**2246-15**

# Workshop on Cosmic Rays and Cosmic Neutrinos: Looking at the Neutrino Sky

*20 - 24 June 2011*

## Probing local UHE sources using CR nuclei and gamma-rays

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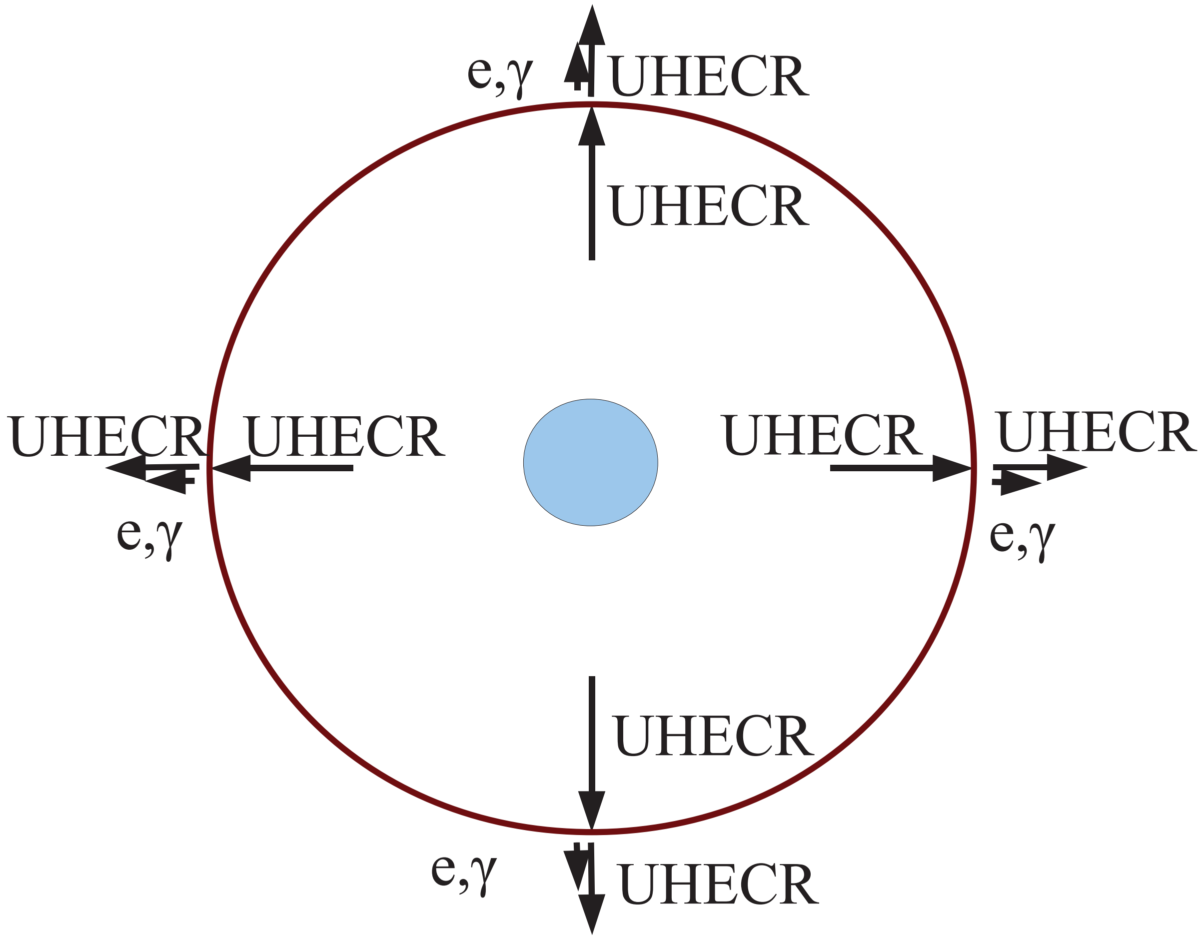
# Probing Local UHECR Sources with Nuclei + Gamma-Rays



# Probing Local UHECR Sources with Nuclei + Gamma-Rays

# Talk Structure

**Using Nuclei**

1) Measurements of UHECR by the PAO

2) Implications for UHECR source requirements

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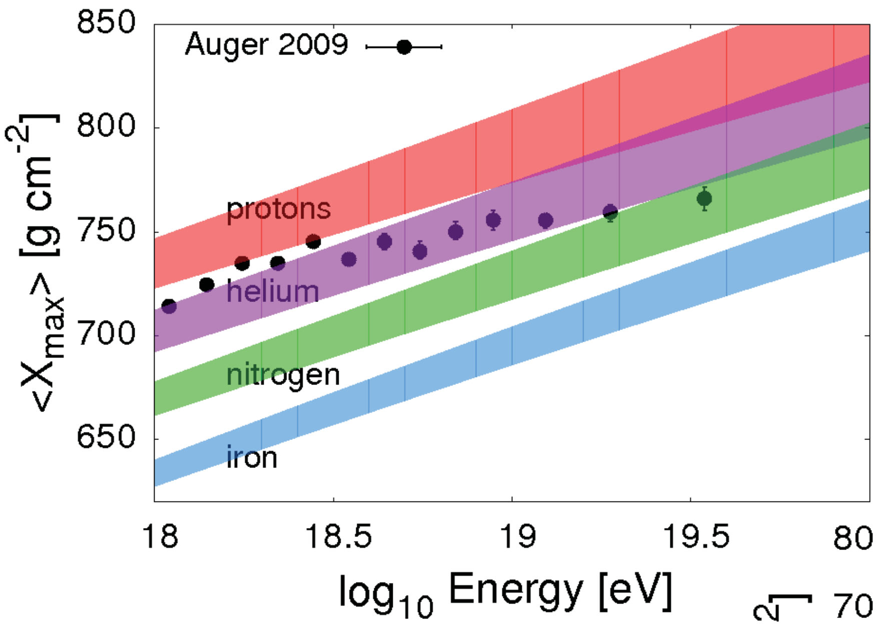
**Using Gamma-Rays**

