

# A Comprehensive Assessment of the Too-Big-to-Fail Problem

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Acknowledgments:

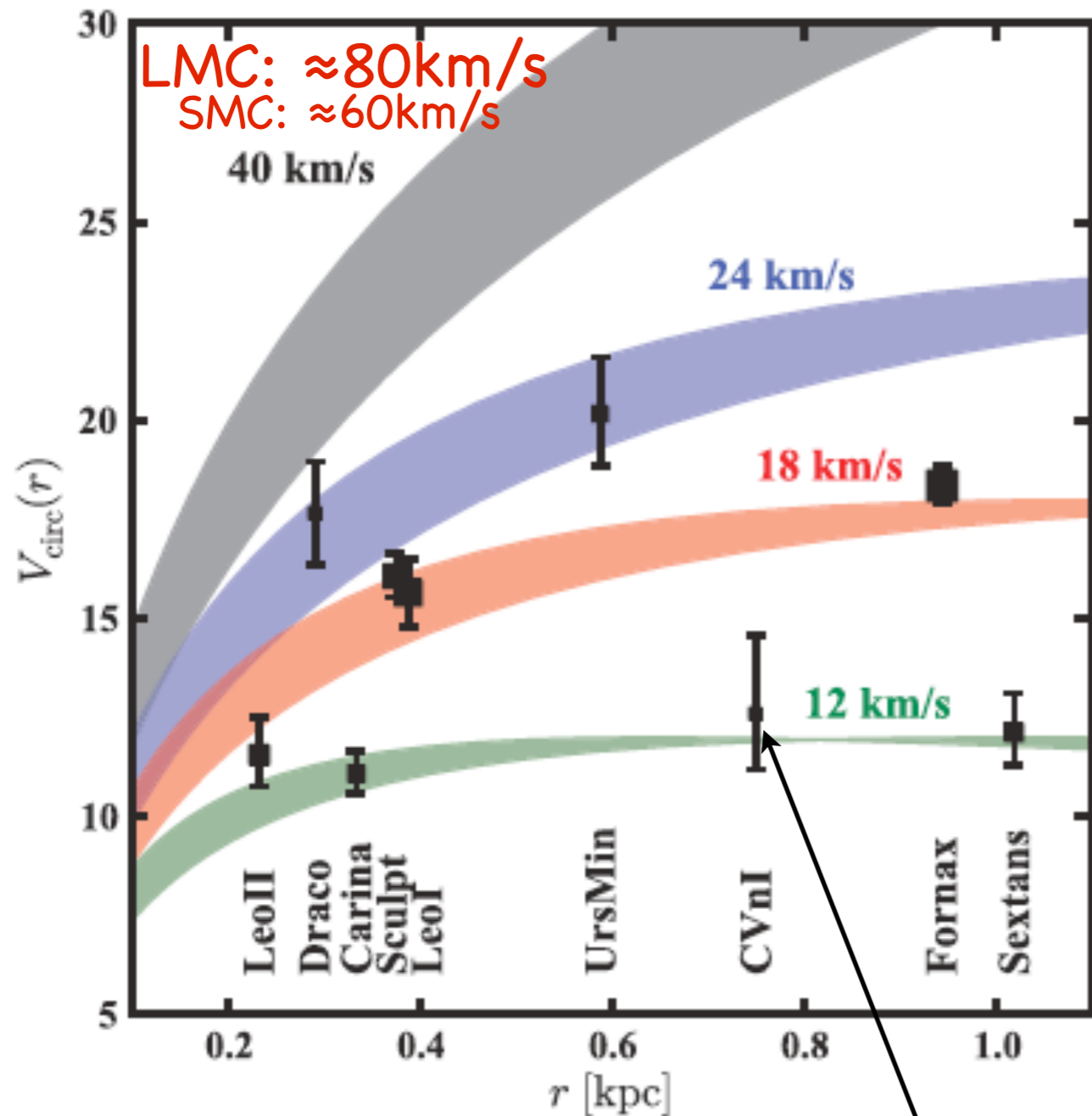
Peter Behroozi, Mike Boylan-Kolchin, Shea Garrison-Kimmel, Erik Tollerud, Andrew Hearin, Duncan Cambell

Workshop on Cosmological Structures, ICTP, Trieste

# Outline

- What is “Too Big To Fail” (TBTf) ?
- Semi-analytical model of dark matter subhaloes
- Severity of TBTf

# TBTF



$V_{\text{circ}}(r_{1/2})$   
Wolf et al. (2010)

Boylan-Kolchin et al. (2012)

## Formulation I:

- simulation: order of 10 subhaloes with  $V_{\text{max}} > 25 \text{ km/s}$
- MW dSphs:  $V_{\text{max}} \leq 25 \text{ km/s}$
- "massive subhalo" formulation

