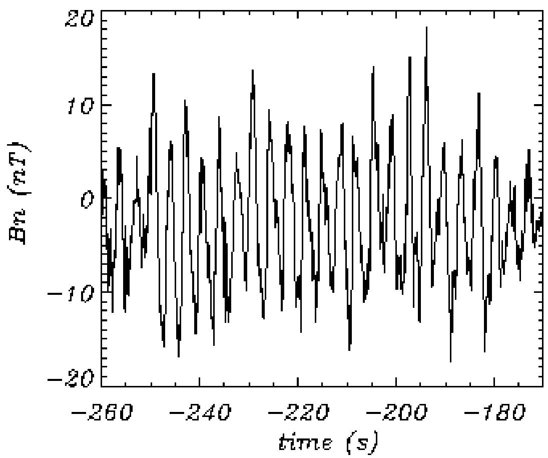
- 2D Hall-MHD simulations

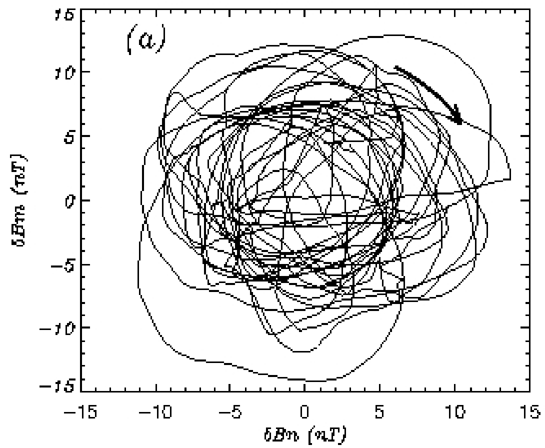
Evidence of DAWs

Quasi-monochromatic dispersive Alfvén waves are commonly observed in the solar wind and in the magnetosheath



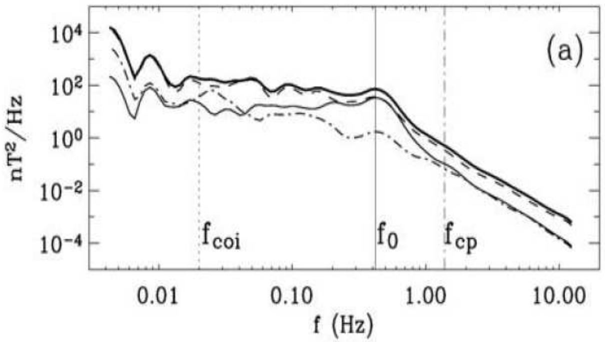
Observation by CLUSTER satellites downstream the quasi-perpendicular bow shock (Alexandrova et al., J. Geophys. Res., (2004, 2006)



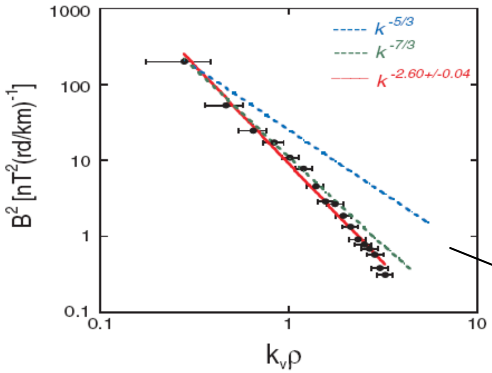


Presence of almost monochromatic left-hand circularly polarized Alfvén waves

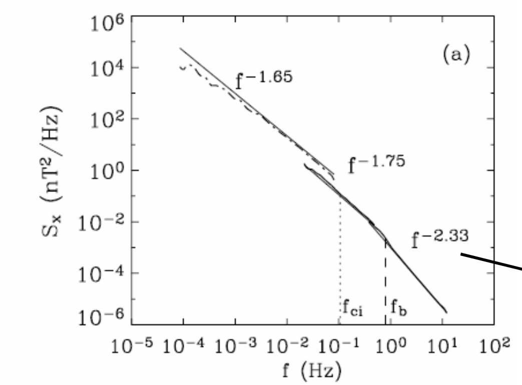
Evidence of turbulence



Magnetic energy spectrum in the magnetosheath downstream of the bow shock (Alexandrova et al., JGR, 2006).



Magnetic energy spectrum in the magnetosheath close to the magnetopause (Sahraoui et al., PRL 2006)



Solar wind turbulent spectrum (Alexandrova et al., 2007)

Space plasmas such as the solar wind or the magnetosheath are turbulent magnetized plasmas with essentially no collisions.

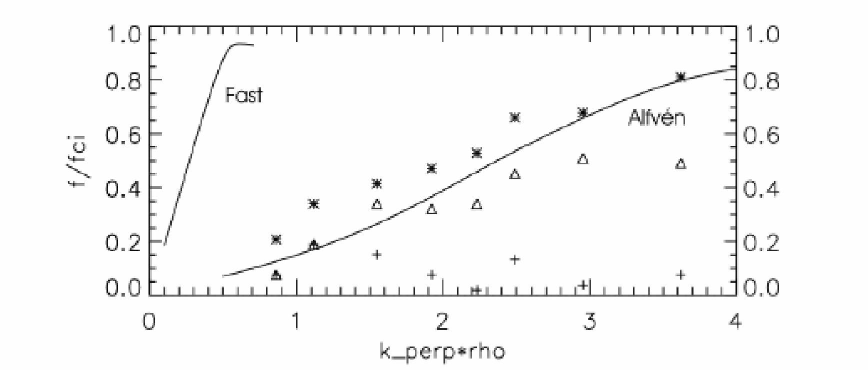
Observed cascade extends beyond the ion Larmor radius: kinetic effects play a significant role.

Here identified as mirror modes using k-filtering technique: modes with essentially zero frequency in the plasma frame

Range of observed frequency power law indices between -2 and -4.5 (Leamon et al. 1998)

RH polarized outward propagating waves (Goldstein et al. JGR 94)

The Alfvén wave cascade develops preferentially perpendicularly to the ambient magnetic field.
Assuming frequencies remain relatively small in comparison with the ion gyrofrequency, the dynamics should be dominated by Kinetic Alfvén waves (and slow modes, but these ones are highly dissipative).
KAWs have been clearly identified using k-filtering technique in the cusp region (Sahraoui et al. AIP, 2007).



Evidence of KAWs

Another medium where KAWs play a fundamental role is the solar corona, where they are believed to mediate the conversion of large-scale modes into heat.

The nature of the fluctuations associated with the power spectrum at frequencies larger than the ion-gyrofrequency in the satellite frame is however not yet established in all situations.

Another issue:

Formation and evolution of small-scale coherent structures (filaments, shocklets, magnetosonic solitons, magnetic holes) observed in various spatial environments :

Typical length scale of the structures:
a few ion Larmor radii.

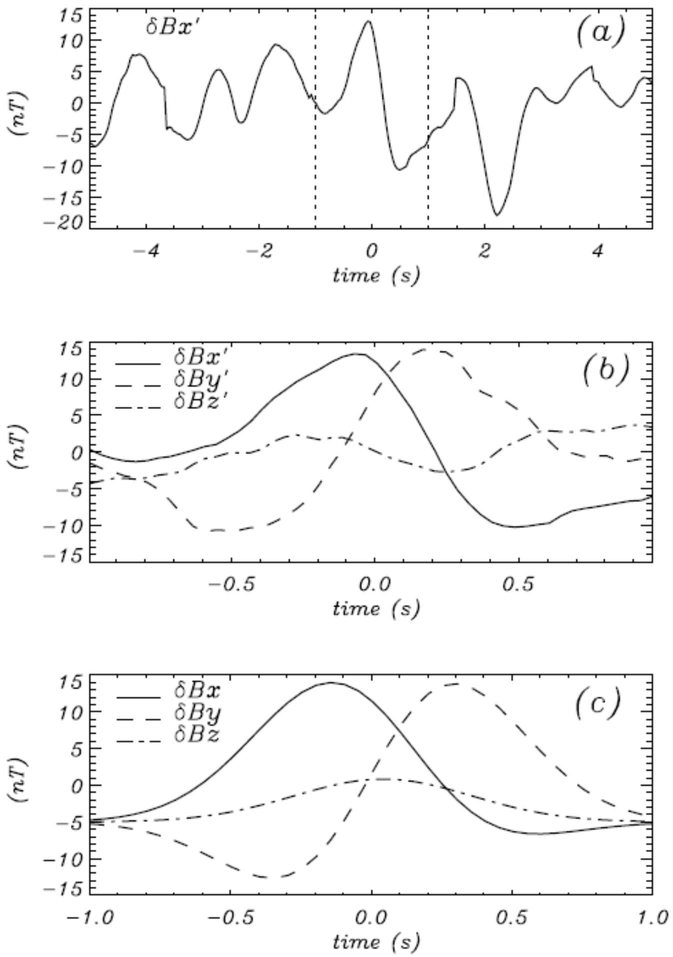


Figure 8. Magnetic field fluctuations, taking $\tau \simeq -420$ s (1755:16 UT) as the origin of time. (a) Fluctuations $\delta B_{x'}$ during 10 s around τ. (b) Fluctuations of the magnetic field components ($\delta B_{x'}$, $\delta B_{y'}$, $\delta B_{z'}$) for the 2-s period around τ. (c) The z-aligned current tube simulation (δB_x, δB_y, δB_z).