3) Their birth + death in regions surrounding the source

4) What current PAO Gamma-Ray limits already tell us about the sources

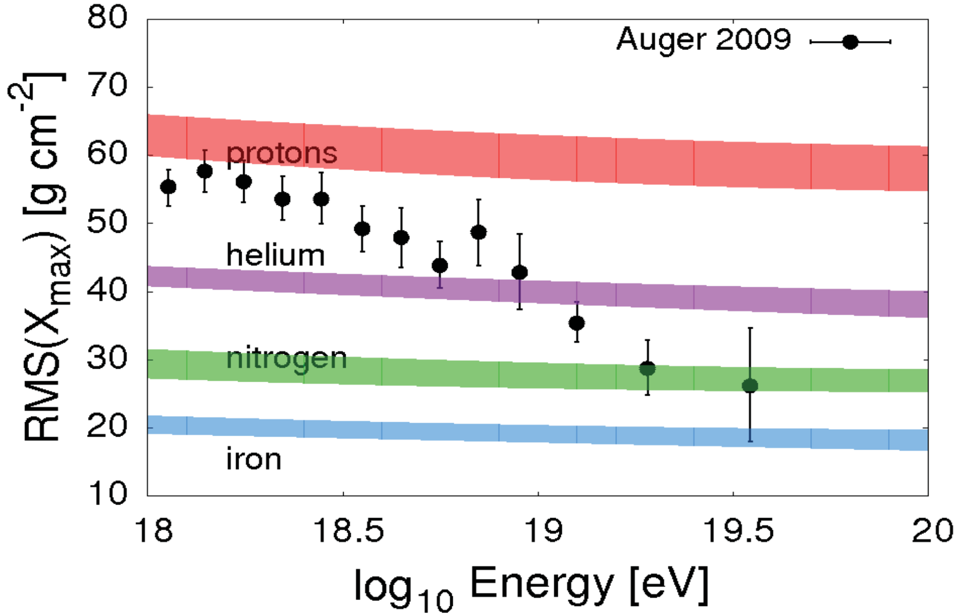
# Part 1: Nuclei

# MEASUREMENTS of UHECR by the PAO



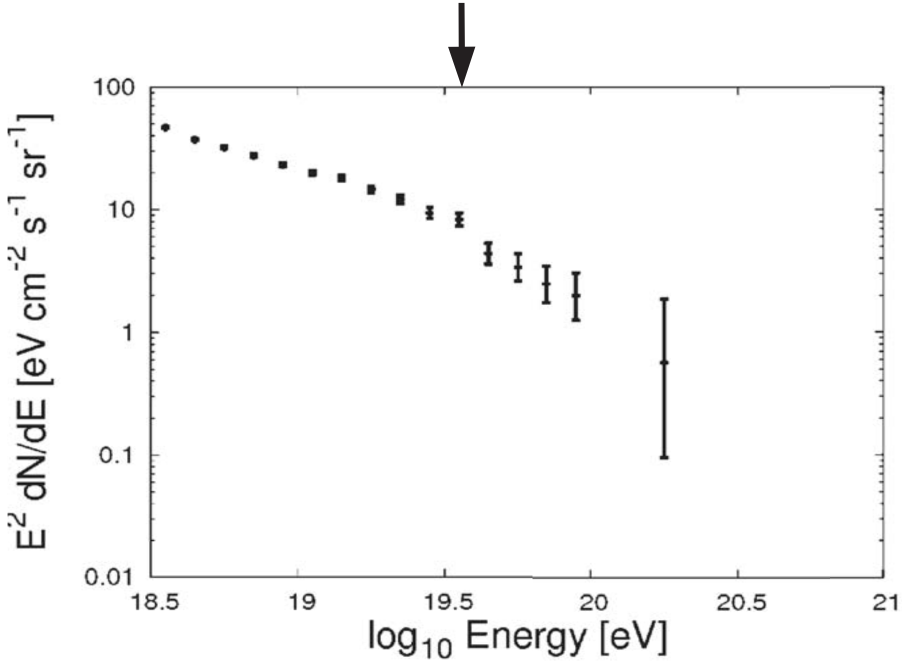

$\leftarrow \langle X_{\max} \rangle$

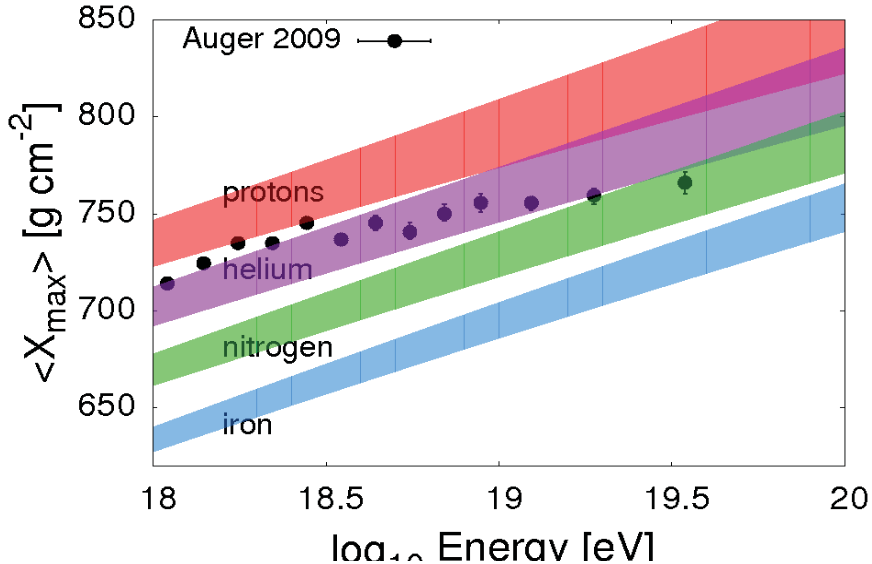
$\mathrm{RMS}(X_{\max})$

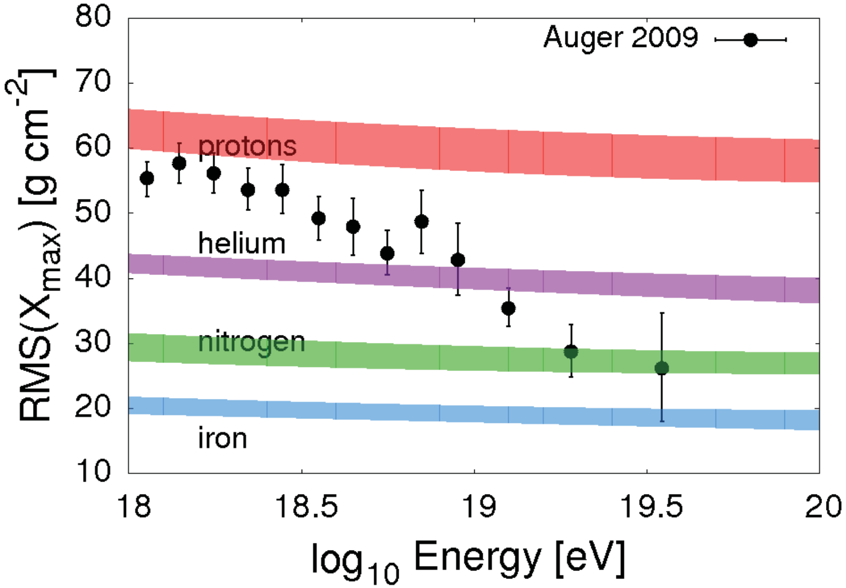


# MEASUREMENTS of UHECR by the PAO

Flux

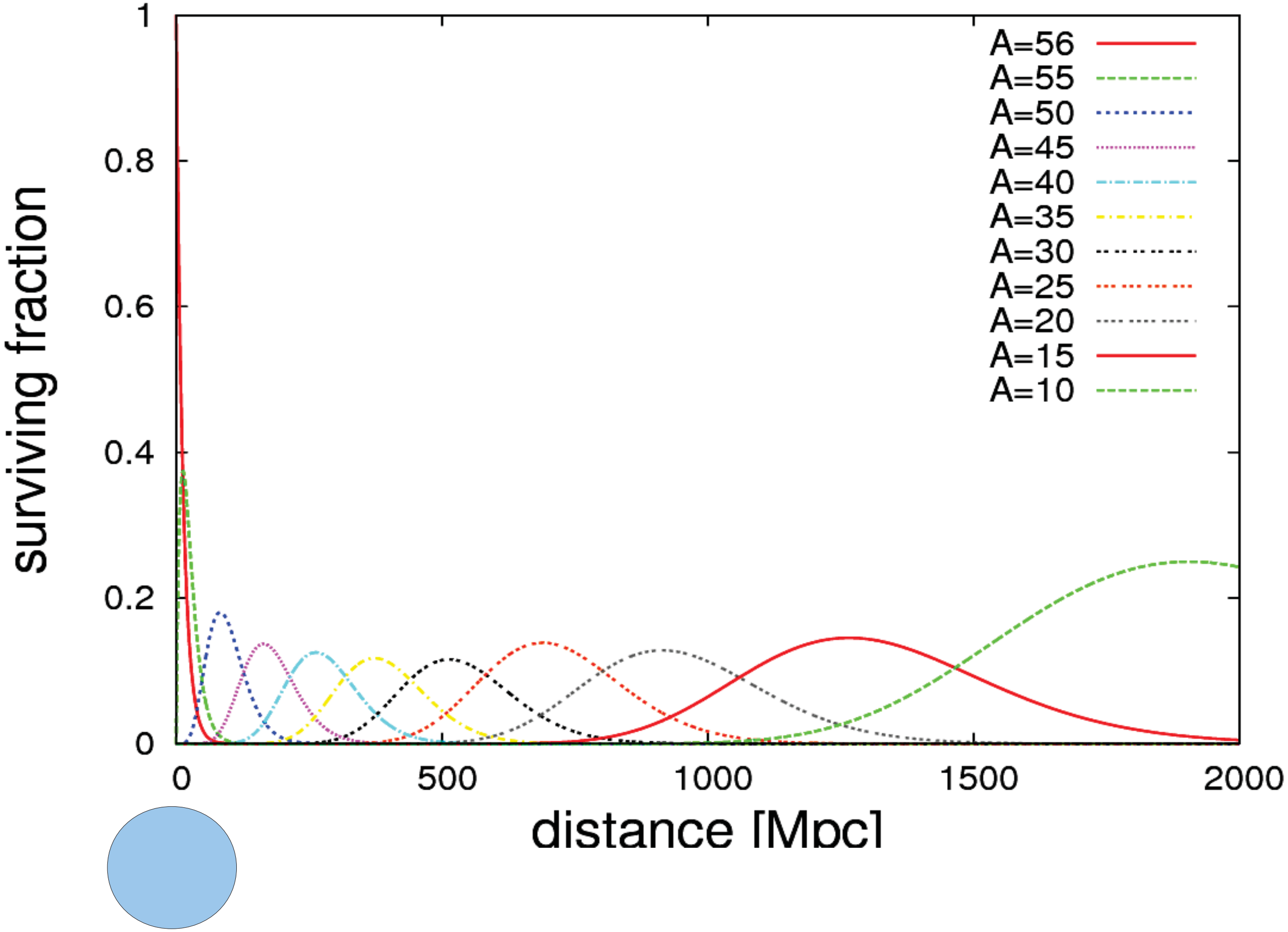


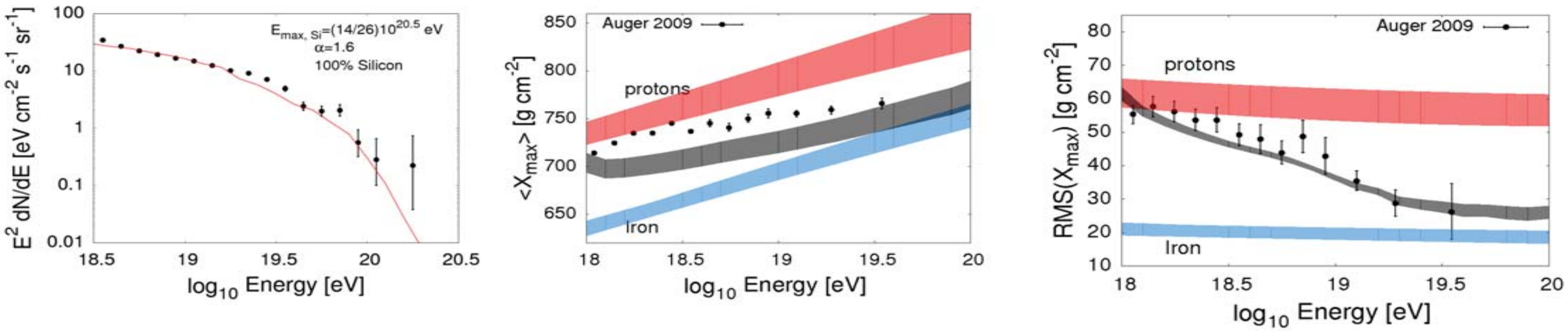
# Nuclei Propagation Away from their Source + their Transmutation