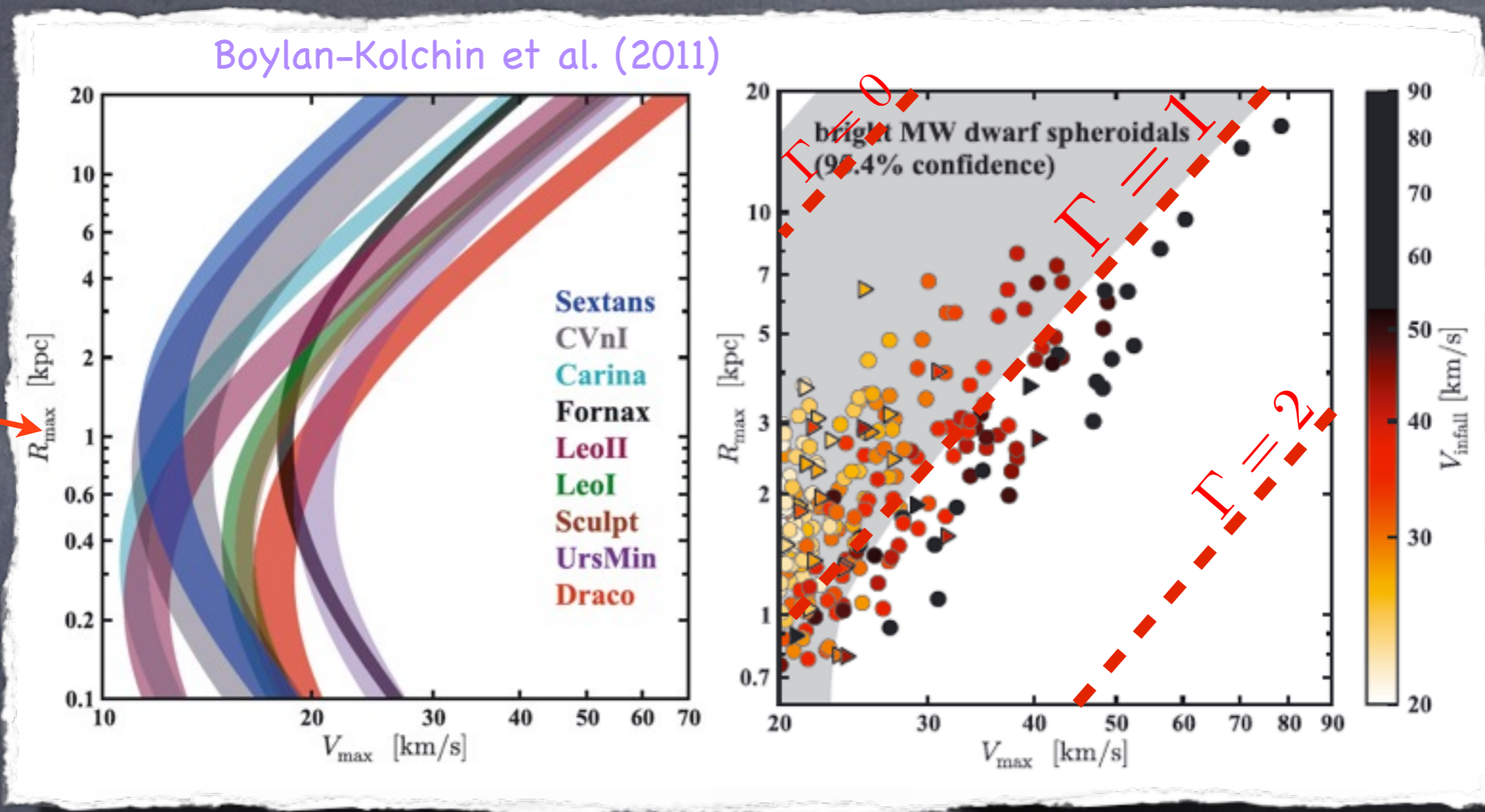
## Formulation II:

a  $V_{\text{max}}$  gap between  $\approx 60 \text{ km/s}$  and  $\approx 25 \text{ km/s}$

# Formulation III:

Boylan-Kolchin et al. (2011)

$R_{\max}$ :  
the  
radius  
at which  
 $V_{\text{circ}}(r)$   
reaches  
 $V_{\max}$



subhalo density proxy  $\Gamma \equiv 1 + \log(0.0014V_{\max}^{2.2}/R_{\max})$   
Purcell & Zentner (2012)

the most massive subhaloes are too dense ( $\Gamma_{\max} > 1$ ) to be consistent with MW dSphs ( $\Gamma < 1$ )

