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Workshop on Cosmic Rays and Cosmic Neutrinos: Looking at the Neutrino Sky

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Physics of cosmic ray acceleration and its implications for the origin of cosmic rays

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PHYSICS OF PARTICLE ACCELERATION AND THE ORIGIN OF COSMIC RAYS

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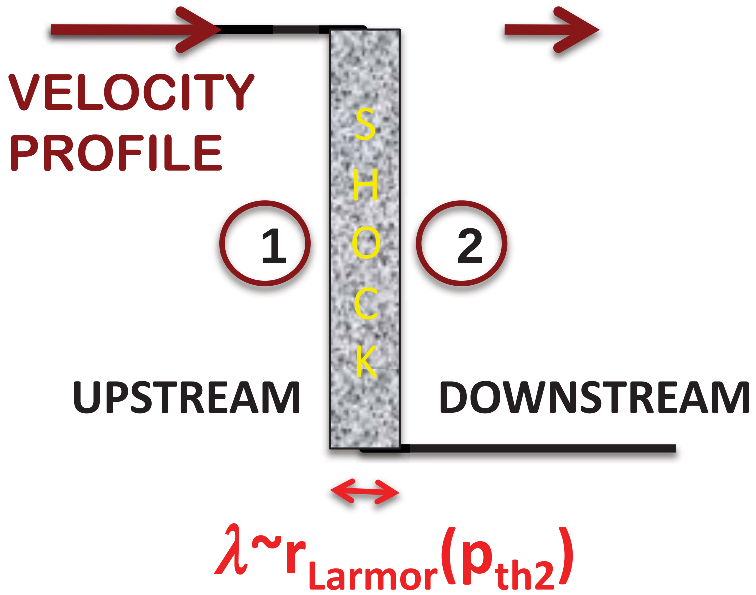
NUSKY – Trieste June 2011

DIFFUSIVE ACCELERATION AT COLLISIONLESS NEWTONIAN SHOCKS

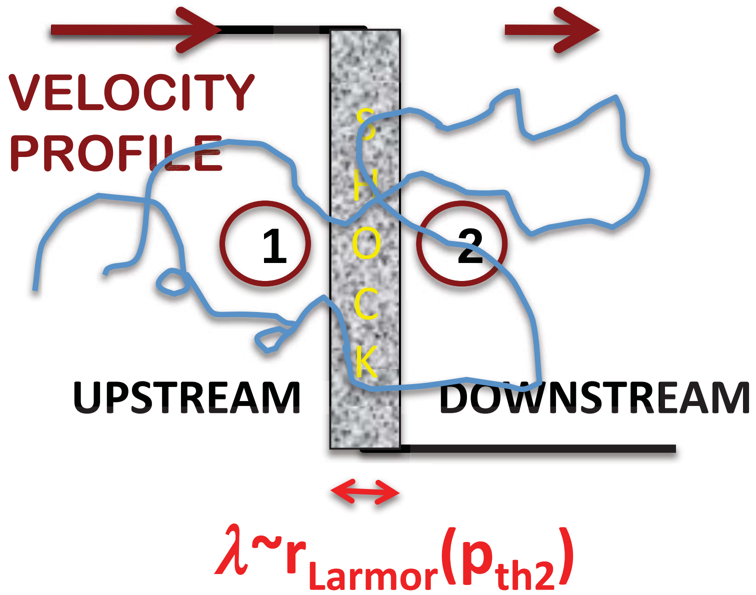


COLLISIONLESS → MEDIATED BY ELECTROMAGNETIC INSTABILITIES

IN GENERAL ONE EXPECTS:

- Different heating for e and p
- Finite thickness of the shock
- Instabilities responsible for the shock formation also responsible for first particles returns (injection)

DIFFUSIVE ACCELERATION AT COLLISIONLESS NEWTONIAN SHOCKS
'test particles'



In test particle theory, all approaches lead to:

- -POWER LAW SPECTRA
- -SLOPE ONLY FUNCTION OF COMPRESSION
- -INDEPENDENT OF D(E)
- -NO CLEAR RECIPE FOR E_{MAX}
- -NO DESCRIPTION OF WHY PARTICLES RETURN TO THE SHOCK (SCATTERING)
- -NO DESCRIPTION OF INJECTION

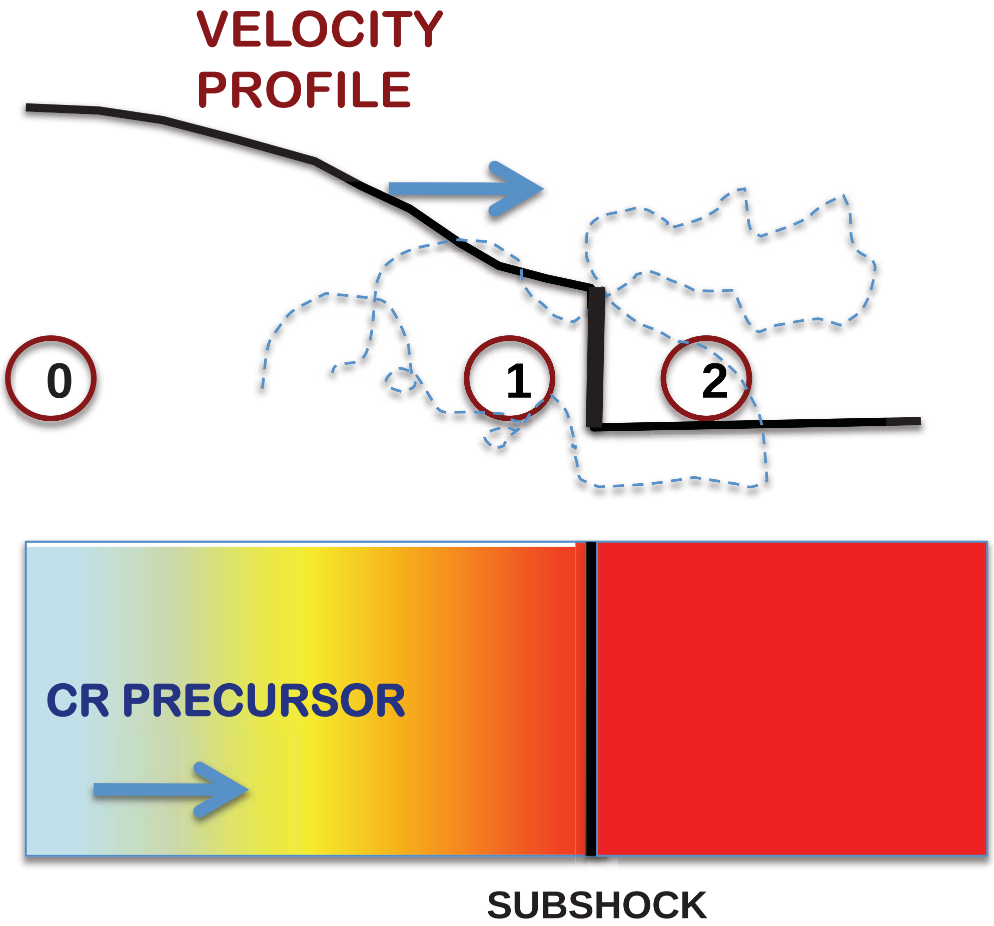
NON LINEAR THEORY

A theory of particle acceleration that allows one to describe:

1. *Dynamical reaction of accelerated particles*
2. *Streaming instability CR-induced B-field*
3. *Dynamical reaction of amplified fields*
4. *Phenomenological recipe for injection (self-regulation of the system)*
5. *Escape of particles from boundaries (Cosmic Rays)*

DIFFUSIVE ACCELERATION AT COLLISIONLESS NEWTONIAN SHOCKS

non linear theory



DENSITY OF ACCELERATED PARTICLES

$$\frac{\partial \rho}{\partial t} = -\frac{\partial (\rho u)}{\partial x}$$

MASS CONSERVATION

$$\frac{\partial (\rho u)}{\partial t} = -\frac{\partial}{\partial x}\left[\rho u^2 + P_g + P_c + P_W\right]$$

MOMENTUM CONSERVATION

$$\frac{\partial}{\partial t}\left[\frac{1}{2}\rho u^2 + \frac{P_g}{\gamma_g - 1}\right] = -\frac{\partial}{\partial x}\left[\frac{1}{2}\rho u^3 + \frac{\gamma_g P_g u}{\gamma_g - 1}\right] - u\frac{\partial}{\partial x}\left[P_c + P_W\right] + \Gamma E_W$$

ENERGY CONSERVATION

$$\frac{\partial f(t,x,p)}{\partial t} + \tilde{u}(x)\frac{\partial f(t,x,p)}{\partial x} = \frac{\partial}{\partial x}\left[D(x,p)\frac{\partial f(t,x,p)}{\partial x}\right] + \frac{p}{3}\frac{\partial f(t,x,p)}{\partial p}\frac{d\tilde{u}(x)}{dx}$$

Closing the system with waves and CR

$$u\frac{\partial P_{\rm g}}{\partial x} + \gamma_{\rm g} P_{\rm g}\frac{{\rm d}u}{{\rm d}x} = (\gamma_{\rm g} - 1)\Gamma E_{\rm W}$$

