Evidence for sterile neutrinos from precision cosmology?

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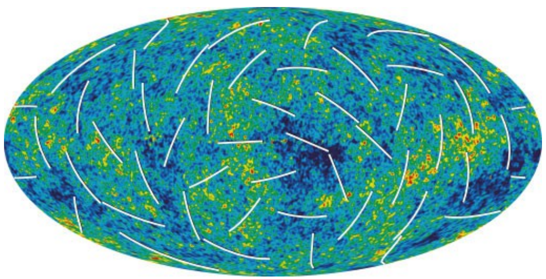
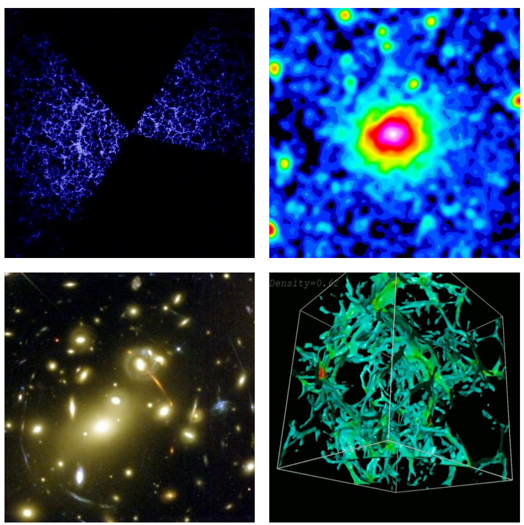
Precision cosmological probes...

Probes of inhomogeneities

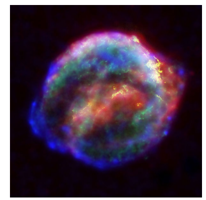


CMB temperature & polarisation anisotropies

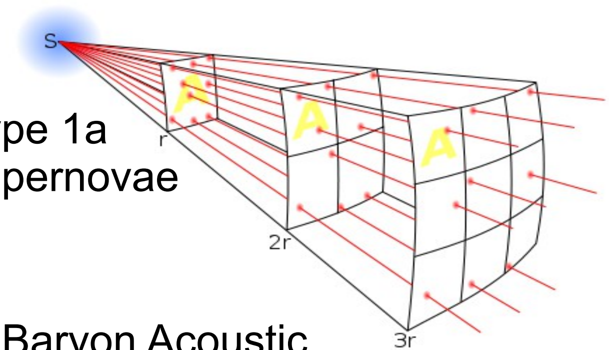
Large-scale matter distribution

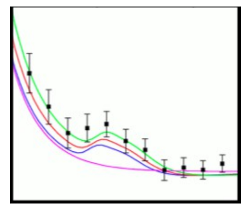


Distance vs redshift



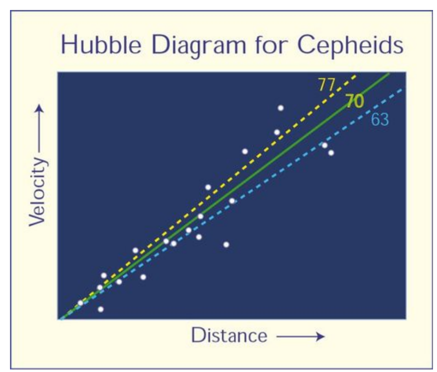
Type 1a supernovae





Baryon Acoustic Oscillation scale

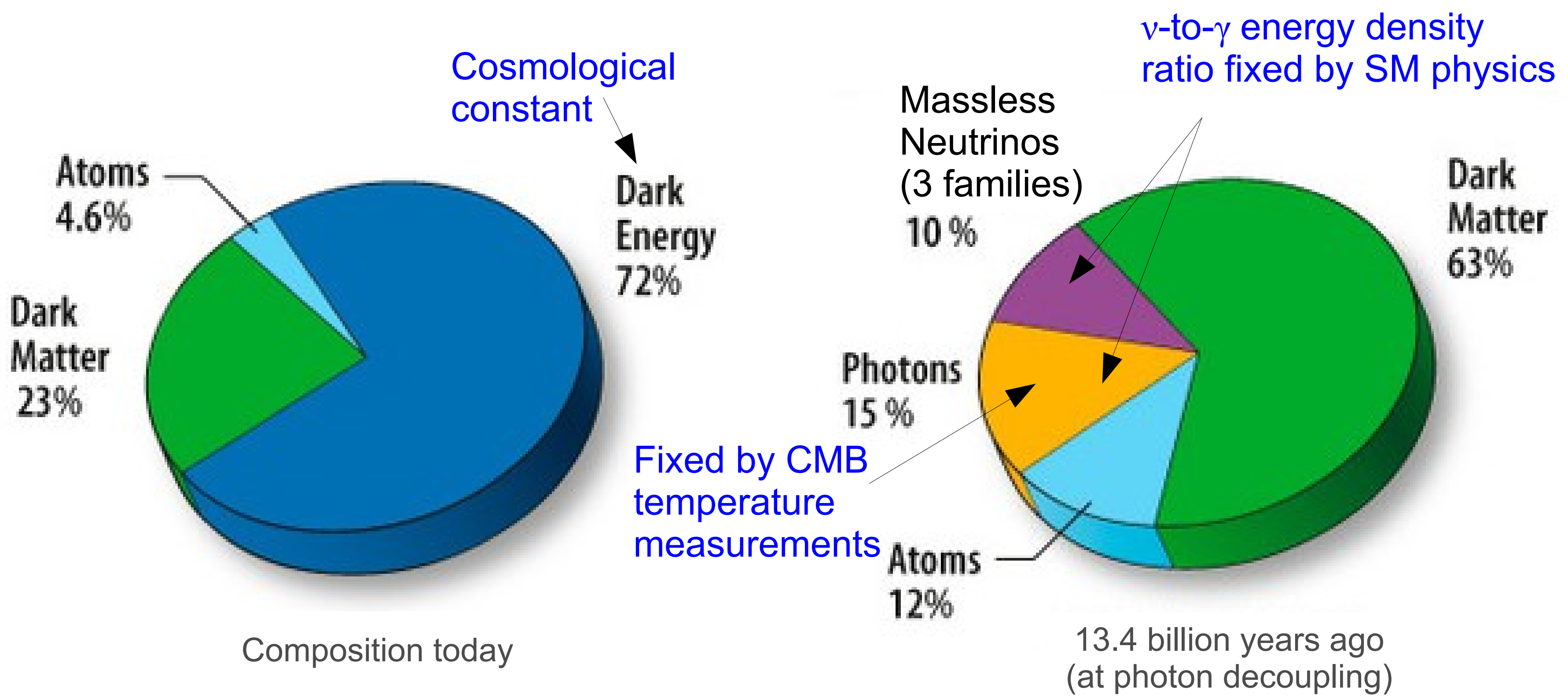
Local Hubble expansion rate



$$H_0 = 100\, h \text{ km/s/Mpc}$$
$$= 73.8 \pm 2.4 \text{ km/s/Mpc}$$

The concordance flat ΛCDM model...

- The simplest model consistent with **present observations**.



Plus flat spatial geometry+initial conditions
from single-field inflation

Neutrino energy density (standard picture)...

- Neutrino decoupling at T ~ O(1) MeV. ← Fixed by weak interactions
- After e^+e^- annihilation (T ~ 0.2 MeV):
  - **Temperature**: $T_\nu = \left(\frac{4}{11}\right)^{1/3} T_\gamma$ ← Assuming instantaneous decoupling
  - **Number density** per flavour: $n_\nu = \frac{6}{4}\frac{\zeta(3)}{\pi^2} T_\nu^3 = \frac{3}{11} n_\gamma$
  - **Energy density** per flavour: $\rho_\nu = \frac{7}{8}\frac{\pi^2}{15} T_\nu^4 = \frac{7}{8}\left(\frac{4}{11}\right)^{4/3} \rho_\gamma$ ← $\frac{3\rho_\nu}{\rho_\gamma} \sim 0.68$

Photon temperature, number density, & energy density

- If massive, then at T << m: $\rho_\nu = m_\nu n_\nu$ → $\Omega_{\nu,0} h^2 = \frac{m_\nu}{94\ \mathrm{eV}}$

Hot dark matter (not within vanilla ΛCDM)

Extending the "neutrino" sector...

- Any particle species whose production is associated with some thermal process and that decoupled while relativistic at relatively late times [T< O(100) MeV] will behave (more or less) like a neutrino as far as cosmological observations are concerned.

$$\sum_i \rho_{\nu,i} + \rho_X = N_{\rm eff}\left(\frac{7}{8}\frac{\pi^2}{15}T_\nu^4\right)$$

$$= (3.046 + \Delta N_{\rm eff})\,\rho_\nu^{(0)}$$

Three SM neutrinos

Other light stuff:
sterile neutrinos,
eV-mass axions,
hidden photons,
etc.

Neutrino
temperature
per definition

Corrections due to
non-instantaneous decoupling,
finite temperature effects in the EM plasma,
and flavour oscillations

Plan...

- Evidence of N_{eff}>3 from **CMB and large-scale structure** observations.

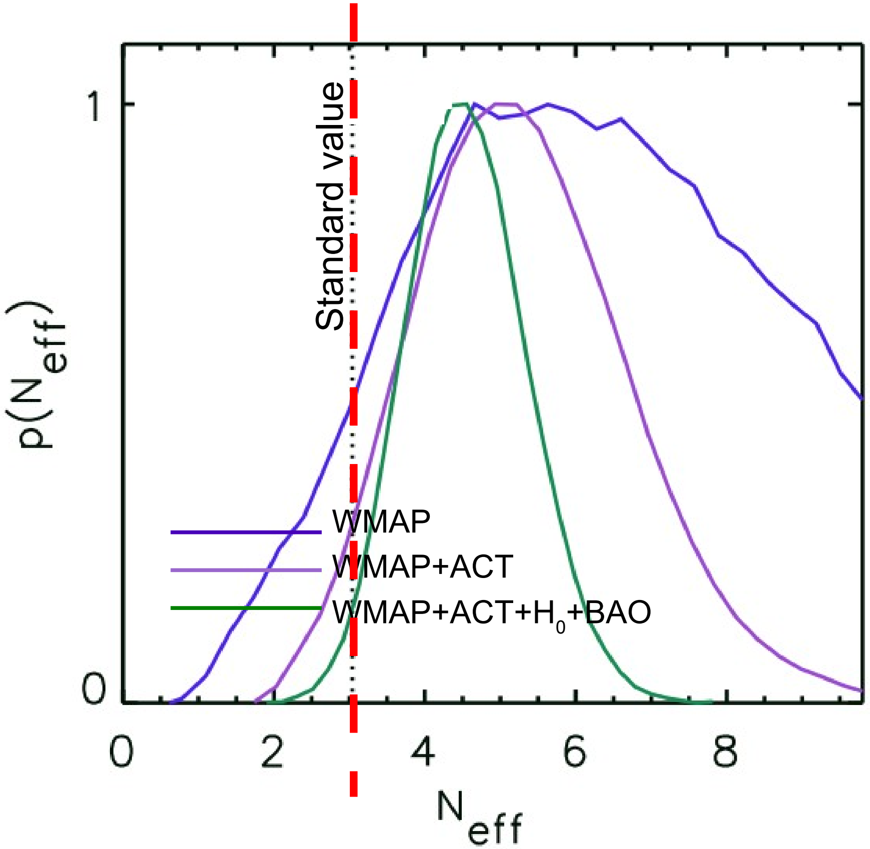
- Connection to the **short baseline sterile neutrino**.