Lorentz factor of nuclei ~conserved

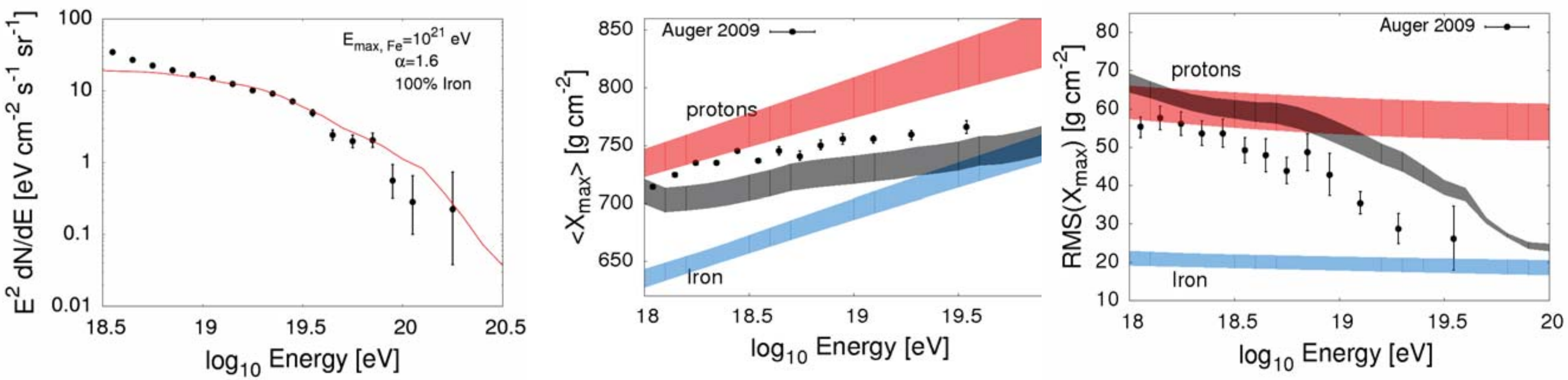


# Nuclei Propagation Away from their Source + their Transmutation

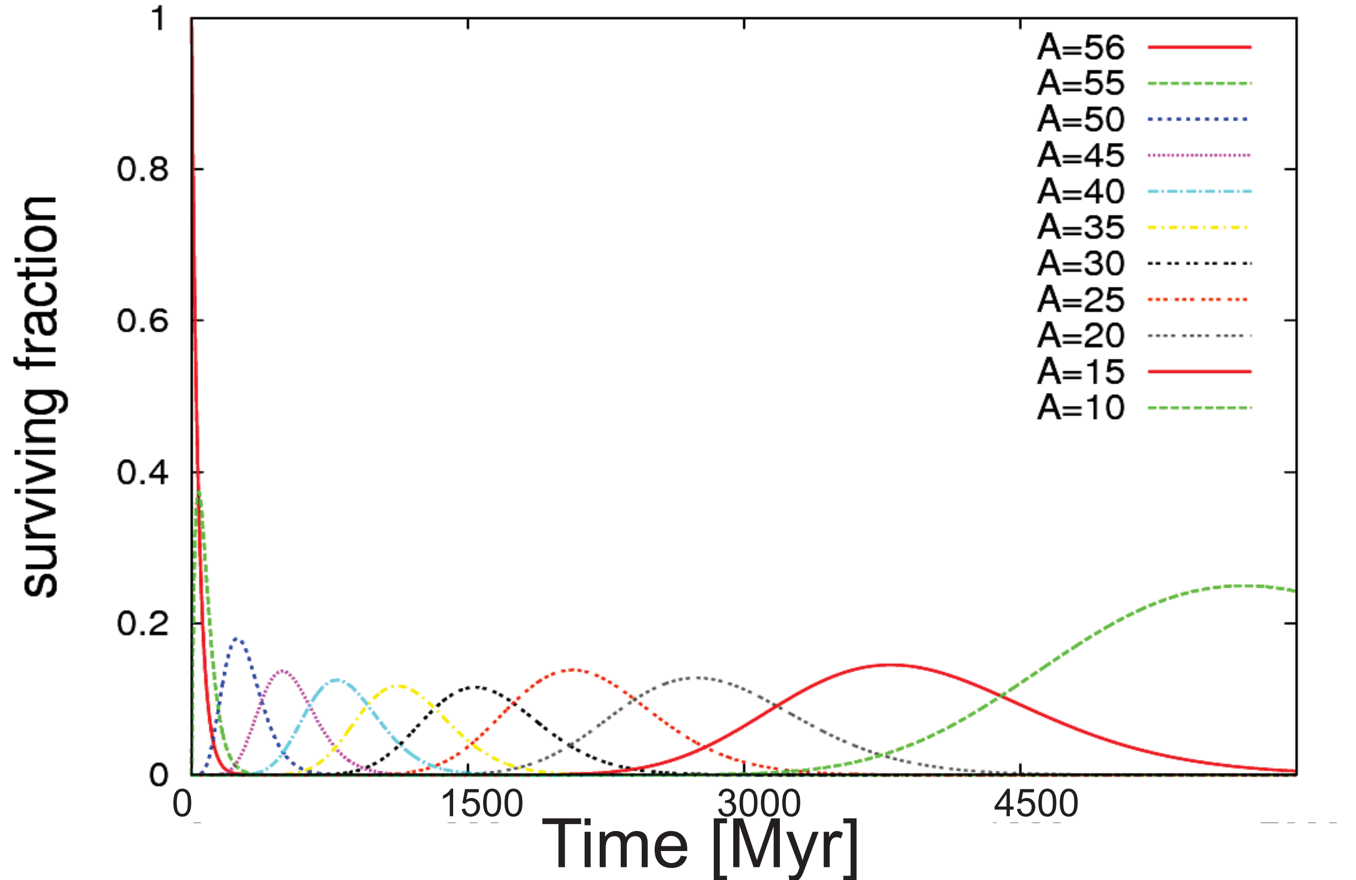
## **Silicon** only?




## **Iron** only?

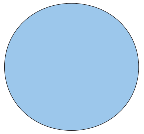


# Nuclei Propagation Within their Source + their Transmutation



# IMPLICATIONS for UHECR Sources

$$f = \frac{t_{\mathrm{trap}}}{t_{\mathrm{int.}}^{\mathrm{CR}\gamma}}$$



$$t_{\mathrm{int.}}^{\mathrm{CR}\gamma} \approx \frac{1}{n_\gamma \sigma_{\mathrm{CR}\gamma} c}$$

$$n_\gamma = \frac{L_\gamma}{c 4\pi R^2 \epsilon_\gamma}$$

$$t_{\mathrm{trap}} \approx \frac{R^2}{2D} = \frac{3R^2}{2R_{\mathrm{Larmor}}}$$

$$f^{\mathrm{CR}\gamma} = \frac{3 L_\gamma \sigma_{\mathrm{CR}\gamma} Z B}{8\pi \epsilon_\gamma E_{\mathrm{CR}}}$$

# IMPLICATIONS for UHECR Sources

$$f^{\mathrm{CR}\gamma} = \frac{3L_\gamma \sigma_{\mathrm{CR}\gamma} Z B}{8\pi \epsilon_\gamma E_{\mathrm{CR}}} = \frac{s_1}{s_2}$$

Photo-disintegration threshold:

$2E_{\mathrm{CR}}\epsilon_\gamma > A m_p c^2 E_{\mathrm{bind.}}$ , where $m_p c^2 E_{\mathrm{bind.}} = 10^{16}\ \mathrm{eV}^2$

Since,

$$L_\gamma[10^{40}\mathrm{erg\ s}^{-1}] = 2\times 10^{41}\mathrm{eV\ cm}^{-1}$$
$$\sigma_{\mathrm{CR}\gamma}[\mathrm{A\ mb}] = \mathrm{A}\times 10^{-27}\ \mathrm{cm}^2$$
$$B[1\ \mathrm{G}] = 300\ \mathrm{eV\ cm}^{-1}$$

$\frac{L_\gamma \sigma_{\mathrm{CR}\gamma} B}{A} = 6\times 10^{16}\ \mathrm{eV}^2$ , ergo.... $f^{\mathrm{CR}\gamma} = 50\frac{Z}{26}$

# IMPLICATIONS for UHECR Sources

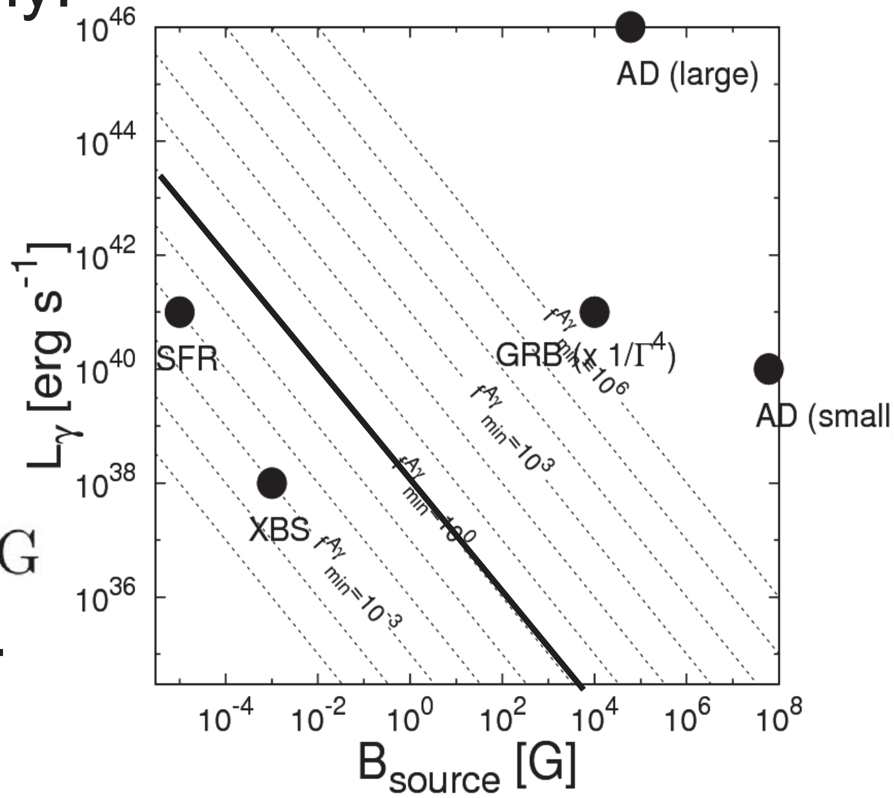
Since, $$\frac{L_\gamma^{\text{Edd.}} \sigma_{\text{CR}\gamma} B^{\text{Edd.}}}{A} = 4 \times 10^{23} \left(\frac{M}{M_\odot}\right)^{1/2} \text{ eV}^2$$

Only heavily sub-Eddington power objects need apply!

