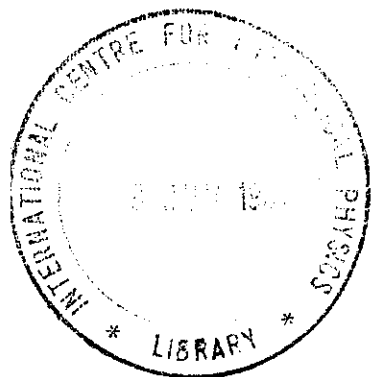


INTERNATIONAL ATOMIC ENERGY AGENCY
UNITED NATIONS EDUCATIONAL, SCIENTIFIC AND CULTURAL ORGANIZATION



INTERNATIONAL CENTRE FOR THEORETICAL PHYSICS
34100 TRIESTE (ITALY) - P.O. B. 500 - MIRAMARE - STRADA COSTIERA 11 - TELEPHONE: 2240-1
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H4.SMR/210 - 33

SPRING COLLEGE ON PLASMA PHYSICS

(25 May - 19 June 1987)

PLASMA PHYSICS PHENOMENA DETECTED AT COMET GIACOBINI-ZINNER

F. SCARF

TRW Space and Technology Group
California, U.S.A.

Plasma Physics Phenomena Detected at Comet Giacobini-Zinner

by

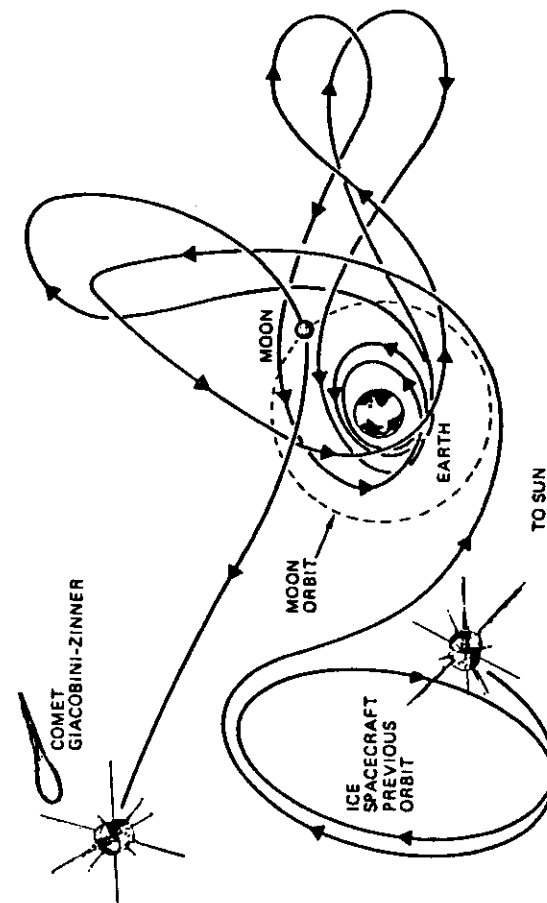
Frederick Scarf

TRW Space and Technology Group

Redondo Beach, California, 90278

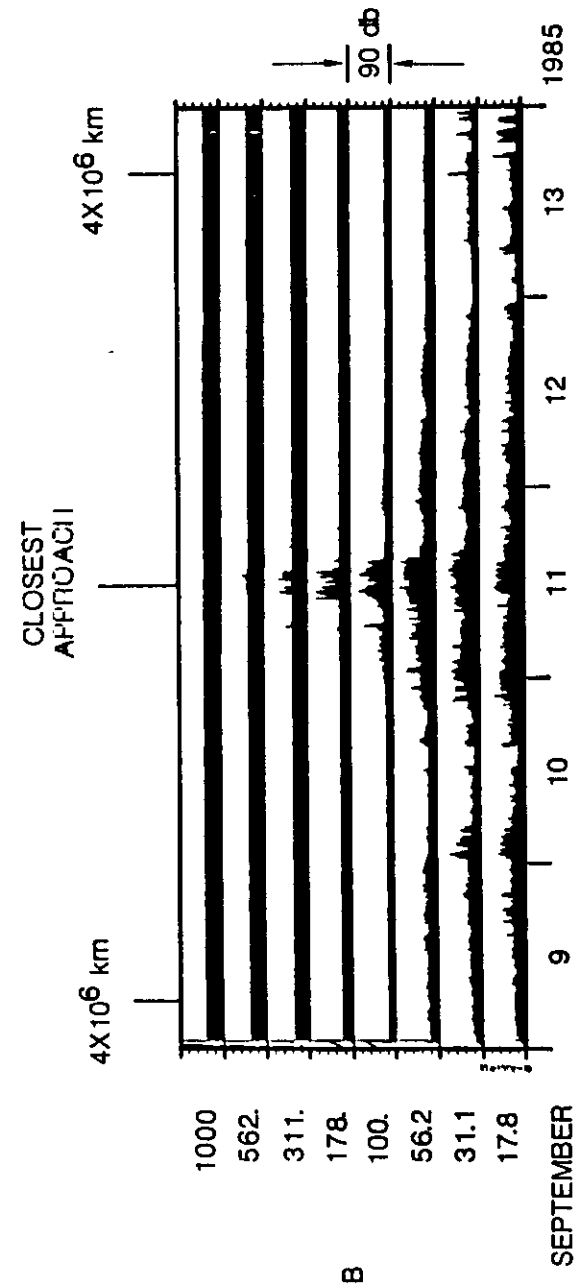
USA

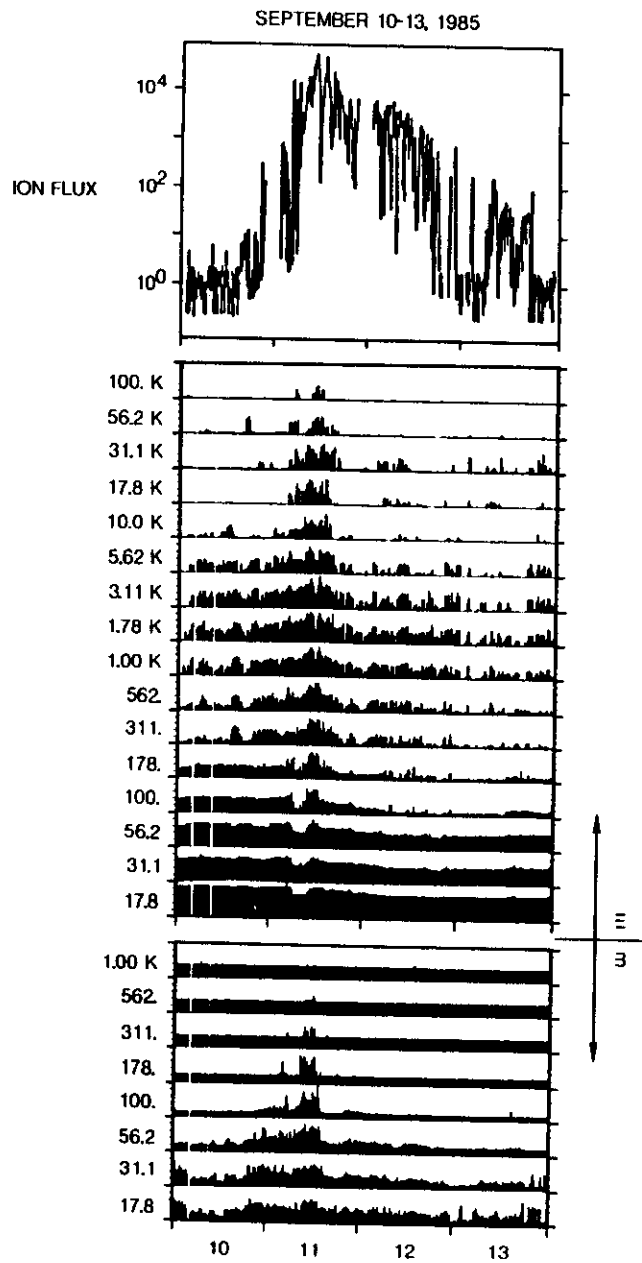
On September 11, 1985, the International Cometary Explorer (ICE) became the first spacecraft to encounter a comet when it traversed the tail of Giacobini-Zinner at a downstream distance of 7800 km. The spacecraft payload included instruments to measure the magnetic field (DC to 1 kHz), the low energy electron distribution functions ($E < 1$ keV), electric field wave phenomena (3 Hz to several MHz), ion composition (thermal and energetic ions), and energetic particle distributions in the range 30 keV to 500 keV. These instruments detected extremely strong interactions that extended over distance scales measured in millions of kilometers. Unique plasma physics phenomena were found to be very important at the comet. The comet interaction processes include pickup and energization of newly-born cometary ions at great distances, effects of mass loading on the solar wind-comet interaction, development of intense plasma turbulence and formation of a diffuse low-Mach number bow shock.