- Bonus slides: **Big bang nucleosythesis**

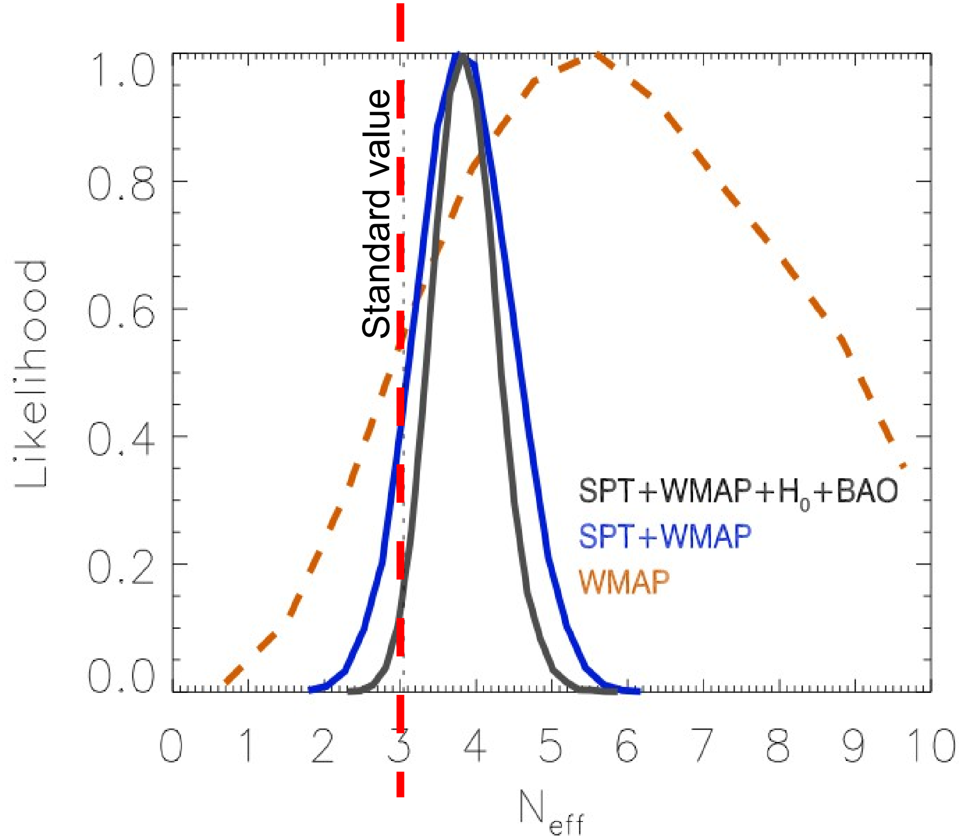
1. CMB+large-scale structure...

Evidence for $N_{eff} > 3$ from CMB+LSS...

- Recent CMB+LSS data appear to prefer $N_{eff} > 3$!



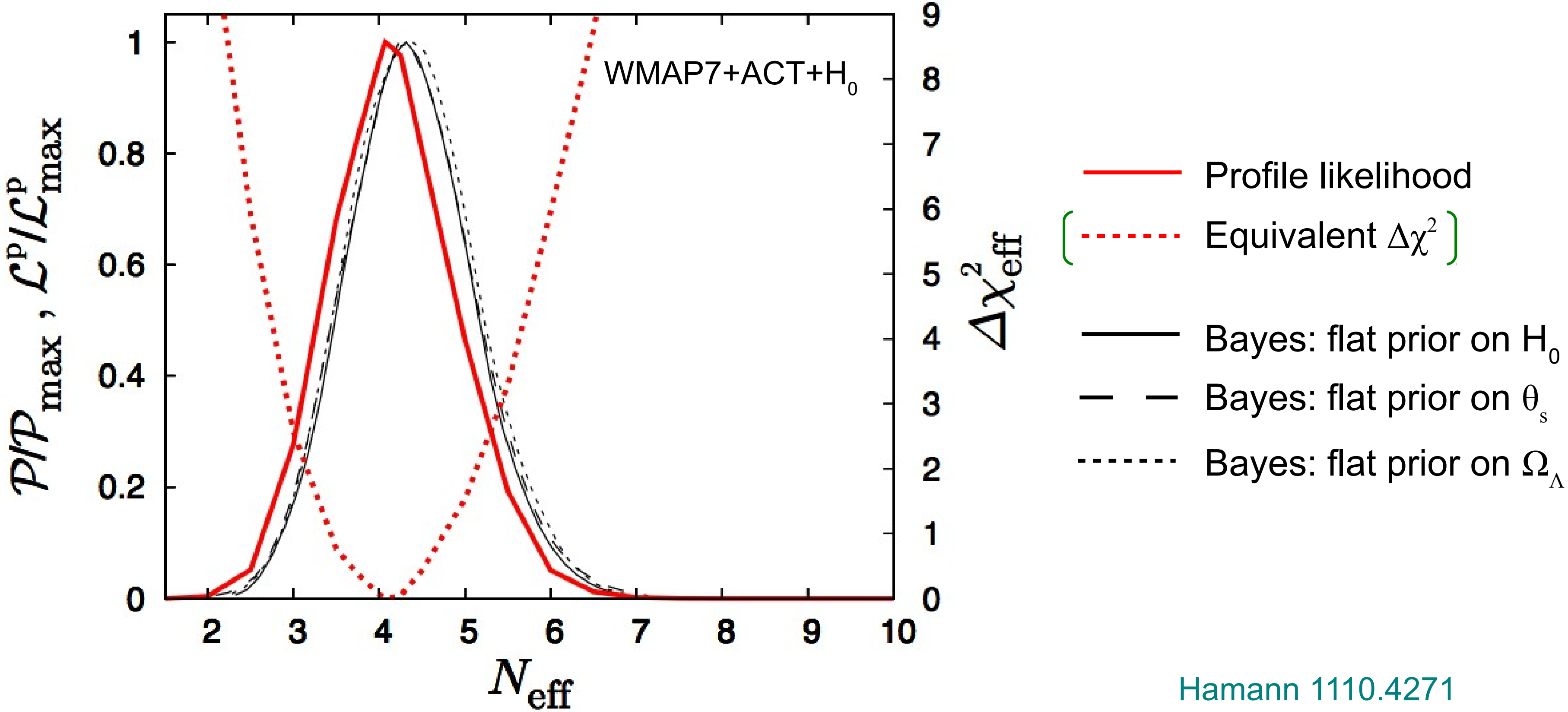
Dunkley et al. [Atacama Cosmology Telescope] 2010



Keisler et al. [South Pole Telescope] 2011

Preference for N_{eff} > 3 is **not** driven by priors...

- The profile likelihood also peaks at $N_{eff} > 3$!



Hamann 1110.4271

Not limited to ACT, SPT...

W-7=WMAP-7

- Many model+data combinations find $N_{eff}>3$ at 95% – 99% C.L.
  - Central value $N_{eff} \sim 4$.

→ = >95% C.L.

Model		Data	N_{eff}
ΛCDM+N_{eff}		W-7+BAO+H_0	$4.34^{+0.86}_{-0.88}$
		W-7+LRG+H_0	$4.25^{+0.76}_{-0.80}$
		W-7+ACT	5.3 ± 1.3
	→	W-7+ACT+BAO+H_0	4.56 ± 0.75
		W-7+SPT	3.85 ± 0.62
	→	W-7+SPT+BAO+H_0	3.85 ± 0.42
	→	W-7+ACT+SPT+LRG+H_0	$4.08^{(+0.71)}_{(-0.68)}$
	→	W-7+ACT+SPT+BAO+H_0	3.89 ± 0.41
		W-7+CL+SPT+BAO+H_0	(< 3.74)
$N_{eff}+f_\nu$		W-7+CMB+BAO+H_0	$4.47^{(+1.82)}_{(-1.74)}$
	→	W-7+CMB+LRG+H_0	$4.87^{(+1.86)}_{(-1.75)}$
$N_{eff}+\Omega_k$		W-7+BAO+H_0	4.61 ± 0.96
	→	W-7+ACT+SPT+BAO+H_0	4.03 ± 0.45
$N_{eff}+\Omega_k+f_\nu$	→	W-7+ACT+SPT+BAO+H_0	4.00 ± 0.43
$N_{eff}+f_\nu+w$		W-7+CMB+BAO+H_0	$3.68^{(+1.90)}_{(-1.84)}$
		W-7+CMB+LRG+H_0	$4.87^{(+2.02)}_{(-2.02)}$
$N_{eff}+\Omega_k+f_\nu+w$	→	W-7+CMB+BAO+SN+H_0	$4.2^{+1.10(+2.00)}_{-0.61(-1.14)}$
	→	W-7+CMB+LRG+SN+H_0	$4.3^{+1.40(+2.30)}_{-0.54(-1.09)}$

Abazajian et al., “Light sterile neutrinos: a white paper”, 2012

Not limited to ACT, SPT... but...

W-7=WMAP-7

- Many model+data combinations find $N_{eff}>3$ at 95% – 99% C.L.
  - Central value N_{eff} ~ 4.

- **One exception**: cluster abundance prefers a more "standard" value:

$$N_{\rm eff} < 3.74 \ (95\,\%\,\mathrm{C.L.})$$

(N_{eff} restricted to ≥ 3)

Burenin & Vikhlinin 2012

→ = >95% C.L.

→ = uses cluster data

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Essentially a WMAP result...

W-7=WMAP-7

- Data combinations are **not** independent.
- All use CMB data from WMAP.

→ Planck will provide an independent check!

→ (red) = >95% C.L.

→ (green) = uses cluster data

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How does it work...

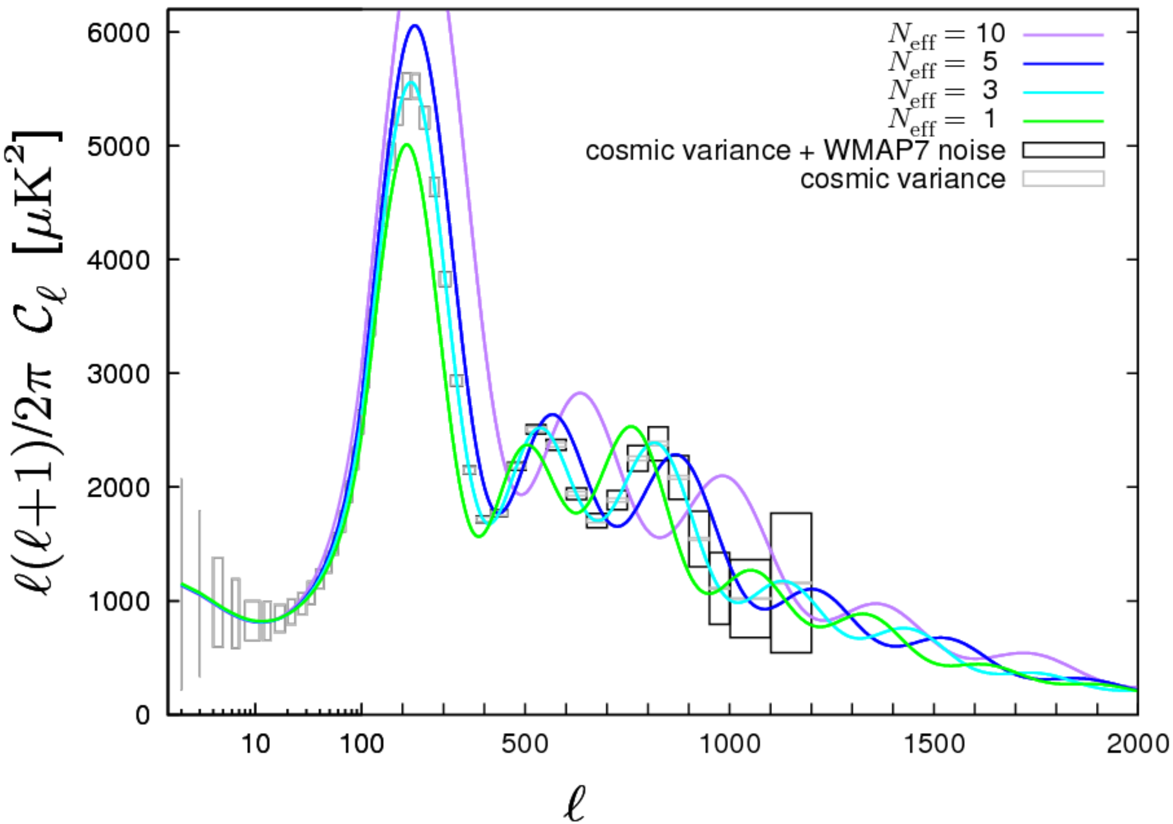


Figure courtesy of J. Hamann

- $N_{\rm eff}$ looks easy to detect..
- But we also use the same data to measure at least 6 other cosmological parameters:

$$(\omega_b, \omega_m, h, A_s, n_s, \tau)$$

baryon density (ω_b), matter density (ω_m), Hubble parameter (h), primordial fluctuation amplitude & spectral index (A_s, n_s), optical depth to reionisation (τ)

- Plenty of parameter degeneracies!