GAS PRESSURE AND WAVES

$$\frac{\partial}{\partial x}\left[\frac{1}{2}\rho u^3 + \frac{\gamma_{\rm g} P_{\rm g} u}{\gamma_{\rm g} - 1} + \frac{\gamma_{\rm c} P_{\rm c}\tilde{u}}{\gamma_{\rm c} - 1} + F_{\rm W} - \bar{D}(x)\frac{\partial E_{\rm c}}{\partial x}\right] = 0$$

$$\frac{\partial F_{\rm W}}{\partial x} = u\frac{\partial P_{\rm W}}{\partial x} + \sigma E_{\rm W} - \Gamma E_{\rm W}$$

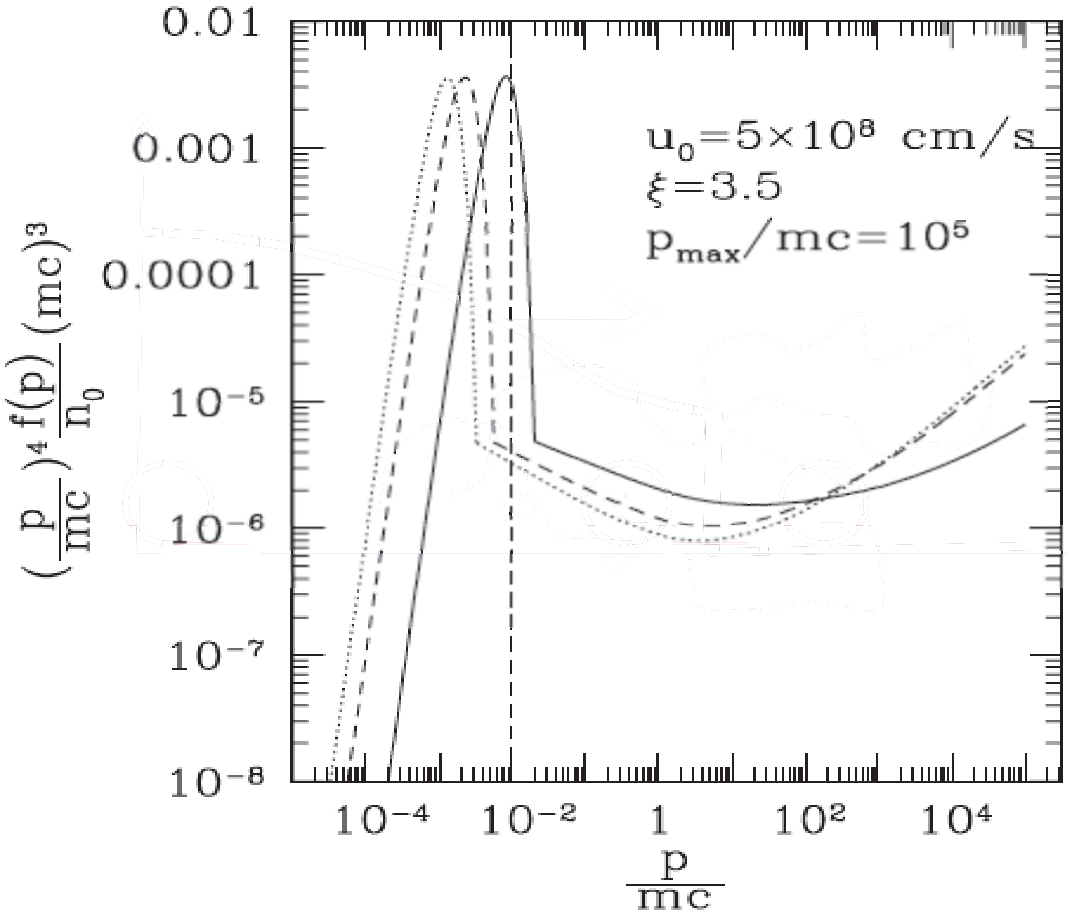
ADVECTION, GROWTH AND DAMPING OF WAVES

$$\sigma E_{\rm W} = v_{\rm A}\frac{\partial P_{\rm c}}{\partial x}$$

ONLY FOR ALFVEN WAVES!!!
AMPLIFICATION OF B-FIELD AS DUE TO CR STREAMING INSTABILITY

DIFFUSIVE ACCELERATION AT COLLISIONLESS NEWTONIAN SHOCKS

non linear theory: BASIC PREDICTIONS



COMPRESSION FACTOR BECOMES FUNCTION OF ENERGY

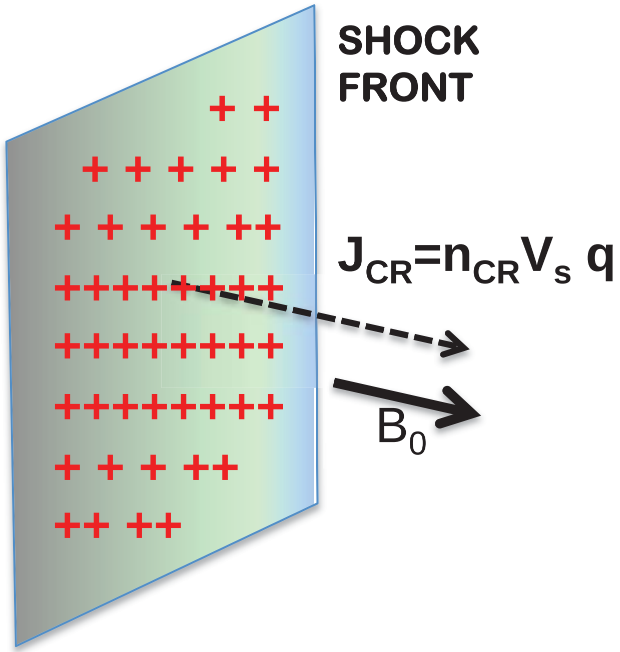
SPECTRA ARE NOT PERFECT POWER LAWS (CONCAVE)

GAS BEHIND THE SHOCK IS COOLER FOR EFFICIENT SHOCK ACCELERATION

SYSTEM SELF REGULATED

EFFICIENT GROWTH OF B-FIELD IF ACCELERATION EFFICIENT

Basics of CR streaming instability



THE UPSTREAM PLASMA REACTS TO THE UPCOMING CR CURRENT BY CREATING A RETURN CURRENT TO COMPENSATE THE POSITIVE CR CHARGE

THE SMALL INDUCED PERTURBATIONS ARE **UNSTABLE** (ACHTERBERG 1983, ZWEIBEL 1978, BELL 1978, BELL 2004, AMATO & PB 2009)

CR MOVE WITH THE SHOCK SPEED (>> V_A). THIS UNSTABLE SITUATION LEADS THE PLASMA TO REACT IN ORDER TO SLOW DOWN CR TO <V_A BY SCATTERING PARTICLES IN THE PERP DIRECTION (B-FIELD GROWTH)

Particle Diffusion ←→ Wave Growth

$$n_{CR} m v_D \rightarrow n_{CR} m v_w \Rightarrow \frac{dP_{CR}}{dt} = \frac{n_{CR} m (v_D - v_w)}{\tau}$$

$$\frac{dP_w}{dt} = \gamma_W \frac{\delta B^2}{8\pi} \frac{1}{v_w}$$

$$\gamma_W = \sqrt{2} \frac{n_{CR}}{n_{gas}} \frac{v_D - v_w}{v_w} \Omega_{cyc}$$

In the ISM this is ~10^{-3} yr^{-1} but close to a shock front the growth can be much larger!!!

δB IS AMPLIFIED BY PARTICLES

MAGNETIC FIELD AMPLIFICATION

SMALL PERTURBATIONS IN THE LOCAL B-FIELD CAN BE AMPLIFIED BY THE SUPER-ALFVENIC STREAMING OF THE ACCELERATED PARTICLES

Particles are accelerated because there is High magnetic field in the acceleration region

High magnetic field is present because particles are accelerated efficiently

Without this non-linear process, no acceleration of CR to High energies (and especially not to the knee!)

BUT…

...MAGNETIC FIELD CAN BE AMPLIFIED BY

1. **RESONANT STREAMING** (Bell 78, Achterberg 83, Zweibel 78)

   Fast generation, fast scattering … saturation?