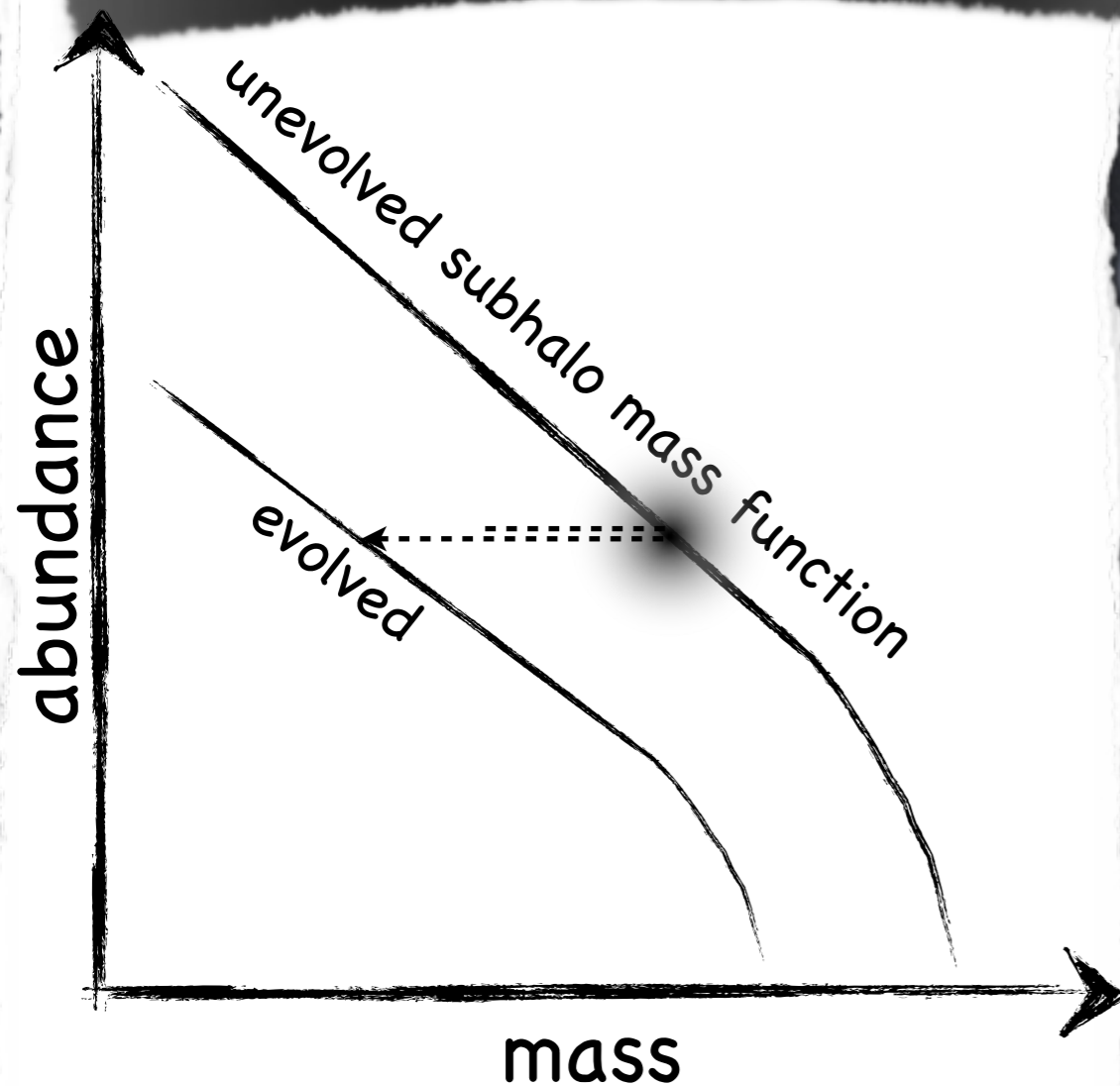
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# Merger Tree (EPS)

Jiang & van den Bosch (2014a)

Parkinson et al. (2008):  $\frac{dN}{dm} (m, z | M_0, z_0) dm$



For more about subhalo evolution:

Jiang & van den Bosch (2014b)

arXiv:1403.6827

## mass evolution:

$$\dot{m} = -A \frac{m}{\tau_{\text{dyn}}} \left( \frac{m}{M} \right)^{0.07}$$

- $P(A)$  : log-normal
- reflects variance in orbital properties & halo concentrations

## disruption:

$$m_{\text{dis}} = m_{\text{acc}} (< \alpha \times r_{s,\text{acc}})$$

- $P(\alpha)$  : log-normal
- $\bar{\alpha} = \bar{\alpha}(m_{\text{acc}}/M_{\text{acc}})$

## structure evolution:

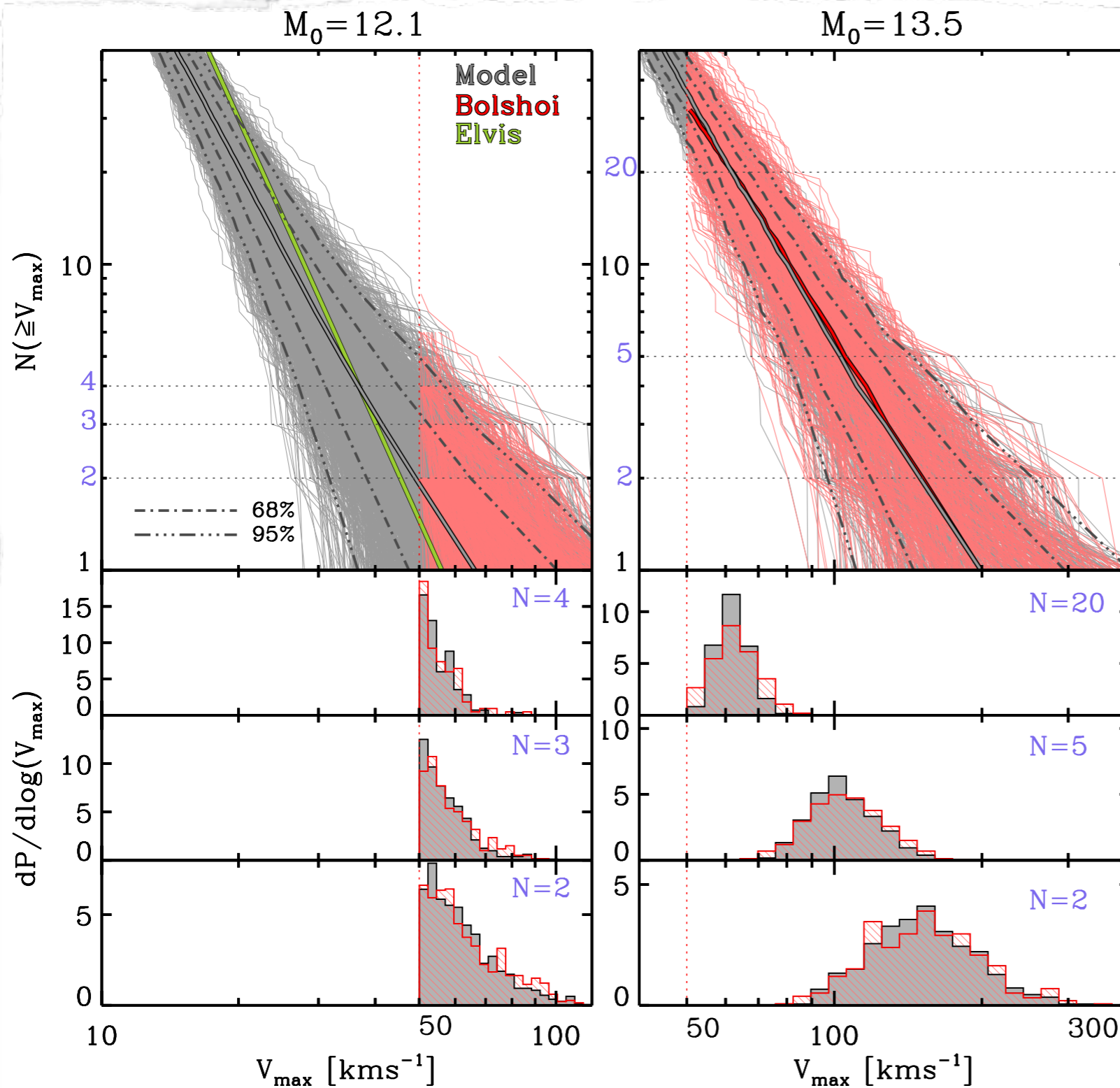
$$V_{\text{max}} = V_{\text{acc}} \times f(m/m_{\text{acc}})$$