What the CMB really probes: equality redshift...

- Ratio of 3rd and 1st peaks sensitive to the redshift of matter-radiation equality via the early ISW effect.

Exact degeneracy between the physical matter density ω_m and N_{eff}.

$$1+z_{eq}=\frac{\omega_m}{\omega_r}\simeq\frac{\omega_m}{\omega_\gamma}\frac{1}{1+0.2271\,N_{eff}}$$

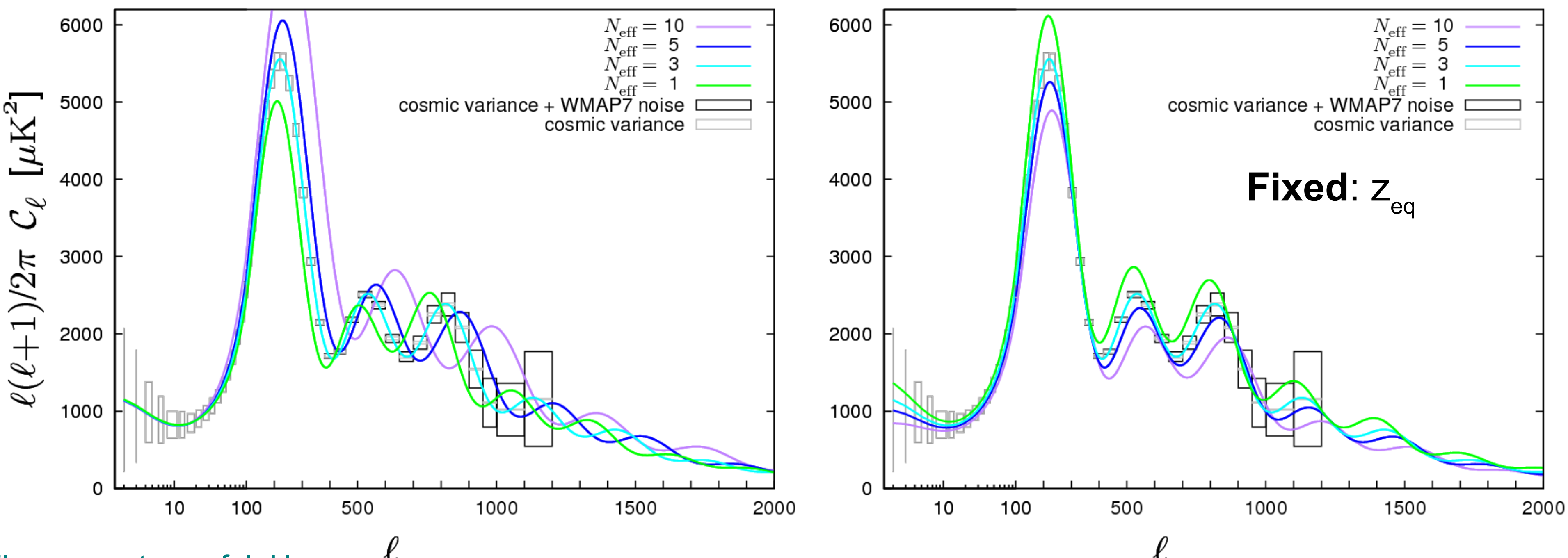


Figure courtesy of J. Hamann

What the CMB really probes: sound horizon...

- Peak positions depend on:

$$\theta_s = \frac{r_s}{D_A}$$

Sound horizon at decoupling (r_s)

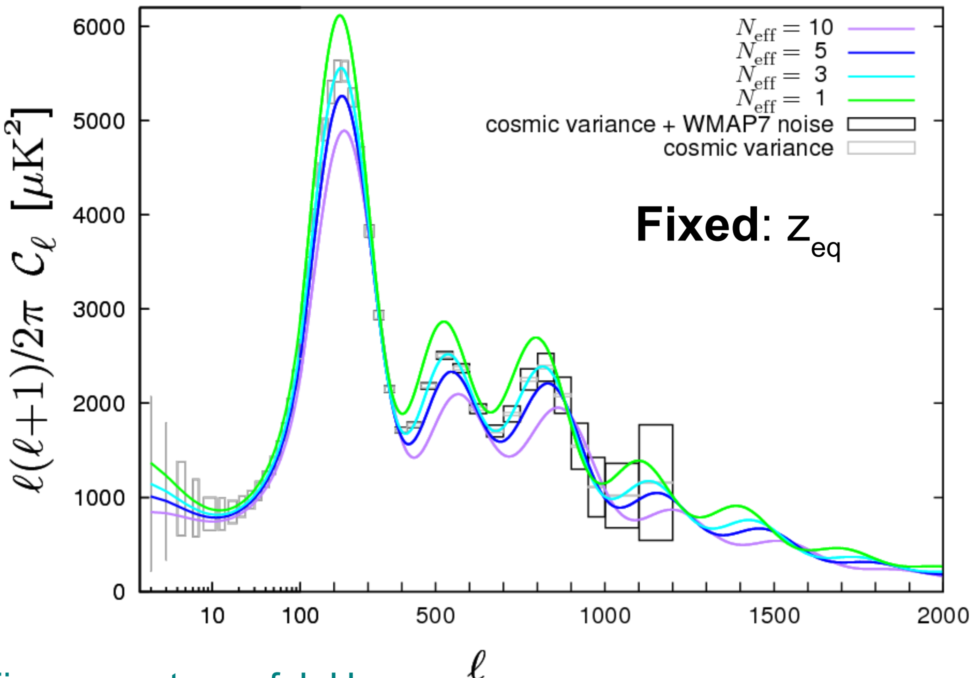
Angular distance to the last scattering surface (D_A)

Flat ΛCDM →

Exact degeneracy between ω_m and the Hubble parameter h.

$$\theta_s \propto \frac{(\omega_m h^{-2})^{-1/2}}{\int_{a_*}^{1} \frac{da}{\sqrt{\omega_m h^{-2} a^{-3} + (1-\omega_m h^{-2})}}}$$

Fixed z_{eq}, ω_b



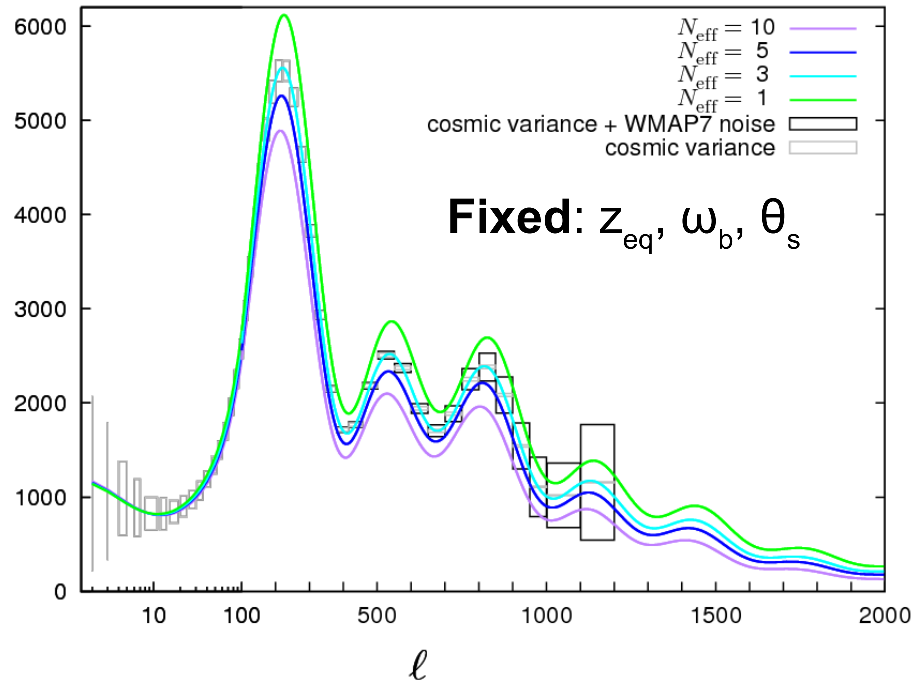
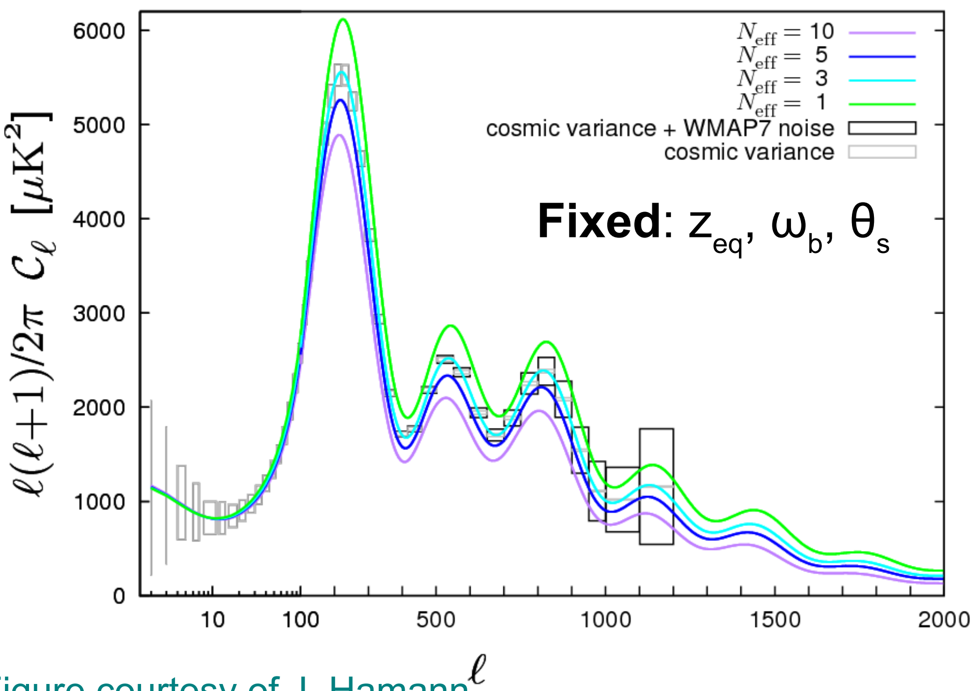


Figure courtesy of J. Hamann

What the CMB really probes: anisotropic stress...