If magnetic + photon luminosity are in equipartition:

$$L_\gamma \approx \beta R^2 B^2$$

Requiring, $B < 4 \times 10^{-5}$ G to ensure safe passage.



# IMPLICATIONS for UHECR Sources

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If magnetic + photon luminosity are in equipartition:

$$L_\gamma \approx \beta R^2 B^2$$

WARNING:

$$R_{\mathrm{Larmor}}(10^{20}\ \mathrm{eV\ Fe}) \approx 0.1\ \mathrm{kpc}$$

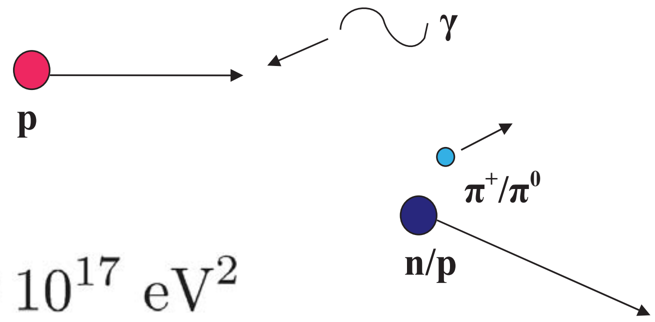
$$t_{\mathrm{acc}} \approx R_{\mathrm{Larmor}}/c\beta^2$$

Requiring, $$B < 4 \times 10^{-5}\ \mathrm{G}$$

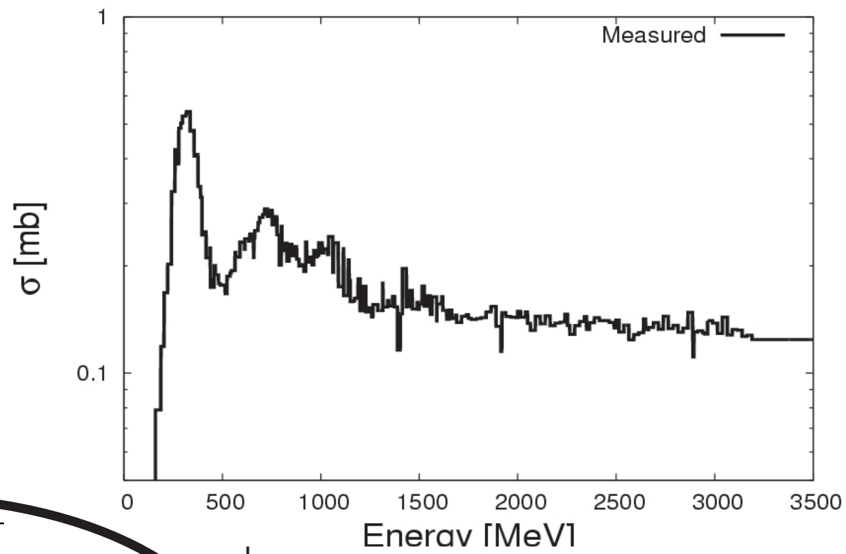
# Part 2: Gamma-Rays

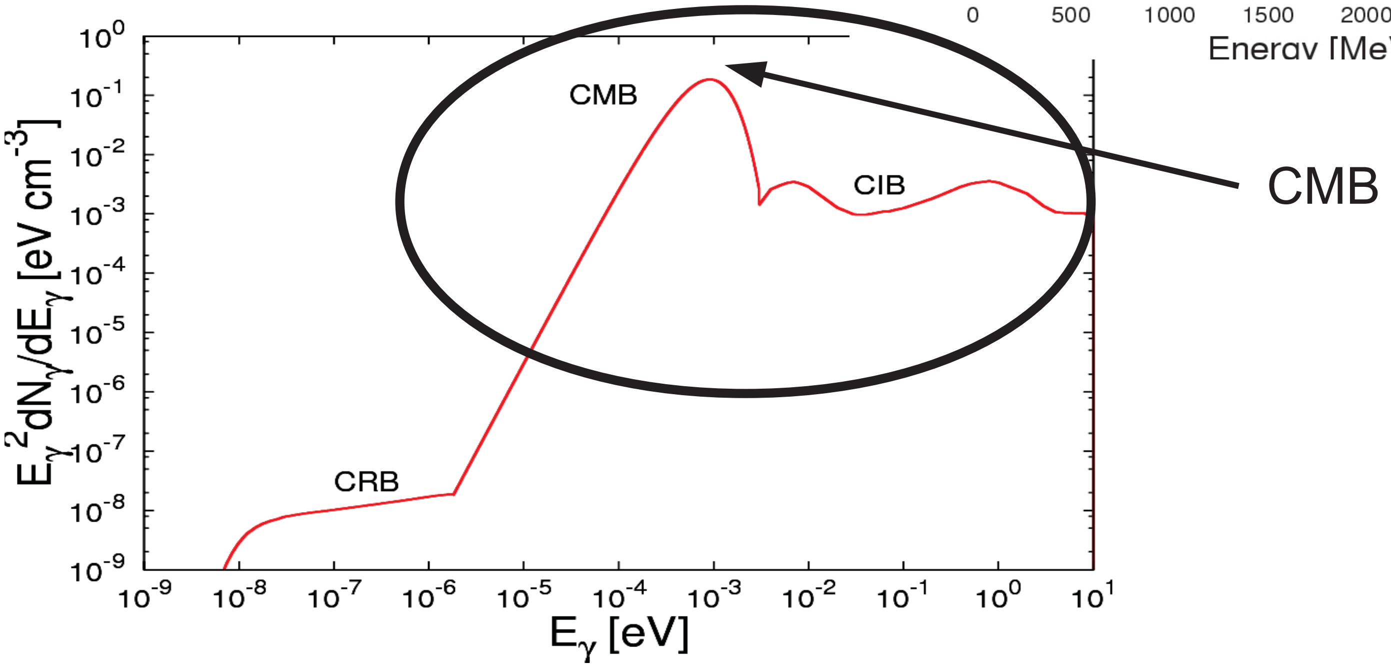
# Gamma-Ray Production

protons




$$s \approx 10^{17}\ \mathrm{eV}^2$$



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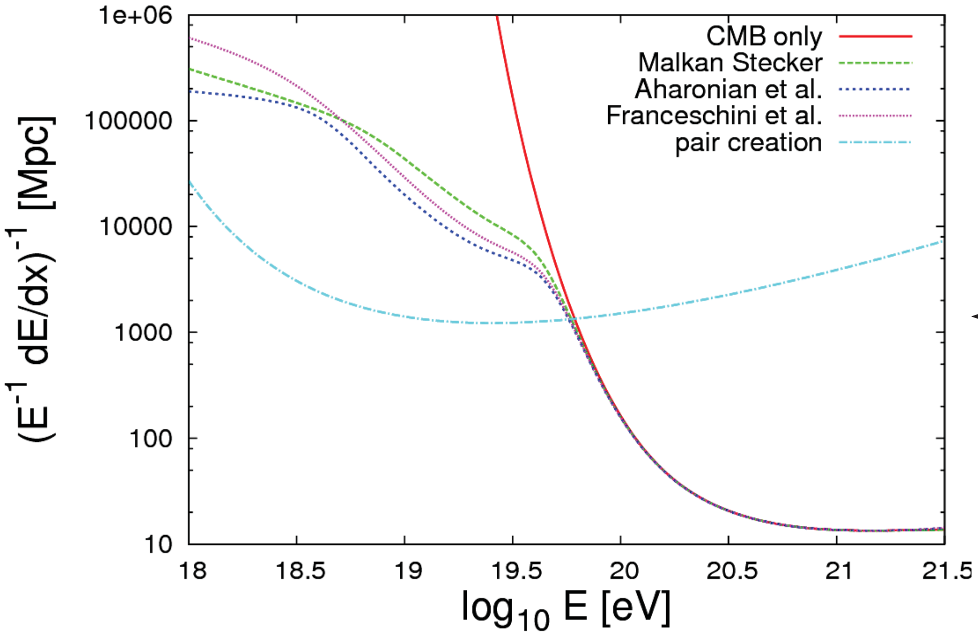