↑  
Penarrubia et al. (2010)

$$V_{\text{acc}} = g(m_{\text{acc}}, c_{\text{acc}})$$

↑  
Zhao et al. (2009)

# Model: Accurate Halo-to-Halo Variance



- benchmark:** Bolshoi simulation
  - 441  $M_0 = 10^{13.5 \pm 0.05} h^{-1} M_{\odot}$
  - 1986  $M_0 = 10^{12.10 \pm 0.01} h^{-1} M_{\odot}$
- model:**
  - 500  $M_0 = 10^{13.5} h^{-1} M_{\odot}$
  - 2000  $M_0 = 10^{12.1} h^{-1} M_{\odot}$

Jiang & van den Bosch, submitted to MNRAS

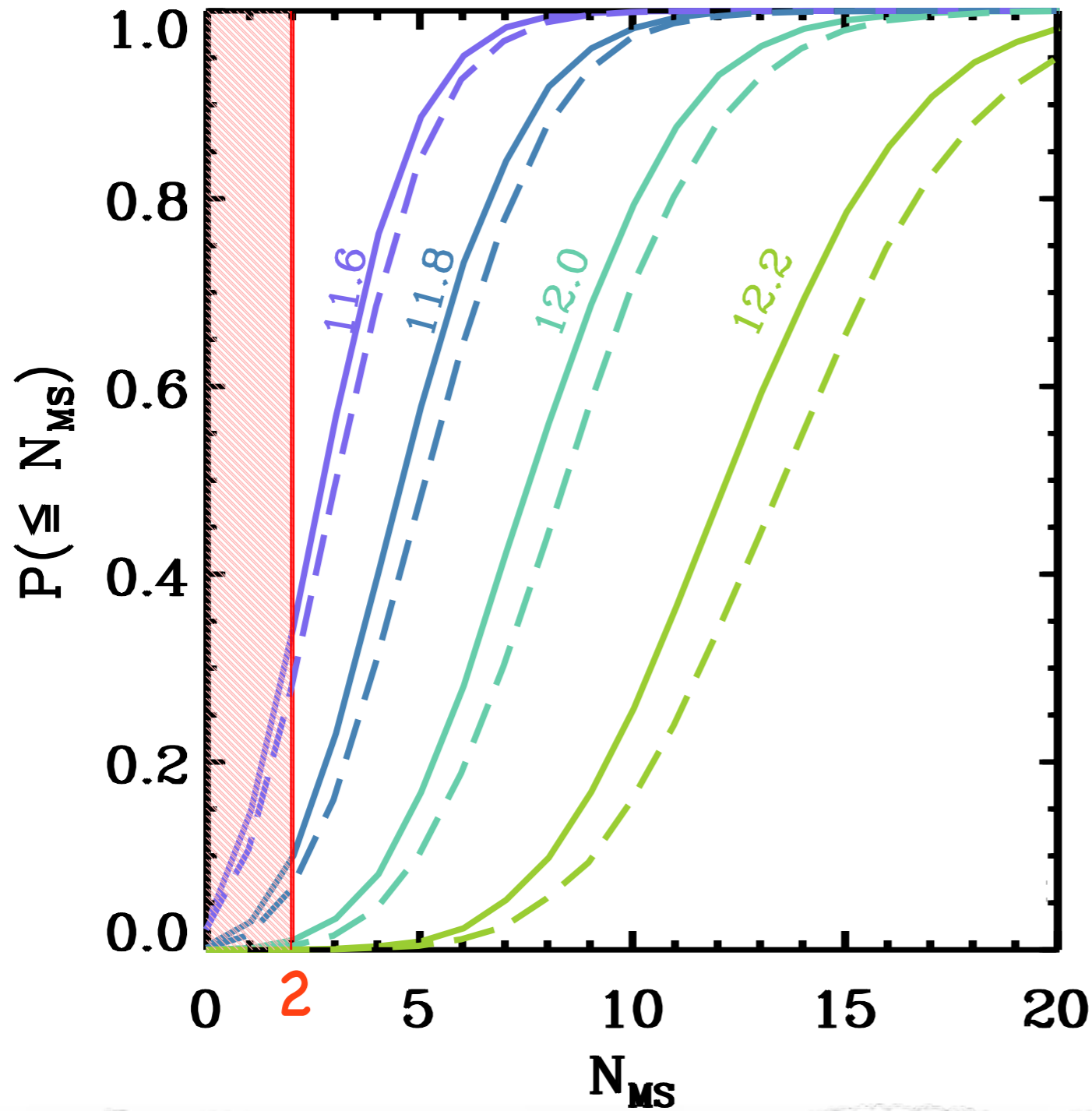
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- **Severity of TBTf**