- Apparent (i.e., not physical) partial degeneracies with primordial fluctuation amplitude A_s and spectral index n_s.

- However, free-streaming particles have **anisotropic stress**.
- First real signature of N_{eff} in the 3rd peak!



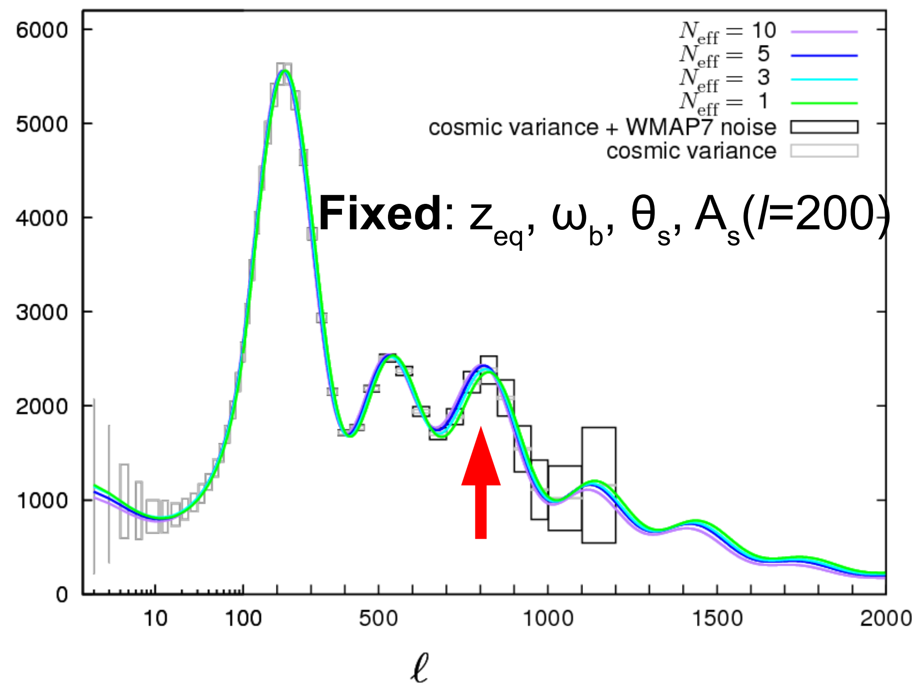
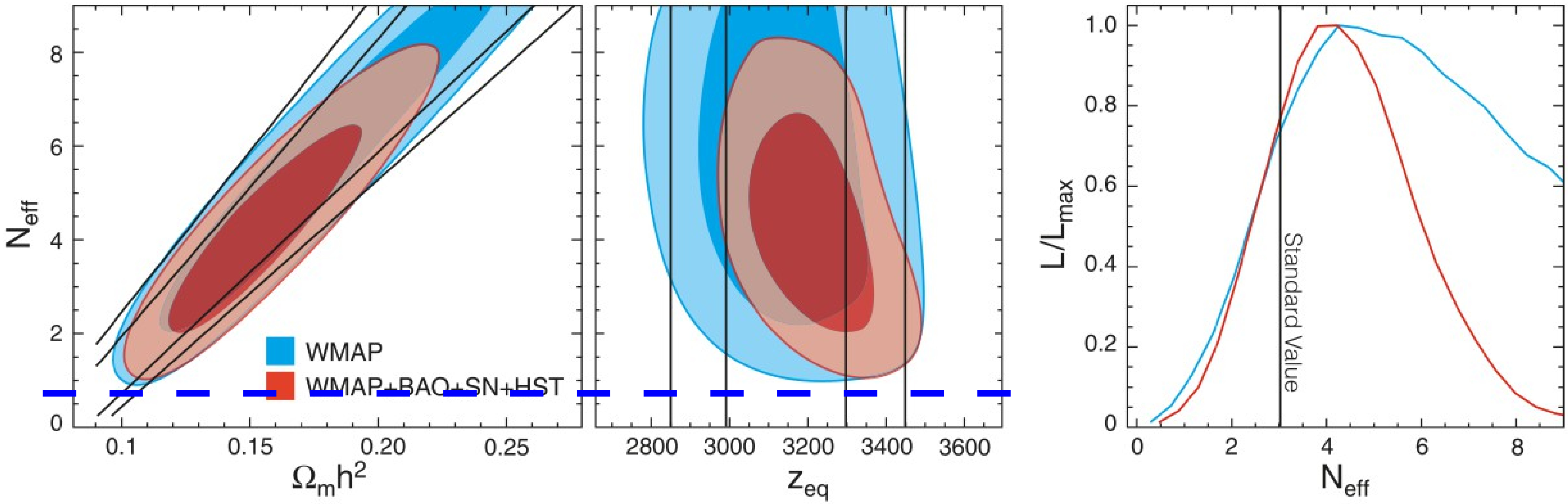


Figure courtesy of J. Hamann

- Measurement of the third peak (since WMAP-5) gives lower limit on N_{eff} from WMAP alone (without supplementary large-scale structure data).



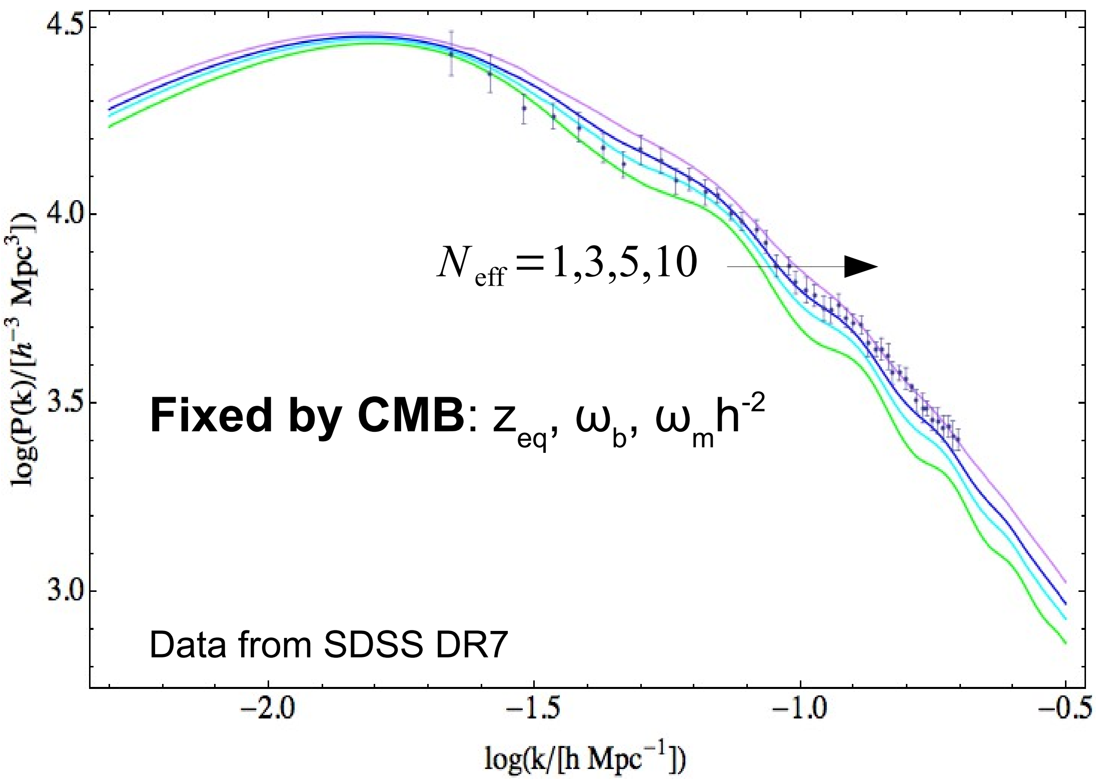
Komatsu et al. [WMAP5] 2008

- **Upper limit** requires combination of WMAP with other observations to break the remaining N_{eff}–ω_m–h parameter degeneracies.
  - Pinning down either ω_m or h will do!

from local (z<0.1) expansion rate measurements

Breaking degeneracies with large-scale structure...

Large-scale matter power spectrum



- The shape of the matter power spectrum is additionally sensitive to the baryon fraction:

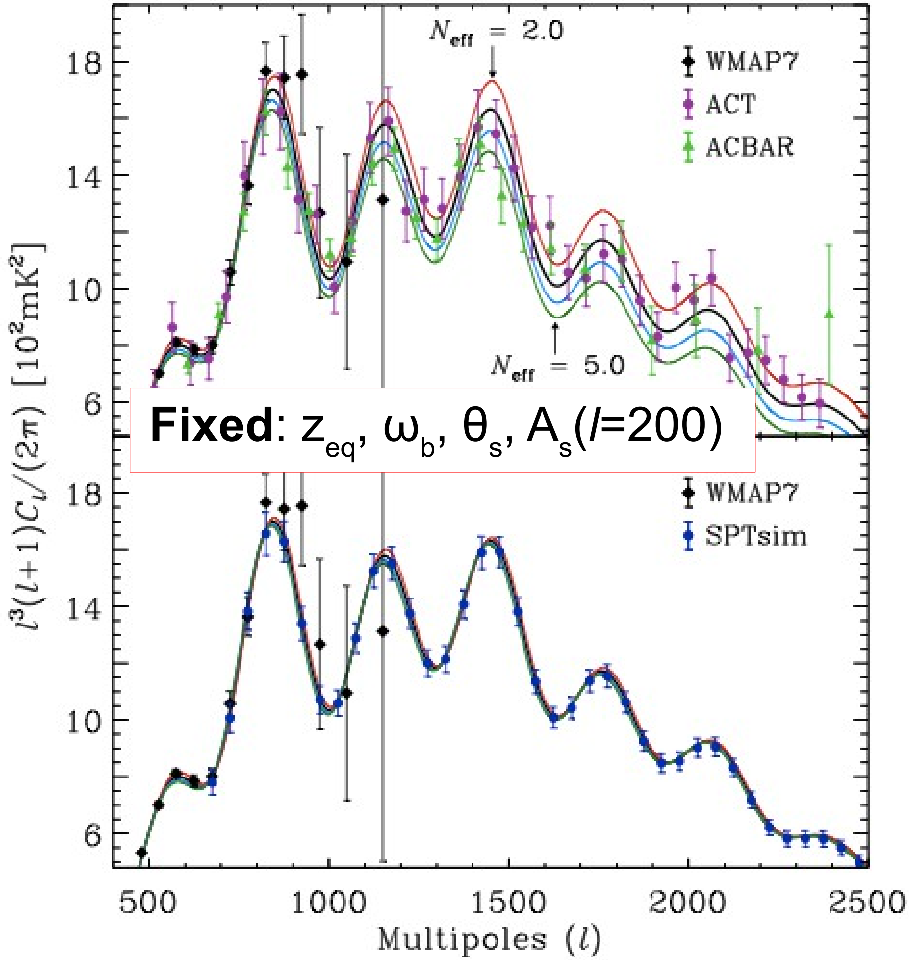
$$f_b \equiv \frac{\omega_b}{\omega_m}$$

Fixed by CMB (ω_b)

- The larger N_{eff}, the smaller f_b → more power at large k.

- (Can partially offset this effect with massive neutrinos.)