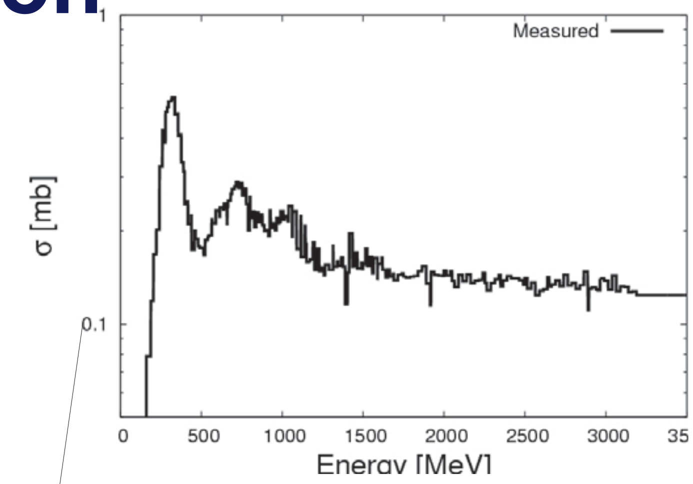


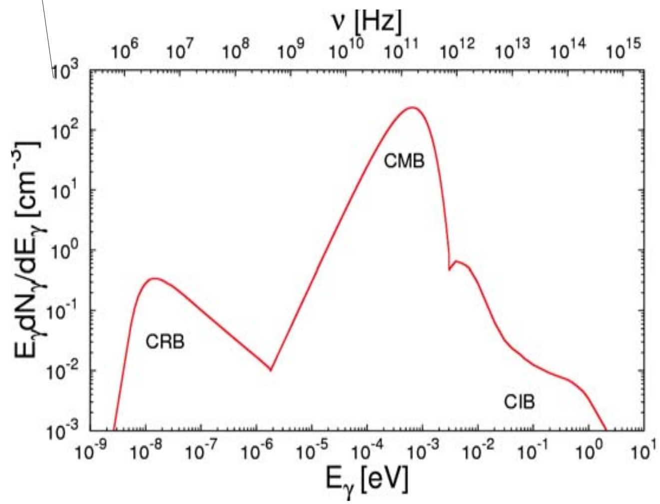

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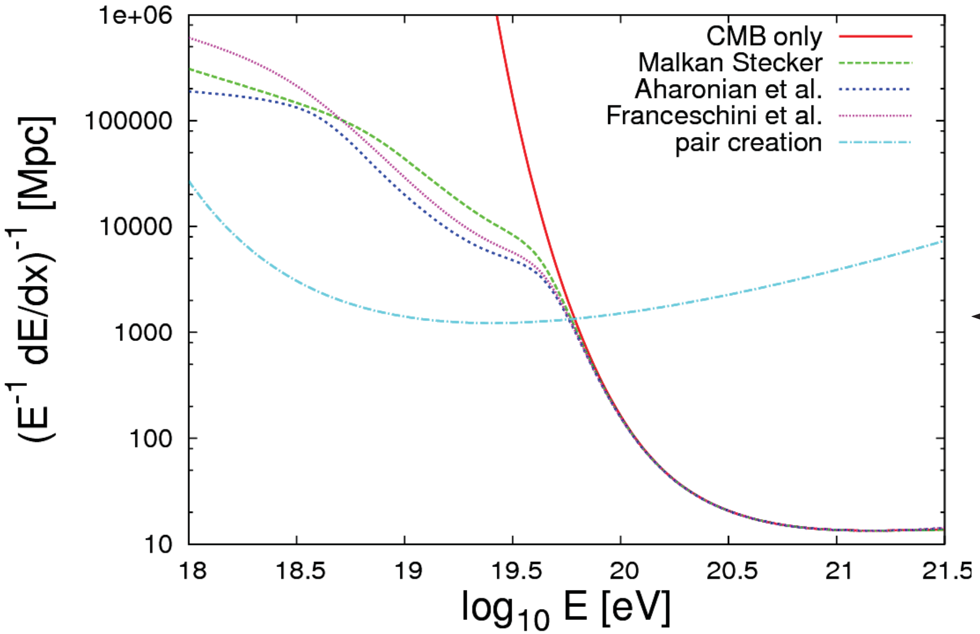
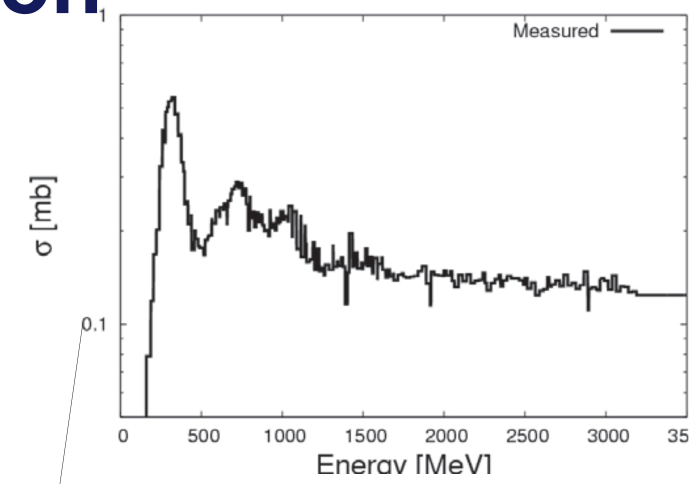
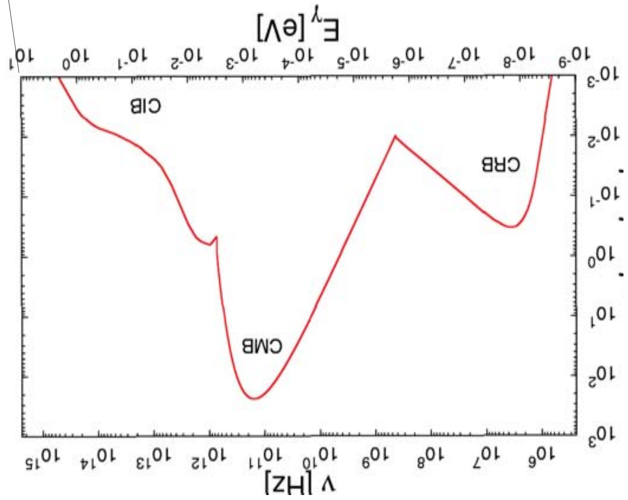
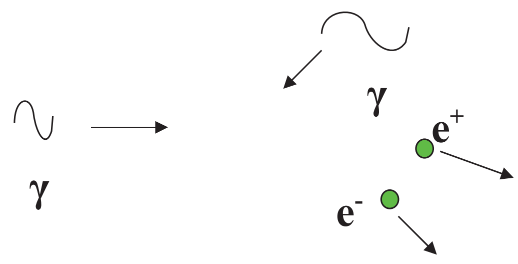
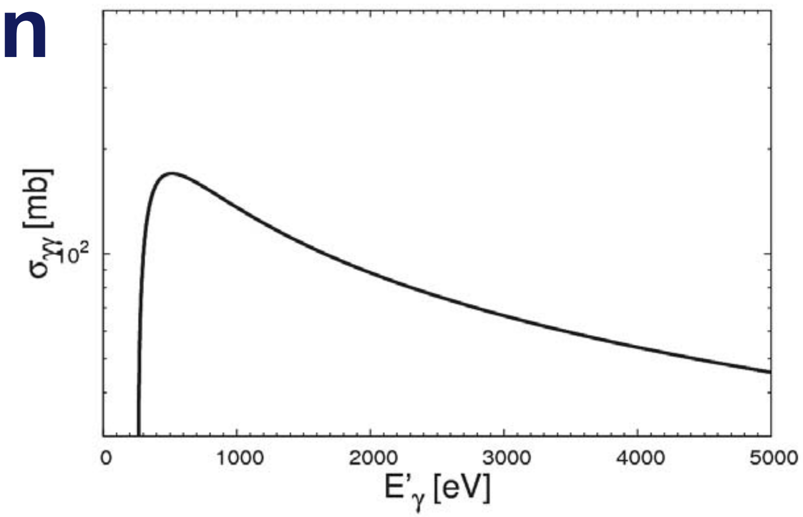