# "Massive Subhalo" Count



definition:

$$V_{\text{acc}} > 30 \text{ km s}^{-1}$$

$$V_{\text{max}} > 25 \text{ km s}^{-1}$$

MW has **2** MSs

Wang et al. (2012):  
lower MW halo mass  
=> significantly  
lower number of MSs

Contemporary MW  
halo mass constraint:

$$M_0 \in [10^{11.7}, 10^{12.2}] h^{-1} M_{\odot}$$

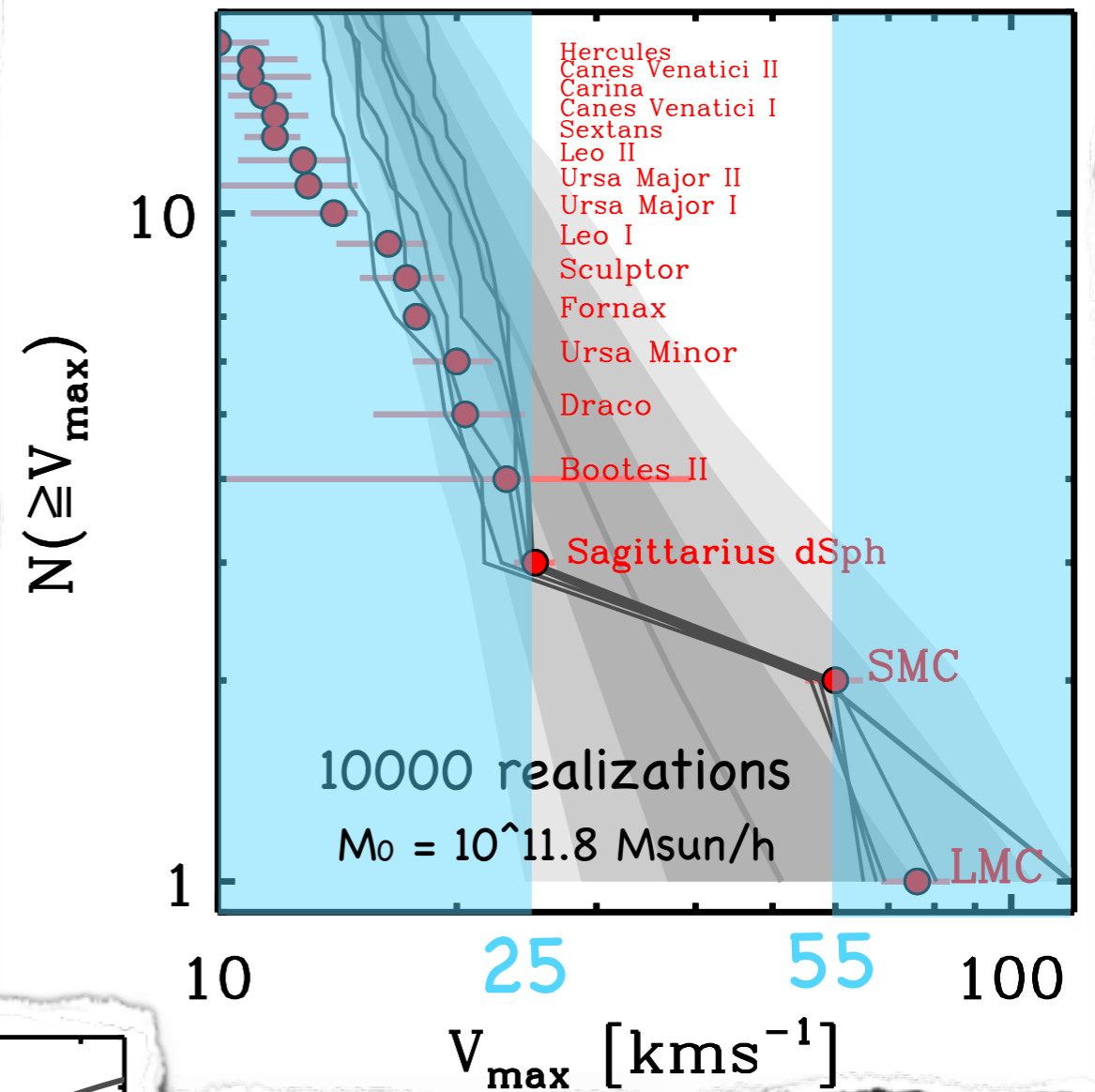
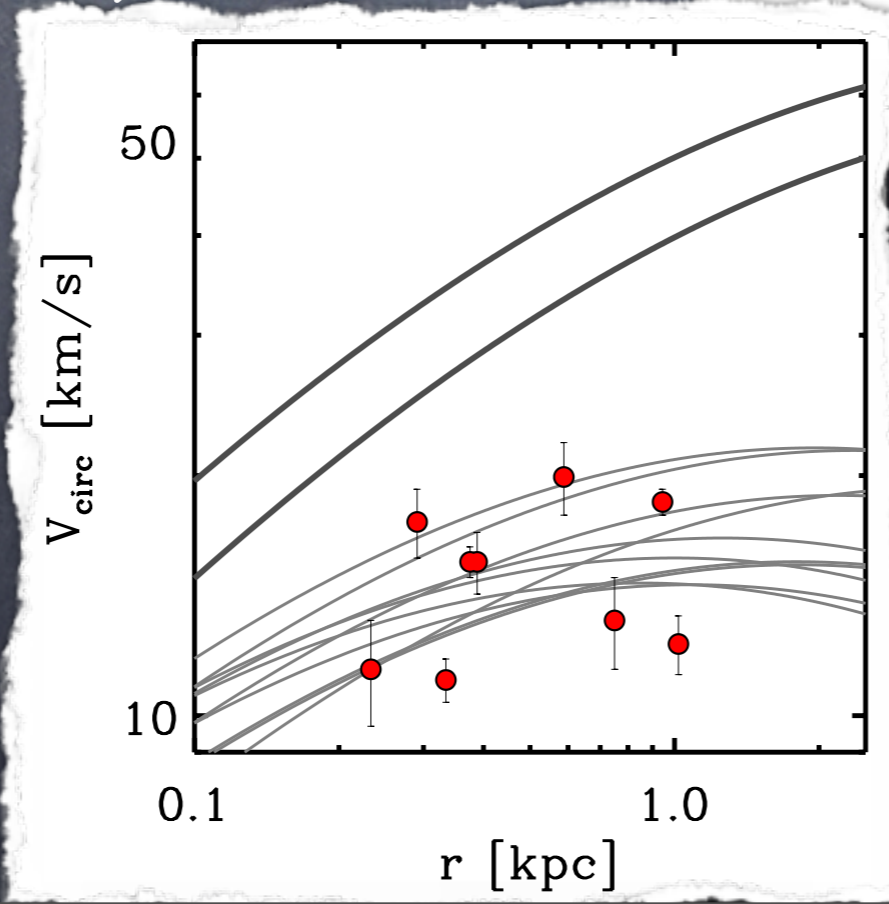
↑  
Kafle et al. (2014)