Breaking degeneracies with the CMB damping tail...



Fixed: z_{eq}, ω_b, θ_s, $A_s(l=200)$

Hou, Keisler, Knox et al. 2011

- **ACT** data available since 2010; **SPT** since 2011; also measured by **Planck**.
- Probe photon diffusion scale:

$$\theta_d = \frac{r_d}{D_A}$$

← Diffusion scale at decoupling

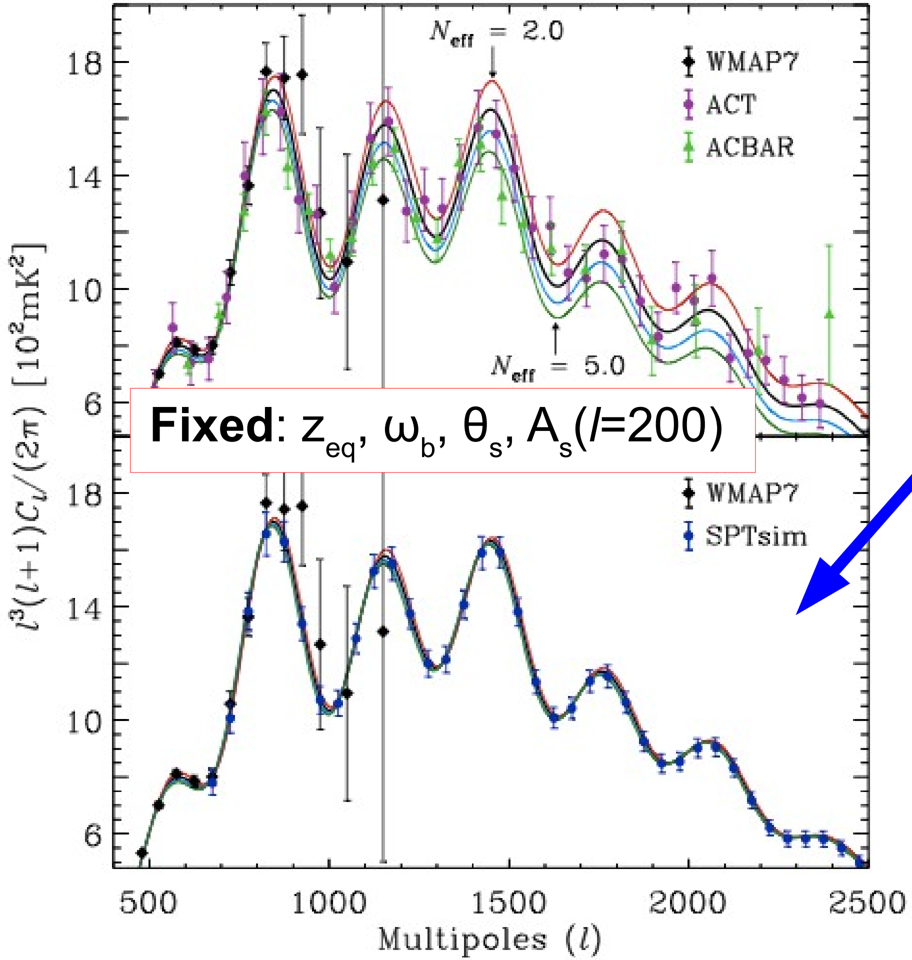
- Combined with sound horizon measurement:

$$\frac{\theta_d}{\theta_s} = \frac{r_d}{r_s} \propto \omega_m^{1/4}$$

← Fixed z_{eq}, ω_b

Breaks (nearly) all N_{eff} degeneracies, robust against low-redshift uncertainties in D_A, e.g., w_{DE}, Ω_k.

Breaking degeneracies with the CMB damping tail...



Hou, Keisler, Knox et al. 2011

- **The N_{eff} – Y_p degeneracy**!
- With ω_b fixed by WMAP, the Helium fraction Y_p determines the free electron density → affects photon diffusion length.
- Current strategy:
  - Either fix Y_p at 0.24 → 0.25.
  - Or apply a BBN consistency relation.
- **Not** an exact degeneracy; can be resolved by Planck.

A quick recap on N_{eff}...

- The WMAP measurement of the acoustic peaks alone does **not** completely constrain N_{eff} because of parameter degeneracies:
  - $N_{eff} - \omega_m$
  - N_{eff} – h (via ω_m – h and $N_{eff} - \omega_m$)
  - $N_{eff} - Y_p$

- Degeneracies can be broken with measurements of
  - CMB damping tail → ω_m (Y_p with Planck)
  - Large-scale structure distribution → ω_m
  - Local Hubble expansion rate → h

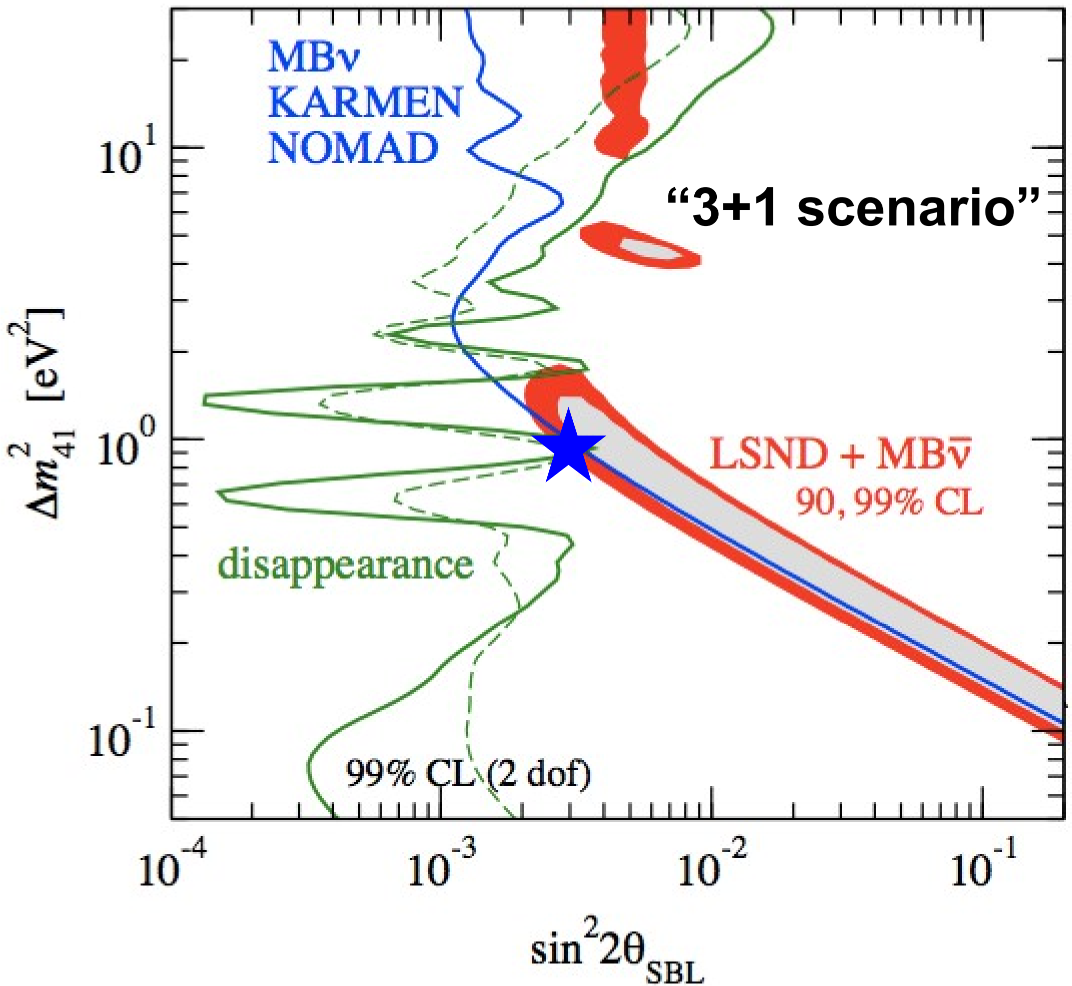
- Preference for N_{eff}>3 appears to be robust against model assumptions.

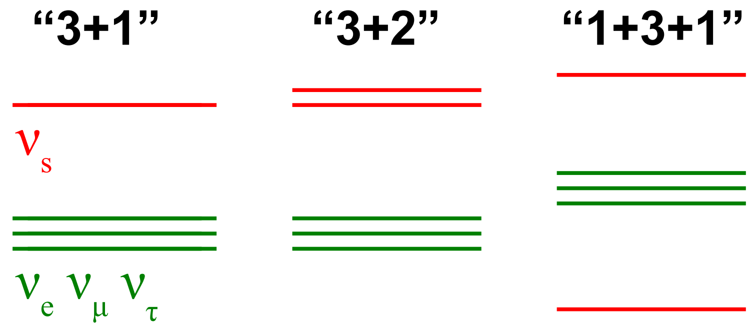
2. Connection to the short baseline sterile neutrino...

Experimental anomalies & the sterile ν interpretation...

- Experiments at odds with the standard **3-neutrino interpretation** of global neutrino oscillation data:
  - LSND ($\bar{\nu}_e$ appearance)
  - MiniBooNE anti-neutrinos ($\bar{\nu}_e$ appearance)
  - Short baseline reactor experiments (re-evaluation of neutrino fluxes) ($\bar{\nu}_e$ disappearance)

- If interpreted as oscillation signals → a 4th (or more) **sterile neutrino** with $\Delta m^2 \sim O(1\ eV^2)$ and $\sin^2 2\theta > 10^{-3}$.