Signature of magnetic filaments
(Alexandrova et al. JGR 2004)

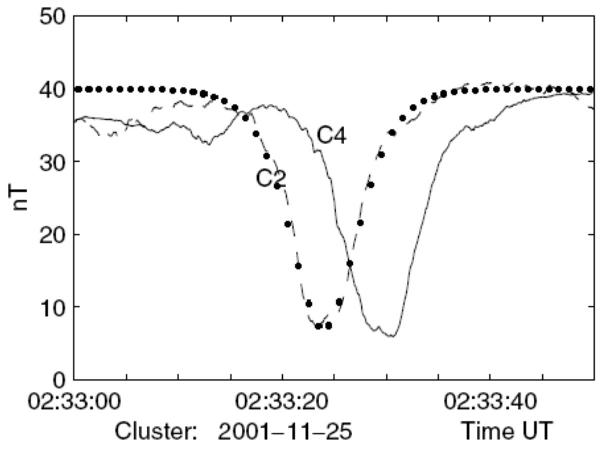
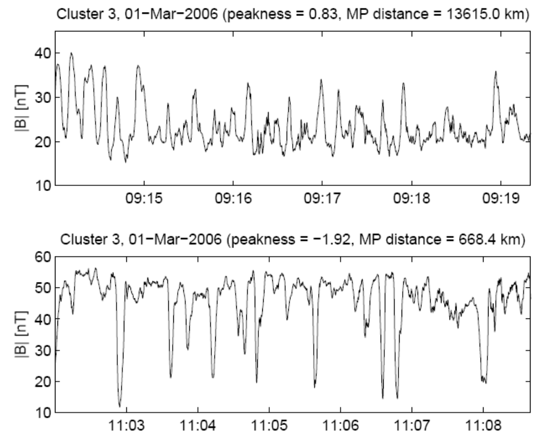


FIG. 1. A large scale soliton observed by Cluster spacecraft C2 (dashed) and C4 (solid) in the total magnetic field. Marked curve shows fit of $b_0 \operatorname{sech}^2[(t - t_0)/\delta t]$ with $b_0 = -33$ nT and $\delta t = 4.4$ s. The soliton moves with velocity $u_0 \approx 250$ km/s and has a width of 2000 km. The position of Cluster satellites was $(-4, 17, 5)\, R_E$ GSE.

Slow magnetosonic solitons
(Stasiewicz et al. PRL 2003)



Mirror structures in the terrestrial magnetosheath
(Soucek et al.JGR 2008)

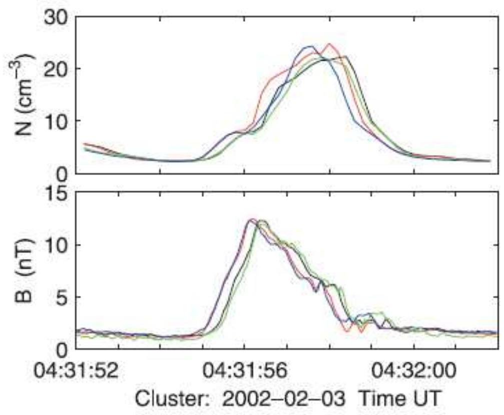


Figure 2. Pulse-like enhancements of the plasma density and magnetic field measured on four Cluster spacecraft: C1–C4, which are color coded in sequence: black, red, green, blue. The measurements represent signatures of fast magnetosonic shocklets moving with supersonic speed in a high-β plasma.

fast magnetosonic shocklets
(Stasiewicz et al. GRL 2003)

Question: How does turbulence develop at dispersive scales?

- Is the transfer suppressed in the parallel direction?
- Are solitonic-type structures generically formed or does weak (or strong) turbulence prevail?
- What kind of structures are formed in the transverse direction?
- What is the origin of the spectral break at the ion Larmor radius scale.

Tentative models for the « dissipation range »: I.

Whistler wave cascade in the parallel direction or magnetosonic wave cascade (and also AW) in the perpendicular direction are proposed as long as β<2.5, using a diffusion equation in wavenumber space with the linear time as the energy transfer time. It leads to a k^{-3} spectrum

$$\left(\frac{\partial E(k)}{\partial t}\right)_{\text{nonlinear}} = \frac{\partial}{\partial k}\left[\frac{\gamma k^4}{4\pi\tau_S(k)}\frac{\partial[k^{-2}E(k)]}{\partial k}\right]$$

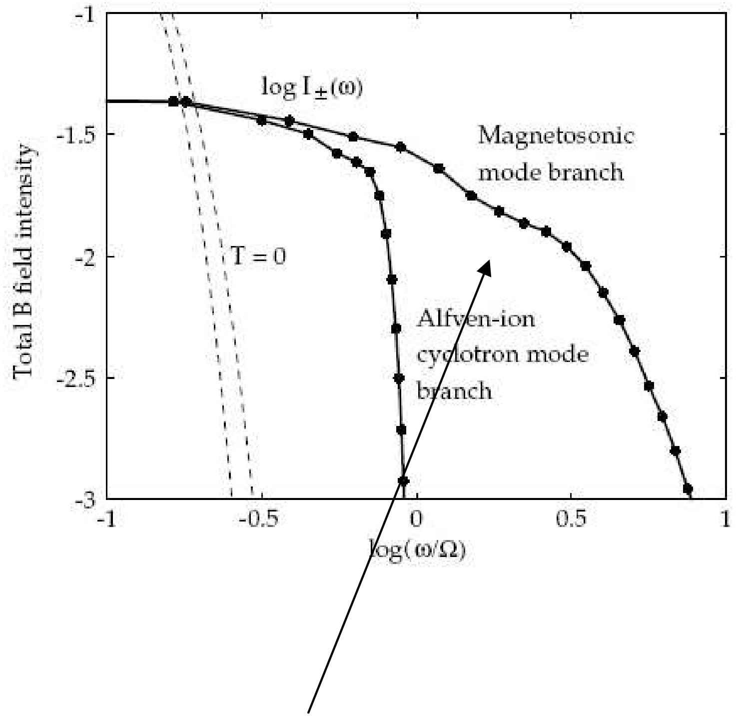
(Leith (1967), Zhou & Matthaeus, JGR 95, 14881 (1990))

(Stawicki, Gary & Li, JGR **106**, 8273 (2001)).

2D PIC simulation of whistler turbulence shows preferential cascade towards perpendicular wavenumbers with steep power laws, <u>and no cascade in 1D</u> (Gary et al. GRL 35, L02104 (2008)).

BUT

Alfvén wave parallel cascade via three-wave decay: transfer from large-scale AW to small-scale ion-cyclotron and magnetosonic whistler wave mediated by ion-sound turbulence (Yoon, PPCF **50** 085007 (2008)).



Weak turbulence of KAW via three-wave decay: inverse cascade if $k_{\perp}\rho_i < 1$, forward cascade otherwise (with a steeper power law). (Voitenko, JPP **60**, 515 (1998)).

Note the knee in the spectrum

Tentative models for the « dissipation range »: II.

Weak turbulence for incompressible Hall MHD:

For $kd_i >> 1$ transfer essentially perpendicular to B_0 : $k_\perp^{-5/2}$

For $kd_i << 1$ transfer exclusively perpendicular to B_0: $k_\perp^{-2}$

(Galtier, JPP **72**, 721 (2006))

2D DNS of compressible HMHD : decaying turbulence shows steepening of the spectrum near the ion-cyclotron scale when the cross-helicity is high (Gosh et al. JGR **101**, 2493 (1996)).

Strong incompressible HMHD simulations of shell model (without mean field): the $k^{-5/3}$ AW cascade steepens to a $k^{-7/3}$ EMHD spectrum when magnetic energy dominates and to a $k^{-11/3}$ spectrum when kinetic energy dominates (Galtier & Buchlin, ApJ **656**, 560 (2007)).

Important role of nonlinearity in the Hall term.

Tentative models for the « dissipation range »: III.

In the MHD range, the AW cascade is essentially transverse to the ambient field.
At the ion Larmor radius, the cascade continues with KAWs ($k_\perp^{-7/3}$), the turbulent fluctuations remaining below the ion cyclotron frequency due to the strong anisotropy.
Gyrokinetic is thus an appropriate tool for the description of this regime.
The range of exponents for power laws could be attributed to:

a. collisionless damping, the true behavior being an exponential fall off, the observed power law being an artifact of instrumental sensitivity?

b. a competition with a dual cascade of entropy modes.

(Howes et al. JGR **113**, A05103 (2008), AIP, CP932 (2007), ApJ **651**, 590 (2006), Apj sup. submitted).

Issues:

It is mentioned that dispersion increases the energy transfer rate, leading to steeper power laws. However it is also known that waves inhibit transfers, leading to shallower spectra (IK spectrum).
$\Rightarrow$ What is the true role of waves and dispersion on nonlinear transfer due to classical steepening phenomena?
This problem is best studied on a simple example containing one kind of wave: DNLS

Is the important scale the ion gyroradius or the ion inertial length?

Does the cascade proceed anisotropically all the way to the electron scale?

Turbulence in DNLS

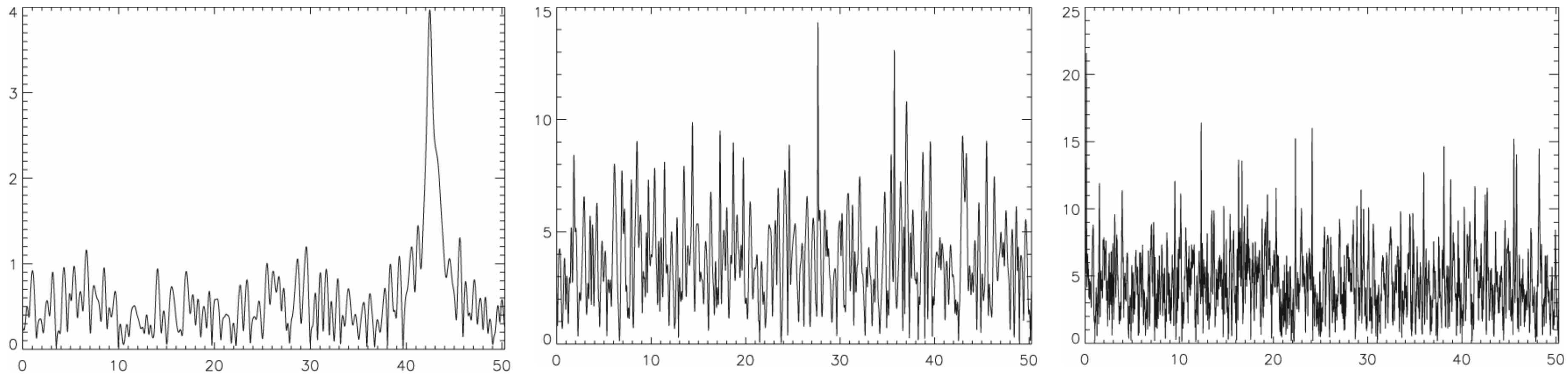
Parallel propagating Alfvén waves can also develop solitonic structures, as seen in the context of DNLS, a large scale 1D reduction:

$$\partial_\tau b + \frac{i}{2R_i}\partial_{\xi\xi} b + \frac{1}{4(1-\beta)}\partial_\xi(|b|^2 b) = 0$$