10,000 realizations for each halo mass

# Vmax Gap

- Vmax (estimates) for MW satellites from
  - Kuhlen et al. (2010)
  - Boylan-Kolchin et al. (2012)
  - Kallivayalil et al. (2013)
- for MW satellites with no published Vmax, use
  - MacConnachie (2012)
  - Rashkov et al. (2012)

$$V_{\max} = 2.2\sigma_{\text{LOS},*}$$



Jiang & van den Bosch,  
submitted to MNRAS

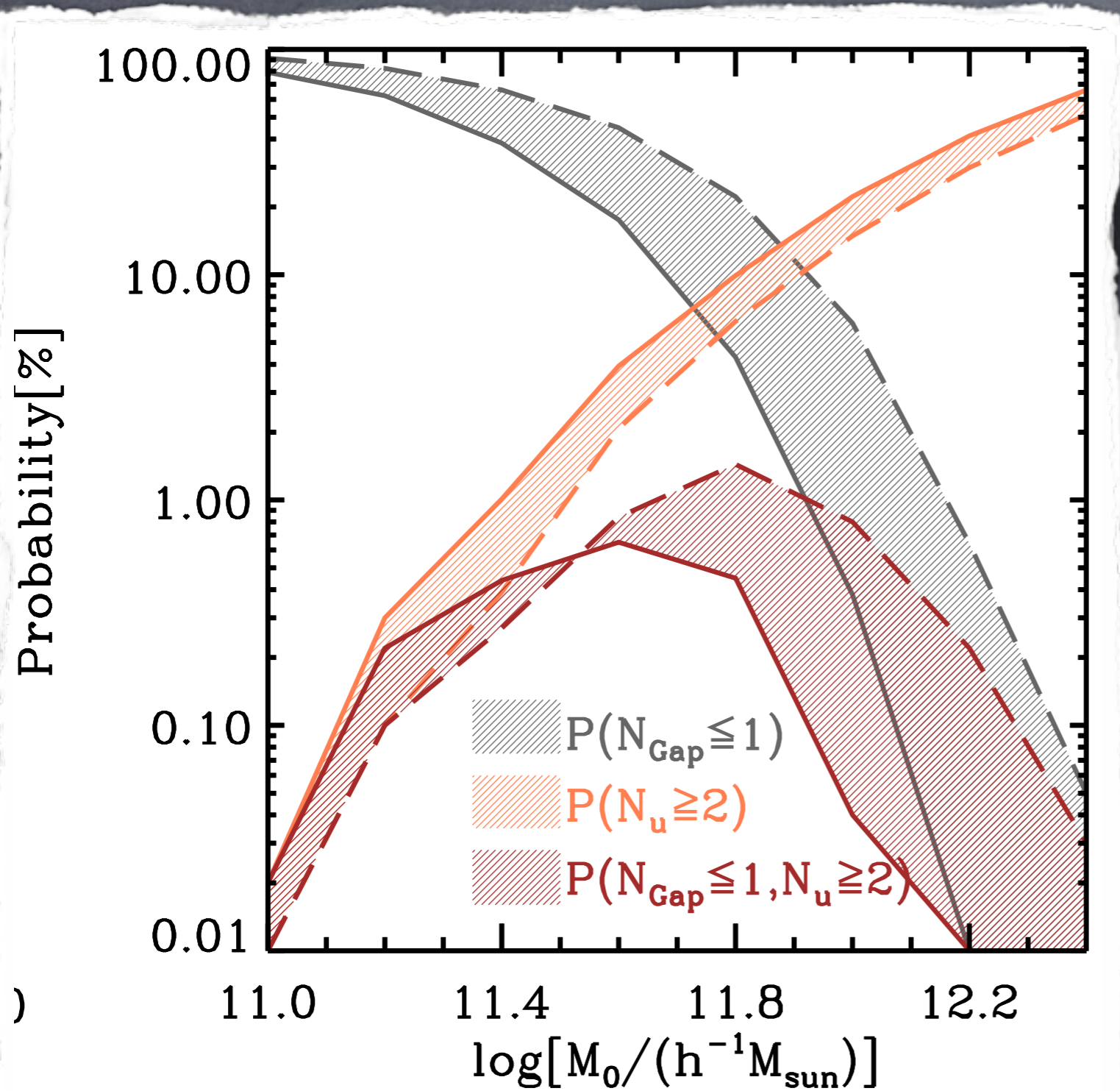
$$V_{\text{circ}}(r | R_{\text{max}}, V_{\text{max}}, \alpha)$$