- "3+1" best-fit: $\Delta m^2_{41} \sim 1 \ \text{eV}^2$

$$\rightarrow \boxed{m_s \sim 1 \ \text{eV}}$$

If lightest neutrino mass ~ 0 eV

Kopp, Maltoni & Schwetz 1103.4570

——— New reactor fluxes

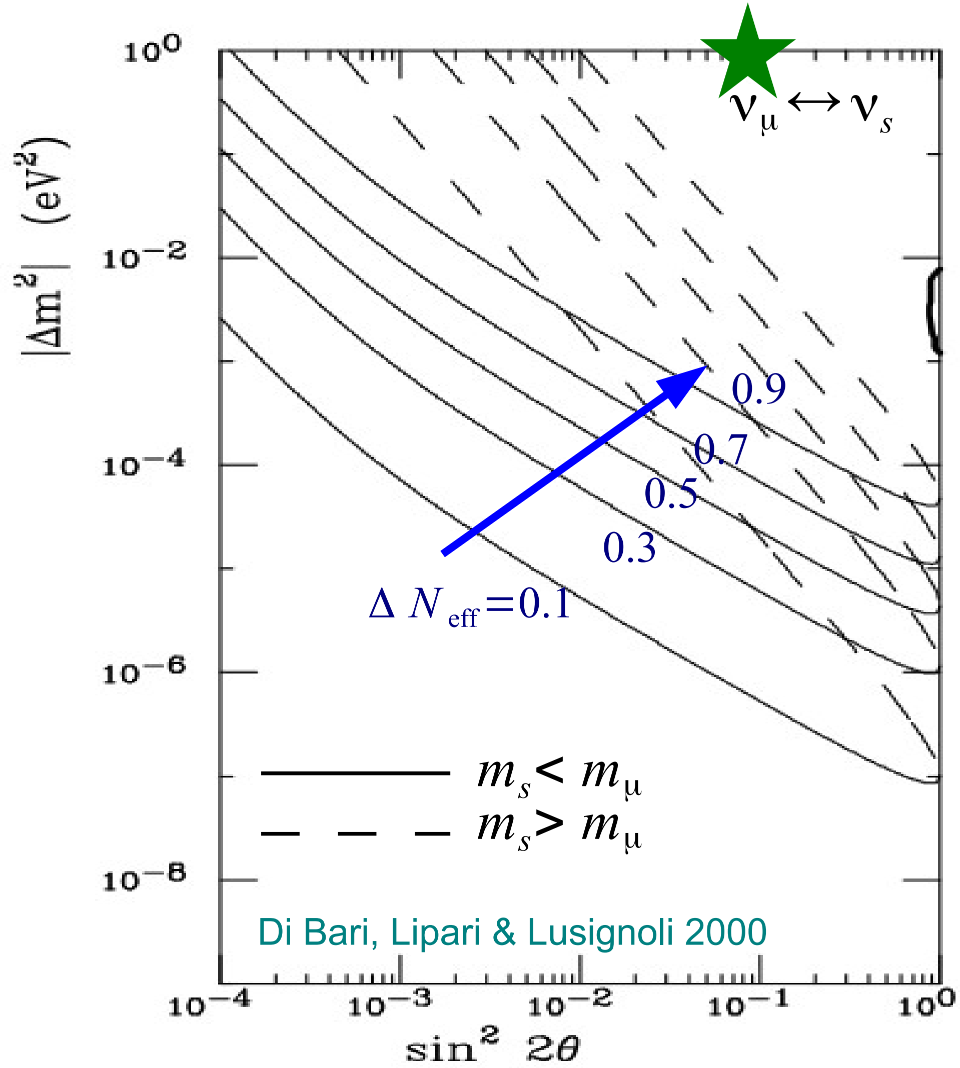
- - - - Old reactor fluxes

Light sterile neutrinos and N_{eff}...

- SBL-preferred Δm^2 and mixing favour the **production and thermalisation** of sterile neutrinos in the early universe via $\nu_\alpha \leftrightarrow \nu_s$ oscillations + ν_α scattering.

→ Can easily produce an excess relativistic energy density of $\Delta N_{eff} \sim 1$.

→ Sterile states have the same temperature as the SM neutrinos.



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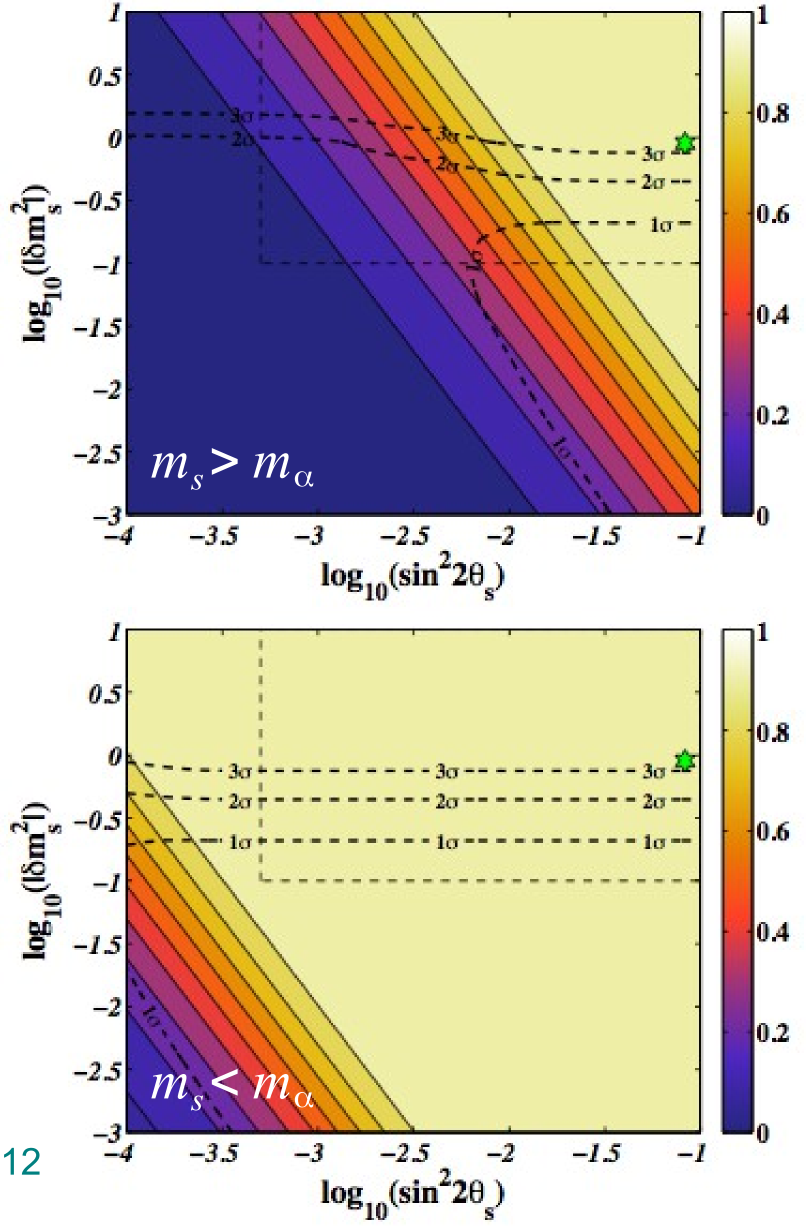
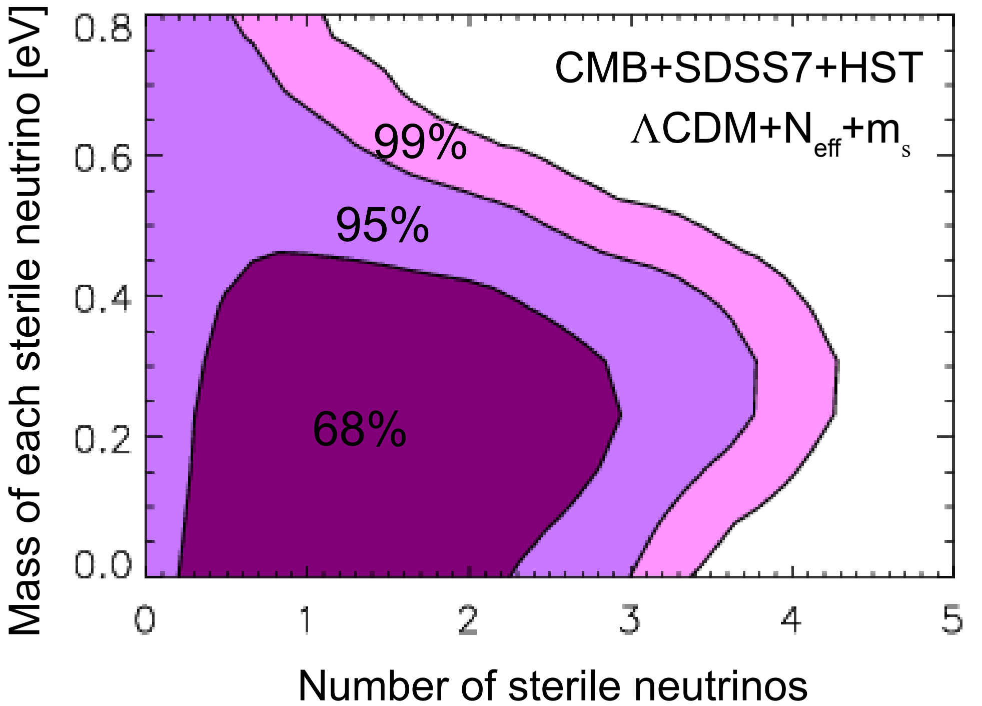


Hannestad, Tamborra & Tram 2012

Can the short baseline sterile neutrino explain $N_{eff} > 3$?

- **Short answer**: Not so easy.
- **Reason**: eV mass neutrinos violate CMB+LSS hot dark matter bounds.



Hamann, Hannestad, Raffelt, Tamborra & Y[3]W 2010

- 3+1 thermalised sterile:

$$m_s < 0.48 \text{ eV} \quad (95\%\, C.I.)$$

Lab best-fit: $m_s \sim 1$ eV

- 3+2 thermalised sterile:

$$m_{s1} + m_{s2} < 0.9 \text{ eV} \quad (95\%\, C.I.)$$

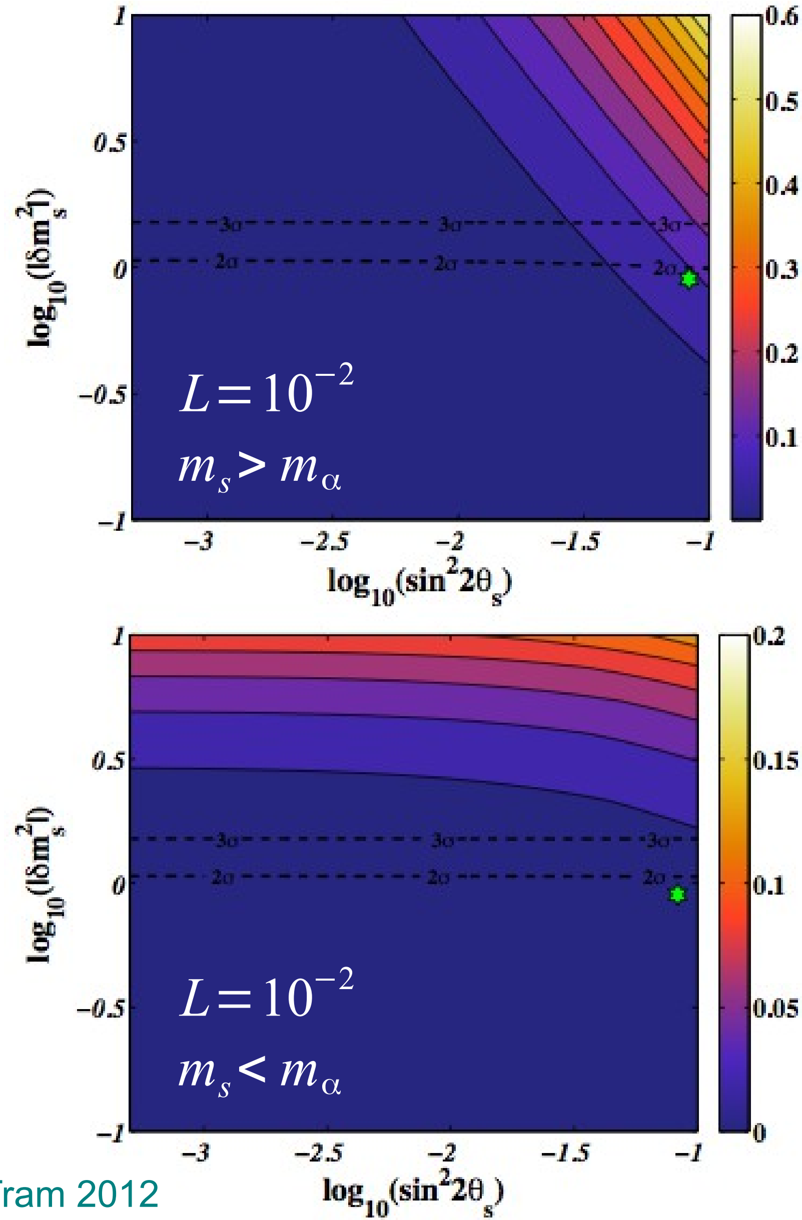
Lab best-fit: $m_{s1} \sim 0.7$ eV, $m_{s2} \sim 0.9$ eV

Is there a way out? Plan A...