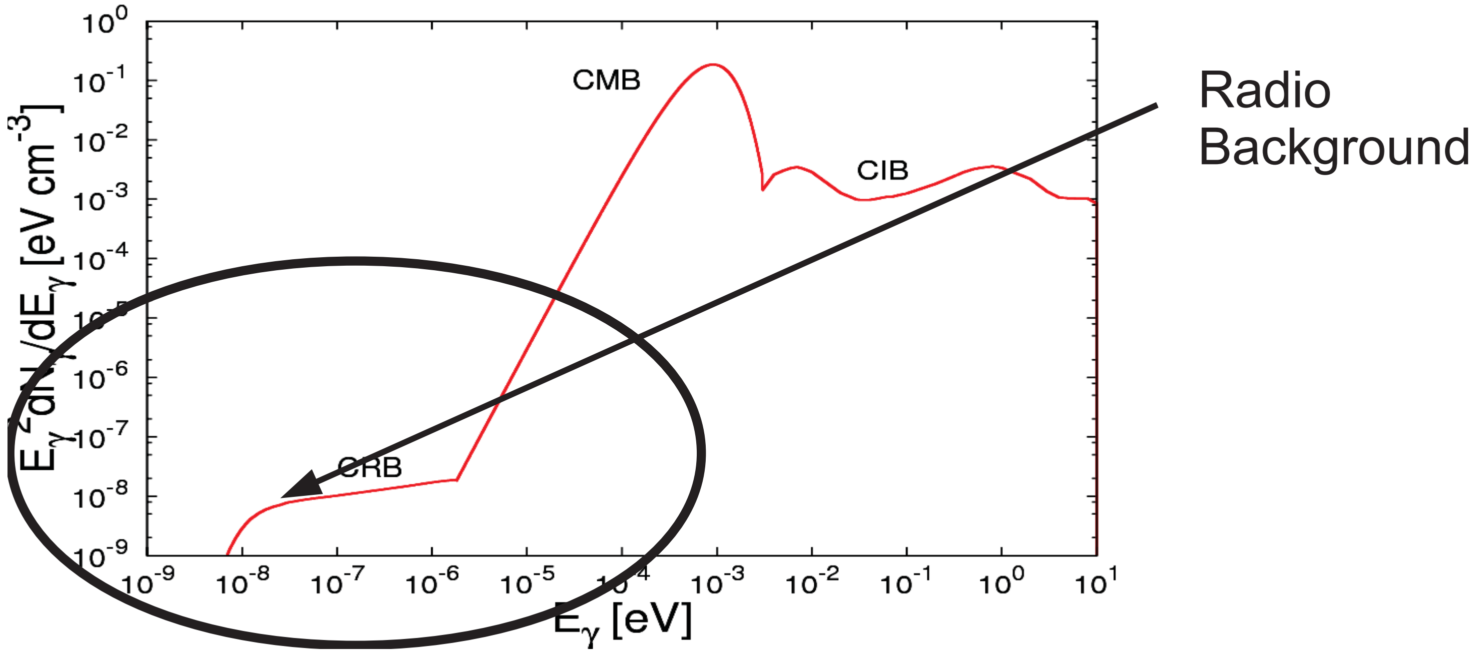



# Gamma-Ray Interaction




$$s \approx 10^{12}\ \mathrm{eV}^2$$



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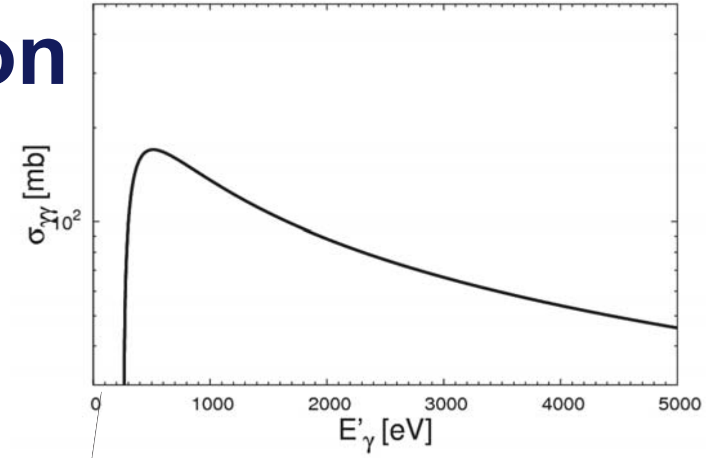
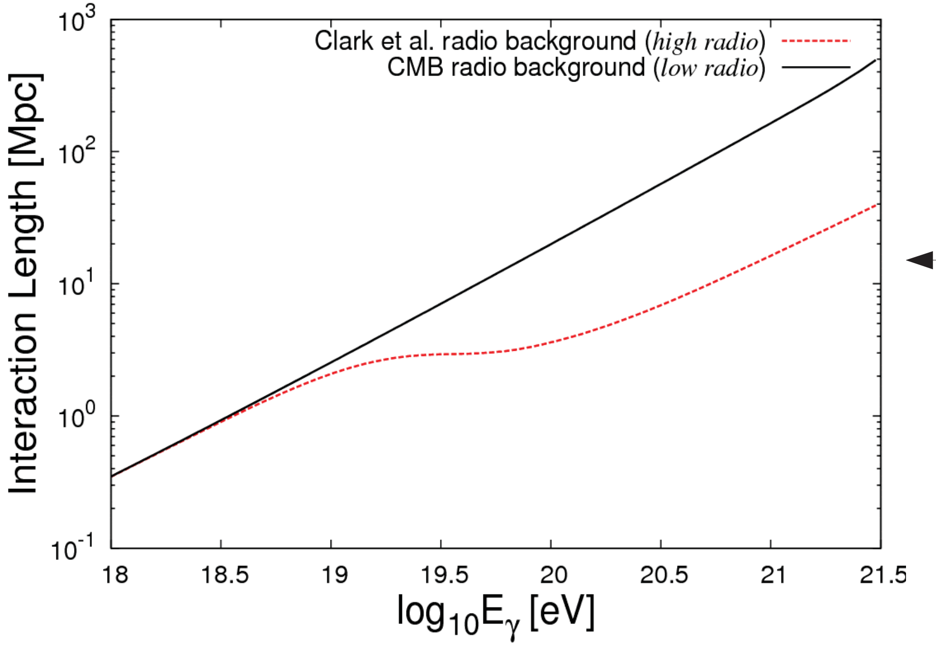
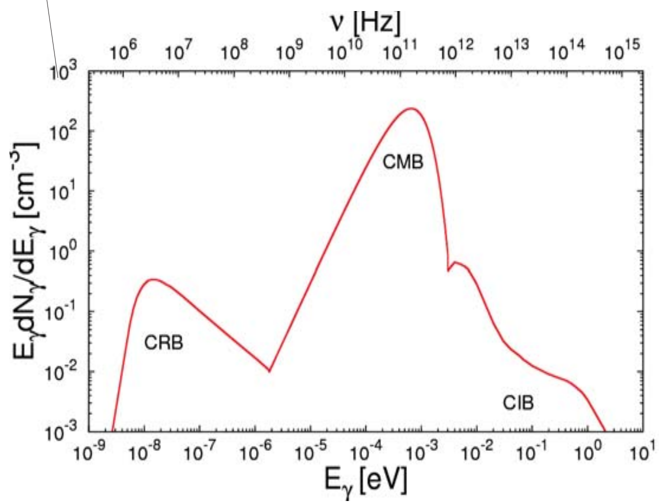
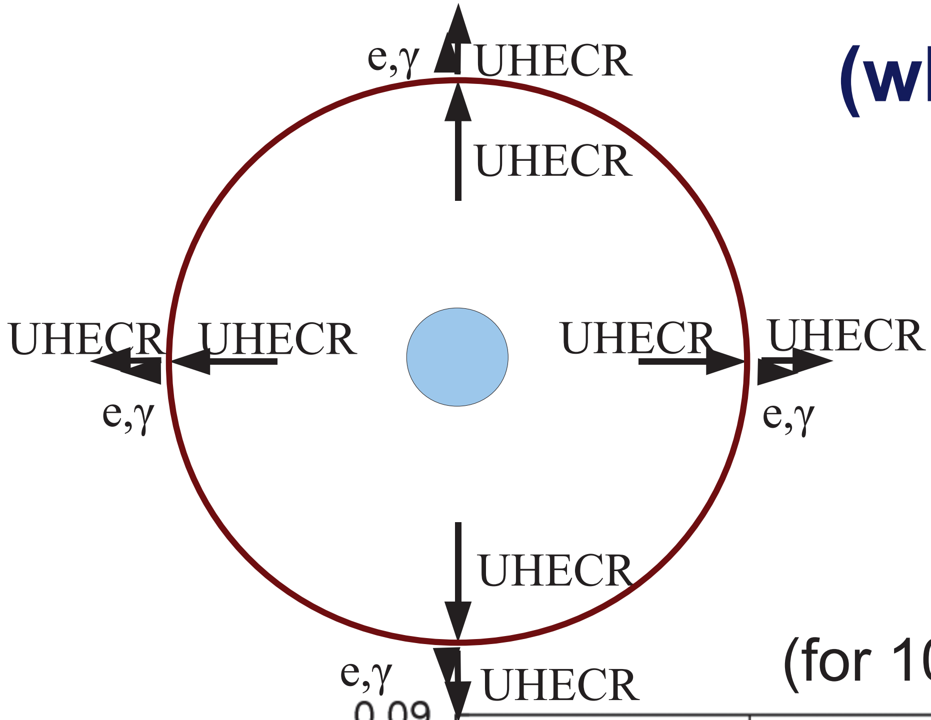


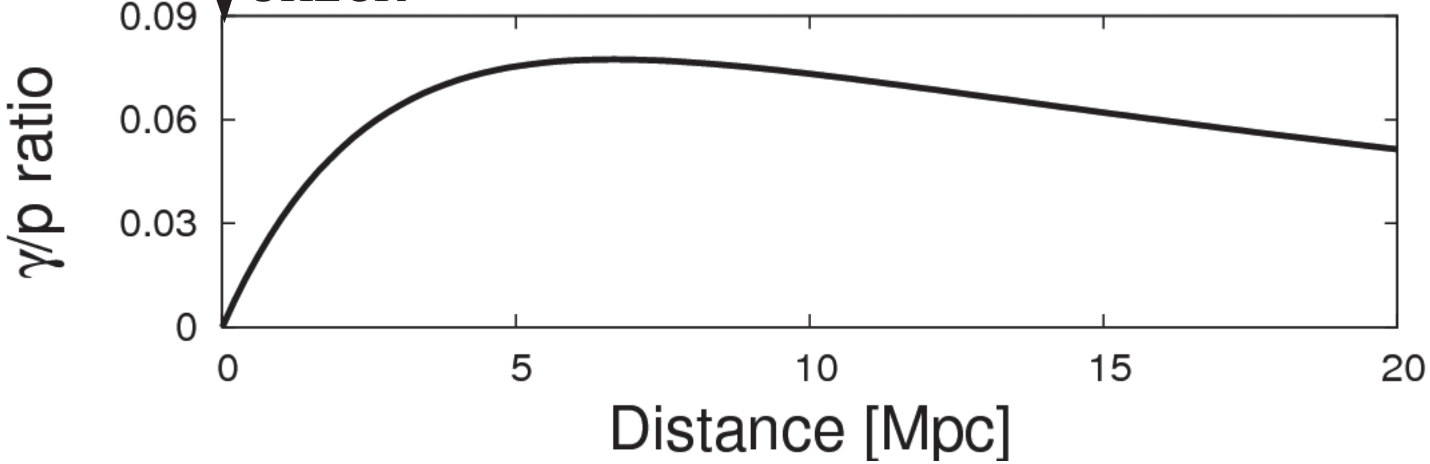



# The Halo Around Heavenly Bodies
## (which acc. UHECR)




(for $10^{19}$eV gamma-rays)