ICE SCIENTIFIC MEASUREMENTS

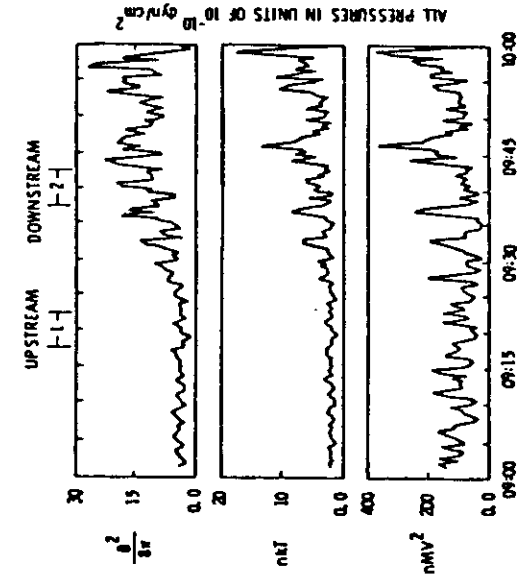
P.I.	INSTRUMENT	SAMPLING PERIOD (SEC)	SPATIAL RESOLUTION (km)
1. SMITH	MAGNETOMETER	1/3	7
2. SCARF	PLASMA WAVES		
	16 CHANNEL E	1	21
	8 CHANNEL B	16	330
3. BAME	PLASMA ELECTRONS	24	500
4. HYND	ENERGETIC PROTONS	32	660
5. STEINBERG	RADIO WAVES	56	1,200
6. HOVESTADT	ENERGETIC ION COMPOSITION	504	10,800
7. OGILVIE	PLASMA IONS	1,270	26,000





85-11

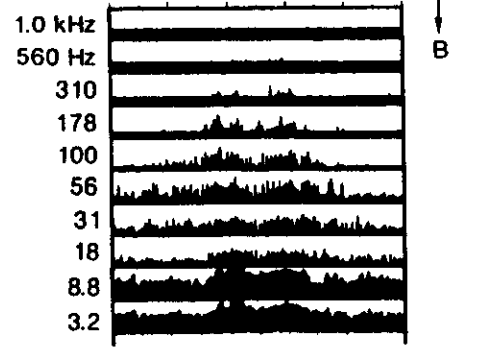
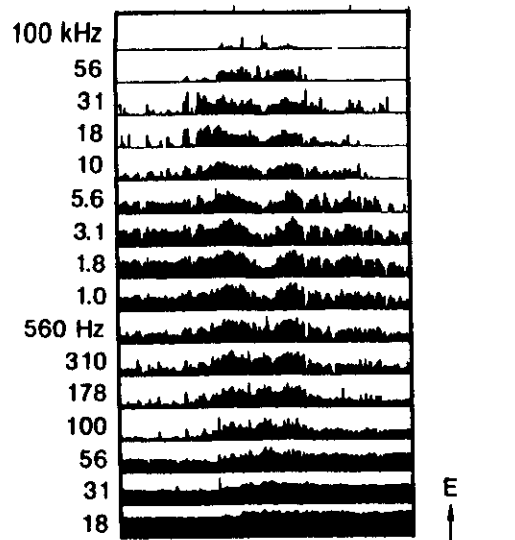
INBOUND: Comet Giacobini-Zinner



Three of the principal pressures (magnetic, solar wind electron and solar wind convective pressures) are shown. The horizontal bars above the data show the analysis intervals corresponding to the results in Table 1.

TABLE 1.