- Suppress sterile neutrino thermalisation using, e.g., a large lepton asymmetry (L >> B ~ 10^{-10}).

  Foot & Volkas 1995

  - Generating a large lepton asymmetry requires new physics.
  - If complete suppression, then $N_{eff} > 3$ must be explained by some other physics (sub-eV thermal axions, hidden photons, etc.?)



Hannestad, Tamborra & Tram 2012

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Is there a way out? Plan B...

- Failing to suppress ν_s thermalisation, exploit parameter degeneracies in the CMB+LSS to **engineer a good fit**.

- No room for play within the ΛCDM model, but extensions of ΛCDM can help to relax the hot dark matter constraint on m_s:

  - Non-standard dark energy equation of state.
  - Modified gravity.
  - Non-flat spatial geometry.
  - Even more massless degrees of freedom.
  - …

Elgarøy & Kristiansen 2011; Hamann, Hannestad, Raffelt & Y[3]W 2011
Giusarma et al. 2012; Motohashi, Starobinsky & Yokoyama 2012;
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  - …

1 x 1 eV sterile neutrino can be reasonably accommodated.

1 x 2eV or 2 x 1 eV is still problematic...

Elgarøy & Kristiansen 2011; Hamann, Hannestad, Raffelt & Y^{3}W 2011
Giusarma et al. 2012; Motohashi, Starobinsky & Yokoyama 2012;
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Is there a way out? Plan B...

Modified gravity scenario to explain accelerated expansion in lieu of dark energy

- **An example**: accommodating 1eV sterile neutrinos with f(R) gravity:

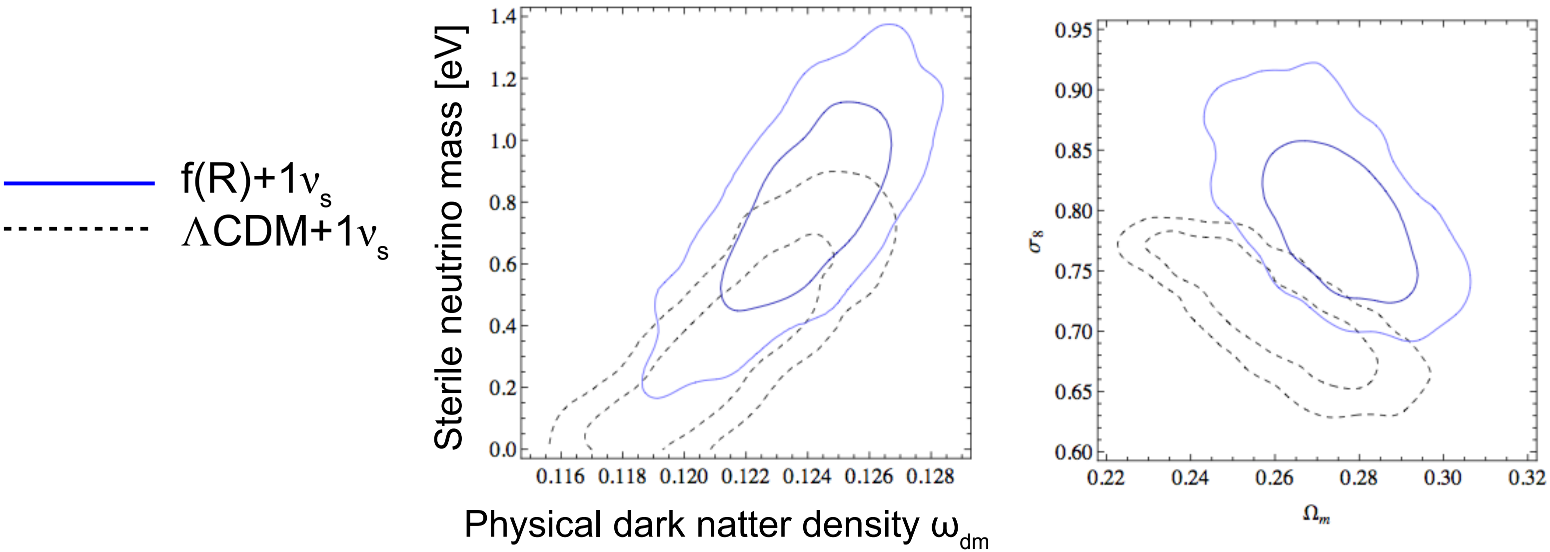


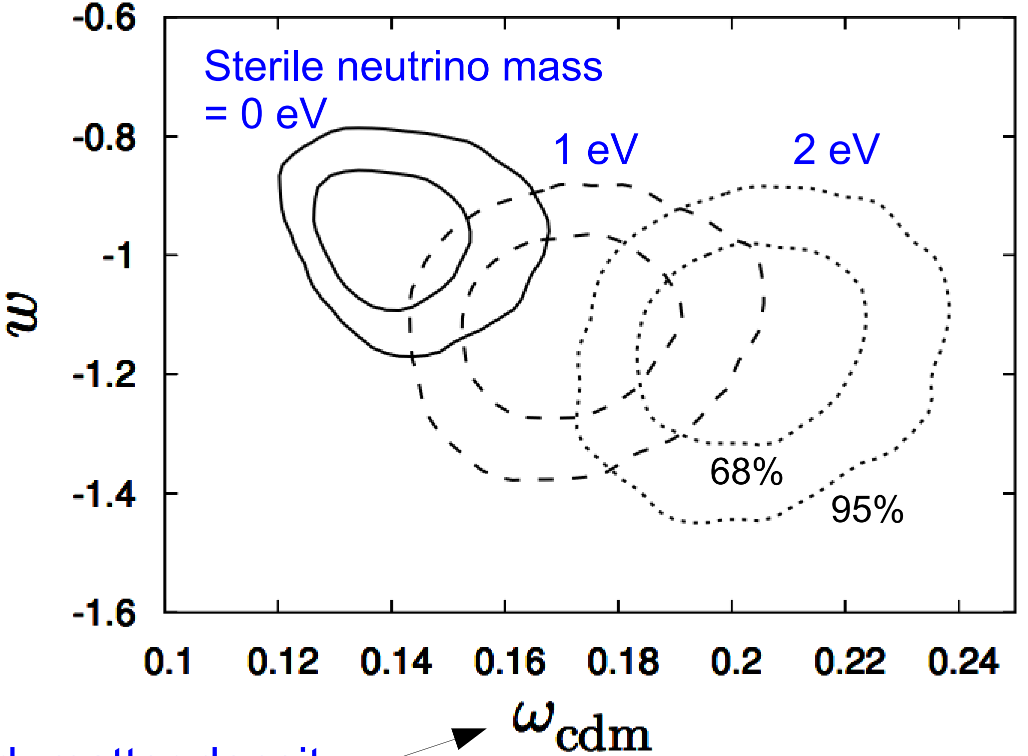
FIG. 1: 1 and 2σ contours of the sterile neutrino mass (left) and σ_8 (right) for the cases with three massless and one massive neutrinos in the ΛCDM model (dashed black) and $f(R)$ gravity (solid blue). $\chi^2_{\rm eff} = 3774.1$ and 3767.0, respectively.

Motohashi, Starobinsky & Yokoyama 2012

Necessary side effects...

- Exploiting parameter degeneracies also implies that other (unrelated) cosmological parameter values will change.

Dark energy equation of state parameter



Cold dark matter density

Hamann, Hannestad, Raffelt & Y[3]W 2011

Planck and N_{eff}...

- If N_{eff} is as large as 4, it will be settled almost immediately by Planck (launched May 14, 2009; public data release early 2013).

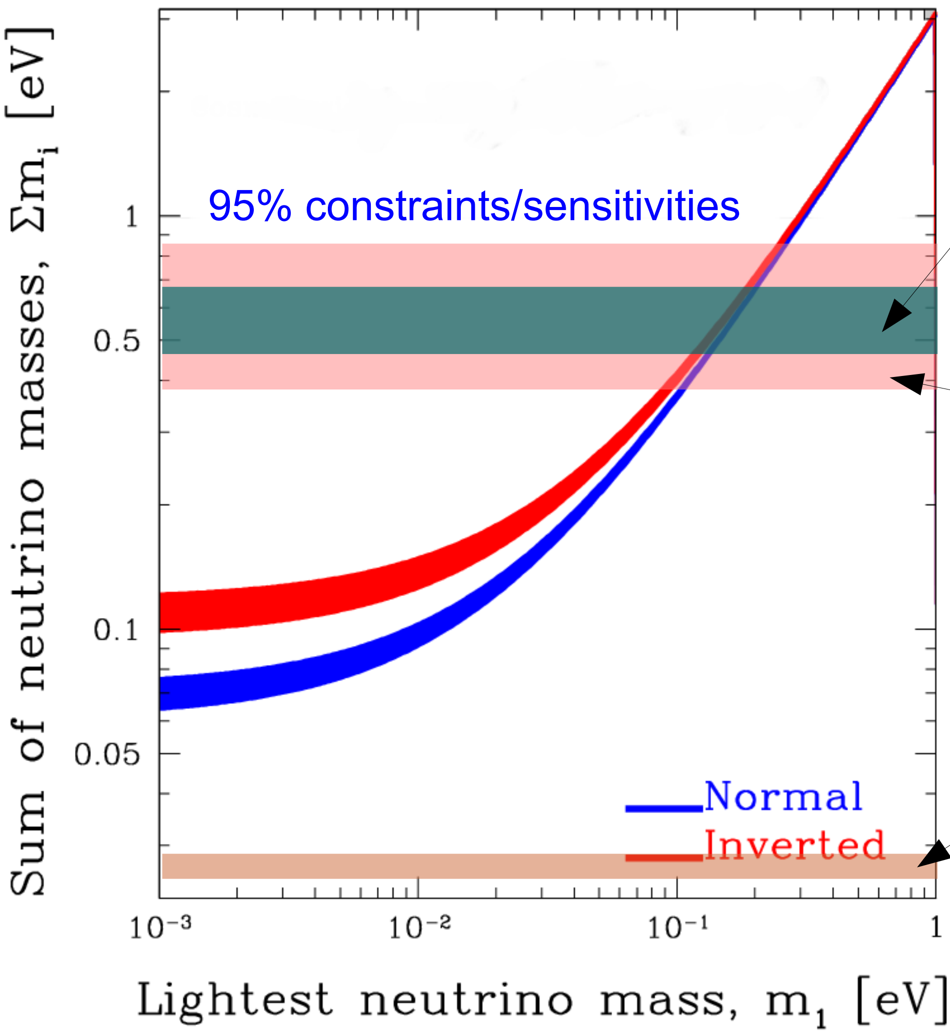
68% sensitivities

Experiment	$f_{\rm sky}$	θ_b	$w_T^{-1/2}$ [μ K']	$w_P^{-1/2}$ [μ K']	ΔN_ν TT	ΔN_ν TT+TE+EE	ΔN_ν (free Y) TT+TE+EE
Planck	0.8	7'	40	56	0.6	0.20	0.24
ACT	0.01	1.7'	3	4	1	0.47	0.9
ACT + Planck					0.4	0.18	0.24
CMBPOL	0.8	4'	1	1.4	0.12	0.05	0.09

Bashinsky & Seljak 2004

Helium fraction
as a free parameter

Planck and neutrino mass...



Current constraints

Band = model complexity

Planck alone (1 year) 2013
(+ current LSS probes → not much improvement)

Perotto et al. 2006

Not quite good enough to rule in or rule out 1 eV sterile neutrinos definitively...