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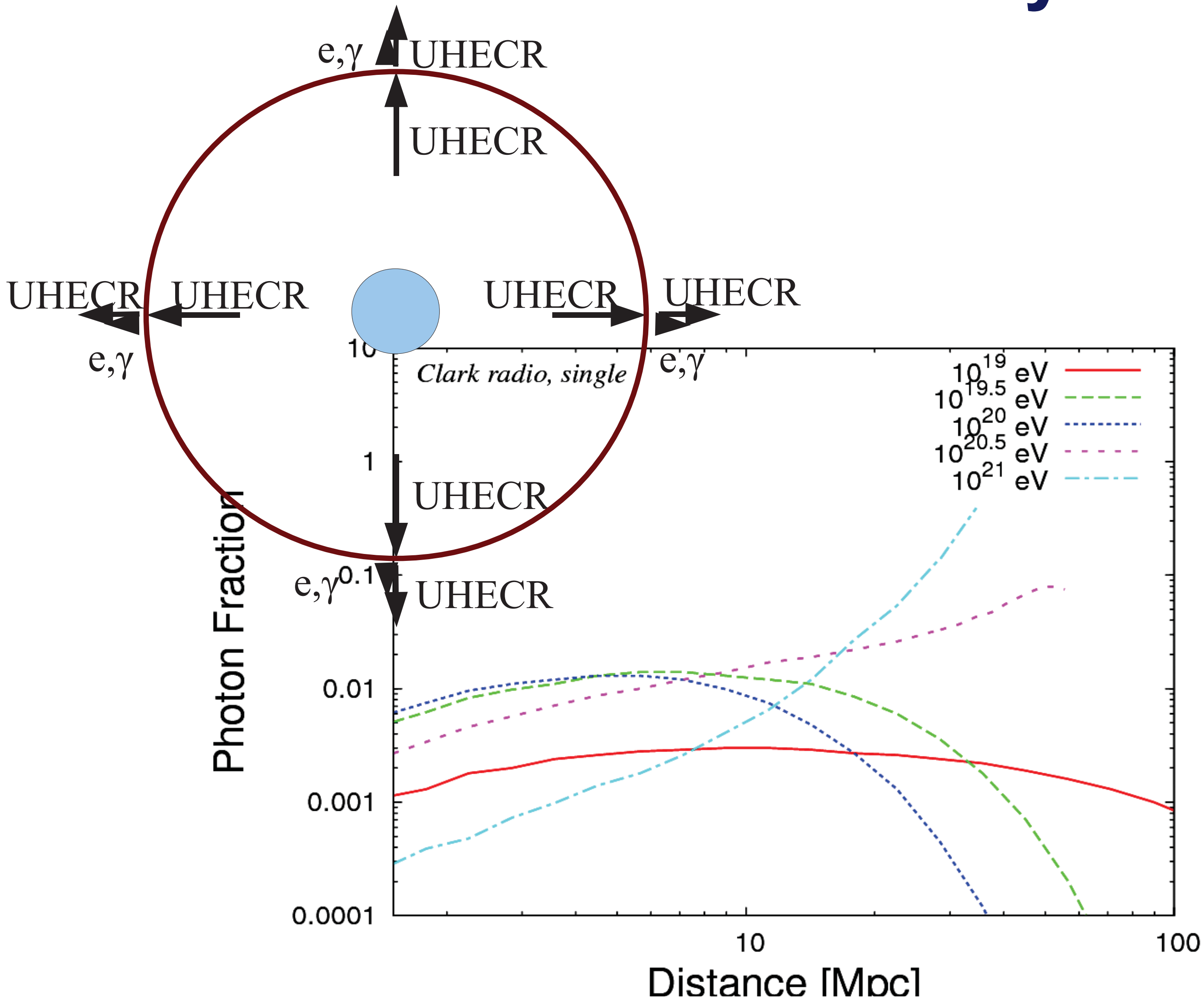
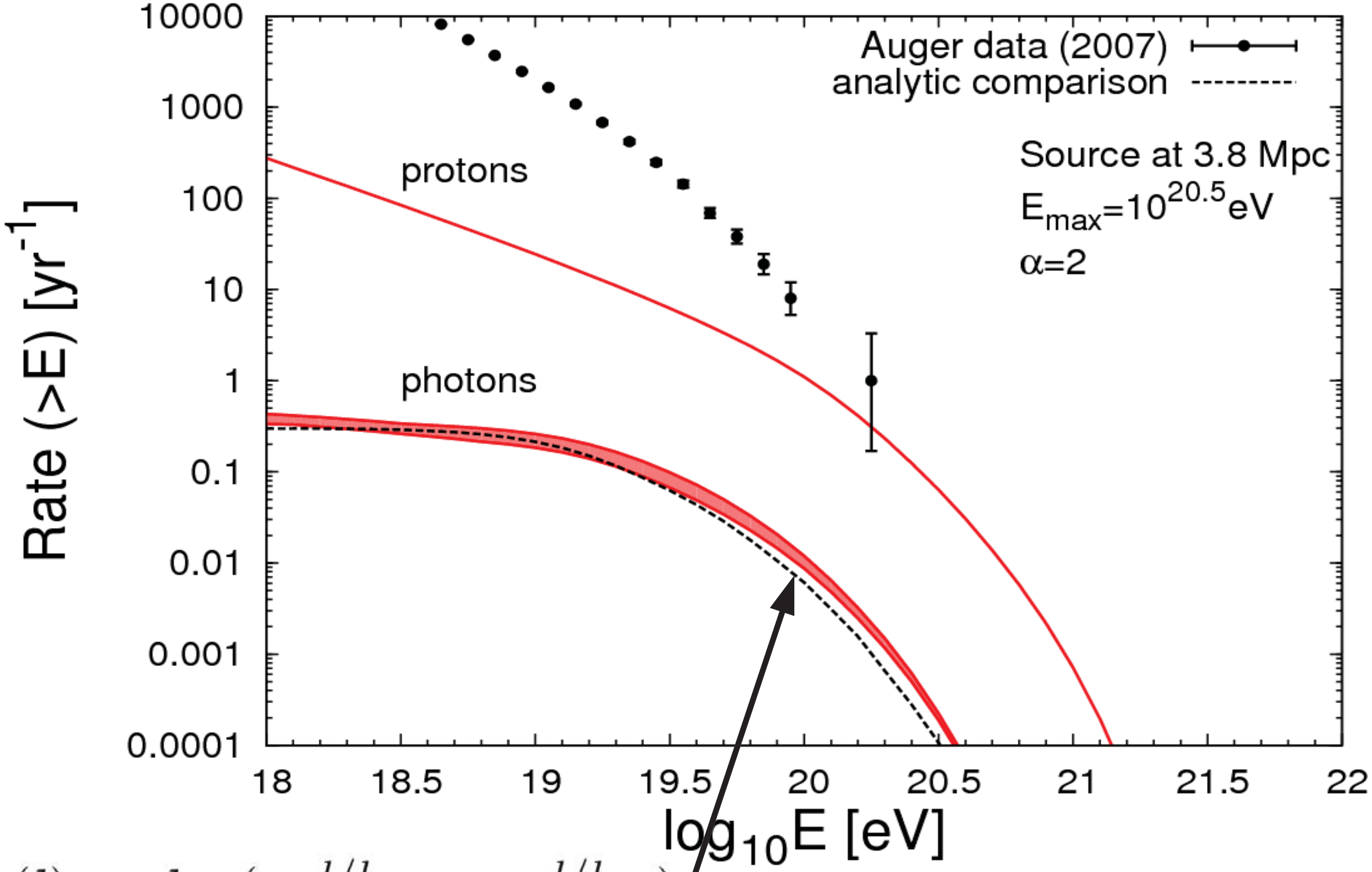


# The Halo Around Cen A




$$\frac{N_\gamma(l)}{N_p(0)} = \frac{l_{\gamma\gamma}(e^{-l/l_{p\gamma}} - e^{-l/l_{\gamma\gamma}})}{(l_{p\gamma} - l_{\gamma\gamma})}$$

# Conclusion

The dominance of nuclei at the highest energies provides useful new information about the nature of UHECR sources

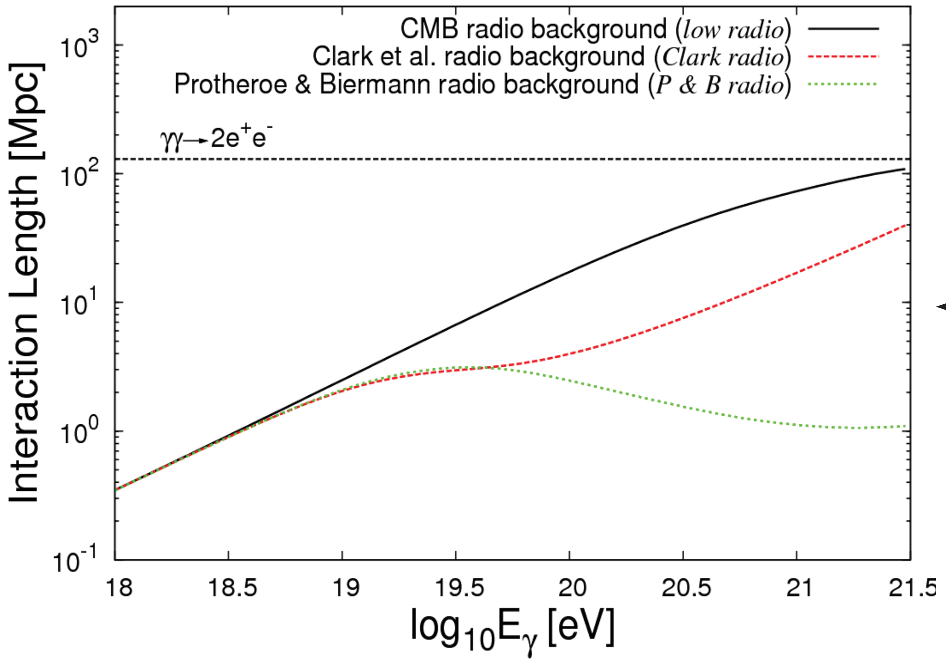
Regions close to luminous objects are excluded as UHECR sources, favouring slow acceleration scenarios