- $\hat{n}$: 0.1210, -0.7089, -0.6948
- θ_{Bn} : 80.1° PERPENDICULAR
- $V_s = |\Delta V \times B_2| / |\Delta B| = 189 \text{ km/s}$
- $V_s = (N_2/N_2 - N_1) \Delta V \cdot \hat{n} = 189 \text{ km/s}$
- $V_1 \cdot \hat{n} = -102, V_1 = 86 \text{ km/s}$
- $C_A = 68, C_S = 106 \text{ km/s}$
- $C_I = 125 \text{ km/s}$
- $M \cong \frac{189}{125} = 1.5 > 1.0$
- A SHOCK

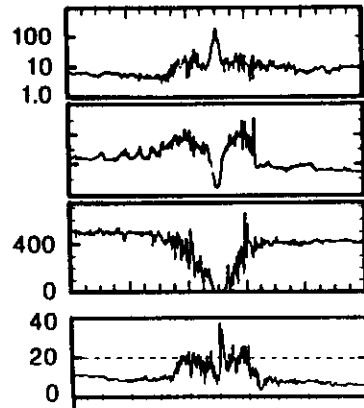


$N_e, \text{ cm}^{-3}$

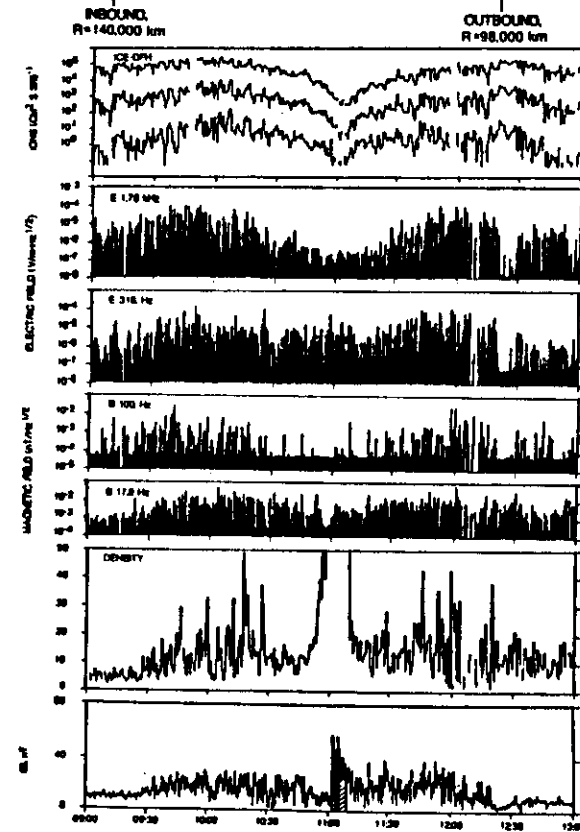
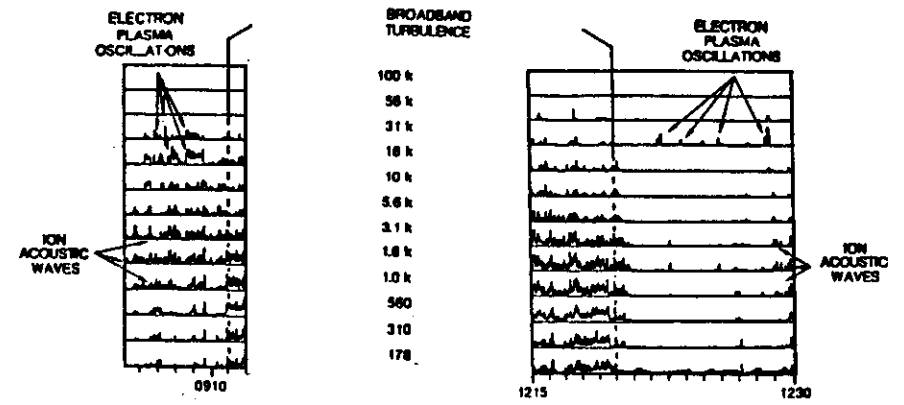
$T_e, 10^5 \text{ }^\circ\text{K}$