Planck+Euclid (lensing+galaxy clustering)
~ 2025

Hamann, Hannestad & Y^3W 2012

Summary...

- Current precision cosmological data show a preference for extra relativistic degrees of freedom (beyond 3 neutrinos).

- Sterile neutrino interpretation of short baseline neutrino anomalies does not quite fit into the simplest picture though...

  - 3+2: **Too many** for BBN

  - 3+1, 3+2: **Too heavy** for CMB/LSS

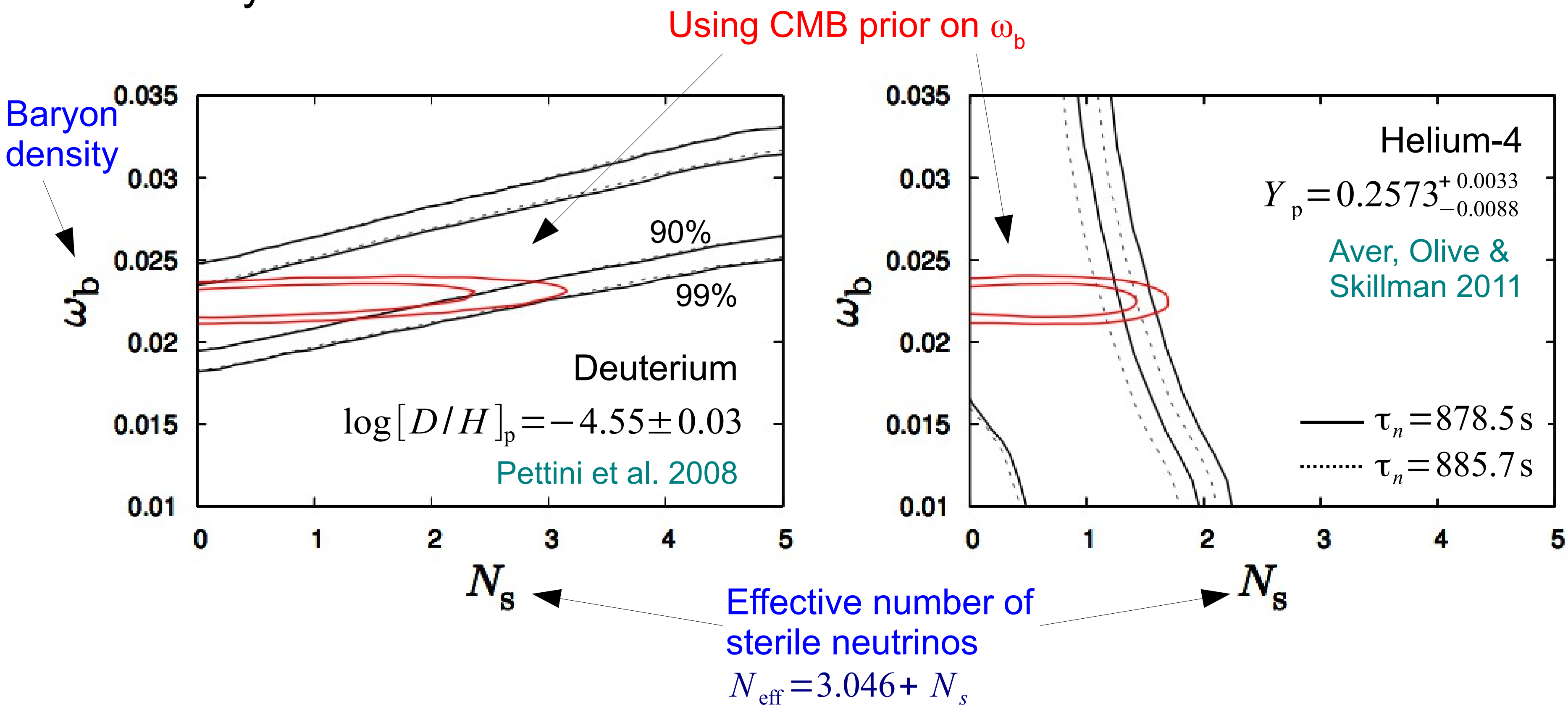
- Non-trivial extensions to ΛCDM can reasonably accommodate 1 x 1 eV fully thermalised sterile neutrino species.

- **Planck with tell** (at least part of the story).

3. Extra slides: BBN

Evidence for $N_{eff} > 3$ from BBN...

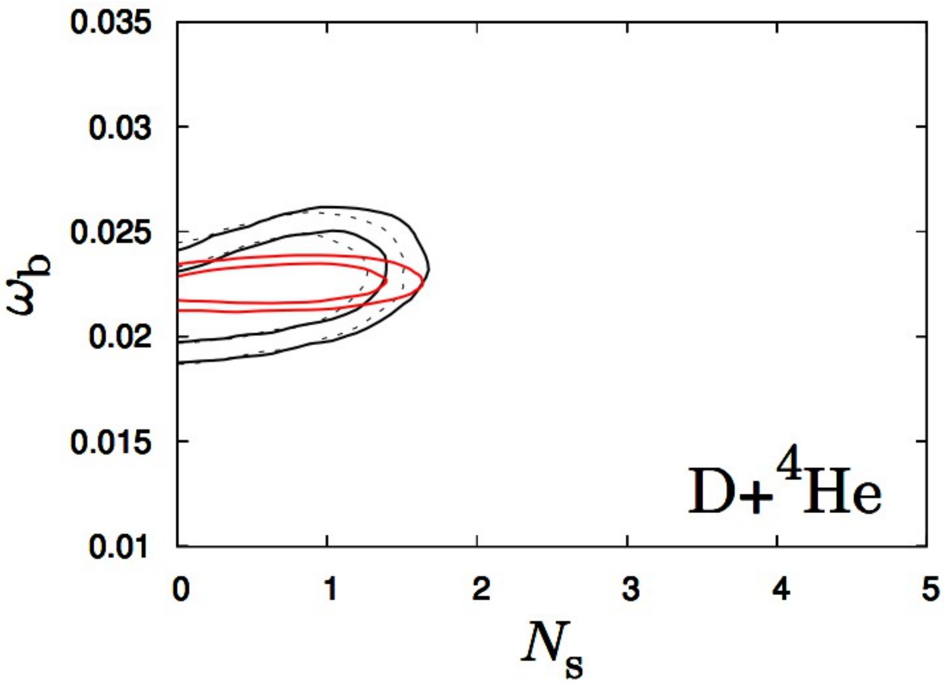
- Light element abundances are sensitive to excess relativistic energy density.

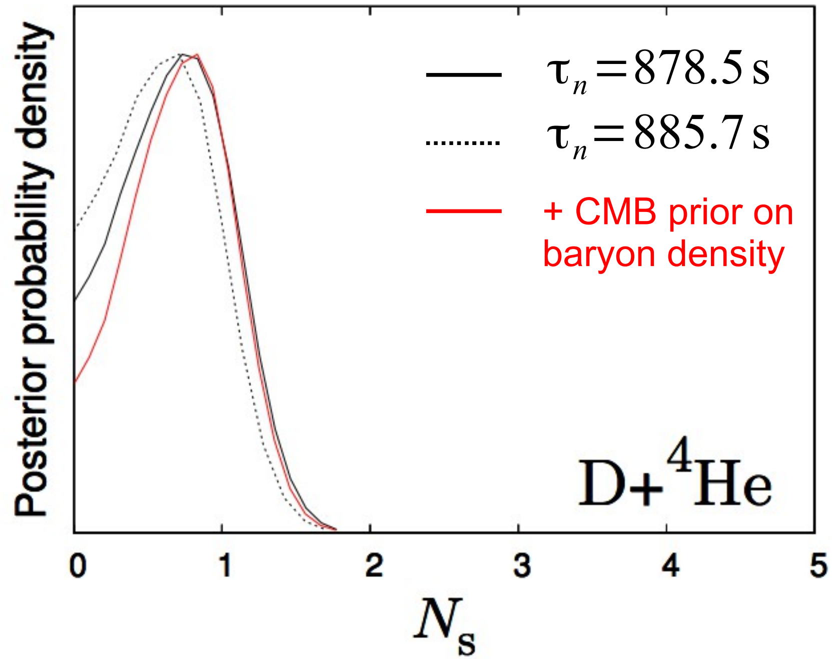


Hamann, Hannestad, Raffelt & Y[3]W 2011

Evidence for $N_{eff} > 3$ from BBN...

- Mild preference for $N_{eff} > 3$ (or $N_s > 0$) from Deuterium+Helium-4.
- But $N_s = 2$ is strongly disfavoured.





Hamann, Hannestad, Raffelt & Y[3]W 2011

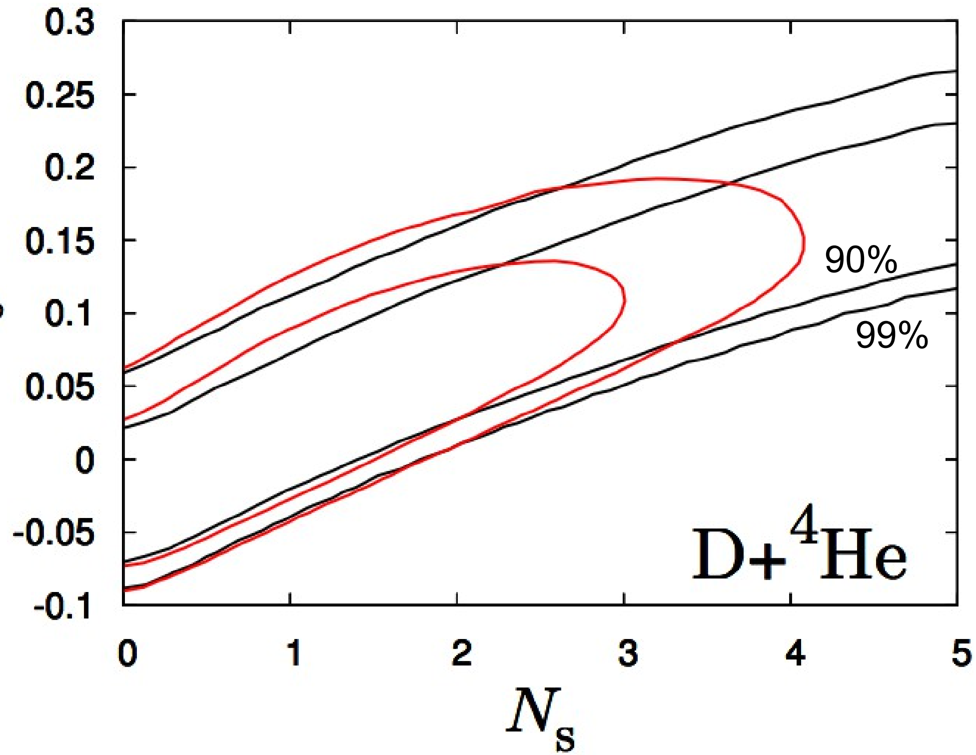
Quick fix: degenerate BBN...

- Introduce a neutrino chemical potential (= O(0.1) lepton asymmetry).
- Then even $N_s = 3$ is allowed by BBN.

Lepton asymmetry

$$L \equiv \frac{n_{\nu_\alpha} - n_{\bar{\nu}_\alpha}}{n_\gamma} = \frac{1}{12\zeta(3)} \left(\frac{T_\nu}{T_\gamma}\right)^3 (\pi^2 \xi + \xi^3)$$

Neutrino chemical potential



Hamann, Hannestad, Raffelt & Y^3W 2011

Question: How to simultaneously get $L = O(0.1)$ and $B = O(10^{-10})$?