2. **NON RESONANT STREAMING** (Bell 04, Amato & PB 09)

   Probably more efficient generation rate but inefficient scattering

3. **SHOCK CORRUGATION** (DOWNSTREAM) Giacalone & Jokipii 07
   Not CR induced!
   It happens downstream only, it does not help with particle acceleration unless perpendicular shock

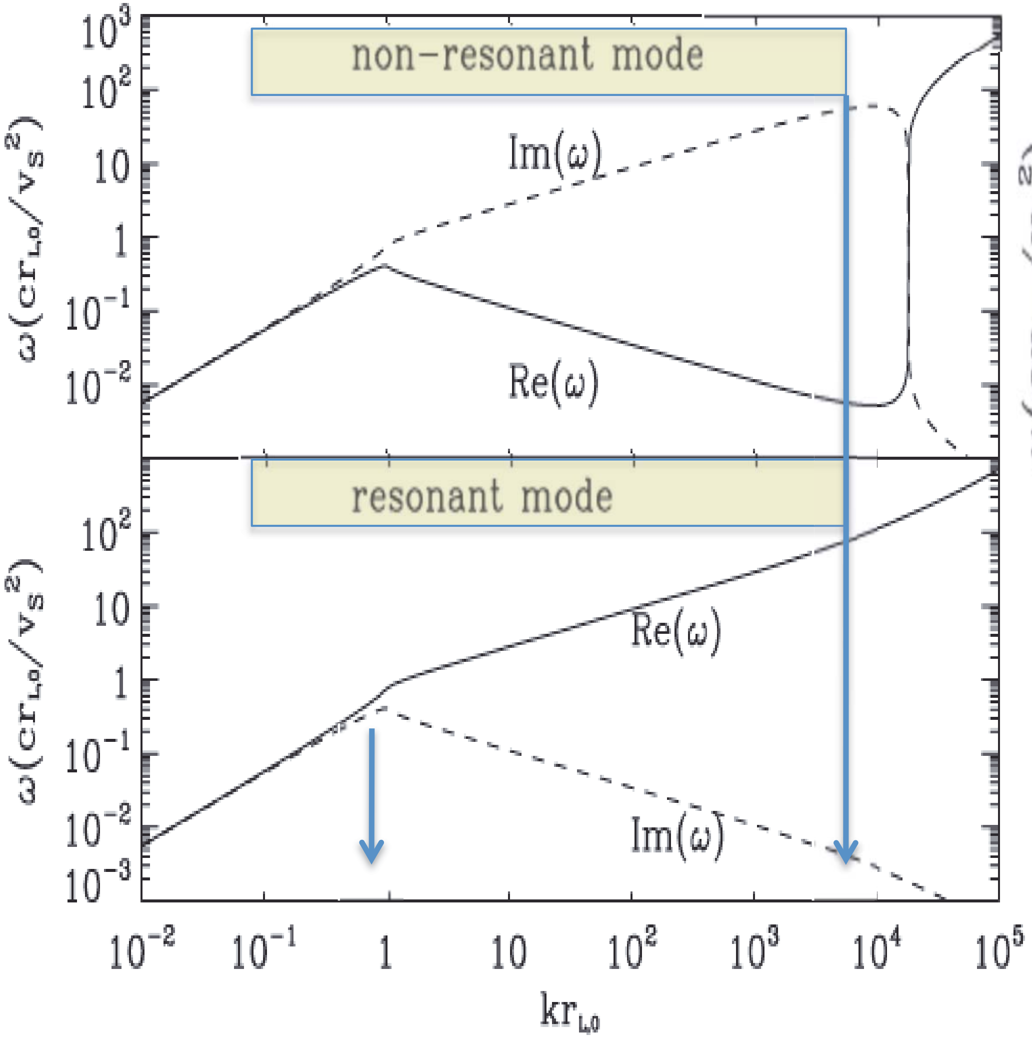
4. **VORTICITY IN THE PRECURSOR** (PB, Matthaeus, et al. 11)

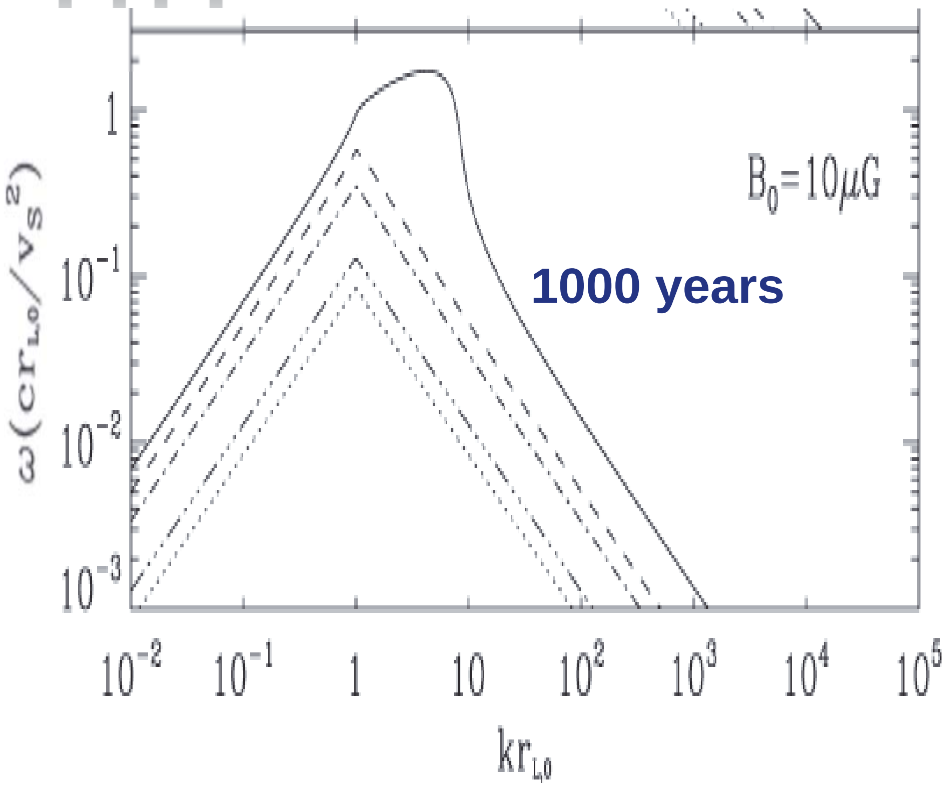
   Potentially very interesting, power on large scales

5. **FIREHOSE INSTABILITY** (Shapiro et al. 98)

   Potentially very interesting, power on large scales

GROWING MODES in CR STREAMING INSTABILITY





NON RESONANT MODES GROW FASTER BUT THEY DO NOT SCATTER PARTICLES EFFECTIVELY UNLESS FAST INVERSE CASCADE

Amato & PB 2009, Bell 2004

SATURATION OF GROWTH

Extremely uncertain. It depends on:

a) Damping (type of waves?)

b) Backreaction of fields on the CR current

c) Coupling between large and small spatial scales

A naïve extrapolation of QLT would lead to:

$$\frac{\delta B^2}{8\pi} = \frac{1}{M_A} \rho V_s^2 \xi_{CR}$$

in the resonant case, upstream (or possibly $\delta B/B \sim 1$ because resonance gets lost)

$$\frac{\delta B^2}{4\pi} = \frac{1}{2} \rho V_s^2 \xi_{CR} \frac{V_s}{c}$$

Estimated analytically from Saturation condition of non resonant Modes (Bell 2004)

X-ray rims and B-field amplification

TYPICAL THICKNESS OF FILAMENTS: ~ 10^{-2} pc

The synchrotron limited thickness is:

$$\Delta x \approx \sqrt{D(E_{max})\tau_{loss}(E_{max})} \approx 0.04\ B_{100}^{-3/2}\ \mathrm{pc}$$