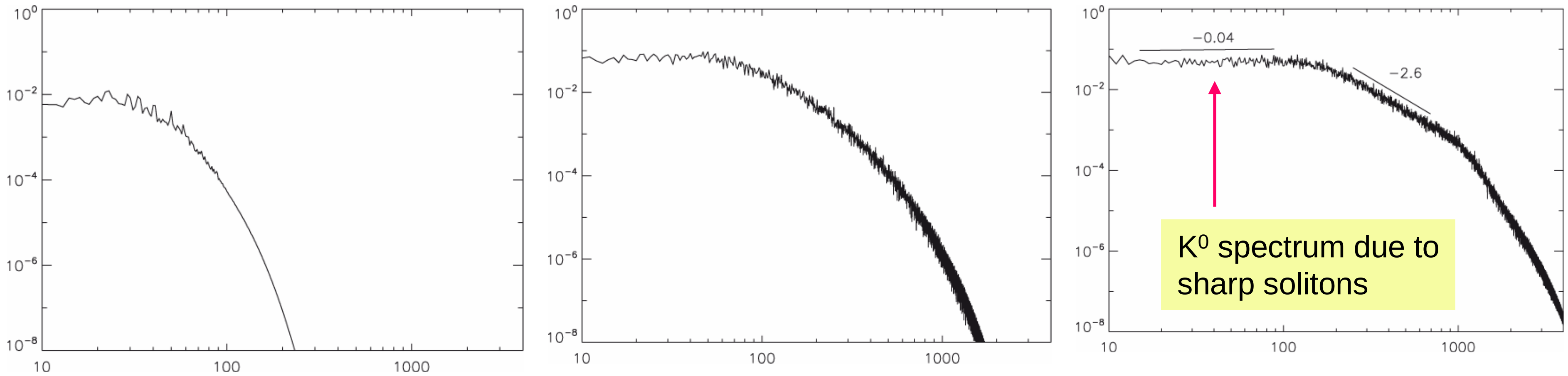
What happens in the presence of external forcing?

Initial condition : soliton
Harmonic forcing at k=50.
Very small dissipation

→ Abrupt transition



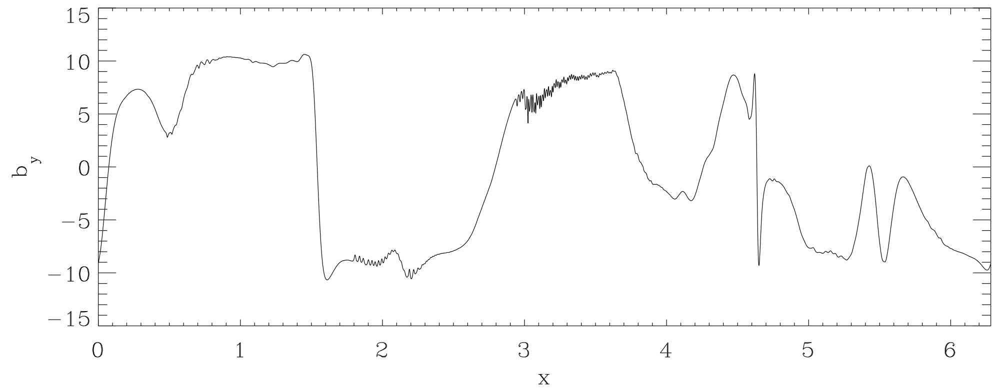
T=100 T=190 T=275

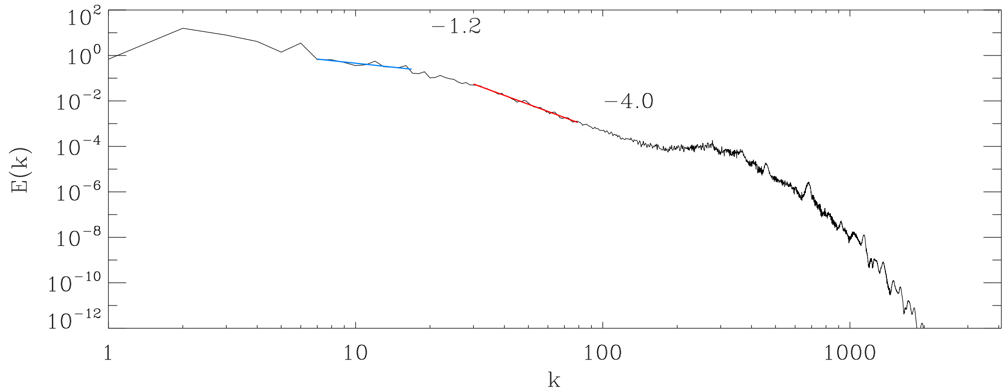


Problem initially investigated by Buti and Nocera, Solar Wind 9, AIP (1999).

Zero intial condition
White noise forcing at k=4
Dispersive scale at k=50

Dissipation: k^2 diffusivity

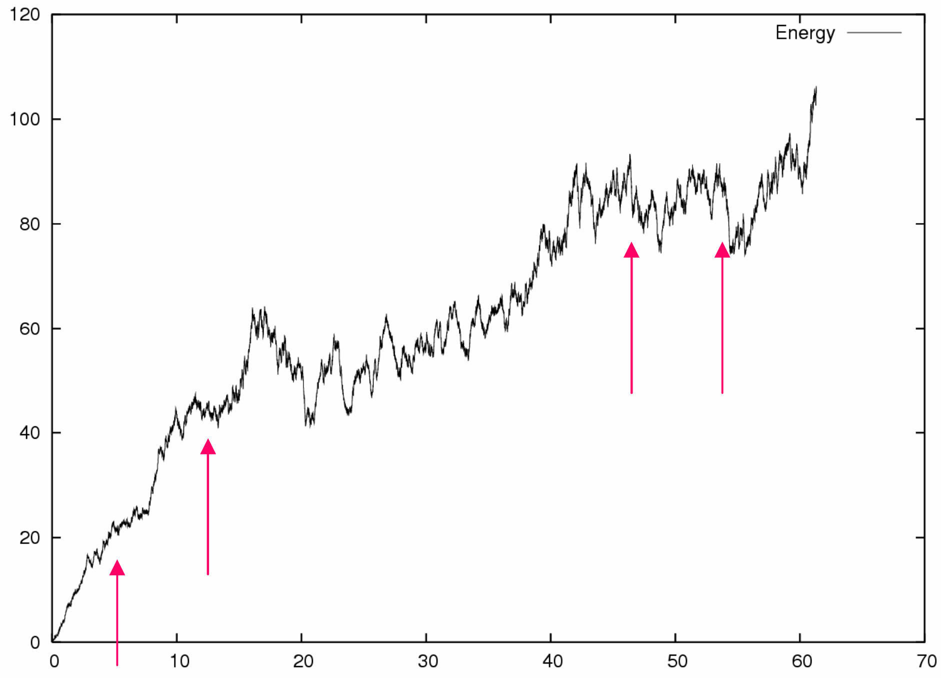


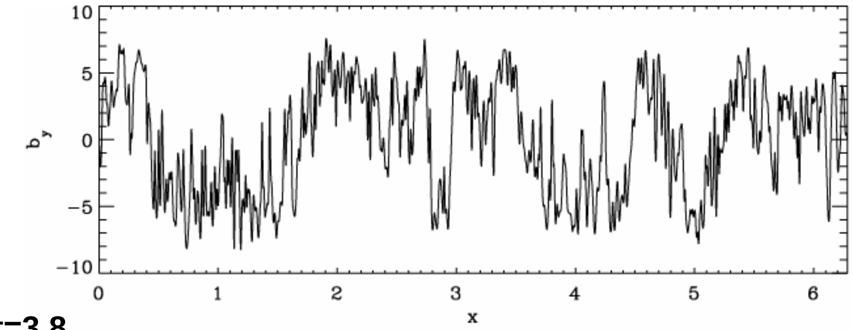


Not turbulent enough

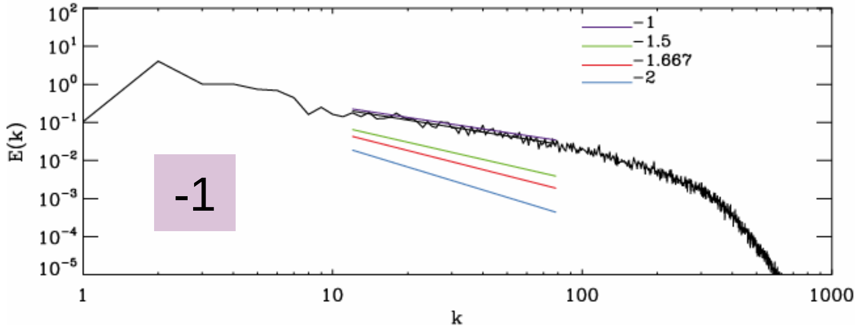
Dissipation: k^8 hyperdiffusivity

Typical evolution of energy vs time.