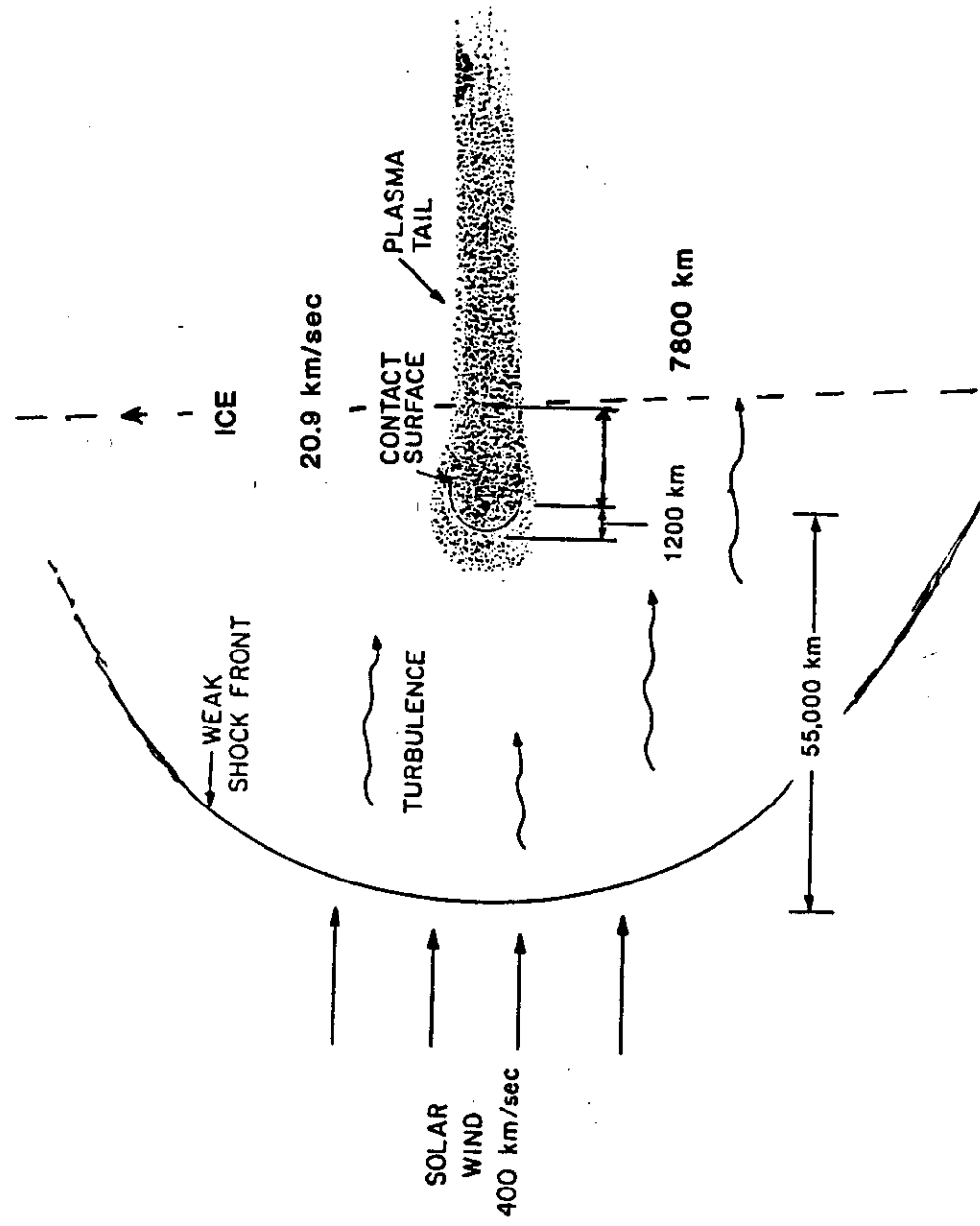
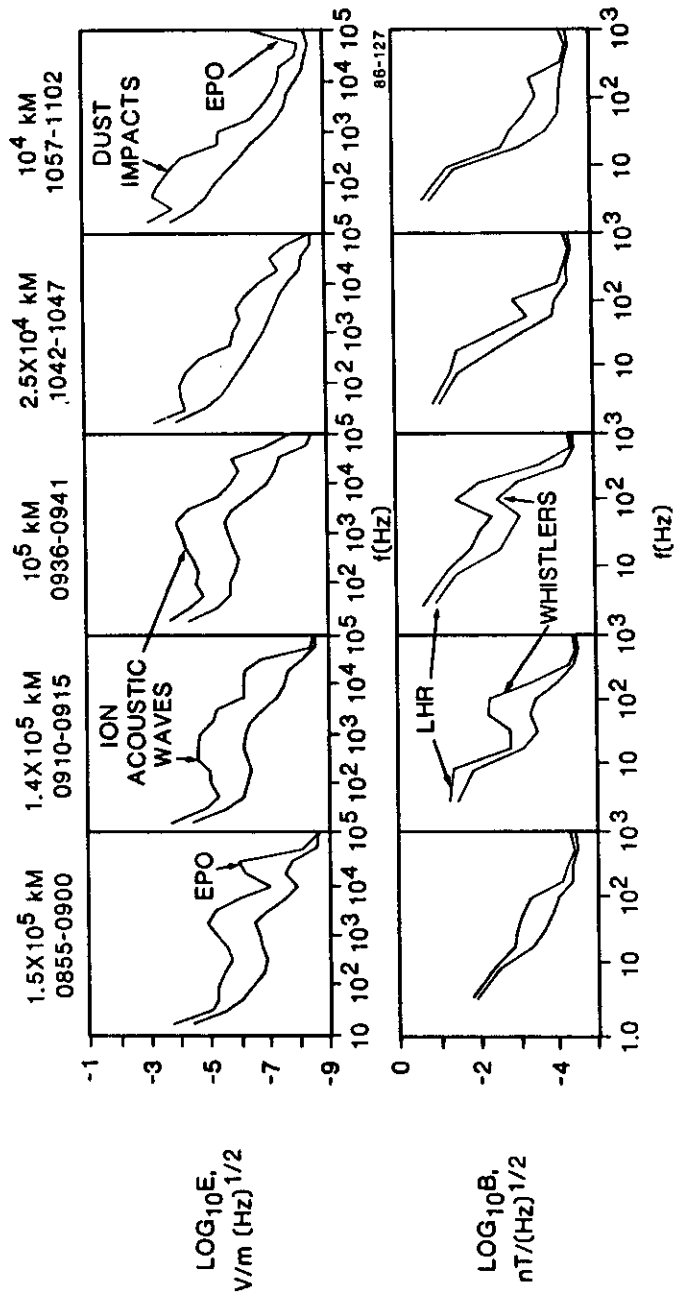
$V, \text{ km/sec}$