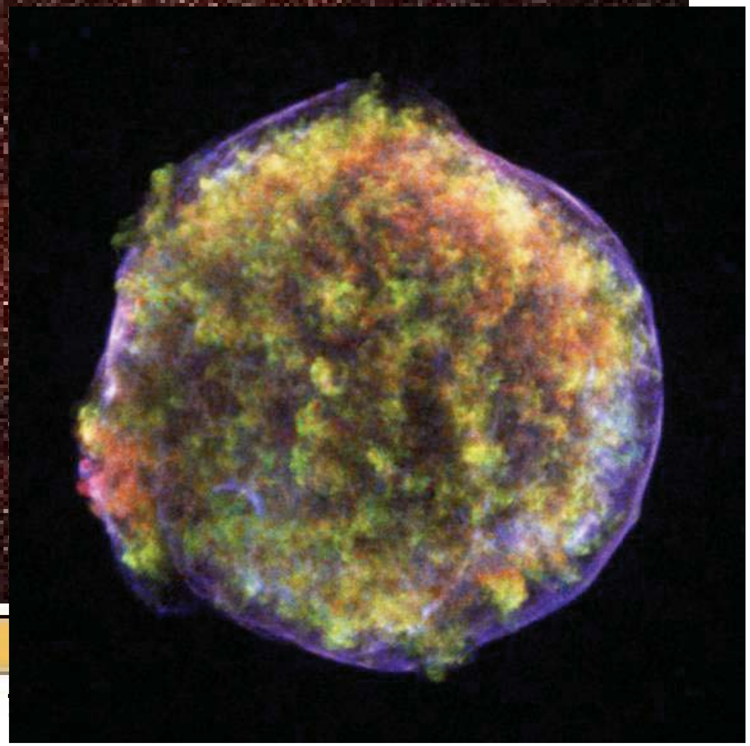
$$B \approx 100\ \mu\mathrm{Gauss}$$

$$E_{max} \approx 10\ B_{100}^{-1/2}\ u_8\ \mathrm{TeV}$$

$$\nu_{max} \approx 0.2\ u_8^2\ \mathrm{keV}$$

In some cases the strong fields are confirmed by time variability of X-rays

Uchiyama & Aharonian, 2007



SPECTRA

THE SPECTRA OF ACCELERATED PARTICLES ARE IN GENERAL CONCAVE AND FLATTER THAN E^{-2} AT HIGH ENERGY

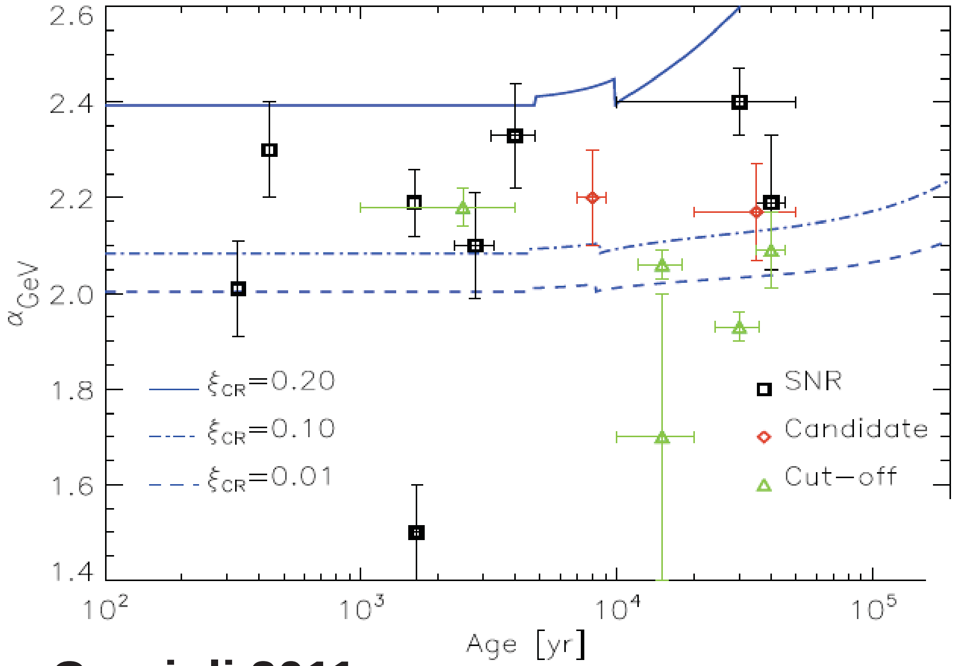
THE MAXIMUM ENERGY WITH B-FIELD AMPLIFICATION REACHS UP TO $\sim 10^{15}$ eV FOR PROTONS (Z TIMES HIGHER FOR NUCLEI)

THESE SPECTRA SHOULD REFLECT IN THE GAMMA RAY SPECTRA (IF DUE TO PP SCATTERING) AND OF NEUTRINOS

BUT THE OBSERVED SPECTRA OF GAMMAS ARE TYPICALLY $\sim E^{-2.3}$

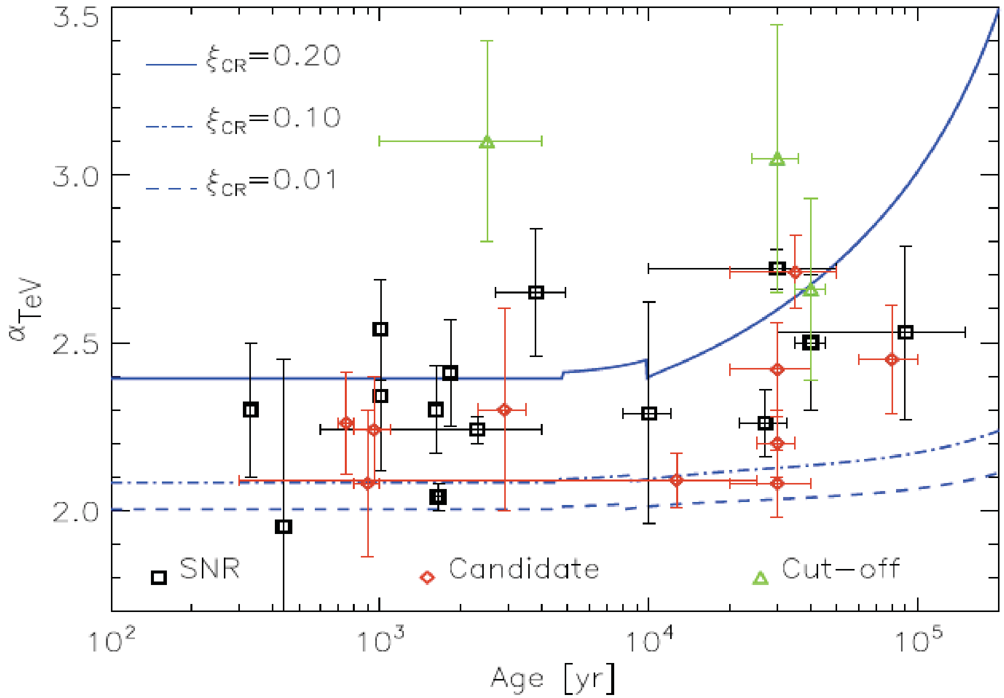
CLEARLY INCOMPATIBLE WITH LEPTONIC MODELS! BUT ALSO NOT COMPATIBLE WITH THE SIMPLEST PREDICTION OF NLDSA

TROUBLE WITH SLOPES ?



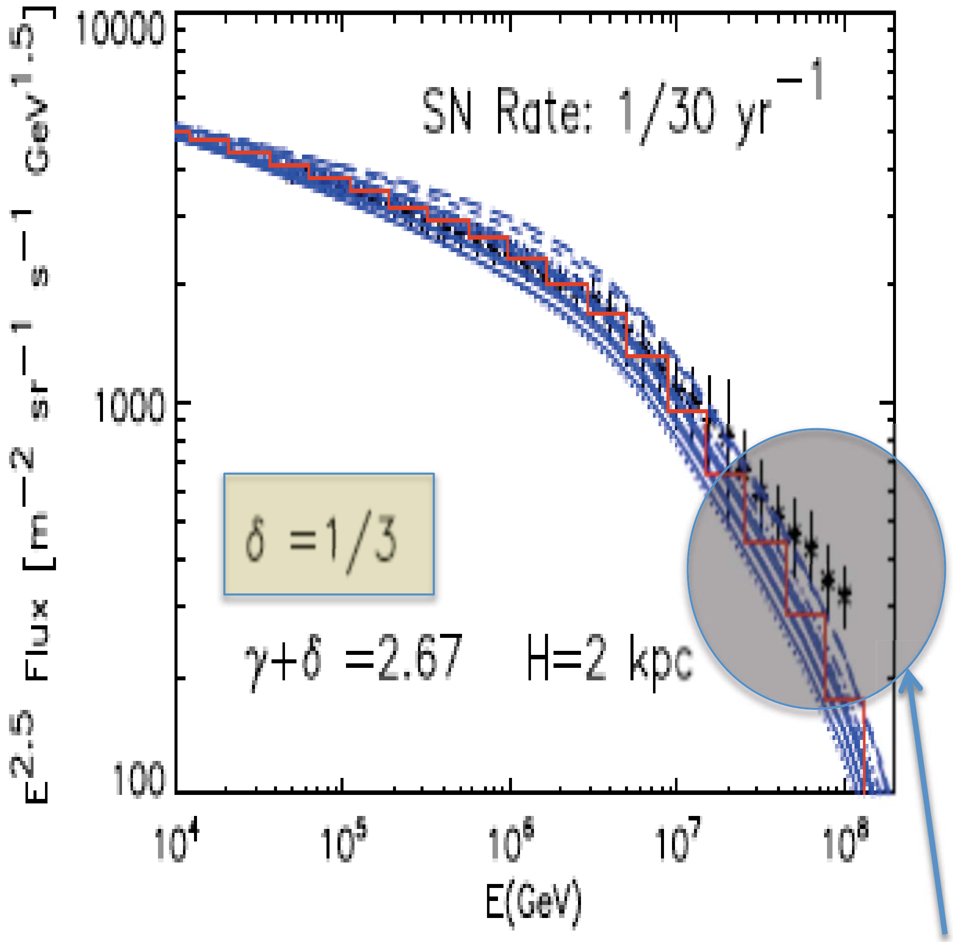
VERY SURPRISING TO SEE THAT THE REQUIRED ACCELERATION EFFIC. ARE HIGH BUT THE SPECTRA ARE STEEP