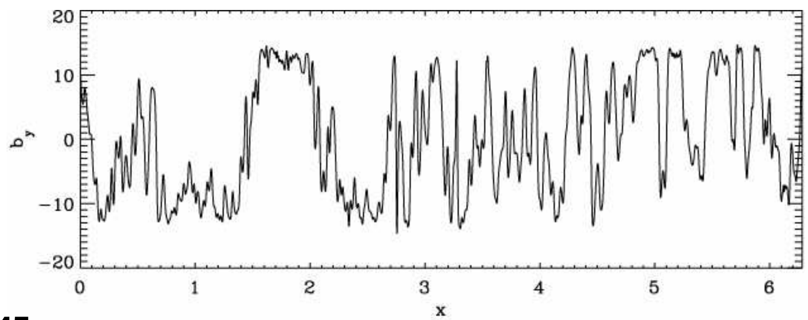


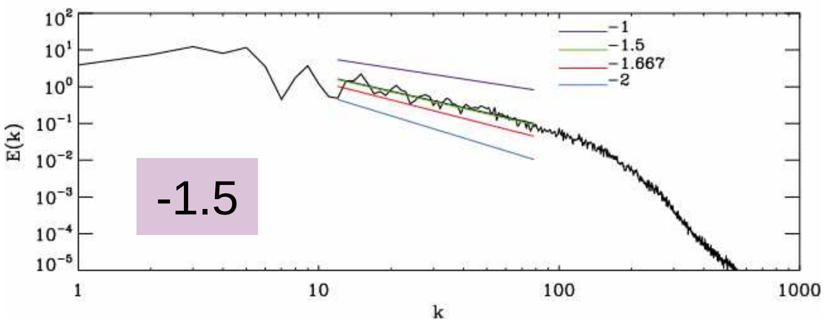


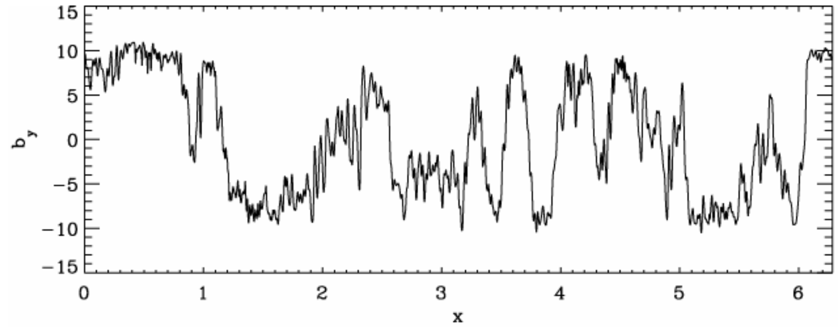
t=3,8



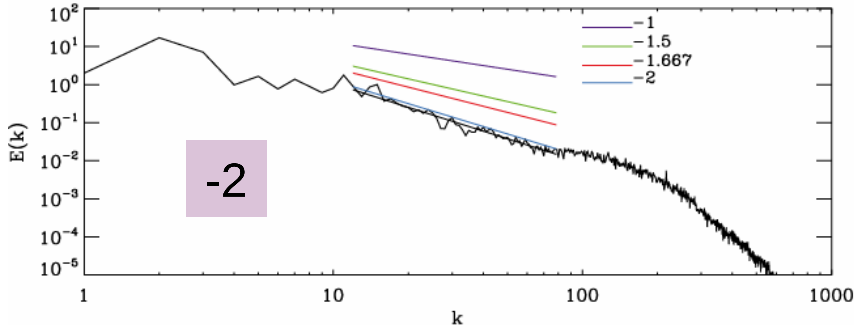
t=45



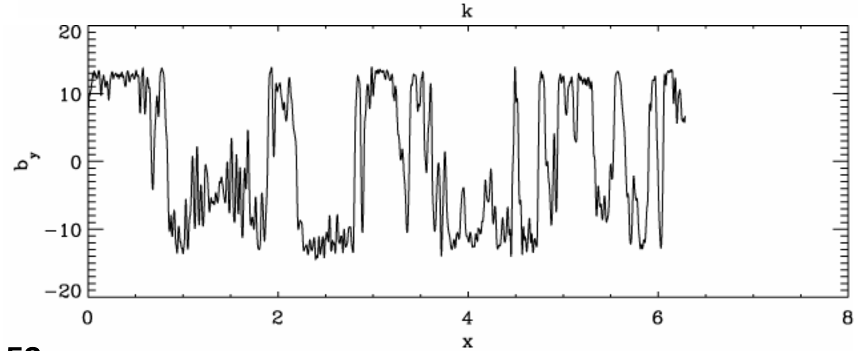


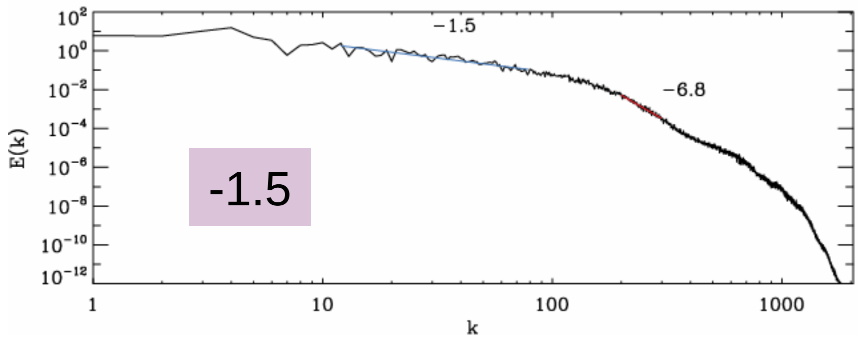


t=12,8



t=53





Phenomenology

K41 for hydrodynamic turbulence

$$\epsilon = \frac{kE_k}{\tau_k} \qquad \tau_k = \tau_{NL} = \frac{1}{\sqrt{k^3 E_k}} \qquad E_k \propto k^{-5/3}$$

Decay time of triple correlations proportional to the turnover time time.

MHD (Kraichnan) $\quad kE_k = |b_k|^2 = |v_k|^2$

$$\epsilon = \frac{kE_k}{\tau_{tr}} \qquad \tau_{tr}\tau_w = \tau_{NL}^2$$

Decay time of triple correlations proportional to the Alfvén time.

non dispersive MHD :

$$\tau_w = \frac{1}{v_A k}$$

$$\tau_{NL} = \frac{1}{\sqrt{k^3 E_k}} \qquad \tau_{tr} = \frac{v_A k}{k^3 E_k} \qquad \epsilon = \frac{1}{v_A} k^3 E_k^2 \qquad E_k \propto k^{-3/2}$$

Does not take into account anisotropy, coherent structures, intermittency

DNLS

Nondispersive scales:

$$\tau_{tr} = \tau_{NL} \equiv \frac{1}{k|b_k|^2} = \frac{1}{k^2 E_k} \qquad \epsilon = k^3 E_k^2 \qquad E_k \propto k^{-3/2}$$

When nonlinearity dominates over dispersion: strongly turbulent

Dispersive scales:

$$\tau_w = k^{-2} \qquad \tau_{NL} \equiv \frac{1}{k|b_k|^2} = \frac{1}{k^2 E_k} \qquad \tau_{tr} = \frac{\tau_{NL}^2}{\tau_w} = \frac{1}{k^2 E_k^2}$$

$$\epsilon_{\mathcal{N}} = E_k k^2 E_k^2 = k^3 E_k^3 \qquad E_k \propto k^{-1}$$

When nonlinear transfer is slowed down by wave dispersion: not very strong turbulence

In the presence of strong shocks a k^{-2} spectrum is expected

Another possibility corresponds to the case where the transfer time is governed by the linear time.

In this case on expects:

$$\varepsilon = \frac{kE(k)}{\tau_w} \qquad \text{and thus} \qquad E(k) \propto k^{-3}$$

Weak turbulence theory (Guyenne et al. In preparation):

Write DNLS as:

$$i\frac{\partial \hat{B}_k}{\partial t} = -\frac{\lambda}{2}\omega_k \hat{B}_k + \int T_k \hat{B}_1 \hat{B}_2 \hat{B}_3^* \, \delta(k_1 + k_2 - k_3 - k) \, dk_1 dk_2 dk_3 \,,$$

For $\langle \hat{B}_k \hat{B}^*_{k'} \rangle = n_k \, \delta(k - k') \,,$ it is possible to derive:

$$\frac{\partial n_k}{\partial t} = -8\pi\lambda \int T_k^2 \left(n_1 n_2 n_3 + n_1 n_2 n_k - n_1 n_3 n_k - n_2 n_3 n_k\right) \times \delta(\omega_1 + \omega_2 - \omega_3 - \omega_k) \, \delta(k_1 + k_2 - k_3 - k) \, dk_1 dk_2 dk_3 \,.$$

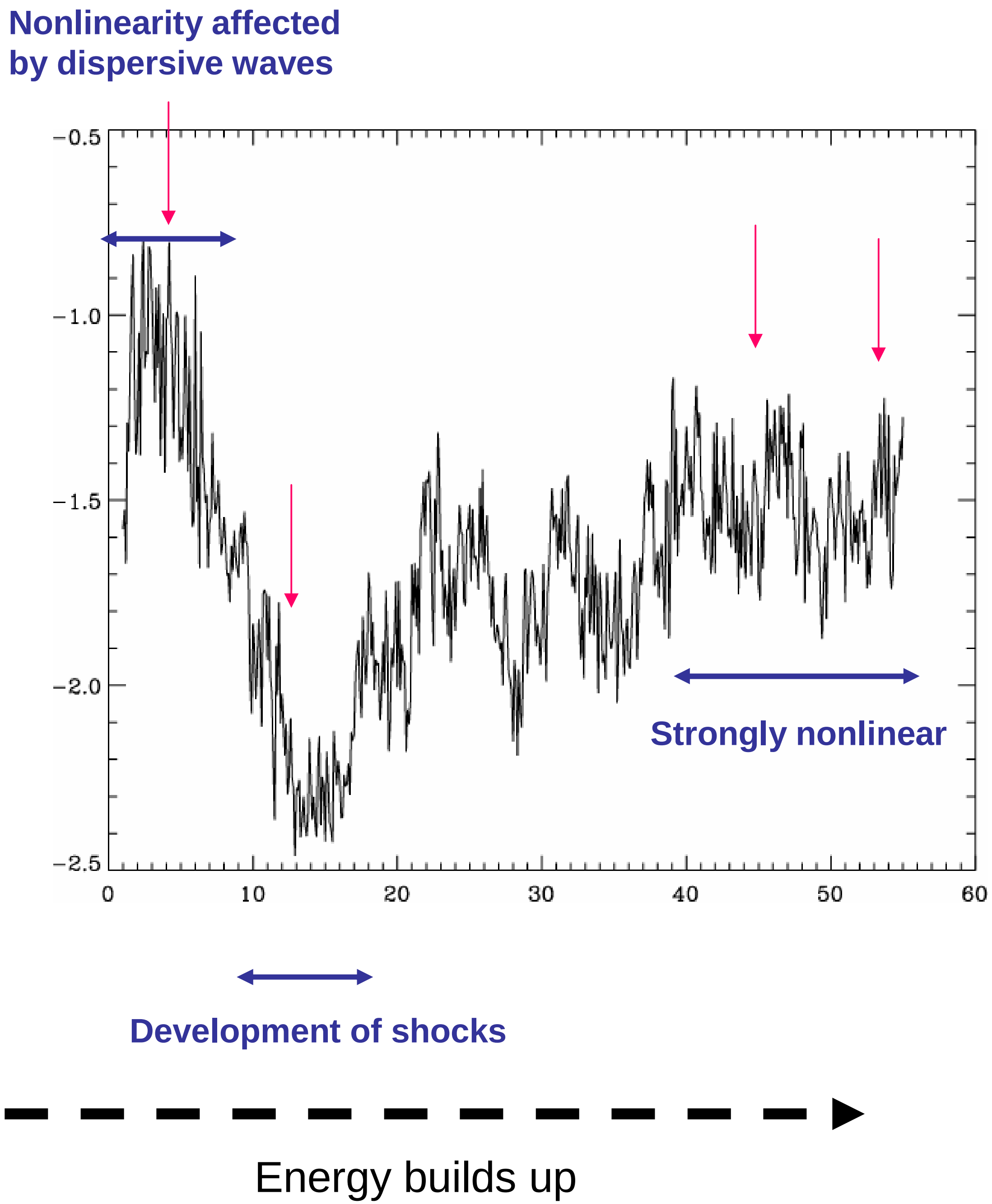
$\omega_k = k^2$ and $T_k = \sigma k$ is the interaction coefficient.