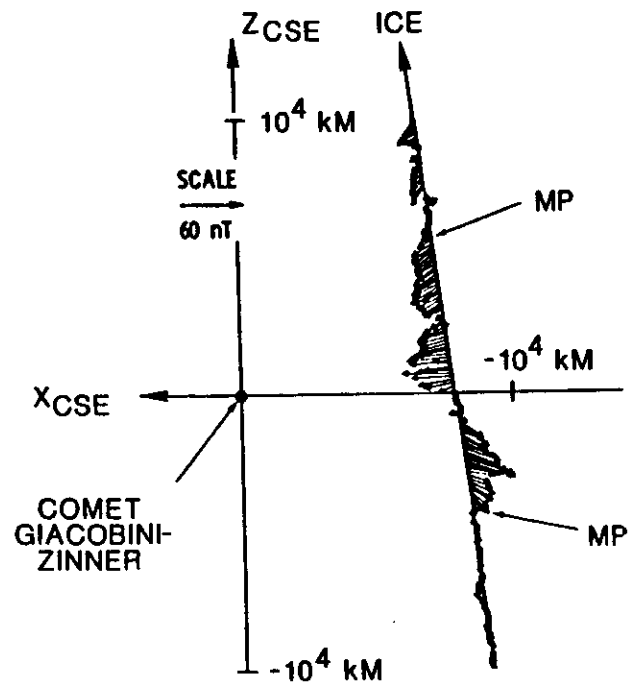
$B, \text{ gamma}$



0600 1600 SEPT 11, 1985







ELECTRON
PRESSURE
(eV cm^{-3})

$\frac{B^2}{3\pi}$
(10^{-8} dynes/cm 2)