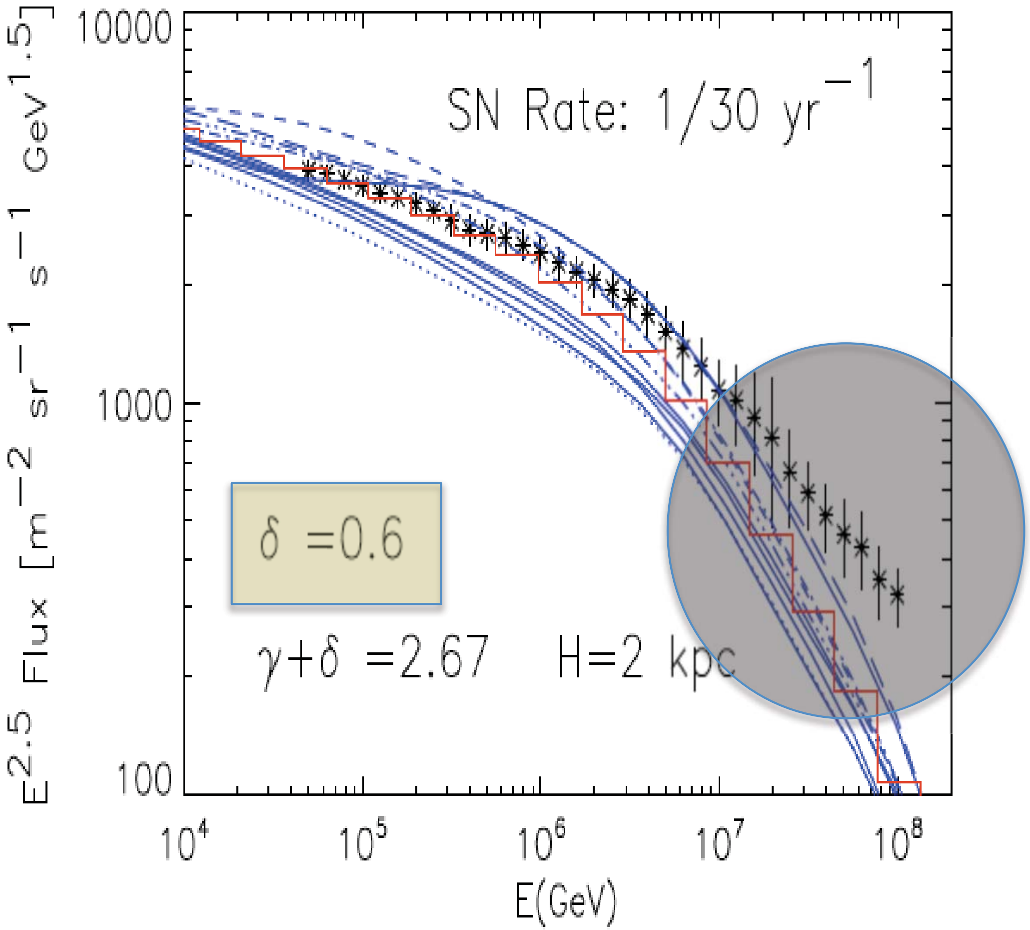
Caprioli 2011



CR spectra and SNRs

Blasi & Amato 2011

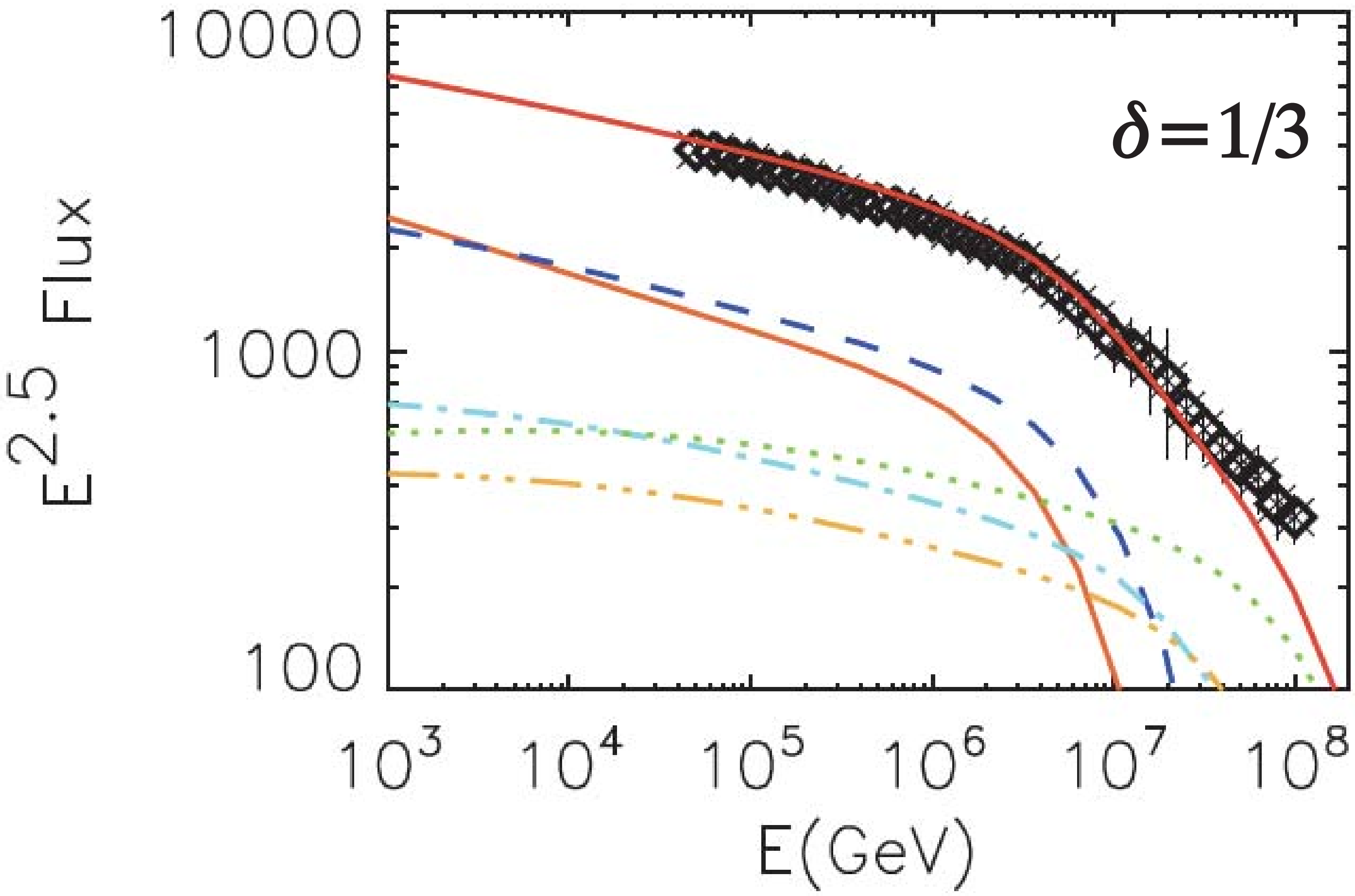




Deficit compensated
by extragalactic CRs

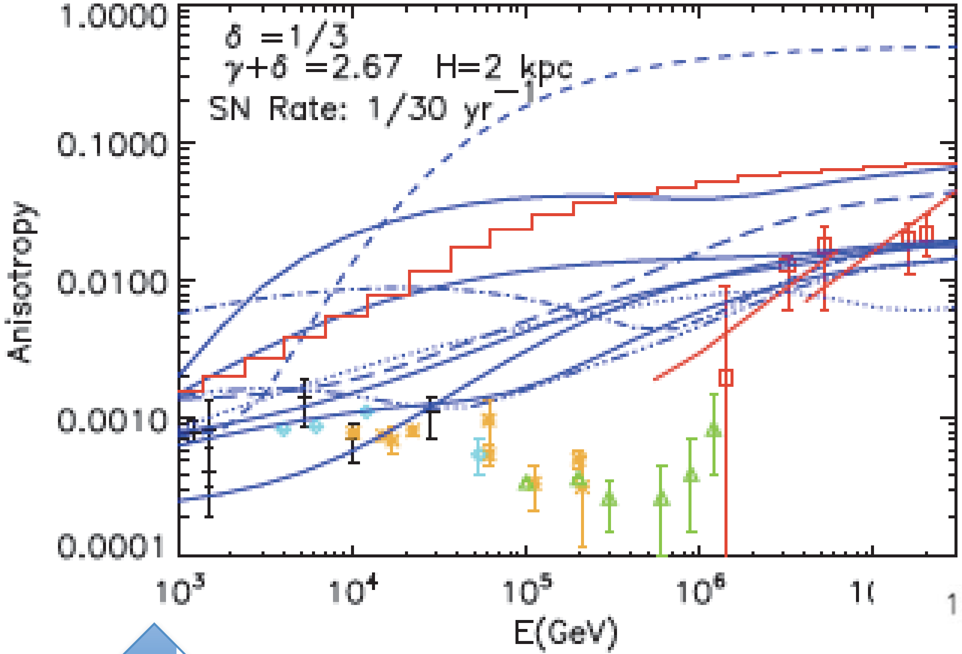
Chemicals and the KNEE

Blasi & Amato 2011



ONLY FOR δ=1/3 SPECTRUM OF He HARDER THAN SPECTRUM OF PROTONS AS A RESULT OF SPALLATION

CR Anisotropy



Naïve expectation:

$$\delta_A = \frac{3}{2^{3/2}} \frac{1}{\pi^{1/2}} \frac{D(E)}{Hc}$$

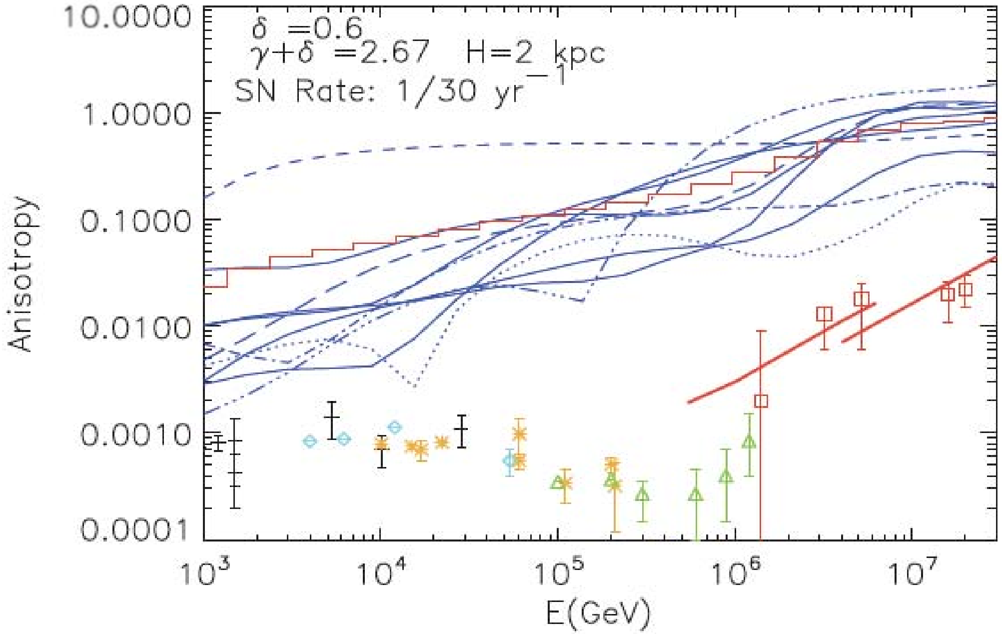
proportional to E^{δ}

δ=1/3

γ~2.35

δ=0.6

γ~2.05



Blasi & Amato 2011

BEYOND THE SIMPLEST APPROACH

1. **DYNAMICAL REACTION OF THE B-FIELD**
   $P_W = B^2/8\pi > P_{gas}$ the eq. of state becomes dominated by B and The compression factor gets smaller → steeper spectra (Caprioli, PB, Amato & Vietri 2008, 2009)

2. **SCATTERING CENTERS WITH LARGE VELOCITY**
   All but trivial (spectra depend on type and helicity of waves) but if $v_W \sim v_A(\delta B) >> v_A$, then:

$$\tilde{r} = \frac{u_1 + v_{A,1}}{u_2 + v_{A,2}} \qquad \alpha = \frac{\tilde{r}+2}{\tilde{r}-1} > 2$$

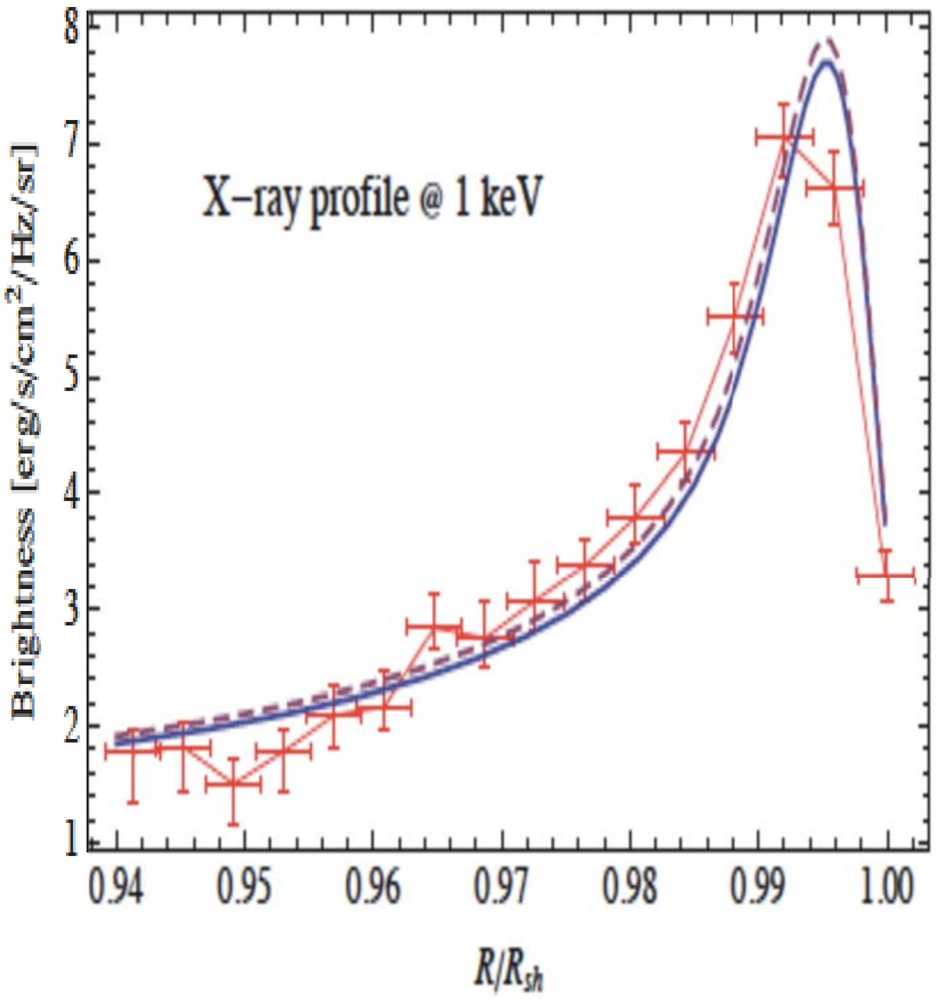
3. **ESCAPE FLUX OF CR IS DIFFERENT FROM THE SPECTRUM OF ACCELERATED PARTICLES** (CAPRIOLI, PB, AMATO 2009)

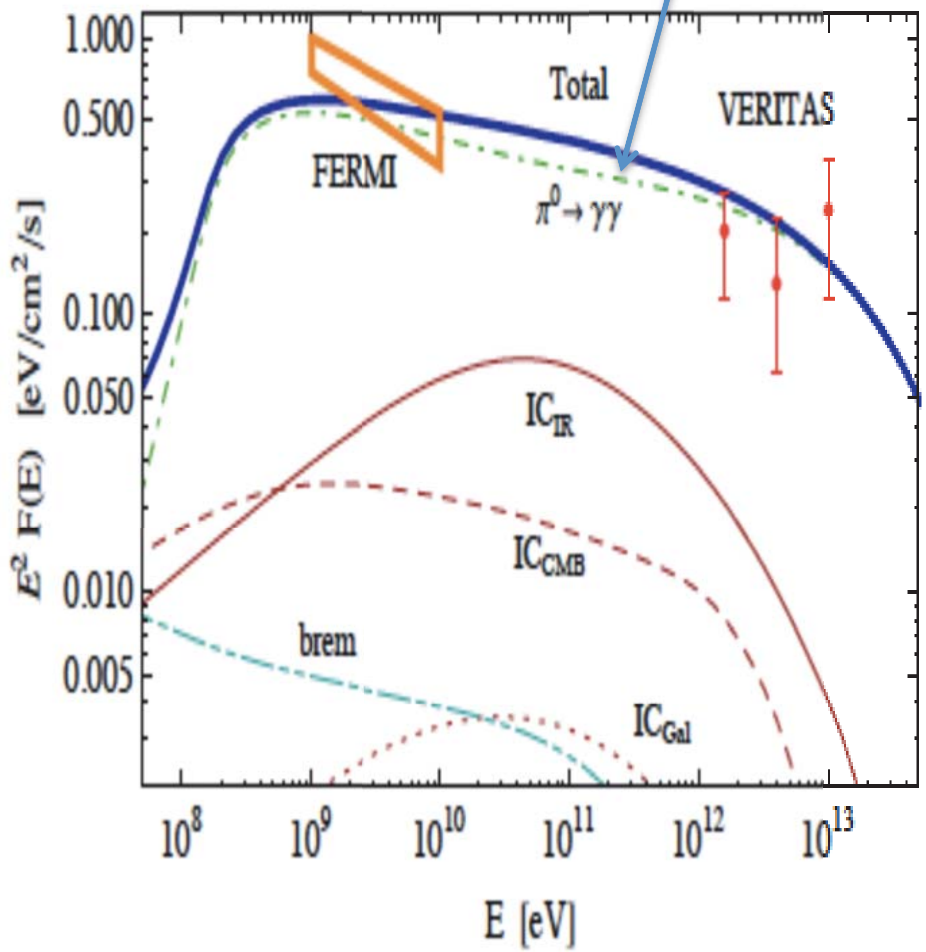
4. **PRESENCE OF NEUTRALS**
   Charge exchange with ions leads to weakening of the shock Strength (PB et al. 2011)