Two power law solutions: k^{-1} for the inverse cascade of wave action $N = \int n_k \, dk$

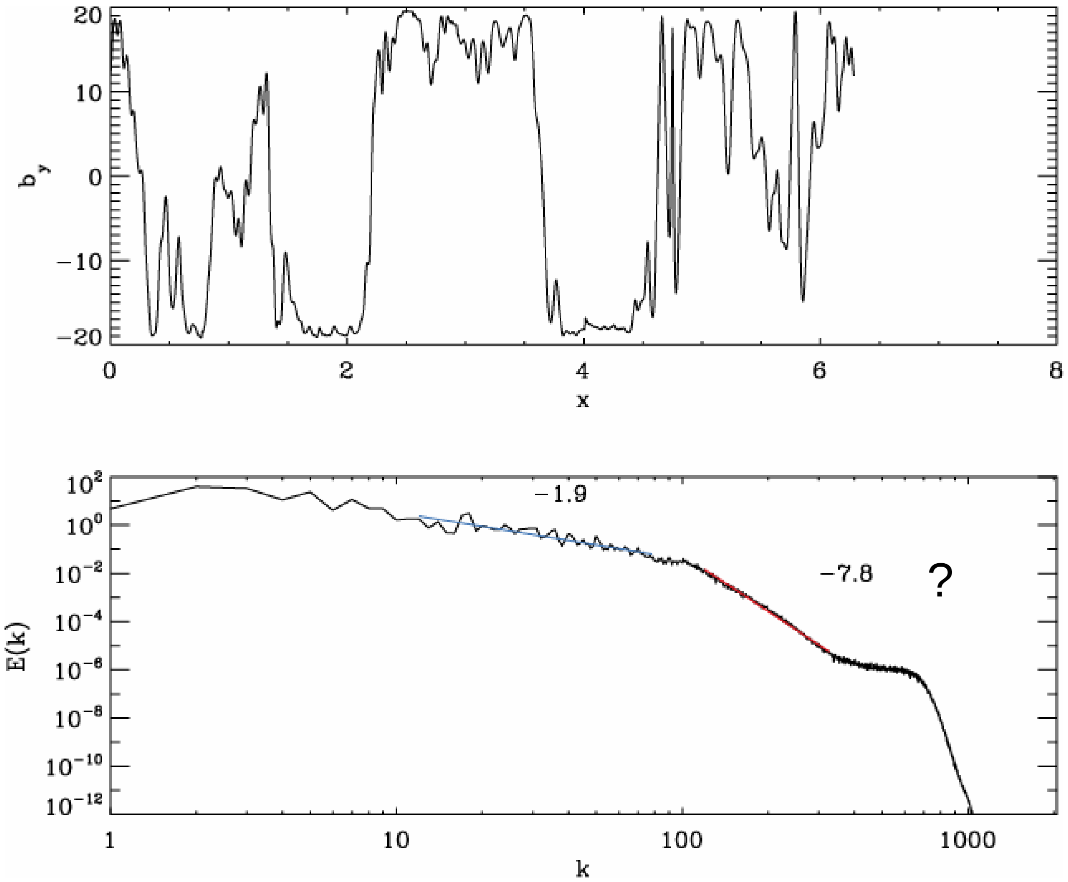
$k^{-5/3}$ for the direct cascade of energy $E = \int \omega_k n_k \, dk$

The latter is not observed in our simulations

Instantaneous slope vs. time:
Non-stationarity



With k^{16} hyperdiffusivity:
Stronger shocks => dominant k^{-2} signature in the spectrum
Shallower slope at much smaller scale, connected by a steep transition region.



What happens with a more refined model including all types of waves (except Langmuir)?
At first let us examine the 1D case within the context of Landau fluids.

Landau fluids

For the sake of simplicity, neglect electron inertia.

Ion dynamics: derived by computing velocity moments from Vlasov Maxwell equations.

$$\partial_t \rho_p + \nabla \cdot (\rho_p u_p) = 0$$

$$\partial_t u_p + u_p \cdot \nabla u_p + \frac{1}{\rho_p} \nabla \cdot \mathbf{p}_p - \frac{e}{m_p}(E + \frac{1}{c} u_p \times B) = 0$$

$$E = -\frac{1}{c}\left(u_p - \frac{j}{ne}\right) \times B - \frac{1}{ne} \nabla \cdot \mathbf{p}_e,$$

$$\partial_t B = -c \nabla \times E$$

$$\rho_r = m_r n_r$$

quasi-neutrality $(n_e = n_p)$

$$j = \frac{c}{4\pi} \nabla \times B$$

$$\mathbf{p}_p = p_{\perp p} \mathbf{n} + p_{\| p} \boldsymbol{\tau} + \mathbf{\Pi}, \text{ with } \mathbf{n} = \mathbf{I} - \widehat{b} \otimes \widehat{b} \text{ and } \boldsymbol{\tau} = \widehat{b} \otimes \widehat{b}, \text{ where } \widehat{b} = B / |B|.$$

Electron pressure tensor is taken gyrotropic (scales >> electron Larmor radius): characterized by the parallel and transverse pressures $p_{\|e}$ and $p_{\perp e}$.

For each particle species,

Perpendicular and parallel pressures

heat flux tensor

$$\partial_t p_\perp + \nabla \cdot (u p_\perp) + p_\perp \nabla \cdot u - p_\perp \hat{b} \cdot \nabla u \cdot \hat{b} + \frac{1}{2}[\mathrm{tr}\ \nabla \cdot \mathbf{q} - \hat{b} \cdot (\nabla \cdot \mathbf{q}) \cdot \hat{b}] = 0$$

$$\partial_t p_\parallel + \nabla \cdot (u p_\parallel) + 2 p_\parallel \hat{b} \cdot \nabla u \cdot \hat{b} + \hat{b} \cdot (\nabla \cdot \mathbf{q}) \cdot \hat{b} = 0$$

Nongyrotropic components (gyroviscous tensor) of the pressure tensor will be evaluated separately by fitting with the linear kinetic theory.

Heat fluxes

$\mathbf{n} = \mathbf{I} - \widehat{b} \otimes \widehat{b}$ and $\boldsymbol{\tau} = \widehat{b} \otimes \widehat{b}$, where $\widehat{b} = B / |B|$

Proton heat flux tensor: $\mathbf{q} = \mathbf{S} + \boldsymbol{\sigma}$ with $\sigma_{ijk} n_{jk} = 0$ and $\sigma_{ijk} \tau_{jk} = 0$.

Nongyrotropic tensor that contributes
at the nonlinear level only

Fluxes of parallel and transverse heat: $S_i^{\|} = q_{ijk} \tau_{jk}$ and $2 S_i^{\perp} = q_{ijk} n_{jk}$.

Parallel heat fluxes of perpendicular and parallel heat $q_{\perp} = S^{\perp} \cdot \widehat{b}$ and $q_{\|} = S^{\|} \cdot \widehat{b}$ are the only contribution to the gyrotropic heat flux tensor.

Write $S^{\perp} = q_{\perp} \widehat{b} + S_{\perp}^{\perp}$ and $S^{\|} = q_{\|} \widehat{b} + S_{\perp}^{\|}$ where the perpendicular heat flux of perpendicular and parallel heat $S_{\perp}^{\perp}$ and $S_{\perp}^{\|}$ are computed in a linearized approximation.

The gyrotropic heat flux components $q_{\perp}$ and $q_{\|}$ obey dynamical equations.

Equations for the parallel and perpendicular (gyrotropic) heat fluxes

$$\begin{cases} \partial_t q_{\|} + \nabla\cdot(q_{\|} u) + 3q_{\|}\widehat{b}\cdot\nabla u\cdot\widehat{b} + 3p_{\|}(\widehat{b}\cdot\nabla)\left(\frac{p_{\|}}{\rho}\right) + \nabla\cdot(\widetilde{r}_{\|\|}\widehat{b}) - 3\widetilde{r}_{\|\perp}\nabla\cdot\widehat{b} + \partial_z R_{\|}^{NG} = 0 \\ \partial_t q_{\perp} + \nabla\cdot(u q_{\perp}) + q_{\perp}\nabla\cdot u + p_{\|}(\widehat{b}\cdot\nabla)\left(\frac{p_{\perp}}{\rho}\right) + \frac{p_{\perp}}{\rho}\left(\partial_x\Pi_{xz} + \partial_y\Pi_{yz}\right) \\ \qquad + \nabla\cdot(\widetilde{r}_{\|\perp}\widehat{b}) + \left((p_{\|} - p_{\perp})\frac{p_{\perp}}{\rho} - \widetilde{r}_{\perp\perp} + \widetilde{r}_{\|\perp}\right)(\nabla\cdot\widehat{b}) + \partial_z R_{\perp}^{NG} = 0 \end{cases}$$

Involve the 4 th rank gyrotropic cumulants $\widetilde{r}_{\|\|}$, $\widetilde{r}_{\|\perp}$, $\widetilde{r}_{\perp\perp}$
expressed in terms of the 4 th rank gyrotropic moments by

$$\widetilde{r}_{\|\|} = r_{\|\|} - 3\frac{p_{\|}^2}{\rho},$$

$$\widetilde{r}_{\|\perp} = r_{\|\perp} - \frac{p_{\perp}p_{\|}}{\rho},$$

$$\widetilde{r}_{\perp\perp} = r_{\perp\perp} - 2\frac{p_{\perp}^2}{\rho}.$$

$R_{\|}^{NG}$ and $R_{\perp}^{NG}$ stand for the nongyrotropic contributions of the fourth rank cumulants.