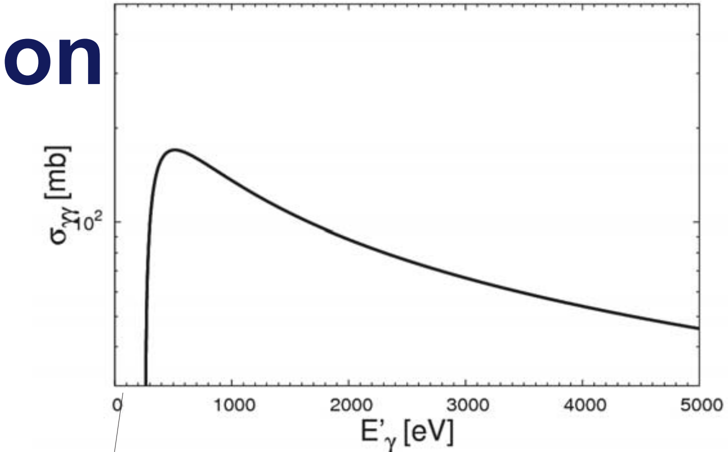
UHE photons can provide a useful probe of local sources

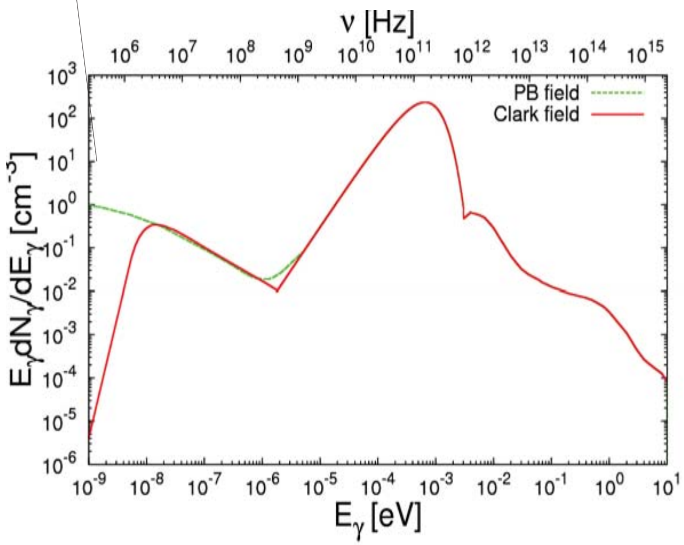
Applied to Cen A we expect an UHE photon in 5 years, if 2 UHECR in the PAO 57 UHECR set originated from Cen A.

# Extra Slides

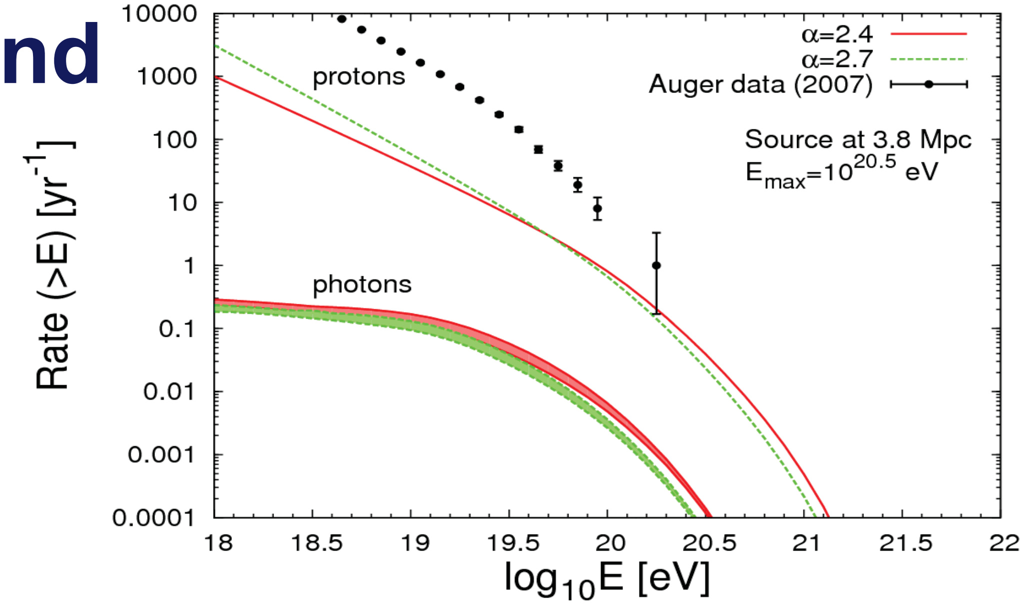
# Gamma-Ray Interaction

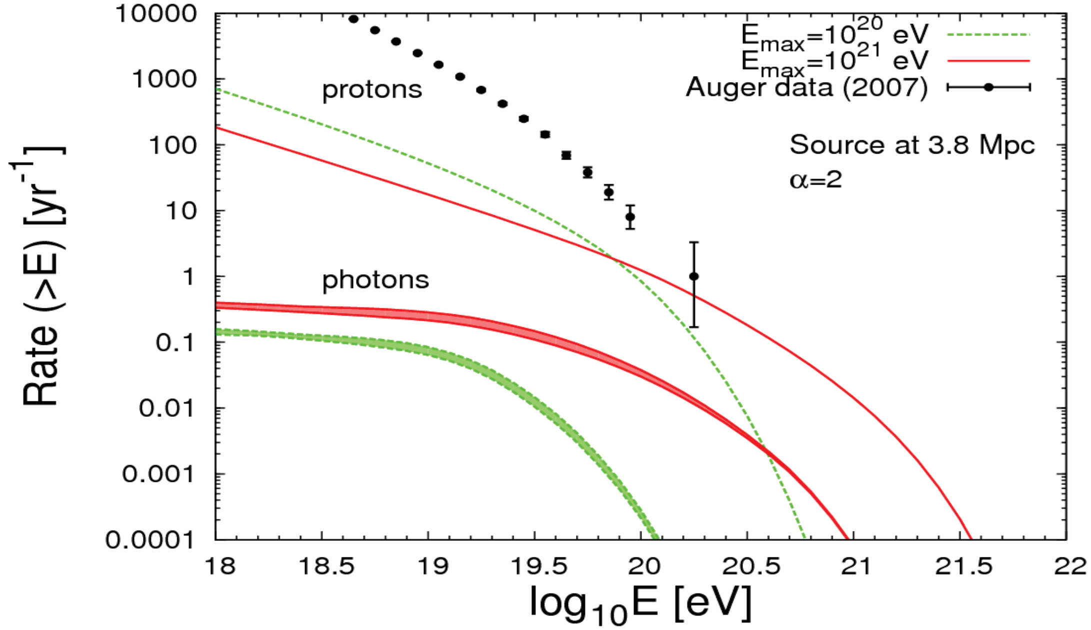


# The Halo Around Cen A



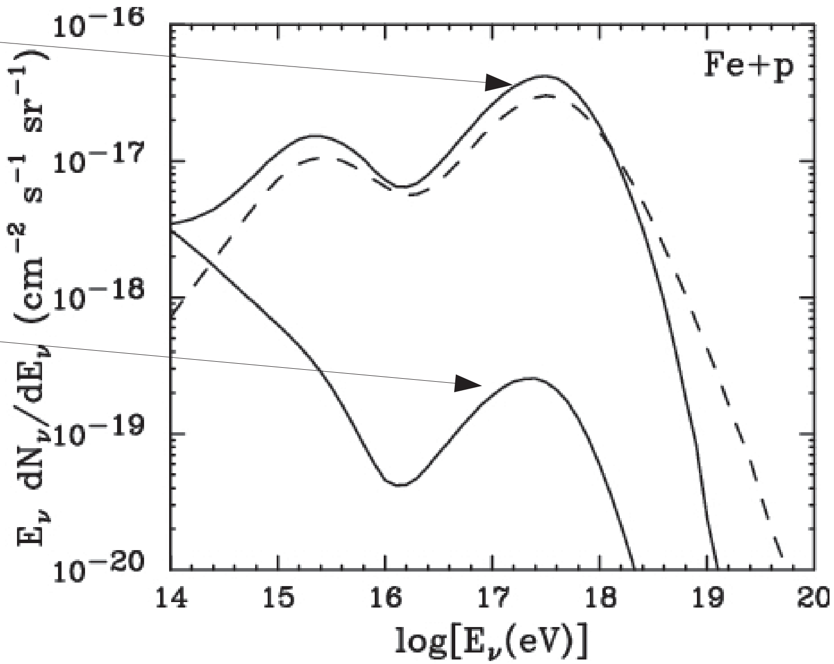



# The Cosmogenic Neutrino Flux

The high energy (>$10^{17}$ eV) flux quoted as the "Guaranteed flux" value

lowest value compatible with all the data

Smaller value obtained since best agreement found for a dominant Fe fraction with **E**$_{max}$=$10^{21}$eV



# Ratio Between Photo-Pion and Photo-Disintegration Rates (2)

with,

$$\sigma_{p,\gamma} = \text{٠.٥ mb}, E_{p,\gamma} = \text{٣١٠ MeV}, \Delta_{p,\gamma} = 100 \text{ MeV}$$

and

$$\sigma_{A_{56},\gamma} = 81 \text{ mb}, E_{A_{56},\gamma} = 18 \text{ MeV}, \Delta_{A_{56},\gamma} = 8 \text{ MeV}$$

therefore

$$R_{A_{56},\gamma}(\Gamma) \approx \frac{\sigma_{A_{56},\gamma}}{\sigma_{p,\gamma}} R_{p,\gamma}(15\Gamma) = 160\, R_{p,\gamma}(15\Gamma)$$