Einasto shape parameter,  
typically 0.18 (Aquarius)

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# MW-consistent fraction as a function of halo mass

10,000 realizations for each halo mass



$N_{\text{Gap}} \leq 1$   
(number of subhaloes in the gap  $\leq 1$ )

$$V_{\text{max}} \in [25, 55] \text{ km s}^{-1}$$

or

$$V_{\text{max}} \in [30, 60] \text{ km s}^{-1}$$

$N_u \geq 2$   
(number of MC analogs  $\geq 2$ )

$$V_{\text{max}} > 55 \text{ km s}^{-1}$$

or

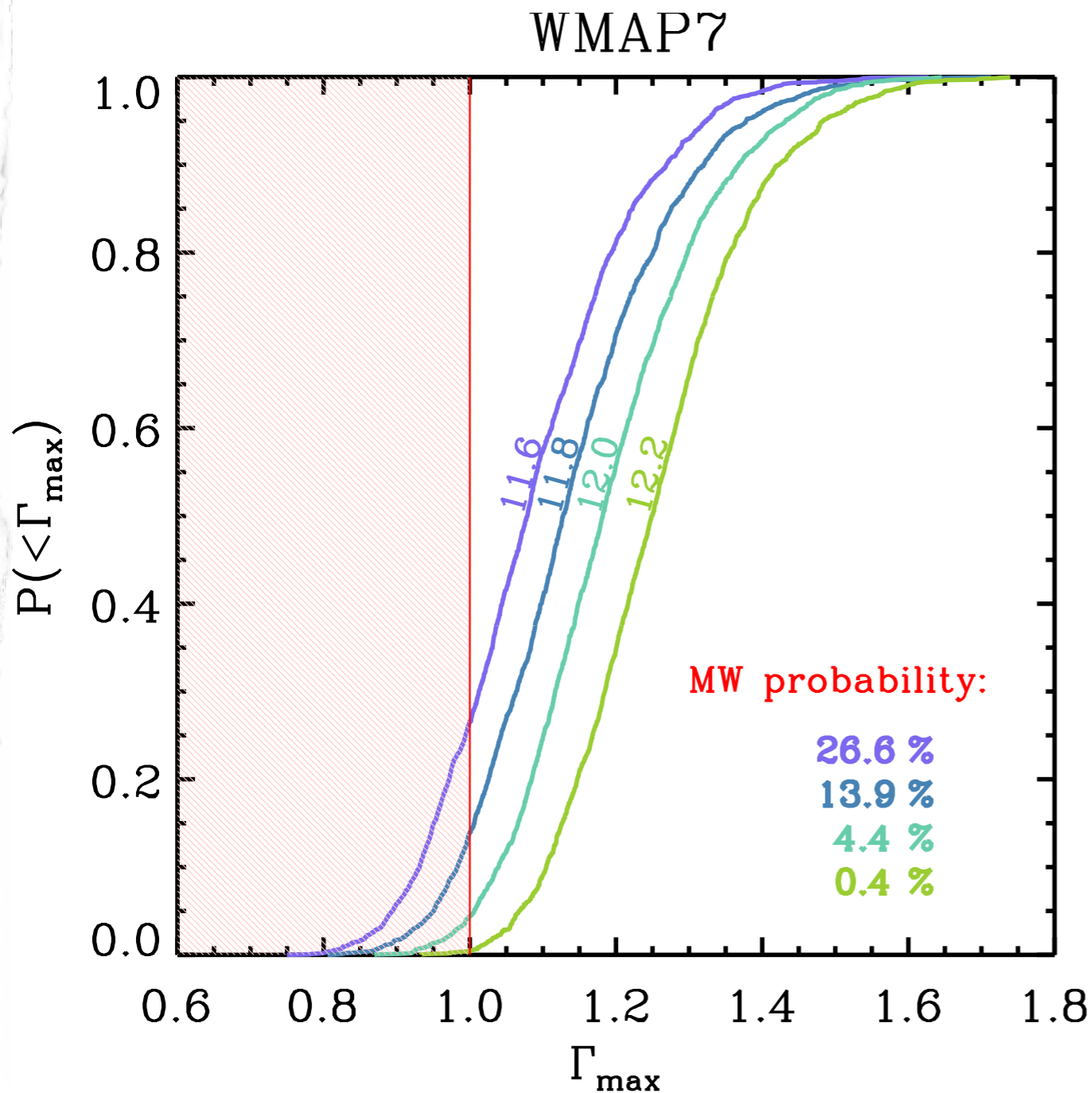
$$V_{\text{max}} > 60 \text{ km s}^{-1}$$

$N_{\text{Gap}} \leq 1$  &  $N_u \geq 2$

probability of having MW-consistent Vmax Gap: always  $< 1\%$

# Subhalo Density

recap: MW-consistent  $\Leftrightarrow \Gamma_{\max} < 1$



- also can be alleviated by lowering MW halo mass
- sensitive to cosmology change