2 main problems:

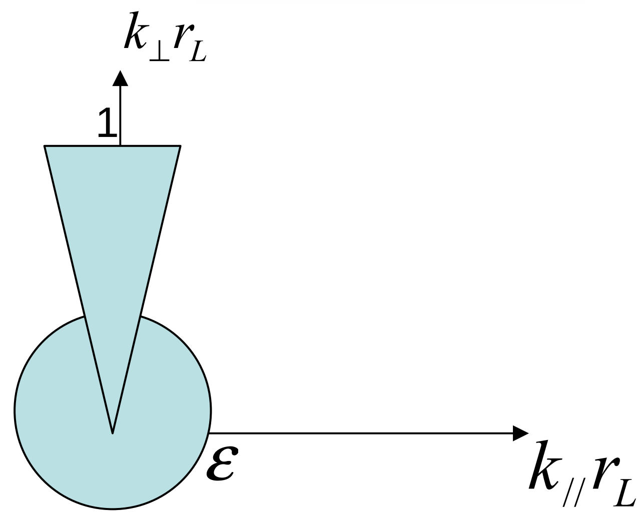
(1) Closure relations are needed to express the 4th order cumulants $\widetilde{r}_{\parallel\parallel}, \widetilde{r}_{\parallel\perp}, \widetilde{r}_{\perp\perp}$
(closure at lowest order also possible, although usually less accurate)

(2) FLR corrections (non-gyrotropic) to the various moments are to be evaluated

The starting point for addressing these points is the linear kinetic theory in the low-frequency limit. $\omega/\Omega \sim \epsilon \ll 1$ (Ω: ion gyrofrequency)

For a unified description of fluid and kinetic scales, FLR-Landau fluids retain contributions of:

- quasi-transverse fluctuations $(k_\parallel/k_\perp \sim \epsilon)$ with $k_\perp r_L \sim 1$
- hydrodynamic scales with $k_\parallel r_L \sim k_\perp r_L \sim \epsilon.$ r_L: ion Larmor radius



CLOSURE RELATIONS are based on linear kinetic theory (near bi-Maxwellian equilibrium) in the low-frequency limit.

For example, for each species, (assuming the ambient magnetic field along the z direction),

$$\widetilde{r}_{\parallel\perp} = \frac{{p_\perp^{(0)}}^2}{\rho^{(0)}}\left[1 - R(\zeta) + 2\zeta^2 R(\zeta)\right]\left[[2b\Gamma_0(b) - \Gamma_0(b) - 2b\Gamma_1(b)]\frac{b_z}{B_0} + b[\Gamma_0(b) - \Gamma_1(b)]\frac{e\Psi}{T_\perp^{(0)}}\right]$$

$\Gamma_n(b) = e^{-b} I_n(b)$, $b = (k_\perp^2 T_\perp^{(0)})/(\Omega^2 m)$, $I_n(b)$ modified Bessel function, $E_z = -\partial_z \Psi$

R is the plasma response function, $\zeta = \frac{\omega}{|k_\parallel| v_{th}}$. (For electrons, $b \approx 0$, $\Gamma_0 \approx 1$, $\Gamma_1 \approx 0$)

It turns out that $\widetilde{r}_{\parallel\perp}$ can be expressed in terms of perpendicular gyrotropic heat flux $q_\perp$ and of the parallel current j_z. One has

$$\widetilde{r}_{\parallel\perp} = \sqrt{\frac{2T_\parallel^{(0)}}{m}}\frac{1 - R(\zeta) + 2\zeta^2 R(\zeta)}{2\zeta R(\zeta)}\left[q_\perp + [\Gamma_0(b) - \Gamma_1(b)]\frac{p_\perp^{(0)} p_\parallel^{(0)}}{\rho^{(0)} v_A^2}\left(\frac{T_\perp^{(0)}}{T_\parallel^{(0)}} - 1\right)\frac{j_z}{en^{(0)}}\right]$$

The approximation consists in replacing the plasma response function R by the three pole Padé approximant $R_3(\zeta) = \dfrac{2 - i\sqrt{\pi}\zeta}{2 - 3i\sqrt{\pi}\zeta - 4\zeta^2 + 2i\sqrt{\pi}\zeta^3}$.

This leads to the approximation $\frac{1 - R(\zeta) + 2\zeta^2 R(\zeta)}{2\zeta R(\zeta)} \approx \frac{i\sqrt{\pi}}{-2 + i\sqrt{\pi}\zeta}$.

(A lower order Padé approximant would overestimates the Landau damping in the large ζ limit).

One finally gets a closure relation in the form of the evolution equation (for each species)

$$[\frac{d}{dt} - \frac{2}{\sqrt{\pi}}\sqrt{\frac{2T_\parallel^{(0)}}{m}}\mathcal{H}_z\partial_z]\widetilde{r}_{\parallel\perp} + \frac{2T_\parallel^{(0)}}{m}\partial_z[q_\perp + [\Gamma_0(b) - \Gamma_1(b)]\frac{p_\perp^{(0)}p_\parallel^{(0)}}{\rho^{(0)}v_A^2}(\frac{T_\perp^{(0)}}{T_\parallel^{(0)}} - 1)\frac{j_z}{en^{(0)}}] = 0,$$

In Fourier space, Hilbert transform $\mathcal{H}_z$ reduces to the multiplication by $i \operatorname{sgn} k_z$.

Improvement: Retain the evolution of the equilibrium state by replacing the (initial) equilibrium pressures and temperatures by the instantaneous fields averaged on space.

In the large-scale limit, $\Gamma_0(0) = 1$ and $\Gamma_1(0) = 0$.

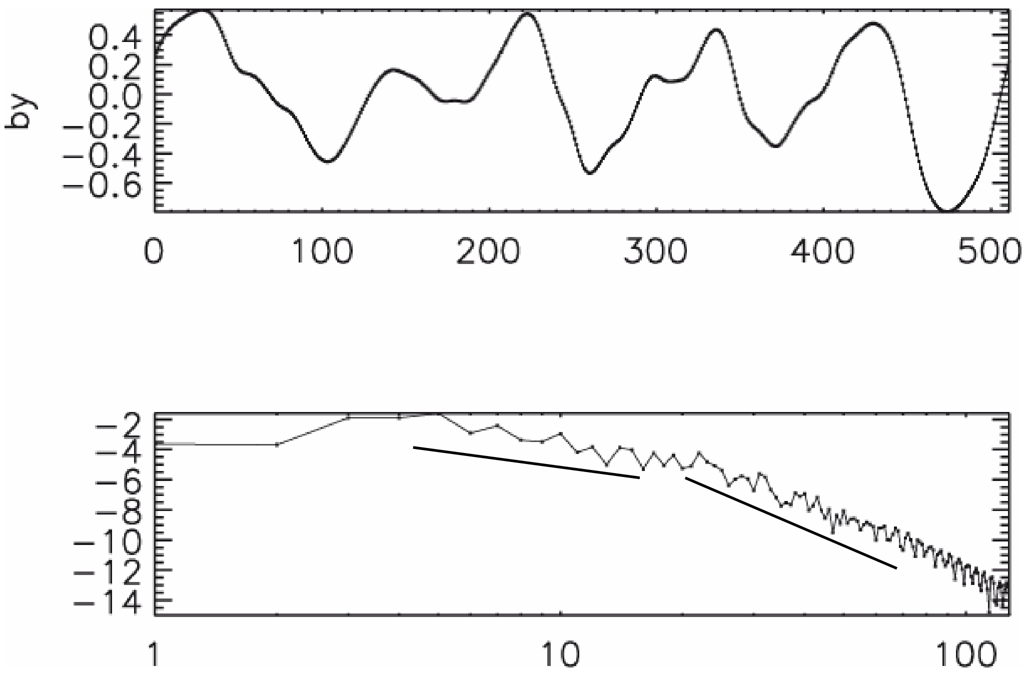
A 1D SIMULATION

with:

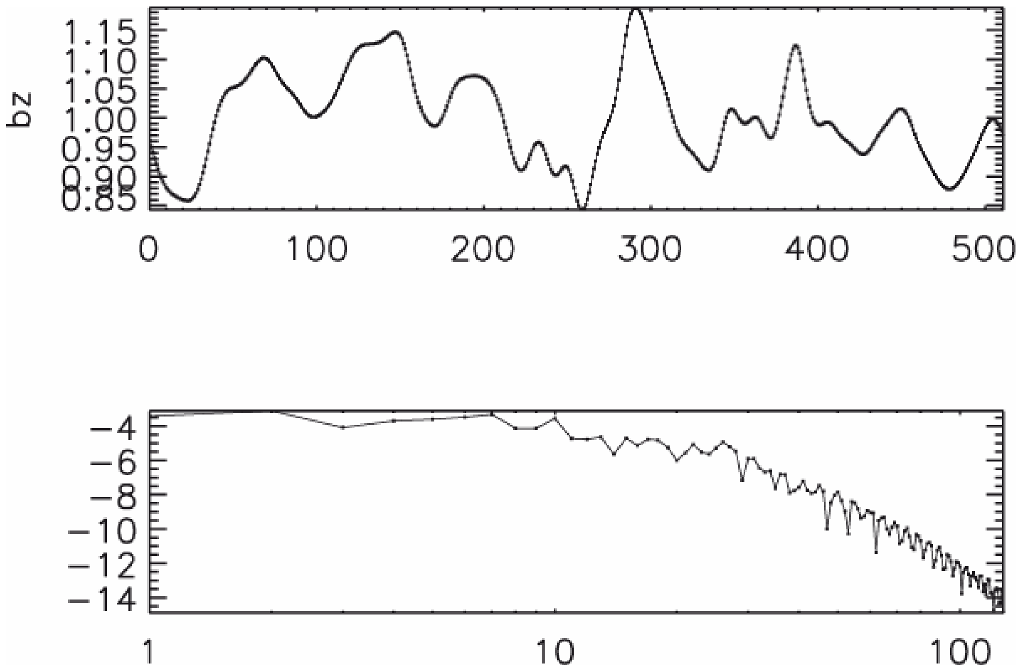
•A small amount of collisions to let parallel and perpendicular pressures tend to the same mean values and thus avoid instabilities.
•Random forcing of the three velocity components between k=2 and 10, peaking at k=5, only on when the total energy falls below prescribed value.
•Angle of propagation: 84°
•β=1, Te/Ti=5
•Size of the domain: 40* 2π
•No extra dissipation