The case of Tycho

Morlino&Caprioli 2011

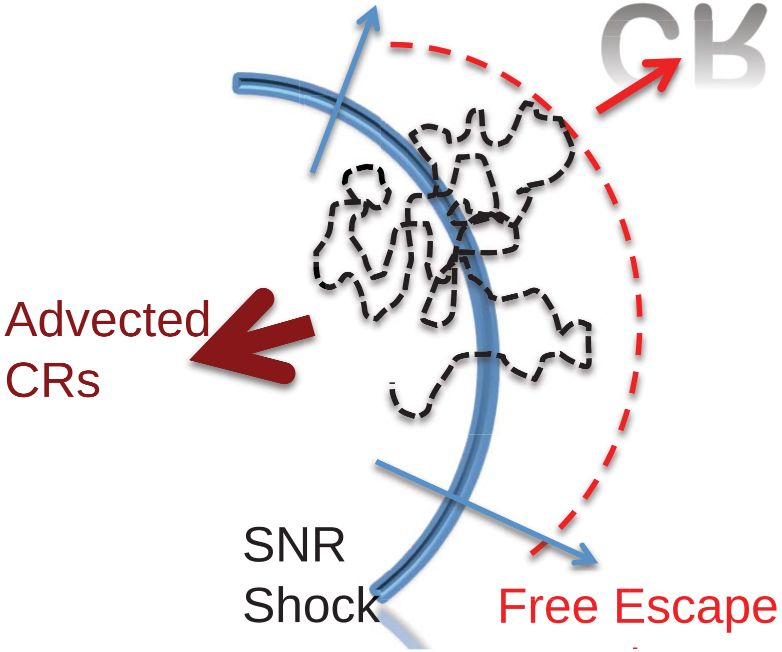
STEEP SPECTRUM BASICALLY IMPOSSIBLE TO EXPLAIN WITH LEPTONS





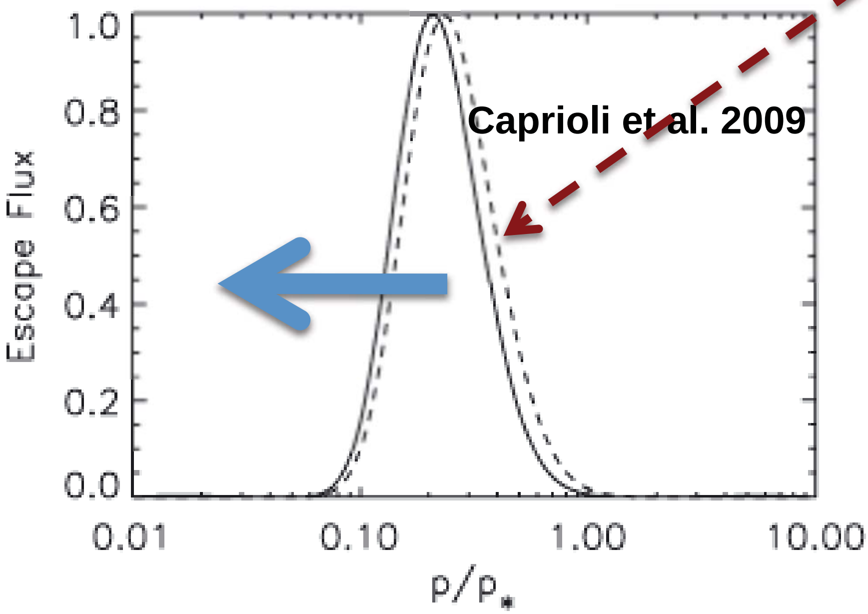
SEE ALSO POSTER 135, 146 FOR FERMI DATA ON TYCHO

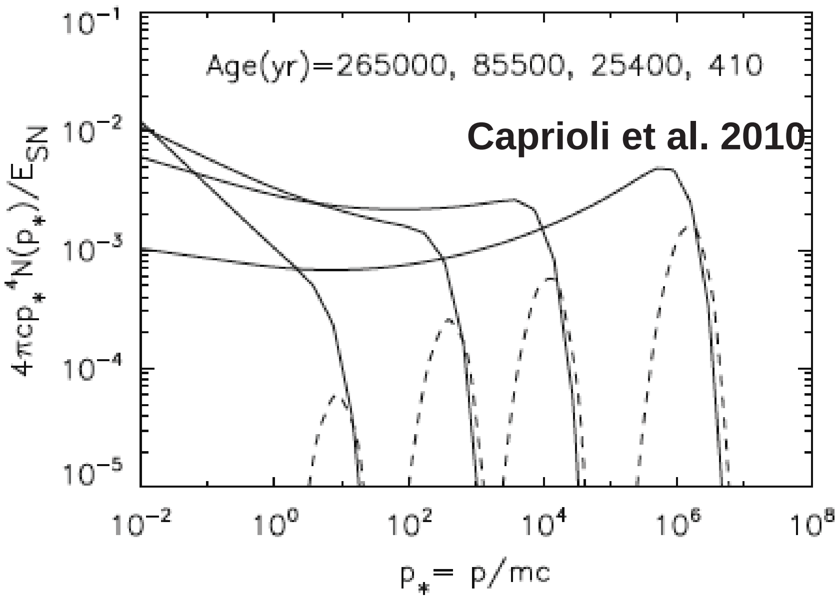
CR ESCAPE



The escape flux can be calculated using the transport equation IF one assumes a free escape boundary surface **(DURING ST PHASE)**

$$\Phi_{esc}(E,x) = D(E)\left(\frac{\partial f(E,x)}{\partial x}\right)_{x=x_{fe}}$$





CR ESCAPE AND CLOUDS

TWO SCENARIOS:

SNR SHOCK ENTERS THE MC

Collisionless shock only involves the small fraction of Ions (low density)

Ion-neutral density kills waves $\rightarrow$ low E_{max}

MC IS ILLUMINATED BY CR FROM SNR

The mc only acts as a target for pp

Gamma ray flux depends on

-Age of SNR

-Diffusion coefficient around the SNR

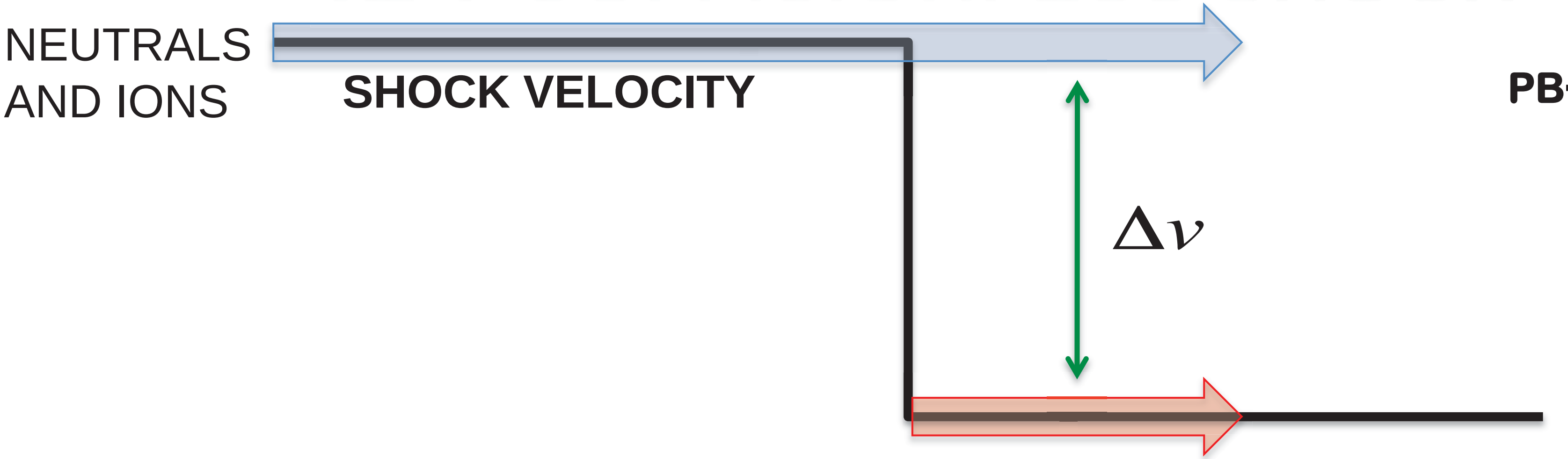
-Escape physics

SOME RECENT PROGRESS AND POSSIBLY FUTURE DEVELOPMENTS

COLLISIONLESS SNR SHOCKS IN PARTIALLY IONIZED MEDIA:

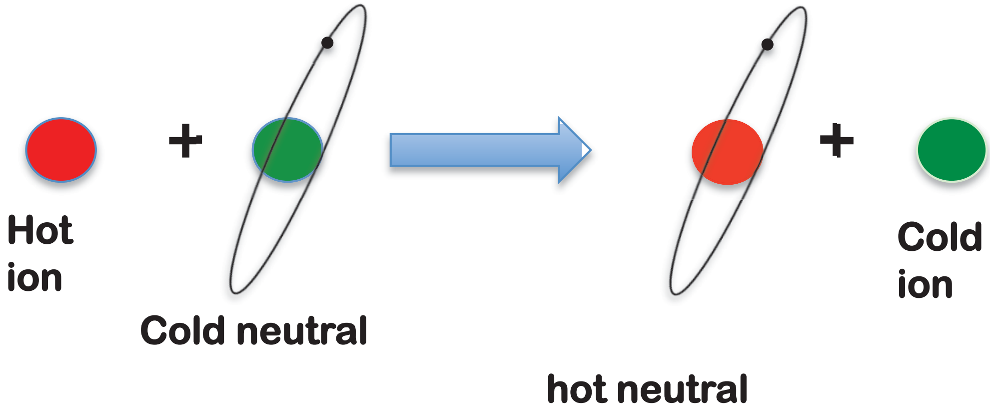
Anomalous width of Balmer lines

SUBTLE ASPECTS OF ACCELERATION AT A COLLISIONLESS SHOCK



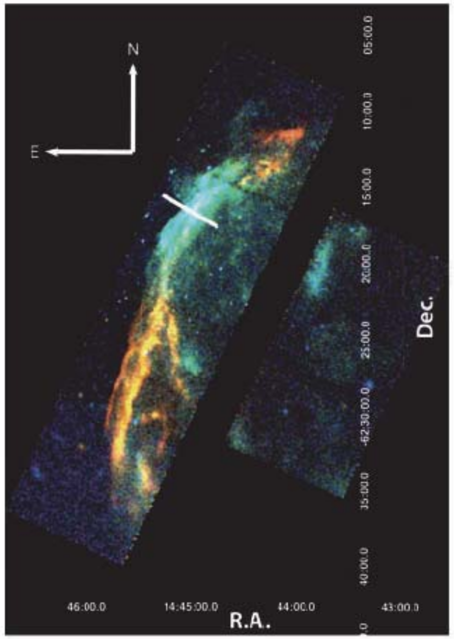
PB+, 2011

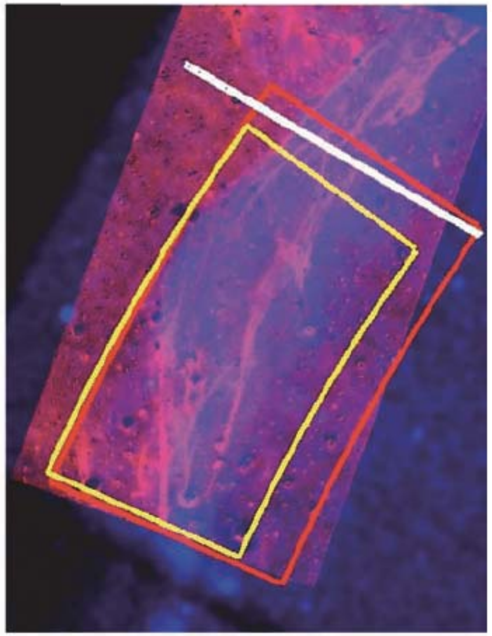
CHARGE EXCHANGE → BROAD BALMER LINE (NEUTRALS THAT MADE CHARGE EXCHANGE) REFLECTING THE TEMPERATURE OF IONS…

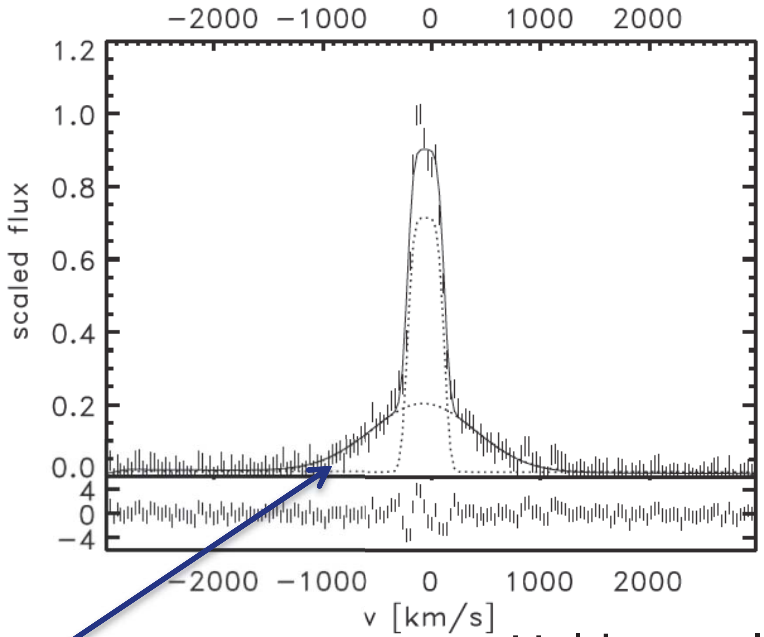


BUT THE LATTER AFFECTED BY EFFICIENT CR ACCELERATION

BROAD BALMER LINES NARROWER THAN FOR UNMODIFIED SHOCKS







Helder et al. 2009

$$W_{broad} = \sqrt{8 \ln 2 \frac{kT_2}{m}} \approx 1.02\ v_{sh}$$

$$W_{broad} = 1100 \pm 63\, km/s \ \rightarrow\ T_2 = 2.3 \pm 0.3\, keV$$

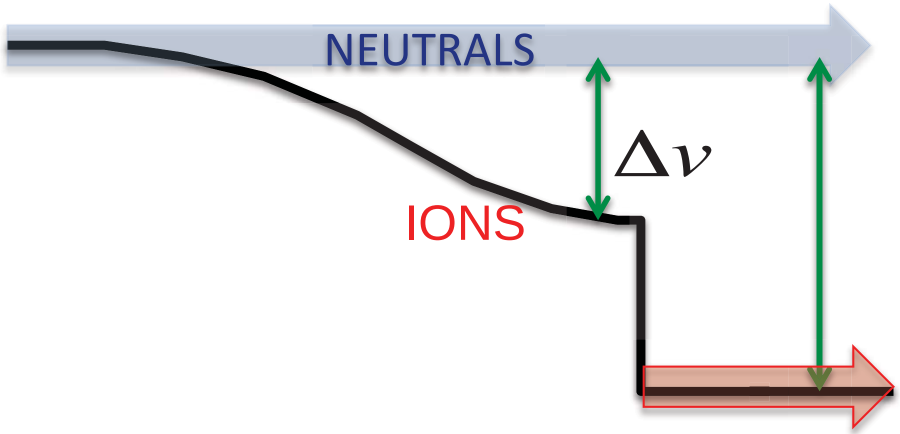
Shock speed from proper motion

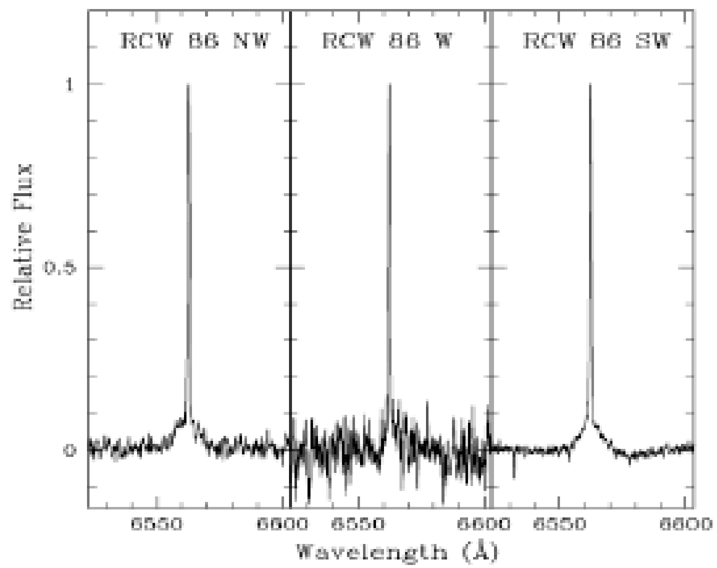
$$v_{shock} = 6000 \pm 2800\, km/s \left(\frac{d}{2.5 \pm .5\, kpc} \right) \left(\frac{\dot{\theta}_{obs}}{0.5 \pm .2''\, yr^{-1}} \right) \rightarrow T_2 = \begin{matrix} 20-150\, keV\ (no\ equilibration) \\ 12-90\, keV\ (equilibration) \end{matrix}$$

INFERRED EFFICIENCY of CR ACCELERATION 50-60% !!! (BUT model dependent)

NARROW BALMER LINES BROADER THAN FOR UNMODIFIED SHOCKS

Sollerman et al. 2003





$$W_{broad} = \sqrt{8\ln 2\frac{kT_0}{m}} \approx 21\ \mathrm{km/s}\left(\frac{T_0}{10^4 K}\right)^{1/2}$$

$$W_n \sim 30-50\ km/s \rightarrow T \sim 2-6\ 10^4 K$$

CHARGE EXCHANGE OCCURS NOW IN THE CR INDUCED PRECURSOR

NARROW BALMER LINE BROADER THAN FOR AN UNMODIFIED SHOCK

CONCLUSIONS

BASIC PRINCIPLES OF ACCELERATION IN SNR WELL POSED – HINT TO END OF GALACTIC CR AT ~FEW 10^{17} eV

BUT HARD TO MOVE AHEAD IN THE DETAILS (WE OBSERVE LARGE SCALES BUT THEY ARE DETERMINED BY VERY SMALL SCALES)

EFFICIENT ACCELERATION $\neq$ BRIGHT GAMMA OR NEUTRINO SOURCE (e.g. HIGH EFF. AND LARGE P_{MAX} FOR A SNII IN TENUOUS BUBBLE)

MAX ENERGY AT THE BEGINNING OF SEDOV: USUALLY INSIDE BUBBLE (NOT EASY TO SEE PEVATRONS UNLESS SNIa)

B-FIELD AMPLIFICATION BUT UNCLEAR DETAILS (SATURATION, SCALES – OBSERVATIONALLY HARD TO ACCESS)

STRONG EVIDENCE FOR STEEP SPECTRA (CAN'T BE LEPTONIC) ~ $E^{-2.2}$ (RECALL ESCAPE SPECTRUM $\neq$ ACCELERATED SPECTRUM)

BIG DEVELOPMENTS FROM BALMER DOMINATED SHOCKS AS INDICATORS OF CR ACCELERATION EFFICIENCY