WMAP7  
 $(\Omega_m, \sigma_8) = (0.266, 0.801)$

Planck  
 $(\Omega_m, \sigma_8) = (0.318, 0.834)$

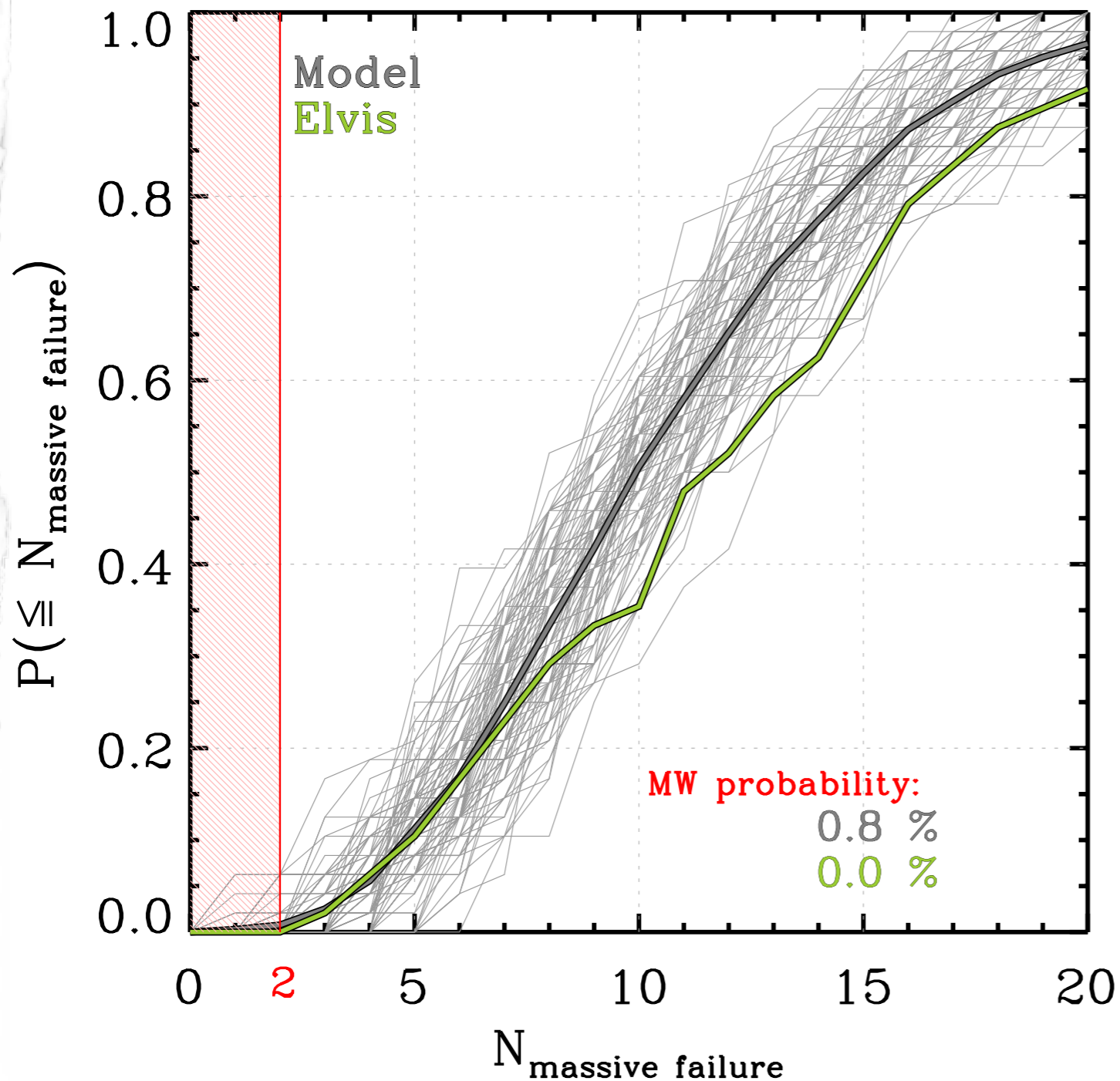
- cosmology comes in mainly via  $R_{\max}$

Jiang & van den Bosch, submitted to MNRAS

# Summary

- **If** TBTF is the missing massive subhaloes:
  - MW-consistent fraction  $<1\%$  for MW-size haloes ( $M_0=12.0$ )
  - reconcilable by lowering MW halo mass, MW-consistent fraction  $\geq 10\%$  for  $M_0=11.8$
  - not very sensitive to cosmology (WMAP7 versus Planck)
- **If** TBTF is the massive subhaloes being too dense:
  - MW-consistent fraction  $<5\%$  for MW-size haloes ( $M_0=12.0$ )
  - reconcilable by lowering MW halo mass, MW-consistent fraction  $\approx 10\%$  for  $M_0=11.8$  (WMAP7)
  - very sensitive to cosmology:  $\approx 3\%$  for  $M_0=11.8$  (Planck)
- **If** TBTF is a  $V_{\max}$  Gap:
  - MW-consistent fraction always  $<1\%$ , irrespective of MW halo mass or cosmology

# Why semi-analytical model? Why not simulations ?



• **ELVIS:**  
48 haloes  
 $M_0 = 10^{12.08 \pm 0.23} h^{-1} M_{\odot}$

• **Model:**  
4800 realizations of  
ELVIS-size haloes  
↕  
100 mock ELVIS suites