A break in the spectrum starts to develop at the dispersive scale

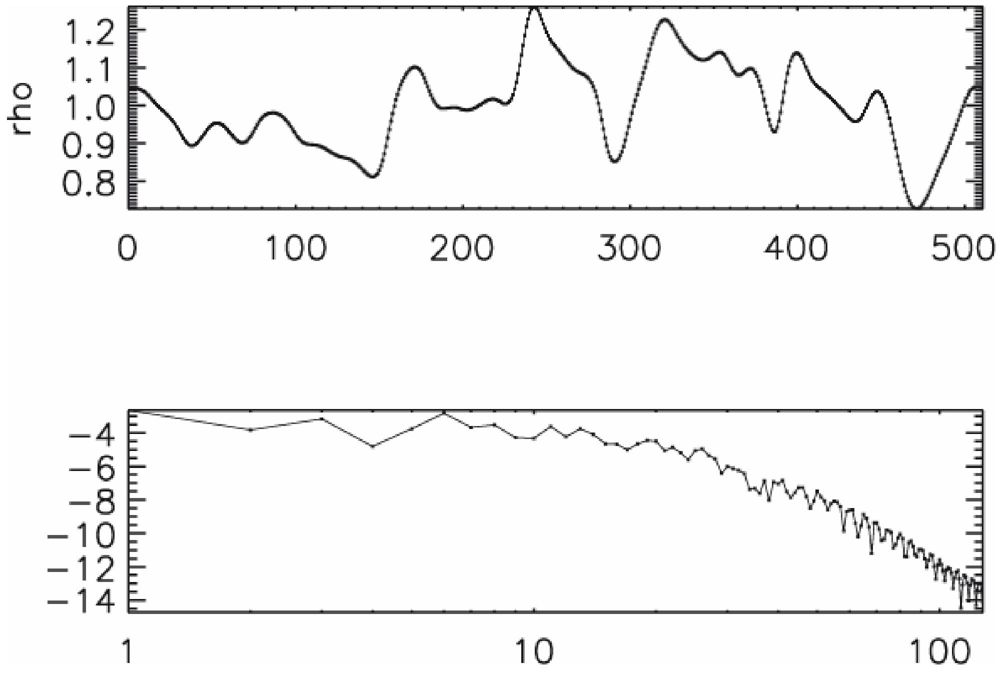
In spite of superimposed turbulence, large-scale structures form.

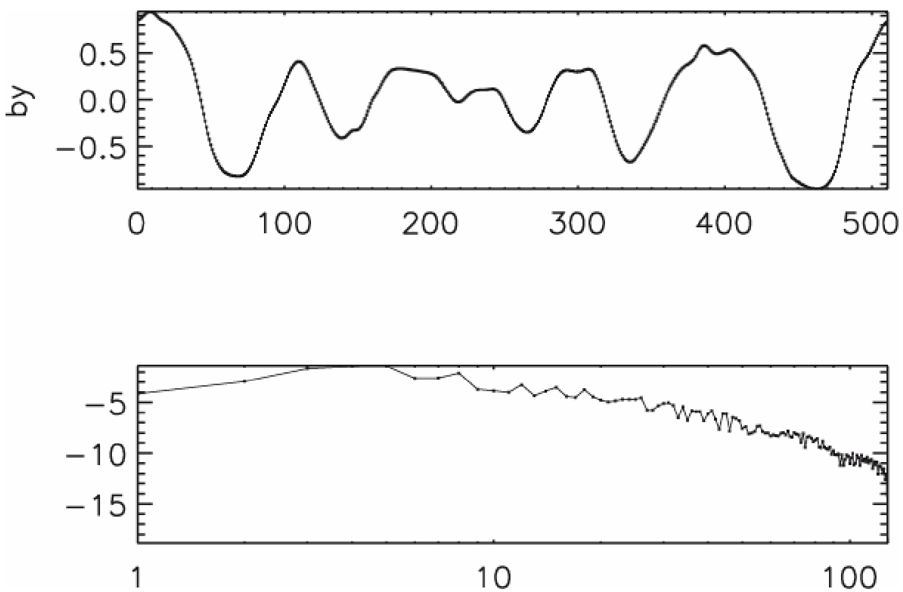


Longitudinal field



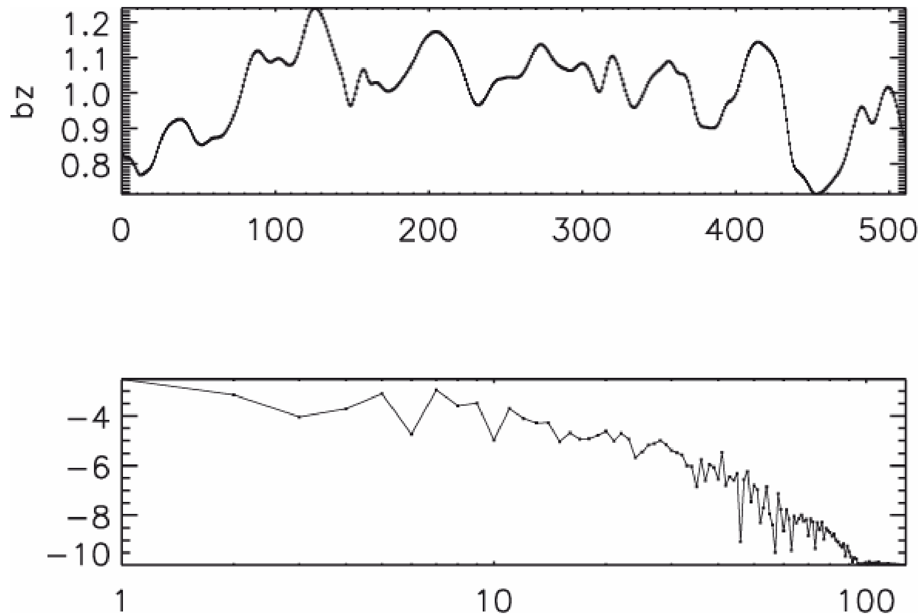
Density

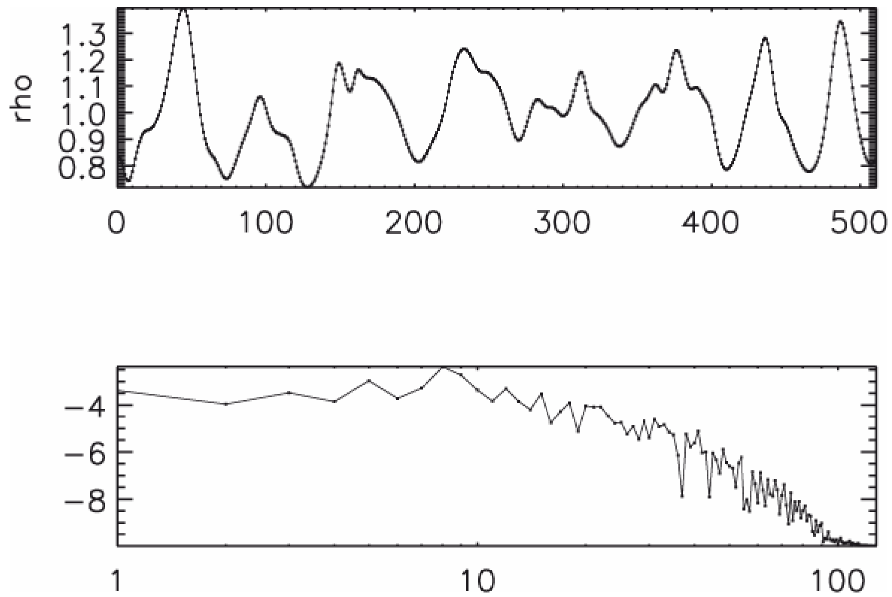




Another case

Forcing only on v_y
Angle of propagation: 87°
β=1, T_e/T_i=50
Size of the domain: 30* 2π





Hall-MHD equations with a polytropic equation of state

$$\partial_t \rho + \nabla.(\rho \mathbf{u}) = 0$$

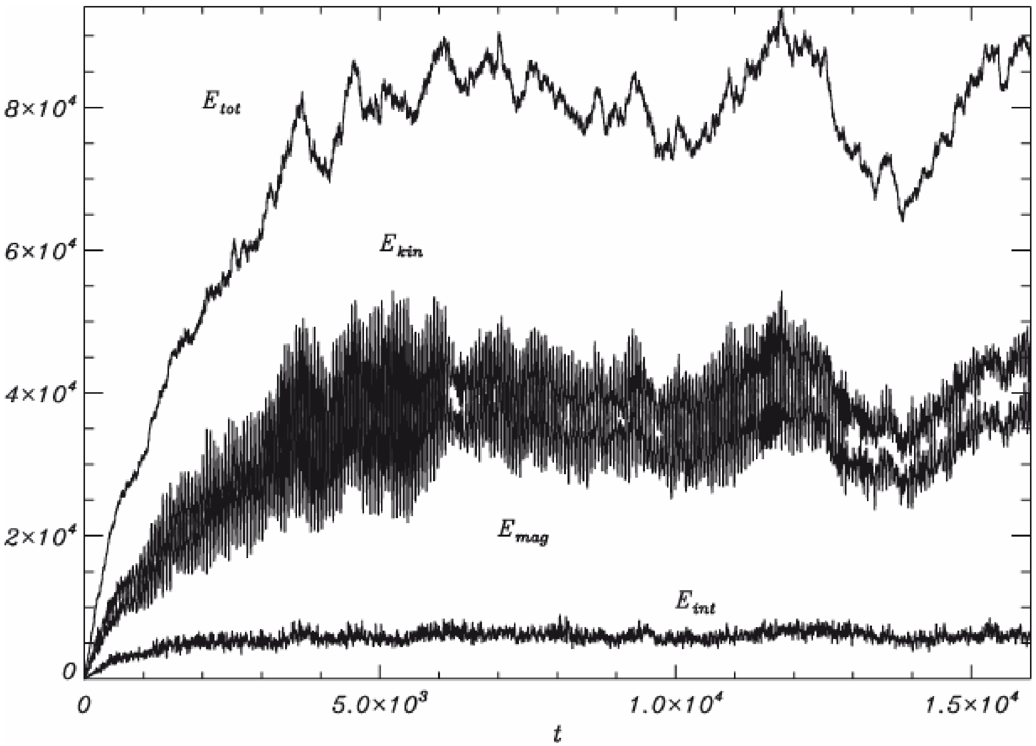
$$\rho(\partial_t \mathbf{u} + \mathbf{u}.\nabla \mathbf{u}) = -\frac{\beta}{\gamma}\nabla \rho^\gamma + (\nabla \times \mathbf{b}) \times \mathbf{b}$$

$$\partial_t \mathbf{b} - \nabla \times (\mathbf{u} \times \mathbf{b}) = -\frac{1}{R_i}\nabla \times (\frac{1}{\rho}(\nabla \times \mathbf{b}) \times \mathbf{b})$$

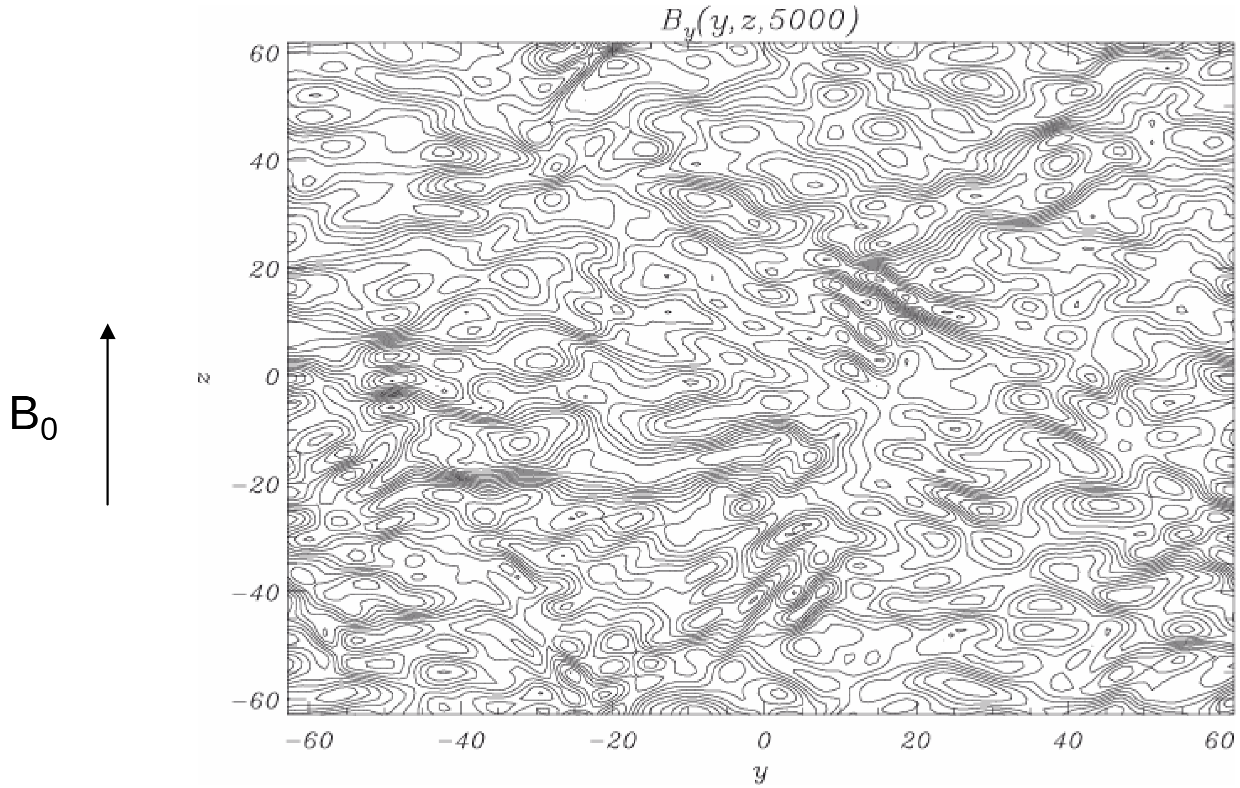
$$\nabla.\mathrm{b} = 0$$

velocity unit: Alfvén speed
length unit : R_i x ion inertial length
time unit: R_i x ion gyroperiod
density unit: mean density
magnetic field unit: ambient field

2D simulations with uniform field in the z-direction
Forcing of the transverse velocity field components at k=2
Parameters: β=1, γ=2, R_i=1, size of domain= 20*2*π
A filter is used to dissipate.

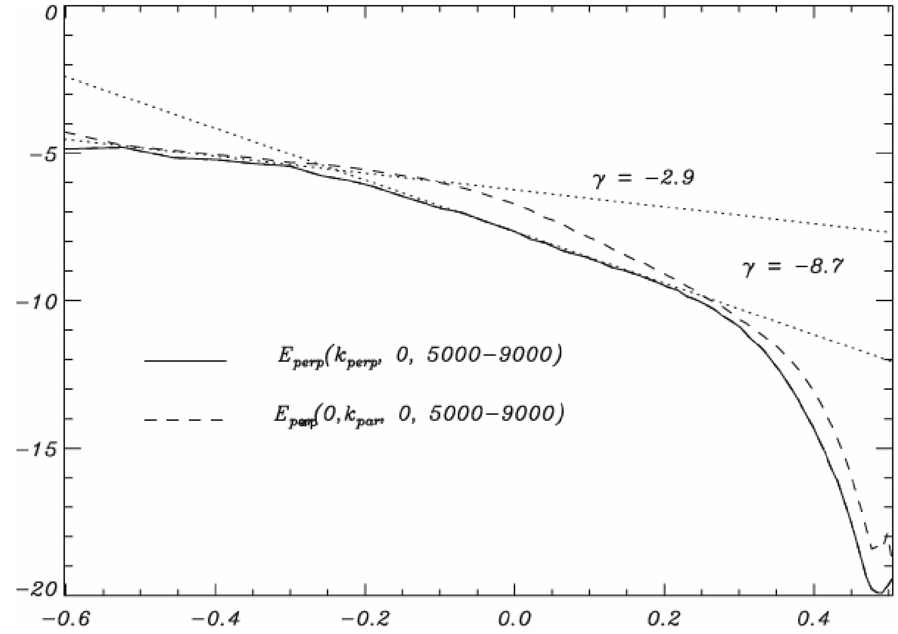


Evolution of the energy as a function of time



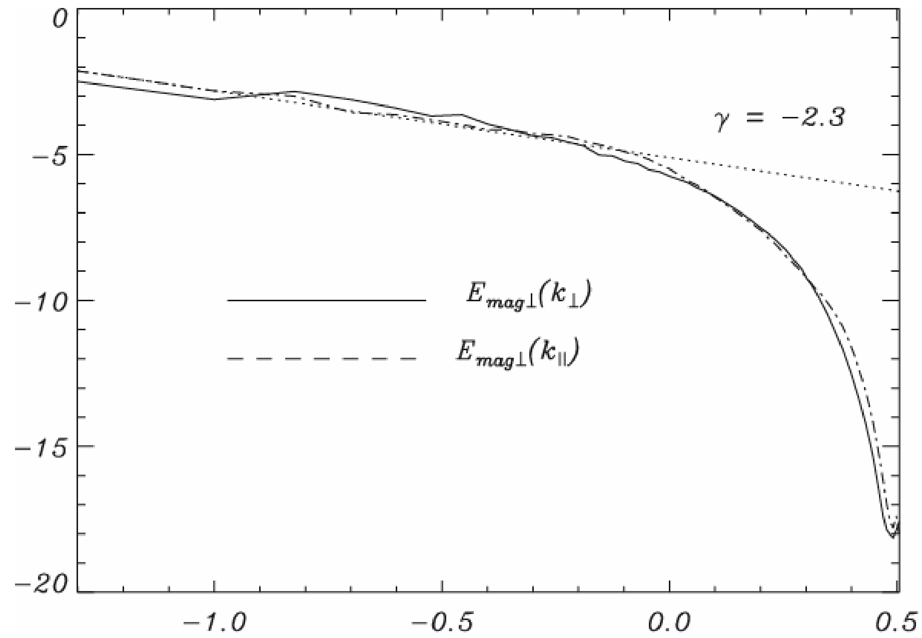
Typical contour plot of By field

Spectra of perpendicular magnetic energy as a function of $k_\perp$ (for $k_{//}$=0): solid and as a function of $k_{//}$ (for $k_\perp$=0): dashed



A break is observed with a steeper slope at small scales. In the paraller direction, the spectrum drops off more quickly but the large-scale range extends to smaller scales.

Spectra of perpendicular magnetic energy as a function of $k_\perp$ (summed over all $k_{//}$): solid and as a function of $k_{//}$ (summed over all $k_\perp$): dashed



Conclusions

Dispersion does not prevent the formation of small scales and the development of a turbulent cascade. Coherent structures nevertheless form.

A break in the spectrum at the dispersive scale is directly observed in DNLS, Landau fluid and Hall-MHD simulations. Further work is needed to identify whether it corresponds to the ion gyro-radius or the ion inertial length.

Further work is also necessary to explore the dynamics of KAWs,
in particular for smaller values of Te/Ti and/or at smaller scales.
Purely kinetic numerical study is probably needed in this case.

Three-dimensional simulations are underway.